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The Second *Fermi* Large Area Telescope Catalog of Gamma-ray Pulsars

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ABSTRACT

This catalog summarizes 117 high-confidence ≥ 0.1 GeV gamma-ray pulsar detections using three years of data acquired by the Large Area Telescope (LAT) on the *Fermi* satellite. Half are neutron stars discovered using LAT data, through periodicity searches in gamma-ray and radio data around LAT unasso-

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ciated source positions. The 117 pulsars are evenly divided into three groups: millisecond pulsars, young radio-loud pulsars, and young radio-quiet pulsars. We characterize the pulse profiles and energy spectra and derive luminosities when distance information exists. Spectral analysis of the off-peak phase intervals indicates probable pulsar wind nebula emission for four pulsars, and off-peak magnetospheric emission for several young and millisecond pulsars. We compare the gamma-ray properties with those in the radio, optical, and X-ray bands. We provide flux limits for pulsars with no observed gamma-ray emission, highlighting a small number of gamma-faint, radio-loud pulsars. The large, varied gamma-ray pulsar sample constrains emission models. *Fermi*'s selection biases complement those of radio surveys, enhancing comparisons with predicted population distributions.

Subject headings: catalogs – gamma rays: observations – pulsars: general – stars: neutron

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1. Introduction

Pulsars have featured prominently in the gamma-ray sky since the birth of gamma-ray astronomy. The Crab and Vela pulsars were the first two sources identified in the 1970’s by SAS-2 (Fichtel et al. 1975) and COS-B (Swanenburg et al. 1981). In the 1990’s the *Compton Gamma-Ray Observatory* brought the pulsar grand total to at least seven, along with three other strong candidates (Thompson 2008). One of these early gamma-ray pulsars, Geminga, was undetected at radio wavelengths (Bignami & Caraveo 1996). Despite the meager number, neutron stars were estimated to represent a sizeable fraction of the EGRET unassociated low-latitude gamma-ray sources (Romani & Yadigaroglu 1995a). The Large Area Telescope (LAT) on the *Fermi Gamma-ray Space Telescope* did not just confirm the expectation: by discovering dozens of radio-quiet gamma-ray pulsars and millisecond pulsars (MSPs, thought to be old pulsars spun up to rapid periods via accretion from a companion, Alpar et al. 1982), the LAT established pulsars as the dominant GeV gamma-ray source class in the Milky Way (Abdo et al. 2010n, The First *Fermi* Large Area Telescope Catalog of Gamma-ray Pulsars, hereafter 1PC).

A pulsar is a rapidly-rotating, highly-magnetized neutron star, surrounded by a plasma-filled magnetosphere. Modeling its emission drives ever-more sophisticated electrodynamic calculations (e.g. Wang & Hirotani 2011; Li et al. 2012; Kalapotharakos et al. 2012b; Pétri 2012). Throughout this paper, we will call pulsars in the main population of the spin period (P) and period derivative (\dot{P}) plane ‘young’ to distinguish them from the much older ‘recycled’ pulsars, including MSPs. All known gamma-ray pulsars, and the most promising candidates to date, are rotation-powered pulsars (RPPs). The LAT has yet to detect significant gamma-ray pulsations from any accretion-powered pulsar or from the magnetars, anomalous X-ray pulsars, and soft gamma repeaters for which the dominant energy source is not electromagnetic braking, but magnetic field decay (Parent et al. 2011).

Here we present 117 gamma-ray pulsars unveiled in three years of on-orbit observations

with *Fermi*. Extensive radio observations by the “Pulsar Timing Consortium” (Smith et al. 2008) greatly enhanced the gamma-ray data analysis. Our analysis of the gamma-ray pulsars is as uniform as is feasible given the widely varying pulsar characteristics. In addition to 1PC, this catalog builds on the 2nd *Fermi* LAT source catalog (Nolan et al. 2012, hereafter 2FGL), which reported pulsations for 83 of the 2FGL sources, included here. An additional 27 pulsars were found to be spatially associated with 2FGL sources and pulsations have since been established for 12 of these, included here. The remaining 22 new pulsars with strong pulsations were either unassociated in 2FGL (pointing to subsequent pulsar discoveries, see Section 3) or were below the 2FGL detection threshold and seen to pulse after 2FGL was completed. We provide our results in FITS files¹ available in the journal electronic supplement as well as on the *Fermi* Science Support Center (FSSC) servers at http://fermi.gsfc.nasa.gov/ssc/data/access/lat/2nd_PSR_catalog/.

2. Observations

Fermi was launched on 2008 June 11, carrying two gamma-ray instruments: the LAT and the Gamma-ray Burst Monitor (Meegan et al. 2009); the latter was not used to prepare this catalog. Atwood et al. (2009) describe *Fermi*’s main instrument, the LAT, and on-orbit performance of the LAT is reported by Abdo et al. (2009f) and Ackermann et al. (2012b). The LAT is a pair-production telescope composed of a 4×4 grid of towers. Each tower consists of a stack of tungsten foil converters interleaved with silicon-strip particle tracking detectors, mated with a hodoscopic cesium-iodide calorimeter. A segmented plastic scintillator anti-coincidence detector covers the grid to help discriminate charged particle backgrounds from gamma-ray photons. The LAT field of view is ~ 2.4 sr. The primary operational mode is a sky survey where the satellite rocks between a pointing above the orbital plane and one below the plane after each orbit. The entire sky is imaged every two orbits (~ 3 hours) and any given point on the sky is observed $\sim 1/6^{\text{th}}$ of the time. Each event classified as a gamma ray in the ground data processing has its incident direction, incident energy (E), and time of arrival recorded in the science data stream.

The LAT is sensitive to gamma rays with energies E from 20 MeV to over 300 GeV, with an on-axis effective area of ~ 8000 cm² above 1 GeV. Multiple Coulomb scattering of the electron-positron pairs created by converted gamma rays degrades the per-photon angular resolution with decreasing energy as $\theta_{68}^2(E) = (3.3)^2(100 \text{ MeV}/E)^{1.56} + (0.1)^2$, averaged over the acceptance for events converting in the front section of the LAT, where θ_{68} is the 68%

¹<http://fits.gsfc.nasa.gov/>

containment radius. The energy- and direction-dependent effective area and PSF are part of the Instrument Response Functions (IRFs). The analysis in this paper used the `Pass7 V6` IRFs selecting events in the “Source” class (Ackermann et al. 2012b).

The data used here to search for gamma-ray pulsars span 2008 August 4 to 2011 August 4. Events were selected with reconstructed energies from 0.1 to 100 GeV and directions within 2° of each pulsar position for pulsation searches (Section 3) and 15° for spectral analyses (Section 6). We excluded gamma rays collected when the LAT was not in nominal science operations mode or the spacecraft rocking angle exceeded 52° , as well as those with measured zenith angles $> 100^\circ$, to greatly reduce the residual gamma rays from the bright limb of the Earth. For PSRs J0205+6449, J1838–0537, and J2215+5135 we did not have timing solutions that were coherent over the full three years. For these pulsars the data sets for pulsation searches and light curve generation only include events within the validity range of the corresponding timing solutions. For the first two pulsars, the data loss is $< 7\%$ but for PSR J2215+5135 it is 60%.

3. Pulsation Discovery

Events recorded by the LAT have timestamps derived from GPS clocks integrated into the satellite’s Guidance, Navigation, and Control (GNC) subsystem, accurate to $< 1 \mu\text{s}$ relative to UTC (Abdo et al. 2009f). The GNC subsystem provides the instantaneous spacecraft position with corresponding accuracy. We compute pulsar rotational phases ϕ_i using TEMPO2 (Hobbs et al. 2006) with the `fermi` plug-in (Ray et al. 2011). The `fermi` plug-in uses the recorded times and spacecraft positions combined with a pulsar timing ephemeris (specified in a TEMPO2 parameter, or ‘par’, file). The timing chain from the instrument clocks through the barycentering and epoch folding software is accurate to better than a few μs for binary orbits, and significantly better for isolated pulsars (Smith et al. 2008). The accuracy of the phase computation is thus determined by the ephemeris. The par file is created using radio or gamma-ray data, or a mix, depending on the LAT pulsar discovery method, as described in the following three subsections.

We required a $\geq 5\sigma$ confidence level detection of modulation in the phase histogram for a pulsar to be included in this catalog, as described below. Gamma-ray pulsar data are extremely sparse, often with fewer than one photon detected in tens of thousands (or in the case of MSPs, millions!) of pulsar rotations. In these circumstances, the favored techniques are unbinned tests for periodic signals. We use the H-test (de Jager et al. 1989; de Jager & Büsching 2010), a statistical test for discarding the null hypothesis that a set of photon phases is uniformly distributed. For N_γ gamma-rays, the H-test statistic is $H \equiv$

$\max(Z_m^2 - 4 \times (m - 1), 1 \leq m \leq 20)$, with $Z_m^2 \equiv \frac{2}{N_\gamma} \sum_{k=1}^m \alpha_k^2 + \beta_k^2$, and α_k and β_k the empirical trigonometric coefficients $\alpha_k \equiv \sum_{i=1}^{N_\gamma} \sin(2\pi k \phi_i)$ and $\beta_k \equiv \sum_{i=1}^{N_\gamma} \cos(2\pi k \phi_i)$. By including a search over a range of harmonics, the H-test maintains sensitivity to light curves with a large range of morphologies (e.g., sharp vs. broad peaks). The sharpness of the peaks in the gamma-ray profile has a large impact on the detectability of the pulsar; in particular, pulsars with narrow, sharp peaks are easier to detect than pulsars with broad peaks covering more of the pulse phase.

Early in the mission most pulsation searches (for example, Abdo et al. 2009a) selected events with arrival directions within a fixed angular distance of the pulsar (the region of interest, or ROI) and a minimum energy cut (E_{\min}). Because of the range of pulsar spectra, fluxes, and levels of diffuse gamma-ray background, combined with the strongly energy-dependent PSF of the LAT, a number of trials must be done over a range of ROI sizes and E_{\min} to optimize the detection significance for each candidate pulsar.

Using the probability that each event originates from the pulsar, computed from a spectral model of the region and the LAT IRFs, the H-test can be extended, using these probabilities as weights (Kerr 2011). This both improves the sensitivity of the H-test and removes the need for trials over event selection criteria. The weights, w_i , representing the probability that the i^{th} event originates from the pulsar are

$$w_i = \frac{dN/dE_{\text{psr}}(E_i, \vec{x}_i)}{\sum_j dN/dE_j(E_i, \vec{x}_i)}, \quad (1)$$

where E_i and \vec{x}_i are the observed energy and position on the sky of the i^{th} event and dN/dE_j is the phase-averaged spectra for the j^{th} source in the ROI (See Section 6). Incorporating the weights yields the weighted H-test, $H \equiv \max(Z_{mw}^2 - 4 \times (m - 1); 1 \leq m \leq 20)$, with

$$Z_{mw}^2 \equiv \frac{2}{N_\gamma} \left(\frac{1}{N_\gamma} \sum_{i=1}^{N_\gamma} w_i^2 \right)^{-1} \sum_{k=1}^m \alpha_{kw}^2 + \beta_{kw}^2 \quad (2)$$

where $\alpha_{kw} = \sum_{i=1}^{N_\gamma} w_i \cos(2\pi k \phi_i)$ and $\beta_{kw} = \sum_{i=1}^{N_\gamma} w_i \sin(2\pi k \phi_i)$. Kerr (2011) provides the probability that a detection is a statistical fluctuation for a given H value, approximated by $e^{-0.4H}$. $H = 36$ (15) corresponds to a 5σ (3σ) detection.

3.1. Using Known Rotation Ephemerides

The first gamma-ray pulsar discovery method, described above, found 61 of the gamma-ray pulsars in this catalog. It uses known rotational ephemerides from radio or X-ray obser-

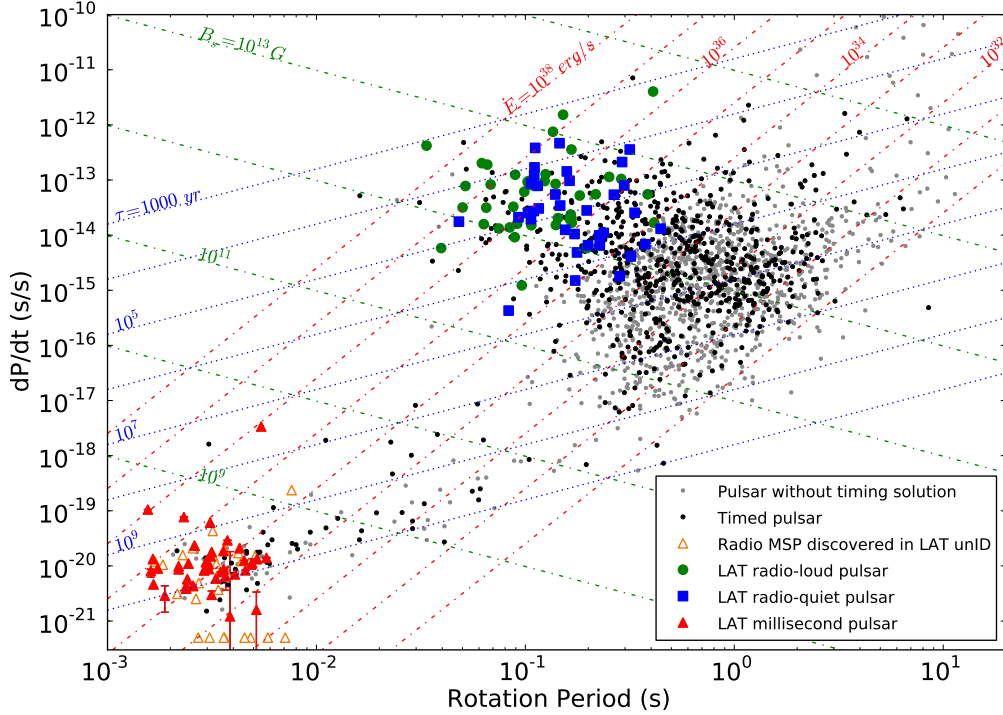


Fig. 1.— Pulsar spindown rate, \dot{P} , versus the rotation period P . Green dots indicate the 42 young, radio-loud gamma-ray pulsars and blue squares show the 35 young, ‘radio-quiet’ pulsars, defined as $S_{1400} < 30 \mu\text{Jy}$, where S_{1400} is the radio flux density at 1400 MHz. Red triangles are the 40 millisecond gamma-ray pulsars. The 710 black dots indicate pulsars phase-folded in gamma rays using rotation models provided by the ‘Pulsar Timing consortium’ for which no significant pulsations were observed. Phase-folding was not performed for the 1337 pulsars outside of globular clusters indicated by gray dots. Orange open triangles indicate radio MSPs discovered at the positions of previously unassociated LAT sources for which we have not yet seen gamma pulsations. We plot them at $\dot{P} \equiv 5 \times 10^{-22}$ when \dot{P} is unavailable. Shklovskii corrections to \dot{P} have been applied to the pulsars with proper motion measurements (see Section 4.3). For clarity, error bars are shown only for the gamma-detected pulsars.

vatories. The 2286 known rotation-powered pulsars (mostly from the ATNF Pulsar Catalog² (Manchester et al. 2005), see Table 1) are all candidate gamma-ray pulsars. Nearly all of these were discovered in radio searches, with a handful coming from X-ray observations.

²<http://www.atnf.csiro.au/research/pulsar/psrcat>

Phase-folding with a radio or X-ray ephemeris is the most sensitive way to find gamma-ray pulsations, since no penalties are incurred for trials in position, P, \dot{P} , or other search parameters. Having a current ephemeris for as many known pulsars as possible is of critical importance to LAT science and is the key goal of the Pulsar Timing Consortium (Smith et al. 2008). EGRET results (Thompson 2008) as well as theoretical expectations indicated that young pulsars with large spindown power³, $\dot{E} > 1 \times 10^{34}$ erg s⁻¹, are the most likely gamma-ray pulsar candidates. Because of their intrinsic instabilities, such as timing noise and glitches, these pulsars are also the most resource intensive to maintain ephemerides of sufficient accuracy. To allow for unexpected discoveries, the Timing Consortium also provides ephemerides for essentially all known pulsars that are regularly timed, spanning the $P\dot{P}$ space of known pulsars (Figure 1). In addition to \dot{E} , the $P\dot{P}$ diagram shows two other physical parameters derived from the timing information: the magnetic field at the neutron star surface, $B_S = (1.5I_0c^3P\dot{P})^{1/2}/2\pi R_{\text{NS}}^3$, assuming an orthogonal rotator with neutron star radius $R_{\text{NS}} = 10$ km and the speed of light in a vacuum, c ; and the characteristic age $\tau_c = P/2\dot{P}$, assuming magnetic dipole braking as the only energy-loss mechanism and an initial spin period much less than the current period. The black dots in Figure 1 show 710 pulsars that we have phase-folded without detecting gamma pulsations, in addition to the 117 gamma-ray pulsars. The locations of all 117 gamma-ray pulsars on the sky are shown in Figure 2.

For known pulsars we use years of radio and/or X-ray time-of-arrival measurements (“TOAs”) to fit the timing model parameters using the standard pulsar timing codes TEMPO (Taylor & Weisberg 1989) or TEMPO2 (Hobbs et al. 2006). In addition to providing a model for folding the gamma-ray data, the radio observations also provide the information needed to measure the absolute phase alignment (after correcting for interstellar dispersion) between the radio or X-ray and gamma-ray pulses, providing key information about the relative geometry of the different emission regions.

3.2. Blind Periodicity Searches

The second method of discovering gamma-ray pulsars, which produced 36 (approximately one-third) of the gamma-ray pulsars in this catalog, involves detecting the rotational period in the LAT data. Both these searches and the radio searches described in the next subsection begin with a target list of candidate pulsars. Some targets are sources known at other wavelengths that are suspected of harboring pulsars. These include supernova rem-

³ $\dot{E} = 4\pi^2 I_0 \dot{P} / P^3$, for which we use $I_0 = 10^{45}$ g cm² as the neutron star moment of inertia.

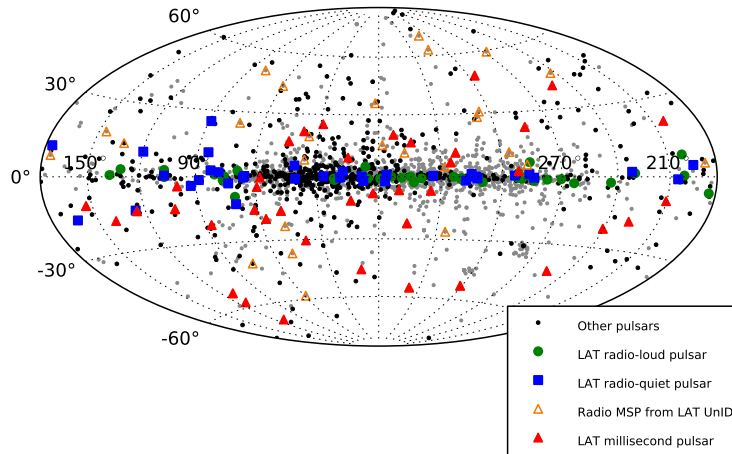


Fig. 2.— Pulsar sky map in Galactic coordinates. The markers are the same as in Figure 1.

nants (SNRs), pulsar wind nebulae (PWNe), compact central objects (CCOs), unidentified TeV sources, and other high-energy sources, mostly along the Galactic plane. Generally, these sources had already been subjected to deep radio searches independent of *Fermi*.

In addition, as the LAT surveys the sky, an increasing number of gamma-ray sources are discovered and characterized that are not associated with previously known objects. Several methods have been used to rank these according to their probabilities of being yet-undiscovered pulsars. Most of these rely on the tendency of gamma-ray pulsars to be non-variable and have spectra that can be fit with exponential cutoffs in the few GeV band (Ackermann et al. 2012a; Lee et al. 2012).

Blind searches for pulsars in gamma rays are challenging, due to the wide pulsar parameter ranges that must be searched and due to the sparseness of the data (a few photons per hour for the brightest sources). This results in very long integration times (months to years) making standard Fast Fourier Transform search techniques computationally prohibitive. New semi-coherent search techniques (Atwood et al. 2006; Pletsch et al. 2012b) have been extremely successful at discovering gamma-ray pulsars with modest computational requirements.

LAT blind search sensitivity depends on a number of parameters: the rotation frequency,

energy spectrum, pulsed fraction, level of diffuse gamma-ray background, event extraction choices (e.g. ROI and E_{\min}), and the accuracy of the position used to barycenter the data. The one-year sensitivity was evaluated using a Monte Carlo study by Dormody et al. (2011). Newer searches (Pletsch et al. 2012b,c) have mitigated dependence on event selection criteria and source localization by weighting events and searching over a grid of positions.

In all, well over one hundred LAT sources have been subjected to blind period searches. Pulsars might have been missed for a few reasons: (1) low pulsed fraction or very high backgrounds, (2) broad pulse profiles (our algorithms detect sharp pulses more easily), (3) high levels of timing noise or glitches, (4) being in an unknown binary system. Most MSPs are in binary systems, where the Doppler shifts from the orbital motion smear the signal. In some cases, multiwavelength observations constrain the orbit and position to make the search more like that of an isolated MSP. Optical studies (Romani & Shaw 2011; Kong et al. 2012; Romani 2012) led to the first discovery of a millisecond pulsar, PSR J1311–3430, in a blind search of LAT data (Pletsch et al. 2012a). Detection of radio pulsations followed shortly (Ray et al. 2013). Even isolated MSP searches require massive computation with fine frequency and position gridding. The Einstein@home⁴ project applies the power of global volunteer computing to this problem.

For the LAT pulsars undetected in the radio (see Section 4.1), or too faint for regular radio timing, we must determine the pulsar timing ephemeris directly from the LAT data. Techniques for TOA determination optimized for sparse photon data have been developed and applied to generate the timing models required for the profile analysis (Ray et al. 2011). This timing provides much more precise pulsar positions than can be determined from the LAT event directions, which is important for multiwavelength counterpart searches. It also allows study of timing noise and glitch behavior (Dormody et al., in prep).

3.3. Radio Pulsar Discoveries Leading to Gamma-ray Pulsations

In the third discovery method that we applied, which yielded 20 of this catalog’s MSPs, unassociated LAT source positions are searched for radio pulsations. When found, the resulting ephemeris enables gamma-ray phase-folding, as in Section 3.1. A key feature of radio pulsar searches is that they are sensitive to binary systems with the application of techniques to correct for the orbital acceleration in short data sets (with durations much less than the binary period, Ransom et al. 2002). This allows for the discovery of binary MSPs, which are largely inaccessible to gamma-ray blind searches, as described above.

⁴<http://einstein.phys.uwm.edu>

Radio searches of several hundred LAT sources by the *Fermi* Pulsar Search Consortium (PSC), an international collaboration of radio observers with access to large radio telescopes, have resulted in the discovery of 47 pulsars, including 43 MSPs and four young or middle-aged pulsars (Ray et al. 2012). As the LAT Collaboration generates internal source lists and preliminary catalogs of gamma-ray sources from the accumulating sky-survey data, these target localizations are provided to the PSC for searching, with rankings of how strongly their characteristics resemble those of gamma-ray pulsars, as described in Section 3.2. This technique was employed during the EGRET era as well, but with modest success, in part due to the relatively poor source localizations. With the LAT, there are many more gamma-ray sources detected and each one is localized to an accuracy that is comparable to, or smaller than, the beam width of the radio telescopes being used. This enables deep searches by removing the need to mosaic a large region. It also facilitates repeated searches of the same source, which is important because discoveries can be missed as a result of scintillation or eclipses in binary systems (e.g., PSR J0101–6422, see Kerr et al. 2012a).

Guided by these ranked lists of pulsar-like gamma-ray sources, the 43 radio MSPs were discovered in a tiny fraction of the radio telescope time that would have been required to find them in undirected radio pulsar surveys. In particular, because the MSP population out to the LAT’s detection limit (~ 2 kpc) is distributed nearly uniformly across the sky, full sky surveys are required, whereas most young pulsar searches have concentrated on the Galactic plane. For comparison, after analyzing thousands of pointings carried out since 2007, the High Time Resolution Universe surveys (Keith 2012; Ng & HTRU Collaboration 2013, and references therein) found 29 new radio MSPs.

Interestingly, the success rate for radio searches of LAT sources in the plane has been much poorer. Only four young pulsars have been discovered, and only one of those turned out to be a gamma-ray pulsar (PSR J2030+3641, Camilo et al. 2012), the others being chance associations. This is probably due to a combination of young pulsars having smaller radio beaming fractions than MSPs (as evidenced by the large number of young, radio-quiet pulsars discovered) and the fact that the Galactic plane has been well surveyed for radio pulsars. The great success of the blind gamma searches in the plane is because young pulsars mainly reside there. Their smaller radio beaming fractions leave a large number of radio-quiet pulsars that can only be discovered in high-energy data.

Once a radio pulsar has been discovered positionally coincident with a LAT source, it must be observed for a substantial period (typically six months to a year or more) to determine a timing model that allows a periodicity search in the LAT data, as described in Section 3.1. In several cases, an initial radio model has allowed discovery of the gamma-ray pulsations, then the LAT data themselves have been used to extend the validity of the

timing model back through the launch of *Fermi*, a few years before the radio discovery. This radio follow up has resulted in the confirmed detection of LAT pulsations from 20 of these MSPs. Five more were detected using data beyond the set described in Section 2. Of the remainder, most will have LAT pulsations detected once their radio timing models are well determined, but a few (e.g., PSR J1103–5403, see Keith et al. (2011)) are likely to be just chance coincidences with the target LAT source.

4. The Gamma-ray Pulsars

The discovery strategies discussed in Section 3 yielded 117 gamma-ray pulsars in three years of data. Of the gamma-ray pulsars in this catalog, roughly half (41 young and 20 MSPs) were known in radio and/or X-rays prior to the launch of *Fermi*. The remaining pulsars were discovered by or with the aid of the LAT, with 36 being young pulsars found in blind searches of LAT data and the remaining being MSPs found in radio searches of unassociated LAT sources. *Fermi* has not only significantly increased the number of known energetic young and millisecond pulsars, but has done so with selection biases complementary to those of previous surveys. The LAT all-sky survey has its greatest sensitivity in regions of the sky away from the Galactic plane (see Section 8.2), increasing the diversity and the uniformity of the sampled neutron star population. As an example, Figure 2 shows the broad range of Galactic latitude of the *Fermi* pulsars.

Table 1. Pulsar varieties

Category	Count	Sub-count	Fraction
Known rotation-powered pulsars (RPPs) ^a	2286		
RPPs with measured $\dot{P} > 0$		1944	
RPPs with measured $\dot{E} > 3 \times 10^{33} \text{ erg s}^{-1}$		552	
Millisecond pulsars (MSPs, $P < 16 \text{ ms}$)	292		
Field MSPs		169	
MSPs in globular clusters		123	
Field MSPs with measured $\dot{E} > 3 \times 10^{33} \text{ erg s}^{-1}$		96	
Globular cluster MSPs with measured $\dot{E} > 3 \times 10^{33} \text{ erg s}^{-1}$		25	
Gamma-ray pulsars in this catalog	117		
Young or middle-aged		77	
Radio-loud gamma-ray ^b		42	36%
Radio-quiet gamma-ray		35	30%
Gamma-ray MSPs (isolated + binary)		(10+30) = 40	34%
Radio MSPs discovered in LAT sources with gamma-ray pulsations ^c	46	34	

^aIncludes the 2193 pulsars, which are all RPPs, in the ATNF Pulsar Catalog (v1.46, Manchester et al. 2005), see <http://www.atnf.csiro.au/research/pulsar/psrcat>, as well as more recent discoveries. D. Lorimer maintains a list of known field MSPs at <http://astro.phys.wvu.edu/GalacticMSPs/>.

^b $S_{1400} > 30 \mu\text{Jy}$, where S_{1400} is the radio flux density at 1400 MHz.

^cOnly 20 of the new radio MSPs showed gamma-ray pulsations when the dataset for this catalog was frozen.

Table 1 summarizes the census of known pulsars, independent of the method by which the pulsars were discovered. Tables 2 and 3 list the characteristics of the 117 gamma-ray pulsars, divided into young and millisecond gamma-ray pulsars, respectively. All have large spindown powers, $\dot{E} > 3 \times 10^{33}$ erg s⁻¹, apparent in Figure 1. The large uncertainties on the two seeming exceptions, PSRs J0610–2100 and J1024–0719, are discussed in Section 6.3.

Pulsar discoveries continue as increased statistics bring light curves above our 5σ detection threshold, improved methods for event selection and blind searches allow increased sensitivity, and multiwavelength studies either detect radio pulsations or constrain the blind-search space for likely pulsar candidates. Table 4 lists a number of LAT pulsars announced since the sample was frozen for the uniform analysis of the present paper.

Table 2. Some parameters of young LAT-detected pulsars

PSR	History	l ($^{\circ}$)	b ($^{\circ}$)	P (ms)	\dot{P} (10^{-15})	\dot{E} (10^{34} erg s $^{-1}$)	S_{1400} (mJy)
J0007+7303	g	119.66	10.46	315.9	357.	44.8	<0.005 ¹
J0106+4855	gu	125.47	−13.87	83.2	0.43	2.9	0.008
J0205+6449	x	130.72	3.08	65.7	190.	2644.	0.045
J0248+6021	r	136.90	0.70	217.1	55.0	21.2	13.7
J0357+3205	gu	162.76	−16.01	444.1	13.1	0.6	<0.004 ¹
J0534+2200	re	184.56	−5.78	33.6	420.	43606.	14.0
J0622+3749	gu	175.88	10.96	333.2	25.4	2.7	<0.012 ²
J0631+1036	r	201.22	0.45	287.8	104.	17.3	0.8
J0633+0632	gu	205.09	−0.93	297.4	79.6	11.9	<0.003 ¹
J0633+1746	xe	195.13	4.27	237.1	11.0	3.3	<0.507 ³
J0659+1414	r	201.11	8.26	384.9	55.0	3.8	3.7
J0729−1448	r	230.39	1.42	251.7	114.	28.2	0.7
J0734−1559	gu	232.06	2.02	155.1	12.5	13.2	<0.005 ⁴
J0742−2822	r	243.77	−2.44	166.8	16.8	14.3	15.0
J0835−4510	re	263.55	−2.79	89.4	125.	690.	1100.
J0908−4913	r	270.27	−1.02	106.8	15.1	49.0	10.0
J0940−5428	r	277.51	−1.29	87.6	32.8	193.	0.66
J1016−5857	r	284.08	−1.88	107.4	80.6	257.	0.46
J1019−5749	r	283.84	−0.68	162.5	20.1	18.4	0.8
J1023−5746	gu	284.17	−0.41	111.5	382.	1089.	<0.030 ⁵
J1028−5819	r	285.06	−0.50	91.4	16.1	83.3	0.36
J1044−5737	gu	286.57	1.16	139.0	54.6	80.2	<0.020 ⁵
J1048−5832	r	287.42	0.58	123.7	95.7	200.	6.5
J1057−5226	re	285.98	6.65	197.1	5.8	3.0	9.5
J1105−6107	r	290.49	−0.85	63.2	15.8	248.	0.75
J1112−6103	r	291.22	−0.46	65.0	31.5	454.	1.4
J1119−6127	r	292.15	−0.54	408.7	4028.	233.	0.8
J1124−5916	r	292.04	1.75	135.5	750.	1190.	0.08
J1135−6055	gu	293.79	0.58	114.5	78.4	206.	<0.030 ⁴
J1357−6429	r	309.92	−2.51	166.2	357.	307.	0.44
J1410−6132	r	312.20	−0.09	50.1	31.8	1000.	6.56 ⁶
J1413−6205	gu	312.37	−0.74	109.7	27.4	81.8	<0.024 ⁵
J1418−6058	gu	313.32	0.13	110.6	169.	494.	<0.029 ¹
J1420−6048	r	313.54	0.23	68.2	82.9	1032.	0.9
J1429−5911	gu	315.26	1.30	115.8	30.5	77.4	<0.021 ⁵
J1459−6053	gu	317.89	−1.79	103.2	25.3	90.9	<0.037 ¹
J1509−5850	r	319.97	−0.62	88.9	9.2	51.5	0.15
J1513−5908	xe	320.32	−1.16	151.5	1529.	1735.	0.94
J1531−5610	r	323.90	0.03	84.2	13.8	91.2	0.6
J1620−4927	gu	333.89	0.41	171.9	10.5	8.1	<0.040 ²
J1648−4611	r	339.44	−0.79	165.0	23.7	20.9	0.58
J1702−4128	r	344.74	0.12	182.2	52.3	34.2	1.1
J1709−4429	re	343.10	−2.69	102.5	92.8	340.	7.3
J1718−3825	r	348.95	−0.43	74.7	13.2	125.	1.3
J1730−3350	r	354.13	0.09	139.5	84.8	123.	3.2

Table 2—Continued

PSR	History	l ($^{\circ}$)	b ($^{\circ}$)	P (ms)	\dot{P} (10^{-15})	\dot{E} (10^{34} erg s $^{-1}$)	S_{1400} (mJy)
J1732–3131	gu	356.31	1.01	196.5	28.0	14.6	<0.015 ¹
J1741–2054	gu	6.43	4.91	413.7	17.0	0.9	0.16
J1746–3239	gu	356.96	–2.18	199.5	6.6	3.3	<0.034 ²
J1747–2958	r	359.31	–0.84	98.8	61.3	251.	0.25
J1801–2451	r	5.25	–0.88	125.0	127.	257.	0.85
J1803–2149	gu	8.14	0.19	106.3	19.5	64.1	<0.024 ²
J1809–2332	g	7.39	–1.99	146.8	34.4	43.0	<0.025 ¹
J1813–1246	gu	17.24	2.44	48.1	17.6	624.	<0.017 ¹
J1826–1256	g	18.56	–0.38	110.2	121.	358.	<0.013 ¹
J1833–1034	r	21.50	–0.89	61.9	202.	3364.	0.071
J1835–1106	r	21.22	–1.51	165.9	20.6	17.8	2.2
J1836+5925	g	88.88	25.00	173.3	1.5	1.1	<0.004 ¹
J1838–0537	gu	26.51	0.21	145.7	465.	593.	<0.017 ⁷
J1846+0919	gu	40.69	5.34	225.6	9.9	3.4	<0.005 ⁵
J1907+0602	g	40.18	–0.89	106.6	86.7	282.	0.0034
J1952+3252	re	68.77	2.82	39.5	5.8	372.	1.0
J1954+2836	gu	65.24	0.38	92.7	21.2	105.	<0.005 ⁵
J1957+5033	gu	84.60	11.00	374.8	6.8	0.5	<0.010 ⁵
J1958+2846	gu	65.88	–0.35	290.4	212.	34.2	<0.006 ¹
J2021+3651	r	75.22	0.11	103.7	95.6	338.	0.1
J2021+4026	g	78.23	2.09	265.3	54.2	11.4	<0.020 ¹
J2028+3332	gu	73.36	–3.01	176.7	4.9	3.5	<0.005 ²
J2030+3641	ru	76.12	–1.44	200.1	6.5	3.2	0.15
J2030+4415	gu	82.34	2.89	227.1	6.5	2.2	<0.008 ²
J2032+4127	gu	80.22	1.03	143.2	20.4	27.3	0.23
J2043+2740	r	70.61	–9.15	96.1	1.2	5.5	9.35
J2055+2539	gu	70.69	–12.52	319.6	4.1	0.5	<0.007 ⁵
J2111+4606	gu	88.31	–1.45	157.8	143.	144.	<0.013 ²
J2139+4716	gu	92.63	–4.02	282.8	1.8	0.3	<0.014 ²
J2229+6114	r	106.65	2.95	51.6	77.9	2231.	0.25
J2238+5903	gu	106.56	0.48	162.7	97.0	88.8	<0.011 ¹
J2240+5832	r	106.57	–0.11	139.9	15.2	21.9	2.7

Note. — Column 2 gives a discovery/detection code: g=gamma-ray blind search, r=radio, u=candidate location was that of an unassociated LAT source, x=X-ray, e=seen by *EGRET*. Columns 3 and 4 give Galactic coordinates for each pulsar. Columns 5 and 6 list the period (P) and its first derivative (\dot{P}), and Column 7 gives the spindown luminosity \dot{E} . The Shklovskii correction to \dot{P} and \dot{E} is negligible for these young pulsars (see Section 4.3). Column 8 gives the radio flux density (or upper limit) at 1400 MHz (S_{1400} , see Section 4.1), taken from the ATNF database except for the noted entries where: (1) Ray et al. (2011); (2) Pletsch et al. (2012b); (3) Geminga: Spoelstra & Hermsen (1984); (4) GBT (this paper); (5) Saz Parkinson et al. (2010); (6) O’Brien et al. (2008); (7) Pletsch et al. (2012c). PSR J1509–5850 should not be confused with PSR B1509–58 (= J1513–5908) observed by instruments on the *Compton Gamma-Ray Observatory*.

Table 3. Some parameters of LAT-detected millisecond pulsars

PSR	Type, history.	l ($^{\circ}$)	b ($^{\circ}$)	P (ms)	\dot{P} (10^{-20})	\dot{E} (10^{-33} erg s $^{-1}$)	S_{1400} (mJy)
J0023+0923	bwru	111.15	-53.22	3.05	1.08	15.1	0.19 ¹
J0030+0451	r	113.14	-57.61	4.87	1.02	3.49	0.60
J0034-0534	br	111.49	-68.07	1.88	0.50	29.7	0.61
J0101-6422	bru	301.19	-52.72	2.57	0.48	12.0	0.28
J0102+4839	bru	124.93	-14.83	2.96	1.17	17.5	0.22 ²
J0218+4232	br	139.51	-17.53	2.32	7.74	243.	0.90
J0340+4130	ru	154.04	-11.47	3.30	0.59	7.9	0.17 ²
J0437-4715	br	253.39	-41.96	5.76	5.73	11.8	149
J0610-2100	bwr	227.75	-18.18	3.86	1.23	8.5	0.40
J0613-0200	br	210.41	-9.30	3.06	0.96	13.2	2.3
J0614-3329	bru	240.50	-21.83	3.15	1.78	22.0	0.60 ³
J0751+1807	br	202.73	21.09	3.48	0.78	7.30	3.2
J1024-0719	r	251.70	40.52	5.16	1.85	5.30	1.5
J1124-3653	bwru	283.74	23.59	2.41	0.58	17.1	0.04 ⁴
J1125-5825	br	291.89	2.60	3.10	6.09	80.5	0.44
J1231-1411	bru	295.53	48.39	3.68	2.12	17.9	0.16 ³
J1446-4701	br	322.50	11.43	2.19	0.98	36.8	0.37
J1514-4946	bru	325.22	6.84	3.58	1.87	16.0	...
J1600-3053	br	344.09	16.45	3.60	0.95	8.05	2.5
J1614-2230	br	352.64	20.19	3.15	0.96	12.1	1.2 ⁵
J1658-5324	ru	334.87	-6.63	2.43	1.10	30.2	0.50 ⁶
J1713+0747	br	28.75	25.22	4.57	0.85	3.53	10.2
J1741+1351	br	37.90	21.62	3.75	3.02	22.7	0.93
J1744-1134	r	14.79	9.18	4.07	0.89	5.20	3.1
J1747-4036	ru	350.19	-6.35	1.64	1.33	116.	1.22 ⁶
J1810+1744	bwru	43.87	16.64	1.66	0.46	39.7	1.89 ¹
J1823-3021A	r	2.79	-7.91	5.44	338.	828.	0.72
J1858-2216	bru	13.55	-11.45	2.38	0.39	11.3	...
J1902-5105	bru	345.59	-22.40	1.74	0.90	68.6	0.90 ⁶
J1939+2134	r	57.51	-0.29	1.56	10.5	1097.	13.9
J1959+2048	bwr	59.20	-4.70	1.61	1.68	160.	0.40
J2017+0603	bru	48.62	-16.03	2.90	0.83	13.0	0.50
J2043+1711	bru	61.92	-15.31	2.38	0.57	15.3	0.17 ⁷
J2047+1053	bru	57.06	-19.67	4.29	2.10	10.5	...
J2051-0827	bwr	39.19	-30.41	4.51	1.28	5.49	2.8
J2124-3358	r	10.93	-45.44	4.93	2.06	6.77	3.6
J2214+3000	bwru	86.86	-21.67	3.12	1.50	19.2	0.85 ³
J2215+5135	bkru	99.46	-4.60	2.61	2.34	51.9	0.47 ¹
J2241-5236	bwru	337.46	-54.93	2.19	0.87	26.0	4.1
J2302+4442	bru	103.40	-14.00	5.20	1.33	3.82	1.2

Note. — Column 2: b=binary, r=radio detected, u=seed position was that of an unassociated LAT source, w=white dwarf companion, k=“redback”. Columns 3 and 4 give the Galactic coordinates, with the rotation period P in column 5. The first period time derivative \dot{P} and the spindown luminosity \dot{E} in Columns 6 and 7 are *uncorrected* for the Shklovskii effect, in this Table. The corrected values are used throughout the rest of the paper, and are listed in Table 6 in Section

4.3. Column 9 gives the radio flux density (or upper limit) at 1400 MHz (Section 4.1), taken from the ATNF database except for the noted entries: (1) Hessels et al. (2011); (2) Bangale et al. (in prep.); (3) Ransom et al. (2011); (4) This paper; (5) Demorest et al. (2010); (6) Kerr et al. (2012b); (7) Guillemot et al. (2012). The three MSPs with no S_{1400} listed scintillate too much to obtain a good flux measurement (PSR J1514–4946), or the radio flux has not yet been measured (PSRs J1858–2216 and J2047+1053).

Table 4. Gamma-ray pulsars not in this catalog

PSRJ	P (ms)	\dot{E} (10^{34} erg s $^{-1}$)	Codes	References
J0307+7443	3.16	2.2	mbr	Ray et al. (2012)
J0737–3039A	22.7	0.59	r	Guillemot et al. (2013)
J1055–6028	99.7	120	r	Hou & Smith (2013)
J1311–3430	2.56	4.9	mbgu	Pletsch et al. (2012a)
J1544+4937	2.16	1.2	mbru	Bhattacharyya et al. (2013)
J1640+2224	3.16	0.35	mbr	Hou & Smith (2013)
J1705–1906	299.0	0.61	r	Hou & Smith (2013)
J1732–5049	5.31	0.37	mr	Hou & Smith (2013)
J1745+1017	2.65	0.53	mbru	Barr et al. (2013)
J1816+4510	3.19	5.2	mbru	Kaplan et al. (2012)
J1824–2452A	3.05	220	mr	Wu et al. (2013); Johnson et al. (in prep.)
J1843–1113	1.85	6.0	mr	Hou & Smith (2013)
J1913+0904	163.2	16	r	Hou & Smith (2013)
J2256–1024	2.29	5.2	mbru	Boyles et al. (2011)
J2339–0533	2.88	2.3	mbru	Ray et al. (in prep.)

Note. — Beyond the 117 pulsars. The above 15 pulsars were discovered in gamma rays as this catalog neared completion. An additional 13 gamma-ray pulsars discovered by the LAT collaboration or by other groups using public LAT data have publications in preparation, for a total of 145 as we go to submission (8 May 2013). We maintain a list at <https://confluence.slac.stanford.edu/display/GLAMCOG/Public+List+of+LAT+Detected+Gamma-Ray+Pulsars>. The codes are: u=discovered in a LAT unassociated source, g=discovered in a gamma-ray blind period search, r=radio detection, m=MSP, b=binary system.

4.1. Radio Intensities

The 1400 MHz flux densities, S_{1400} , of the young LAT-detected pulsars are listed in Table 2, and in Table 3 for the MSPs. Figure 3 shows how they compare with the overall pulsar population. Whenever possible, we report S_{1400} as given in the ATNF Pulsar Catalog. For radio-loud pulsars with no published value at 1400 MHz, we extrapolate to S_{1400} from measurements at other frequencies, assuming $S_\nu \propto \nu^\alpha$, where α is the spectral index. For most pulsars α has not been measured, and we use an average value $\langle \alpha \rangle = -1.7$, a middle ground between -1.6 from Lorimer et al. (1995) and -1.8 from Maron et al. (2000). For those pulsars with measured spectral indices, we use the the published value of α for the extrapolation. In the Table notes, we list those pulsars for which we have extrapolated S_{1400} from another frequency and/or used a value of α other than -1.7 for the extrapolation.

Table 2 also reports upper limits on S_{1400} for blind search pulsars that have been observed, but not detected, at radio frequencies. We define these upper limits as the sensitivity of the observation given by the pulsar radiometer equation (Eq. 7.10 on page 174 of Lorimer & Kramer 2004) assuming a minimum signal-to-noise ratio of 5 for a detection and a pulse duty cycle of 10%. We mention here the unconfirmed radio detections of Geminga and PSR J1732–3131 at low radio frequencies, consistent with their non-detection above 300 MHz (Malofeev & Malov 1997; Maan et al. 2012).

All pulsars discovered in blind searches have been searched deeply for radio pulsations (Saz Parkinson et al. 2010; Ray et al. 2011, 2012), and four of the 36 have been detected (Camilo et al. 2009; Abdo et al. 2010m; Pletsch et al. 2012b). In 1PC we labeled the young pulsars by how they were discovered (radio-selected vs. gamma-ray selected), whereas we now define a pulsar as ‘radio-loud’ if $S_{1400} > 30 \mu\text{Jy}$, and ‘radio-quiet’ if the measured flux density is lower, as for the two pulsars with detections of very faint radio pulsations, or if no radio detection has been achieved. The horizontal line in Figure 3 shows the threshold. This definition favors observational characteristics instead of discovery history. Of the four radio-detected blind-search pulsars, two remain radio-quiet whereas the other two could in principle have been discovered in a sensitive radio survey. The diagonal line in Figure 3 shows a possible alternate threshold at pseudo-luminosity $100 \mu\text{Jy-kpc}^2$, for reference. Three of the four have pseudo-luminosities lower than for any previously known young pulsar, and comparable to only a small number of MSPs.

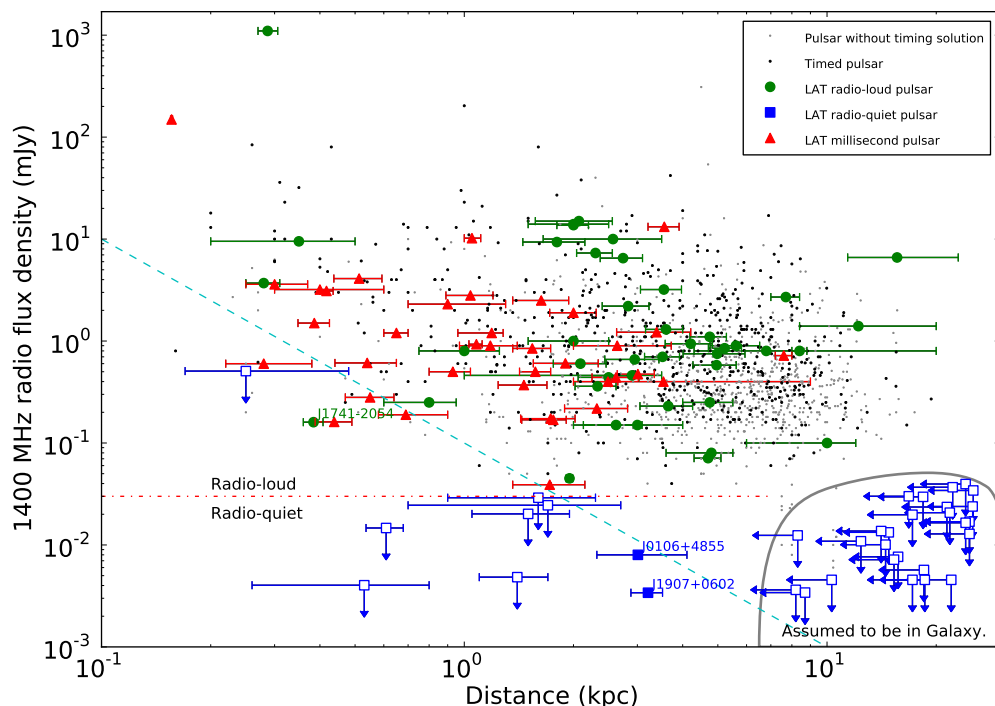


Fig. 3.— Radio flux density at 1400 MHz versus pulsar distance. Markers are as in Figure 1, except that blue open squares show pulsars discovered in gamma-ray blind period searches for which no radio signal has been detected. The horizontal line at $30 \mu\text{Jy}$ is our convention for distinguishing radio “loud” from radio “quiet” pulsars. The diagonal line shows a threshold in pseudo-luminosity of $100 \mu\text{Jy} \cdot \text{kpc}^2$. Four gamma-discovered pulsars have been detected at radio frequencies: two are radio-quiet and are labeled. Of the two that are radio loud, one is labeled while PSR J2032+4127 is in the cloud of points. The pulsars at lower-right are assigned distance limits along the Milky Way’s rim in Figure 4.

4.2. Distances

Converting measured pulsar fluxes to emitted luminosities L_γ (detailed in Section 6.3) requires the distances to the sources. Knowing the distances also allows mapping neutron star distributions relative to the Galaxy’s spiral arms, as in Figure 4, or evaluating their scale height above the plane. Several methods can be used to estimate pulsar distances; however, the methods vastly differ in reliability. Deciding which method to use can be subjective. Tables 5 and 6 list the distance estimates that we adopt, the methods with which these estimates were acquired, and the appropriate references.

The most accurate distance estimator is the annual trigonometric parallax. Unfortunately, parallax can only be measured for relatively nearby pulsars, using X-ray or optical images, radio interferometric imaging, or accurate timing. For 14 *Fermi* pulsars a parallax has been measured. We rejected two with low-significance ($< 2\sigma$). For the remaining 12 pulsars we consider this the best distance estimate. One caveat when converting parallax measurements to distances is the Lutz-Kelker effect, an overestimate of parallax values (and hence underestimate of distances) that must be corrected for the larger volume of space traced by smaller parallax values (Lutz & Kelker 1973). We use the Lutz-Kelker corrected distance estimates determined by Verbiest et al. (2012).

The dispersion measure (DM) is by far the most commonly used pulsar distance estimator. DM is the column density of free electrons along the path from Earth to the pulsar, in units of pc cm^{-3} . The electrons delay the radio pulse arrival by $\Delta t = \text{DM}(p\nu^2)^{-1}$ where ν is the observation frequency in MHz and $p = 2.410 \times 10^{-4} \text{ MHz}^{-2} \text{ pc cm}^{-3} \text{ s}^{-1}$. Given a model for the electron density n_e in the various structures of our Galaxy, integrating $\text{DM} = \int_0^d n_e dl$ along the line of sight dl yields the distance d for which DM matches the radio measurement. In this work we use the NE2001 model (Cordes & Lazio 2002), available as off-line code. To estimate the distance errors we re-run NE2001 twice, using $\text{DM} \pm 20\%$, as the authors recommend. The measured DM uncertainty is much lower than this, but this accommodates unmodelled electron-rich or poor regions. This yields distance uncertainties less than 30% for many pulsars. Nevertheless, significant discrepancies with the true pulsar distances along some lines of sight still occur. As examples, the DM distances for PSR J2021+3651 (Abdo et al. 2009d) and PSR J0248+6021 (Theureau et al. 2011) may be more than three times greater than the true distances.

For some pulsars, an absorbing hydrogen column density N_{H} below 1 keV has been obtained (see Section 9.1). Comparing N_{H} with the total hydrogen column density for that line of sight obtained from 21 cm radio surveys yields a rough distance estimate. The Doppler shift of neutral hydrogen (H I) absorption or emission lines measured from clouds on the line of sight, together with a Galactic rotation model as described in Section 4.3, can give

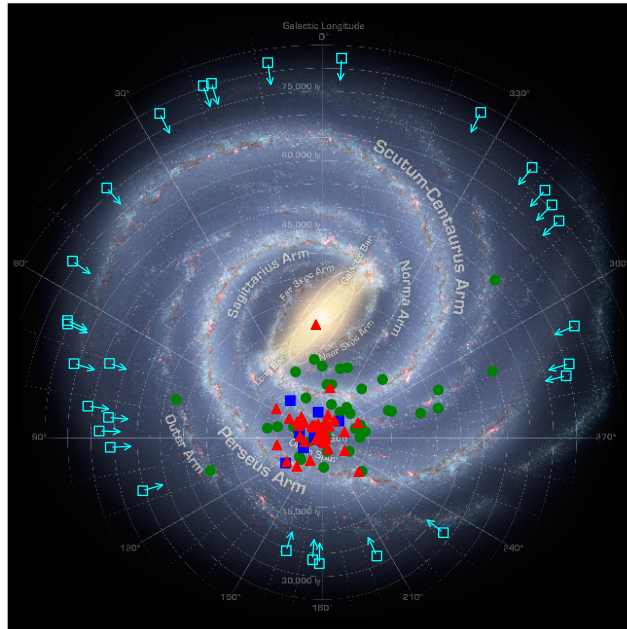


Fig. 4.— Gamma-ray pulsar positions projected onto the Milky Way model of Reid et al. (2009). The pulsar that appears to be coincident with the Galactic center, PSR J1823–3021A in the globular cluster NGC 6624, lies well above the Galactic plane. Distance uncertainties are not shown, for clarity, however they can be quite large especially for the more distant objects. The open squares with arrows indicate the lines of sight toward pulsars for which no distance estimates exist, placed at the distances where 95% of the electron column density has been integrated in the NE2001 model. The markers are the same as in Figure 1.

“kinematic” distances to the clouds. The pulsar distance is then constrained if there is evidence that the pulsar is in one of the clouds, or between some of them. Associations can be uncertain and these distance estimates can be controversial.

With the growing number of gamma-ray pulsars not detected at radio wavelengths, and thus without a DM, and the difficulties of the other methods, we face an ever-growing pulsar distance problem. We have 26 objects with no distance estimates, compared to nine in 1PC. To mitigate this, we determine a maximum distance by *assuming* that the pulsar is within the Galaxy. We define the Galaxy edge as the distance for a given line of sight where the NE2001 DM reaches its maximum value (“DMM” in Table 5, illustrated in Figure 4).

Table 5. Distance estimates for young LAT-detected pulsars

Pulsar Name	Distance (kpc)	Method	Reference
J0007+7303	1.4 ± 0.3	K	Pineault et al. (1993)
J0106+4855	$3.0^{+1.1}_{-0.7}$	DM	Pletsch et al. (2012b)
J0205+6449	1.95 ± 0.04	KP	Xu et al. (2006)
J0248+6021	2.0 ± 0.2	K	Theureau et al. (2011)
J0357+3205	< 8.2	DMM	...
J0534+2200	2.0 ± 0.5	O	Trimble (1973)
J0622+3749	< 8.3	DMM	...
J0631+1036	1.0 ± 0.2	O	Zepka et al. (1996)
J0633+0632	< 8.7	DMM	...
J0633+1746	$0.25^{+0.23}_{-0.08}$	P	Verbiest et al. (2012)
J0659+1414	0.28 ± 0.03	P	Verbiest et al. (2012)
J0729–1448	3.5 ± 0.4	DM	Morris et al. (2002)
J0734–1559	< 10.3	DMM	...
J0742–2822	2.1 ± 0.5	DM	Janssen & Stappers (2006)
J0835–4510	$0.29^{+0.02}_{-0.02}$	P	Dodson et al. (2003)
J0908–4913	2.6 ± 0.9	DM	Hobbs et al. (2004a)
J0940–5428	3.0 ± 0.5	DM	Manchester et al. (2001)
J1016–5857	$2.9^{+0.6}_{-1.9}$	K	Ruiz & May (1986)
J1019–5749	$6.8^{+13.2}_{-2.5}$	DM	Kramer et al. (2003)
J1023–5746	< 16.8	DMM	...
J1028–5819	2.3 ± 0.3	DM	Keith et al. (2008)
J1044–5737	< 17.2	DMM	...
J1048–5832	2.7 ± 0.4	DM	Johnston et al. (1995)
J1057–5226	0.3 ± 0.2	O	Mignani et al. (2010b)
J1105–6107	5.0 ± 1.0	DM	Kaspi et al. (1997)
J1112–6103	$12.2^{+7.8}_{-3.8}$	DM	Manchester et al. (2001)
J1119–6127	8.4 ± 0.4	K	Caswell et al. (2004)
J1124–5916	$4.8^{+0.7}_{-1.2}$	X	Gonzalez & Safi-Harb (2003)
J1135–6055	< 18.4	DMM	...
J1357–6429	$2.5^{+0.5}_{-0.4}$	DM	Lorimer et al. (2006)
J1410–6132	$15.6^{+7.4}_{-4.2}$	DM	O’Brien et al. (2008)
J1413–6205	< 21.4	DMM	...
J1418–6058	1.6 ± 0.7	O	Yadigaroglu & Romani (1997)
J1420–6048	5.6 ± 0.9	DM	Weltevrede et al. (2010)
J1429–5911	< 21.8	DMM	...
J1459–6053	< 22.2	DMM	...
J1509–5850	2.6 ± 0.5	DM	Weltevrede et al. (2010)
J1513–5908	4.2 ± 0.6	DM	Hobbs et al. (2004a)
J1531–5610	$2.1^{+0.4}_{-0.3}$	DM	Kramer et al. (2003)
J1620–4927	< 24.1	DMM	...
J1648–4611	5.0 ± 0.7	DM	Kramer et al. (2003)
J1702–4128	4.8 ± 0.6	DM	Kramer et al. (2003)
J1709–4429	2.3 ± 0.3	DM	Johnston et al. (1995)
J1718–3825	3.6 ± 0.4	DM	Manchester et al. (2001)
J1730–3350	$3.5^{+0.4}_{-0.5}$	DM	Hobbs et al. (2004b)
J1732–3131	0.6 ± 0.1	DM	Maan et al. (2012)

Table 5—Continued

Pulsar Name	Distance (kpc)	Method	Reference
J1741–2054	0.38 ± 0.02	DM	Camilo et al. (2009)
J1746–3239	< 25.3	DMM	...
J1747–2958	4.8 ± 0.8	X	Gaensler et al. (2004)
J1801–2451	$5.2^{+0.6}_{-0.5}$	DM	Hobbs et al. (2004b)
J1803–2149	< 25.2	DMM	...
J1809–2332	1.7 ± 1.0	K	Oka et al. (1999)
J1813–1246	< 24.7	DMM	...
J1826–1256	< 24.7	DMM	...
J1833–1034	4.7 ± 0.4	K	Gupta et al. (2005); Camilo et al. (2006)
J1835–1106	2.8 ± 0.4	DM	D’Amico et al. (1998)
J1836+5925	0.5 ± 0.3	X	Halpern et al. (2002)
J1838–0537	< 24.1	DMM	...
J1846+0919	< 22.0	DMM	...
J1907+0602	3.2 ± 0.3	DM	Abdo et al. (2010m)
J1952+3252	2.0 ± 0.5	K	Greidanus & Strom (1990)
J1954+2836	< 18.6	DMM	...
J1957+5033	< 14.5	DMM	...
J1958+2846	< 18.5	DMM	...
J2021+3651	$10.0^{+2.0}_{-4.0}$	O	Hessels et al. (2004)
J2021+4026	1.5 ± 0.4	K	Landecker et al. (1980)
J2028+3332	< 17.2	DMM	...
J2030+3641	3.0 ± 1.0	O	Camilo et al. (2012)
J2030+4415	< 15.7	DMM	...
J2032+4127	3.7 ± 0.6	DM	Camilo et al. (2009)
J2043+2740	1.8 ± 0.3	DM	Ray et al. (1996)
J2055+2539	< 15.3	DMM	...
J2111+4606	< 14.8	DMM	...
J2139+4716	< 14.1	DMM	...
J2229+6114	$0.80^{+0.15}_{-0.20}$	K	Kothes et al. (2001)
J2238+5903	< 12.4	DMM	...
J2240+5832	7.7 ± 0.7	O	Theureau et al. (2011)

Note. — The best known distances of the 77 young pulsars detected by *Fermi*. The methods are: K – kinematic method; P – parallax; DM – from dispersion measure using the Cordes & Lazio (2002) NE2001 model; X – from X-ray measurements ; O – other methods. For DM, the reference gives the DM measurement. For the 26 pulsars with no distance estimate, DMM is the distance to the Galaxy’s edge, taken as an upper limit, determined from the maximum NE2001 DM value for that line of sight.

Table 6. Millisecond pulsar distances and spindown Doppler corrections

PSR	d (pc)	Method	Ref ^b	μ (mas yr ⁻¹)	Ref ^c	\dot{P}^{int} (10 ⁻²¹)	\dot{P}^{shk} (10 ⁻²¹)	\dot{P}^{gal} (10 ⁻²¹)	\dot{E}^{int} (10 ³³ erg s ⁻¹)	ξ (%)
J0023+0923	690 ⁺²¹⁰ ₋₁₀	DM	(1)							
J0030+0451	280 ⁺¹⁶⁰ ₋₆₀	P	(2)	5.7 ± 1.1	(1)	10.7 ± 0.1	0.11	-0.60	3.64 ± 0.02	-5
J0034-0534	540 ± 100	DM	(3)	31.0 ± 9.0	(2)	2.9 ± 1.4	2.37	-0.31	17.3 ± 8.6	41
J0101-6422	550 ⁺⁹⁰ ₋₈₀	DM	(4)	15.6 ± 1.7	(3)	4.4 ± 0.2	0.84	-0.39	10.1 ± 0.5	9
J0102+4839	2320 ⁺⁵⁰⁰ ₋₄₃₀	DM	(1)							
J0218+4232	2640 ⁺¹⁰⁸⁰ ₋₆₄₀	DM	(5)	5.0 ± 6.0	(2)	76.9 ± 0.9	0.37	0.09	243.2 ± 2.8	0.6
J0340+4130	1730 ± 300	DM	(1)							
J0437-4715	156 ± 1	P	(2)	141.3 ± 0.1	(4)	14.1 ± 0.3	43.59	-0.40	2.9 ± 0.1	75
J0610-2100	3540 ⁺⁵⁴⁶⁰ ₋₁₀₀₀	DM	(6)	18.2 ± 0.2	(5)	1.2 ^{+17.0} _{-1.1}	11.00	0.10	0.8 ^{+11.7} _{-0.8}	90
J0613-0200	900 ⁺⁴⁰⁰ ₋₂₀₀	P	(2)	10.8 ± 0.2	(6)	8.7 ^{+0.3} _{-0.2}	0.77	0.08	12.0 ^{+0.5} _{-0.2}	9
J0614-3329	1900 ⁺³⁵⁰ ₋₂₀₀	DM	(7)							
J0751+1807	400 ⁺²⁰⁰ ₋₁₀₀	P	(2)	6.0 ± 2.0	(7)	7.7 ± 0.1	0.12	-0.02	7.2 ± 0.1	1
J1024-0719	390 ± 40	DM ^a	(8)	59.9 ± 0.2	(5)	1.6 ^{+1.8} _{-1.4}	17.37	-0.44	0.4 ^{+0.5} _{-0.4}	92
J1124-3653	1720 ⁺⁴³⁰ ₋₃₆₀	DM	(1)							
J1125-5825	2620 ± 370	DM	(9)							
J1231-1411	440 ± 50	DM	(7)	62.2 ± 4.7	(8) ^a	6.5 ± 2.9	15.15	-0.41	5.1 ± 2.3	70
J1446-4701	1460 ± 220	DM	(10)							
J1514-4946	940 ± 120	DM	(4)							
J1600-3053	1630 ⁺³¹⁰ ₋₂₇₀	DM	(11)	7.2 ± 0.3	(6)	8.6 ^{+0.2} _{-0.1}	0.74	0.15	7.3 ± 0.1	9
J1614-2230	650 ± 50	P	(12)	36.5 ± 0.2	(9)	3.0 ± 0.5	6.65	-0.003	3.8 ± 0.6	69
J1658-5324	930 ⁺¹¹⁰ ₋₁₃₀	DM	(4)							
J1713+0747	1050 ⁺⁶⁰ ₋₅₀	P	(2)	6.30 ± 0.01	(10)	8.28 ^{+0.03} _{-0.02}	0.46	-0.21	3.42 ± 0.01	3
J1741+1351	1080 ⁺⁴⁰ ₋₅₀	P	(13)	11.71 ± 0.01	(11)	29.1 ± 0.1	1.35	-0.20	21.76 ^{+0.04} _{-0.05}	4
J1744-1134	417 ± 17	P	(11)	21.02 ± 0.03	(6)	7.0 ± 0.1	1.82	0.08	4.11 ± 0.04	21
J1747-4036	3390 ± 760	DM	(4)							
J1810+1744	2000 ⁺³¹⁰ ₋₂₈₀	DM	(1)							
J1823-3021A	7600 ± 400	O	(14)							
J1858-2216	940 ⁺²⁰⁰ ₋₁₃₀	DM	(15)							
J1902-5105	1180 ± 210	DM	(4)							
J1939+2134	3560 ± 350	DM	(16)	0.80 ± 0.02	(12)	105.5 ± 0.1	0.01	-0.36	1096.6 ± 0.5	-0.3
J1959+2048	2490 ⁺¹⁶⁰ ₋₄₉₀	DM	(17)	30.4 ± 0.6	(13)	8.1 ^{+0.7} _{-1.8}	9.00	-0.25	76.3 ^{+6.4} _{-17.1}	52
J2017+0603	1570 ± 150	DM	(18)							
J2043+1711	1760 ⁺¹⁵⁰ ₋₃₂₀	DM	(1)	13.0 ± 2.0	(14)	4.3 ± 0.6	1.72	-0.35	12.7 ^{+1.6} _{-1.8}	24
J2047+1053	2050 ⁺³²⁰ ₋₂₉₀	DM	(15)							

Table 6—Continued

PSR	d (pc)	Method	Ref ^b	μ (mas yr ⁻¹)	Ref ^c	\dot{P}^{int} (10 ⁻²¹)	\dot{P}^{shk} (10 ⁻²¹)	\dot{P}^{gal} (10 ⁻²¹)	\dot{E}^{int} (10 ³³ erg s ⁻¹)	ξ (%)
J2051–0827	1040 ± 150	DM	(19)	7.3 ± 0.4	(15)	12.6 ± 0.1	0.61	-0.47	5.43 ± 0.05	1
J2124–3358	300 ⁺⁷⁰ ₋₅₀	P	(2)	52.3 ± 0.3	(6)	11.2 ^{+2.3} _{-1.6}	9.83	-0.46	3.7 ^{+0.8} _{-0.5}	46
J2214+3000	1540 ± 180	DM	(7)							
J2215+5135	3010 ⁺³³⁰ ₋₃₇₀	DM	(1)							
J2241–5236	510 ± 80	DM	(20)							
J2302+4442	1190 ⁺⁹⁰ ₋₂₃₀	DM	(18)							

Note. — Columns 2 and 3 give the distances for the 40 MSPs detected by *Fermi*, and the method used to find them: P – parallax; DM – from dispersion measure using the Cordes & Lazio (2002) NE2001 model; O – other methods. For DM, the references in Column 4 give the DM measurement. For the 20 MSPs with a proper motion measurement, it is listed in Column 5, obtained from the reference in Column 6. Columns 7 to 10 list the intrinsic \dot{P}^{int} , the contributions of the Shklovskii effect \dot{P}^{shk} and of acceleration due to the Galactic rotation, \dot{P}^{gal} , and the intrinsic spindown power \dot{E}^{int} . The relative correction ξ is defined from $\dot{E}^{\text{int}} = \dot{E}(1 - \xi)$.

^aProper motion for J1231–1411 from this work: $\mu_\alpha \cos \delta = -60.5(44)$ mas yr⁻¹, $\mu_\delta = -14.3(82)$ mas yr⁻¹. For J1024–0719, the timing parallax distance measurement gives negative spindown and we use instead the DM distance following Espinoza et al. (2013).

^bDistance and DM references: (1) Hessels et al. (2011); (2) Verbiest et al. (2012); (3) Bailes et al. (1994); (4) Kerr et al. (2012a); (5) Hobbs et al. (2004b); (6) Burgay et al. (2006); (7) Ransom et al. (2011); (8) Hotan et al. (2006); (9) Bates et al. (2011); (10) Keith et al. (2012); (11) Verbiest et al. (2009); (12) Lassus (2013); (13) Freire et al. (in prep.); (14) Kuulkers et al. (2003); (15) Ray et al. (2012); (16) Cognard et al. (1995); (17) Arzoumanian et al. (1994); (18) Cognard et al. (2011); (19) Doroshenko et al. (2001); (20) Keith et al. (2011).

^cProper motion references: (1) Abdo et al. (2009e); (2) Hobbs et al. (2005); (3) Kerr et al. (2012a); (4) Deller et al. (2009); (5) Espinoza et al. (2013); (6) Verbiest et al. (2009); (7) Nice et al. (2005); (8) This work; (9) Lassus (2013); (10) Splaver et al. (2005); (11) Freire et al. (in prep.); (12) Cognard et al. (1995); (13) Arzoumanian et al. (1994); (14) Guillemot et al. (2012); (15) Lazaridis et al. (2011).

4.3. Doppler corrections

Many pulsar characteristics, including some listed in Tables 2 and 3, depend on the *intrinsic* spin period P^{int} and spindown rate \dot{P}^{int} . The Doppler shift of the *observed* period is $P = (1 + v_{\text{R}}/c)P^{\text{int}}$, where v_{R} is the pulsar’s radial velocity along the unit vector \mathbf{n}_{10} from the solar system. The Doppler correction to \dot{P} is obtained by differentiating the equation and separating the effects of the system’s proper motion (Shklovskii 1970) from the acceleration due to Galactic rotation:

$$\dot{P}^{\text{int}} = \dot{P} - \dot{P}^{\text{shk}} - \dot{P}^{\text{gal}} \quad (3)$$

with

$$\dot{P}^{\text{shk}} = \frac{1}{c} \mu^2 d P = k \left(\frac{\mu}{\text{mas yr}^{-1}} \right)^2 \left(\frac{d}{\text{kpc}} \right) \left(\frac{P}{\text{s}} \right) \quad (4)$$

and

$$\dot{P}^{\text{gal}} = \frac{1}{c} \mathbf{n}_{10} \cdot (\mathbf{a}_1 - \mathbf{a}_0) P \quad (5)$$

where $k = 2.43 \times 10^{-21}$ for pulsar distance d and proper motion transverse to the line of sight μ . The Galactic potential model of Carlberg & Innanen (1987) and Kuijken & Gilmore (1989) provides the accelerations \mathbf{a}_1 of the pulsar and \mathbf{a}_0 of the Sun. Since the constant k is small, the corrections are negligible for the young gamma-ray pulsars, which all have $\dot{P} > 10^{-16}$. However, for MSPs $\mu^2 d$ can be large enough that \dot{P}^{int} differs noticeably from the *observed* \dot{P} values and quantities derived from \dot{P} will also be affected.

From the literature we compiled proper motion measurements for 243 pulsars, all but one of which also have distance estimates and \dot{P} measurements, and we calculated \dot{P}^{int} and its uncertainties for those 242 pulsars. Of these, 69 have $P < 30$ ms, and 20 are gamma-ray MSPs, listed in Table 6. The magnitude of the correction is $\xi = (\dot{P} - \dot{P}^{\text{int}})/\dot{P}$, or, equivalently, $\dot{E}^{\text{int}} = \dot{E}(1 - \xi)$ since $\dot{E} \propto \dot{P}$. For $|\xi|$ greater than a few percent, $\dot{P}^{\text{shk}} > |\dot{P}^{\text{gal}}|$ and corrected \dot{E}^{int} is *less* than the observed value. Hence flagging gamma-ray pulsar candidates that have large \dot{E} selects some with lower \dot{E}^{int} , but we would not miss candidates by neglecting proper motion. For large Doppler corrections, the Galactic term is negligible, $\dot{P}^{\text{shk}} \gg \dot{P}^{\text{gal}}$, and we calculate only the uncertainty due to \dot{P}^{shk} . Unless otherwise noted, throughout this paper we use \dot{P}^{int} from Table 6 to replace \dot{P} and the derived quantities (\dot{E} , τ , et cetera) in the Figures and Tables. In Section 6.3 we discuss the Doppler correction’s effect on the gamma-ray luminosity for a few cases.

5. Profile Characterization

5.1. Gamma-ray and Radio Light Curves

Appendix A contains a small sample of gamma-ray pulse profiles (Figures A-1 to A-8), overlaid with the radio profiles when available, and showing the fits described in Section 5.2, below. All pulse profiles are provided in the online material. We display gamma-ray light curves by computing a weighted histogram of gamma-ray rotational phases ϕ . The error on the i^{th} histogram bin containing N_i photons is estimated as $\sigma_i^2 = 1 + \sum_{j=0}^{N_i} w_j^2$, using the weights w_j defined in Eq. 1. The “1” term mitigates a bias towards low histogram levels and σ_i values caused by background-dominated bins. For these bins, the typical photon weight is very low, but the large weights of rare background photons near the pulsar position (within the point spread function) can substantially increase the bin level. The additional term compensates for the spread due to the presence or absence of a single such photon. We choose the number of bins in the histogram according to the weighted H-test: 25 bins ($H < 100$); 50 bins ($100 < H < 1000$); and 100 bins ($H > 1000$).

We estimate the background contribution from diffuse sources and neighboring point sources by computing the expectation value of w under the hypothesis that the photon does not originate from the pulsar: $b \equiv \int_0^1 dw w \times [1 - f(w)] \approx \sum_{j=0}^{N_\gamma} w_j - \sum_{j=0}^{N_\gamma} w_j^2$, where $f(w)$ is the probability distribution of the photon weights and we have used a Monte Carlo approximation to evaluate the integral. The corresponding background level for a weighted histogram, shown as a horizontal dashed line in the gamma-ray light curves, is b/N_{bins} . The dominant error in this quantity arises from systematic errors in the normalization of the diffuse background. We estimate this contribution by increasing/reducing the overall normalization of the background by 6% (see Section 6.1 for discussion.)

For pulsars with known radio profiles, we also display by preference the flux density at 1400 MHz. When 1400 MHz data are unavailable or highly scattered by the ISM, we include lower- or higher-frequency profiles, noting the frequency and provenance of the profile on the figures (Appendix A, and the online material).

The propagation of radio pulses through the dispersive ISM is delayed by $\Delta t \propto DM \nu^{-2}$. This delay, as well as Roemer-type delays associated with the configuration of the telescope relative to the solar system barycenter, are accounted for by TEMPO2, allowing precise alignment of radio and gamma-ray light curves. The absolute time at which $\phi \equiv 0$ is indicated in the pulsar par files, provided with the online material, by the parameters TZRMJD (giving the time of arrival of the zero phase), TZRFREQ (giving the frequency for which this time is correct), and TZRSITE (encoding the radio telescope/site of arrival).

The zero of phase—the fiducial phase—is ideally the rotational phase when the magnetic axis, the spin axis, and the line of sight lie in the same plane. For some pulsars, this phase can be identified by fitting the rotating vector model (Radhakrishnan & Cooke 1969) to radio polarization position angle versus phase. However, radio pulses often cover too narrow a phase interval to constrain fits using the rotating vector model.

If radio emission from the polar cap is symmetric, the peak intensity can also be used as a proxy for the fiducial phase, and we adopt this approach here for radio-loud pulsars. In detail, this approach can fail if there is appreciable asymmetry in the radio beam or profile evolution with frequency. Additionally, for pulsars with radio interpulses and for some MSPs, we observe emission from the field lines from both poles and must choose with which to associate the fiducial phase. In these cases, we generally choose the hemisphere furthest separated in phase from the gamma rays, consistent with an interpretation of gamma-ray emission arising from the outer magnetosphere. Finally, some pulsars (e.g., J0034–0534 shown in Figure A-1) exhibit a clear double symmetry in the radio light curve. In these cases, we choose a fiducial point near the point of symmetry. We note which prescription we have followed in Tables 7 and 8 with a ‘p’ (fiducial point at peak intensity), ‘h’ (fiducial point from hemisphere opposite to peak intensity), ‘s’ (fiducial point placed at point of symmetry rather than peak). For radio-quiet pulsars, we put the first gamma-ray peak (identified by looking for sharp rises and bridge emission) at $\phi = 0.1$ for display purposes.

5.2. Gamma-ray Light Curve Fitting

Generally, the gamma-ray light curve of a pulsar can be represented by a wrapped probability density function (pdf) of $\phi \in [0, 1)$. A compact approximation of the true pdf for photon phases can be constructed as a linear combination of N unimodal, possibly skew distributions:

$$f(\phi) = \sum_{i=1}^N n_i g_i(\phi) + \left(1 - \sum_{i=1}^N n_i\right), \quad (6)$$

where each of the g_i is an individually normalized distribution and $1 - \sum_{i=1}^N n_i \leq 1$ is a uniform distribution representing an unpulsed component. To explicitly enforce normalization, we use spherical polar coordinates lying within the unit sphere as internal parameters. With three components, e.g., the normalizations n_i are given in terms of the internal parameters χ_i as $n_1 = \sin \chi_1 \cos \chi_2$; $n_2 = \sin \chi_1 \sin \chi_2 \cos \chi_3$; $n_3 = \sin \chi_1 \sin \chi_2 \sin \chi_3$.

Although generalized wrapped distributions exist, analytic forms for such distributions are typically unavailable. Instead we adopt well-known pdfs, viz. the Gaussian (normal)

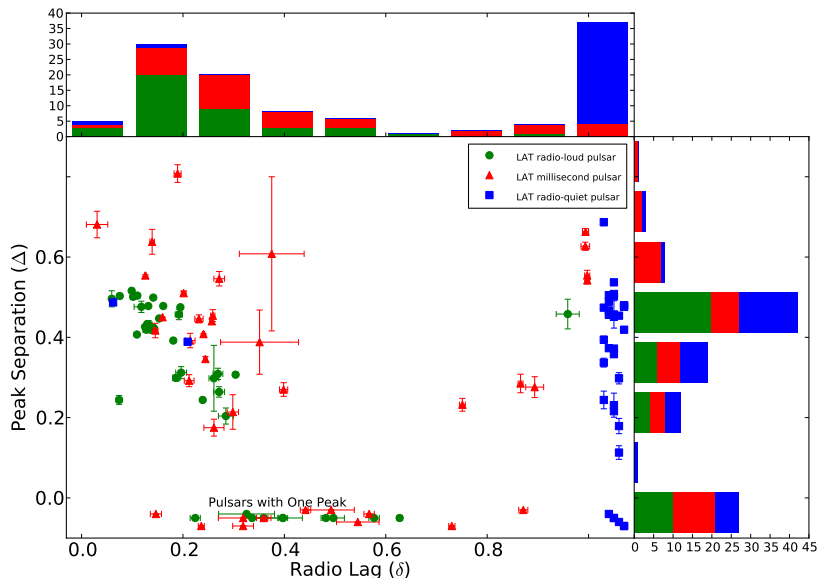


Fig. 5.— Phase lag δ of the gamma peak relative to the fiducial phase versus the phase separation Δ between the gamma-ray peaks. The artificial staggering of the points along the horizontal axis (single-peaked pulsars) and the right-hand vertical axis (pulsars with no radio detection) is to enhance clarity. The markers are the same as in Figure 1, as is the color code of the histograms projected onto the axes.

and Lorentzian (Cauchy) distributions, by wrapping them onto a circle:

$$g(\phi) = \sum_{i=-\infty}^{+\infty} g'(\phi + i), \quad (7)$$

where g' is defined on the real line and g on the circle. For practical reasons this sum must be truncated, and we define a truncated, wrapped distribution as

$$g_T(\phi, N) = \sum_{i=-N}^{+N} g'(\phi + i) + \left(1 - \int_{-N}^{+N} g'(x) dx\right), \quad (8)$$

i.e., we approximate the tails as a uniform distribution. In the fits discussed below, $N \equiv 10$.

Because the peaks of gamma-ray light curves may have a caustic origin, asymmetric distributions are needed to model their fast rise and slow fall. We generalize the symmetric Gaussian and Lorentzian distributions by matching two distributions with differing width parameters, σ_1 and σ_2 , at the maximum, x_0 . Defining $z \equiv (x - x_0)/\sigma$ with $\sigma = \sigma_1$ if $z \leq 0$

and $\sigma = \sigma_2$ otherwise, the functional forms of the resulting distributions on the real line are

$$g'(x) = \frac{2}{\pi(\sigma_1 + \sigma_2)(1 + z^2)} \text{ (Lorentzian),} \quad (9)$$

$$g'(x) = \sqrt{\frac{2}{\pi}} \frac{\exp(\frac{-z^2}{2})}{\sigma_1 + \sigma_2} \text{ (Gaussian),} \quad (10)$$

We employ maximum likelihood to determine the best-fit parameters of the mixture distribution. If w_i is a probability that a photon originates from the pulsar, then the logarithm of the likelihood is

$$\log \mathcal{L} = \sum_{j=1}^{N_\gamma} \log [w_i f(\phi_i) + (1 - w_i)]. \quad (11)$$

For large data sets ($N_\gamma > 10^4$) we speed up the computation by binning $f(\phi)$ to 512 values.

Most LAT light curves can be modelled with good fidelity by one or two two-sided narrow peak distributions and a single broad bridge component. From these fits, we determine the following quantities: N_{peaks} , the number of non-bridge components; δ , the offset of the mode of the leading peak from the fiducial phase (see above); and Δ , the difference between the modes of the leading and trailing peak (for $N_{\text{peaks}} > 1$). These parameters appear in Tables 7 and 8. The strong correlation between Δ and δ in Figure 5, as well as the dependence on spindown power (Figure 6) are discussed in Section 10. The peak widths are included in the online material.

The uncertainty on δ is estimated by combining in quadrature the statistical uncertainty on the position of the relevant gamma-ray peak with that incurred from uncertainty in the DM. The statistical uncertainty naturally includes a contribution from uncertainty in the timing solution (the TEMPO TRES quantity) which serves to smear out light curve features by a characteristic width $\delta\phi \approx \text{TRES}/P$. The statistical uncertainty on Δ is determined by the sum in quadrature of the position uncertainty of the relevant gamma-ray peaks.

The representation of the light curve (Gaussian vs. Lorentzian, presence or absence of additional components) affects the results for both δ and Δ . In addition to the uncertainties above, we estimate a “model” uncertainty of ~ 0.01 in phase.

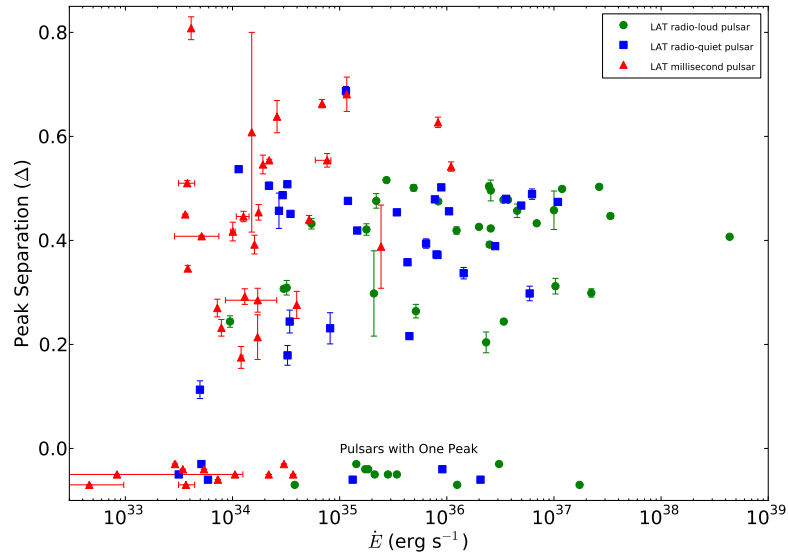


Fig. 6.— The phase separation Δ between gamma-ray peaks versus the Shlovskii-corrected spindown power. The markers are the same as in Figure 1.

Table 7. Pulse shape parameters of young LAT-detected pulsars

PSR ^a	Peaks	Radio lag δ	Shift Method	gamma-peak separation Δ	Off-peak definition ϕ
J0007+7303	2	...		0.216 ± 0.005	0.50 – 0.86
J0106+4855	2	0.062 ± 0.002	p	0.487 ± 0.003	0.14 – 0.50, 0.75 – 0.01
J0205+6449	2	0.075 ± 0.004	p	0.503 ± 0.004	0.64 – 0.99
J0248+6021	1	0.336 ± 0.017	p	...	0.61 – 0.18
J0357+3205	1	0.39 – 0.90
J0534+2200	2	0.109 ± 0.001	o	0.407 ± 0.001	0.71 – 0.99
J0622+3749	2	...		0.457 ± 0.034	0.37 – 0.52, 0.69 – 0.89
J0631+1036	1	0.497 ± 0.022	s	...	0.63 – 0.17
J0633+0632	2	...		0.476 ± 0.003	0.24 – 0.52, 0.67 – 1.00
J0633+1746	2	...		0.508 ± 0.001	0.85 – 0.95
J0659+1414	1	0.224 ± 0.010	p	...	0.40 – 0.04
J0729–1448	1	0.577 ± 0.010	p	...	0.70 – 0.47
J0734–1559	1	0.28 – 0.84
J0742–2822	1	0.627 ± 0.005	p	...	0.74 – 0.01, 0.01 – 0.48
J0835–4510	3	0.129 ± 0.001	p	0.433 ± 0.001	0.81 – 0.03
J0908–4913	2	0.102 ± 0.005	p	0.501 ± 0.006	0.66 – 0.04, 0.17 – 0.54
J0940–5428†	1	0.451 ± 0.035	s	...	0.59 – 0.19
J1016–5857	2	0.143 ± 0.003	s	0.423 ± 0.004	0.64 – 0.04
J1019–5749	1	0.482 ± 0.010	p	...	0.74 – 0.36
J1023–5746	2	...		0.474 ± 0.002	0.76 – 0.02
J1028–5819	2	0.195 ± 0.001	p	0.475 ± 0.001	0.75 – 0.08
J1044–5737	2	...		0.373 ± 0.004	0.56 – 0.98
J1048–5832	2	0.125 ± 0.001	s	0.426 ± 0.001	0.65 – 0.02
J1057–5226	3	0.304 ± 0.003	sh	0.307 ± 0.004	0.72 – 0.14
J1105–6107	2	0.110 ± 0.001	s	0.504 ± 0.006	0.74 – 0.04, 0.22 – 0.47
J1112–6103	2	0.192 ± 0.007	p	0.457 ± 0.013	0.79 – 0.04, 0.31 – 0.56
J1119–6127	2	0.285 ± 0.015	p	0.204 ± 0.020	0.59 – 0.18
J1124–5916	2	0.141 ± 0.003	p	0.499 ± 0.004	0.69 – 0.05
J1135–6055	1	0.49 – 0.92
J1357–6429	1	0.359 ± 0.028	p	...	0.64 – 0.13
J1410–6132	2	0.959 ± 0.023	p	0.458 ± 0.037	0.04 – 0.24, 0.55 – 0.85
J1413–6205	2	...		0.372 ± 0.003	0.57 – 0.01
J1418–6058*	2	...		0.467 ± 0.003	0.65 – 0.93
J1420–6048	2	0.196 ± 0.011	s	0.312 ± 0.015	0.63 – 0.12
J1429–5911	2	...		0.479 ± 0.004	0.28 – 0.40
J1459–6053*	1	0.59 – 0.91
J1509–5850	2	0.271 ± 0.011	p	0.264 ± 0.013	0.64 – 0.14
J1513–5908	1	0.325 ± 0.055	p	...	0.53 – 0.15
J1531–5610†	1	0.413 ± 0.035	s	...	0.57 – 0.23
J1620–4927	2	...		0.231 ± 0.030	0.51 – 0.95
J1648–4611	2	0.261 ± 0.010	p	0.298 ± 0.082	0.64 – 0.17
J1702–4128	1	0.397 ± 0.038	p	...	0.57 – 0.17
J1709–4429	2	0.239 ± 0.001	p	0.244 ± 0.002	0.72 – 0.07
J1718–3825	1	0.397 ± 0.009	s	...	0.64 – 0.09
J1730–3350	2	0.128 ± 0.007	p	0.419 ± 0.007	0.63 – 0.04, 0.20 – 0.38

Table 7—Continued

PSR ^a	Peaks	Radio lag δ	Shift Method	gamma-peak separation Δ	Off-peak definition ϕ
J1732–3131	2	...		0.419 ± 0.002	0.59 – 0.95
J1741–2054	2	0.074 ± 0.006	s	0.244 ± 0.011	0.44 – 0.96
J1746–3239	2	...		0.179 ± 0.019	0.44 – 0.02
J1747–2958	2	0.181 ± 0.003	p	0.392 ± 0.005	0.64 – 0.08
J1801–2451	2	0.060 ± 0.005	p	0.496 ± 0.020	0.13 – 0.48, 0.61 – 0.97
J1803–2149	2	...		0.394 ± 0.009	0.59 – 0.03
J1809–2332	2	...		0.358 ± 0.002	0.55 – 0.93
J1813–1246	2	...		0.489 ± 0.010	0.74 – 0.98
J1826–1256	2	...		0.480 ± 0.001	0.67 – 0.98
J1833–1034	2	0.153 ± 0.002	p	0.447 ± 0.004	0.69 – 0.10
J1835–1106	2	0.139 ± 0.006	p	0.421 ± 0.011	0.64 – 0.04
J1836+5925	2	...		0.537 ± 0.006	0.76 – 0.92
J1838–0537*	2	...		0.298 ± 0.014	0.51 – 0.01
J1846+0919	2	...		0.244 ± 0.022	0.43 – 0.95
J1907+0602	2	0.209 ± 0.003	p	0.389 ± 0.004	0.69 – 0.12
J1952+3252	3	0.161 ± 0.002	p	0.478 ± 0.003	0.71 – 0.06
J1954+2836	2	...		0.456 ± 0.004	0.67 – 0.02
J1957+5033	1	0.46 – 0.93
J1958+2846	2	...		0.454 ± 0.004	0.65 – 0.04
J2021+3651	2	0.132 ± 0.001	p	0.478 ± 0.001	0.73 – 0.99
J2021+4026	2	...		0.687 ± 0.009	0.24 – 0.40
J2028+3332	2	...		0.451 ± 0.003	0.57 – 0.97
J2030+3641	2	0.269 ± 0.010	p	0.309 ± 0.014	0.67 – 0.18
J2030+4415	2	...		0.505 ± 0.007	0.65 – 0.01
J2032+4127	2	0.099 ± 0.001	p	0.516 ± 0.001	0.22 – 0.55, 0.68 – 1.00
J2043+2740	2	0.132 ± 0.007	p	0.432 ± 0.010	0.63 – 0.05
J2055+2539*	2	...		0.113 ± 0.017	0.39 – 0.89
J2111+4606*	2	...		0.337 ± 0.011	0.53 – 0.01
J2139+4716	1	0.25 – 0.90
J2229+6114	2	0.187 ± 0.007	p	0.299 ± 0.008	0.68 – 0.10
J2238+5903	2	...		0.502 ± 0.002	0.68 – 0.02
J2240+5832	2	0.118 ± 0.014	p	0.476 ± 0.014	0.70 – 0.05, 0.16 – 0.46

^aA dagger (†) means the pulse profile fit is unreliable. A star (*) means that systematic offset uncertainties from the radio timing residuals (TRES variable in TEMPO) are between 10 and 19 milliseconds.

Note. — Column 2 gives the gamma-ray peak multiplicity. Columns 3 and 5 give the gamma-radio phase lag δ and separation Δ between the gamma-ray peaks. Column 4 gives the method used to define the radio fiducial phase: ‘p’ – peak radio intensity; ‘s’ – point of symmetry in the radio profile; ‘h’ – opposite hemisphere (0.5 phase shift from ‘p’ or ‘s’ point). PSR J0534+2200 (the Crab pulsar) is exceptional, and we align its profile with the low frequency component observed by Moffett & Hankins (1996) and denote this method ‘o’ for ‘other’. Column 6 gives the off-peak interval definition from Section 7.

6. Spectral Analyses

Most models of pulsar gamma-ray emission predict that the spectrum in the LAT energy range should be dominated by curvature radiation, in the radiation-reaction limited regime (e.g., Muslimov & Harding 2004). This mechanism predicts that pulsar spectra should be exponentially cut off near energies of a few GeV. The detection of pulsed emission above 100 GeV from the Crab (see Sections 8.4 and 10) suggests that, for some pulsars, either an additional component becomes dominant above the cutoff or a different mechanism, e.g., inverse Compton scattering, may be responsible for the LAT emission. For the purposes of this catalog, we assume that all gamma-ray pulsars have a cutoff spectrum but test for any deviations from this model and report the pulsars that exhibit a different spectral shape.

GeV emission from pulsars is largely modulated at the rotational period, with the emission concentrated in one or more narrow peaks. Depending on the viewing geometry and emission model, a pulsar can have a 100% duty cycle and significant magnetospheric emission can also exist away from the peaks. Young pulsars power PWNe, some of which are detected with the LAT, spatially overlapping their associated pulsars. PWN-like emission could be confused with the magnetospheric signal but would not be modulated at the rotational period. Significant contamination from a PWN-like background source needs to be accounted for to properly study the pulsar emission.

With this goal, we first analyzed the off-peak phase intervals as described in Section 7 and characterized significantly detected off-peak emission, to identify constant magnetospheric emission or PWN-like background emission. In this section, we use the data with all rotation phases to characterize the magnetospheric emission, removing background PWN-like contamination, and report the phase-averaged spectra.

6.1. Spectral Method

For spectral analysis we used the data set described in Section 2 for phase-averaged fits. We evaluated the point-source detection significance and the spectral parameters for each LAT pulsar in this catalog with an analysis similar to that performed for 2FGL using the standard LAT Science Tools package⁵. The spectrum of each LAT pulsar was modeled as a power law with an exponential cutoff,

$$\frac{dN}{dE} = K \left(\frac{E}{E_0} \right)^{-\Gamma} \exp \left(-\frac{E}{E_{\text{cut}}} \right)^b. \quad (12)$$

⁵<http://fermi.gsfc.nasa.gov/ssc/data/analysis/scitools/overview.html>

Table 8. Pulse shape parameters of LAT-detected millisecond pulsars

PSR ^a	Peaks	Radio lag δ	Shift Method	gamma-peak separation Δ	Off-peak definition ϕ
J0023+0923	2	0.375 ± 0.064	p	0.608 ± 0.192	0.60 – 0.80
J0030+0451	2	0.160 ± 0.001	p	0.450 ± 0.001	0.72 – 0.08
J0034–0534	2	0.866 ± 0.005	s	0.285 ± 0.023	0.28 – 0.75
J0101–6422	2	0.145 ± 0.005	sh	0.417 ± 0.018	0.85 – 0.09
J0102+4839	2	0.259 ± 0.004	p	0.454 ± 0.015	0.83 – 0.17, 0.32 – 0.59
J0218+4232	2	0.351 ± 0.077	s	0.388 ± 0.080	0.84 – 0.02, 0.02 – 0.23
J0340+4130	2	0.751 ± 0.006	p	0.232 ± 0.016	0.11 – 0.62
J0437–4715	1	0.442 ± 0.010	p	...	0.60 – 0.13
J0610–2100	1	0.236 ± 0.006	h	...	0.98 – 0.20
J0613–0200	2	0.261 ± 0.020	p	0.175 ± 0.021	0.60 – 0.10
J0614–3329	2	0.126 ± 0.002	s	0.554 ± 0.002	0.40 – 0.52
J0751+1807	3	0.398 ± 0.008	s	0.270 ± 0.017	0.77 – 0.03, 0.03 – 0.32
J1024–0719	1	0.492 ± 0.046	p	...	0.88 – 0.37
J1124–3653	2	0.298 ± 0.011	s	0.214 ± 0.043	0.88 – 0.19
J1125–5825†	1	0.645 ± 0.002	p	...	0.74 – 0.48
J1231–1411	3	0.241 ± 0.002	p	0.408 ± 0.002	0.72 – 0.12
J1446–4701	1	0.319 ± 0.021	p	...	0.65 – 0.21
J1514–4946	2	0.214 ± 0.009	s	0.392 ± 0.018	0.69 – 0.15
J1600–3053	1	0.147 ± 0.011	p	...	0.51 – 0.09
J1614–2230	2	0.201 ± 0.005	p	0.510 ± 0.005	0.78 – 0.14
J1658–5324	1	0.359 ± 0.014	s	...	0.64 – 0.26
J1713+0747	1	0.319 ± 0.049	p	...	0.67 – 0.00, 0.00 – 0.19
J1741+1351	1	0.730 ± 0.006	p	...	0.91 – 0.59
J1744–1134	2	0.189 ± 0.007	sh	0.808 ± 0.022	0.05 – 0.10
J1747–4036	2	0.031 ± 0.021	p	0.681 ± 0.033	0.24 – 0.44, 0.73 – 0.88
J1810+1744	2	0.894 ± 0.018	p	0.276 ± 0.026	0.23 – 0.68
J1823–3021A	2	0.993 ± 0.009	p	0.627 ± 0.010	0.08 – 0.55
J1858–2216†	1	0.727 ± 0.011	p	...	0.16 – 0.66
J1902–5105	2	0.994 ± 0.006	p	0.663 ± 0.008	0.70 – 0.93
J1939+2134	2	0.997 ± 0.004	p	0.542 ± 0.009	0.08 – 0.47, 0.59 – 0.95
J1959+2048	2	0.997 ± 0.005	p	0.554 ± 0.013	0.19 – 0.38
J2017+0603	2	0.212 ± 0.010	p	0.292 ± 0.015	0.60 – 0.12
J2043+1711	2	0.231 ± 0.008	p	0.446 ± 0.010	0.73 – 0.10
J2047+1053	1	0.567 ± 0.010	p	...	0.16 – 0.49
J2051–0827	1	0.545 ± 0.042	p	...	0.80 – 0.28
J2124–3358	1	0.871 ± 0.009	p	...	0.01 – 0.51
J2214+3000	2	0.271 ± 0.010	p	0.546 ± 0.018	0.09 – 0.24, 0.63 – 0.73
J2215+5135	2	0.257 ± 0.004	p	0.440 ± 0.008	...
J2241–5236	3	0.139 ± 0.004	p	0.638 ± 0.031	0.60 – 0.67
J2302+4442	2	0.244 ± 0.003	s	0.346 ± 0.006	0.70 – 0.18

^aA dagger (†) means the pulse profile fit is unreliable.

Note. — Column 2 gives the gamma-ray peak multiplicity. Columns 3 and 5 give the phase lag δ and peak separation Δ . Column 4 gives the method used to define the radio fiducial phase: ‘p’ – peak radio intensity; ‘s’ – point of symmetry in the radio profile; ‘h’ – opposite hemisphere (0.5 phase shift from ‘p’ or ‘s’ point). The off-peak phase interval (Section 7) is in column 6.

The four parameters are the normalization factor (K), the photon index at low energy (Γ), the cutoff energy (E_{cut}), and a parameter representing the sharpness of the cutoff (b) which is fixed to 1 in the default fit. The energy E_0 at which K is defined is arbitrary; thus, we used the 2FGL pivot energy when it exists and 1 GeV otherwise. In the likelihood analysis we constrained Γ to be between 0 and 5 and E_{cut} to be between 0.1 and 100 GeV.

We constructed models including all 2FGL sources within 20° of each pulsar but only the spectral parameters of point sources within 8° were left free. For sources known to be significantly extended we used the same spatial templates as 2FGL and fixed the spectral parameters to the 2FGL values. For pulsars with no 2FGL counterpart and an unassociated source within 0.1° of the pulsar, we moved the source to the timing position; otherwise, we added a new point source to the model at the timing position. The study of the off-peak phase interval in Section 7 yielded four pulsars with wind-like emission (‘W’-type), and four pulsars where the origin of the off-peak emission is unidentified (‘U’-type) but the emission appears spatially extended without evidence of a spectral cutoff, likely due to poorly modeled diffuse emission. We added a new point source to the model for each of these eight pulsars. We used the same models for diffuse gamma-ray emission as 2FGL to account for the Galactic, isotropic, and Earth limb emission: `gal_2yearp7v6_v0.fits`, `iso_p7v6source.txt`, `limb_2year_P76_source_v0_smooth.txt`, and `limb_smooth.fits`. These are available from the *Fermi* Science Support Center⁶.

For each pulsar, we selected a $20^\circ \times 20^\circ$ square region, centered on the timing position, for a binned maximum likelihood `gtlike` analysis using the *pyLikelihood* python module included with the *Fermi* Science Tools. The best-fit parameters are obtained by maximizing the log-likelihood surface that represents the input model using the MINUIT2 fitting engine⁷. The statistical uncertainties on the parameters were estimated from the quadratic development of the log-likelihood around the best fit. We first performed a fit with a weak convergence criterion and evaluated the point-source detection significance (the “Test Statistic” TS , Mattox et al. 1996) for each point source that had free parameters, except the pulsar of interest. We removed all point sources with $TS < 2$ and re-optimized the fit with a stricter convergence criterion.

We report the maximum likelihood values of Γ and E_{cut} from the phase-averaged analysis and TS values in Tables 9 and 10 for young and millisecond pulsars, respectively. In addition, we report the integrated photon and energy fluxes in the 0.1 to 100 GeV energy band (F_{100}

⁶<http://fermi.gsfc.nasa.gov/ssc/data/access/lat/BackgroundModels.html>

⁷http://project-mathlibs.web.cern.ch/project-mathlibs/sw/5_15_04/Minuit2/html

and G_{100} , respectively), obtained from the fits as

$$F_{100} = \int_{100 \text{ MeV}}^{100 \text{ GeV}} \frac{dN}{dE} dE, \quad G_{100} = \int_{100 \text{ MeV}}^{100 \text{ GeV}} E \frac{dN}{dE} dE. \quad (13)$$

Statistical uncertainties on F_{100} and G_{100} are obtained using derivatives with respect to the primary parameters and the covariance matrix obtained from the fitting process.

We tested the validity of modeling the pulsar spectrum as a power law with a simple exponential cutoff shape (PLEC1, Eq. 12 with $b \equiv 1$) by repeating the analysis using a pure power-law shape (PL), and a power law with a more general exponential cutoff shape leaving the b parameter free (PLEC). For a number of pulsars, a PLEC1 spectral model is not significantly better than a PL. We identified these by computing $TS_{\text{cut}} \equiv 2\Delta \log(\textit{likelihood})$ (comparable to a χ^2 distribution with one degree of freedom) between the models with and without the cutoff. We say that the PLEC1 model is not significantly preferred over the PL model for pulsars with $TS_{\text{cut}} < 9$, listed in Tables 9 and 10. Similarly, we calculated $TS_{\text{bfree}} \equiv 2\Delta \log(\textit{likelihood})$ between the PLEC1 and PLEC models, also listed in the Tables.

In all cases where PLEC is significantly preferred over the PLEC1 model ($TS_{\text{bfree}} \geq 9$), the maximum likelihood value of b is significantly less than 1, indicating a sub-exponential cutoff. As noted by Abdo et al. (2010o) and Celik & Johnson (2011), a sub-exponential cutoff is a functional form which approximates the superposition of several PLEC1 models with varying values of Γ and E_{cut} as different regions of the pulsar magnetosphere cut across our line of sight. This is further supported by analysis of PSR J1057–5226 for which the b parameter in a PLEC fit is consistent with 1 and the spectral parameters show very little variation with phase (Abdo et al. 2010c). Thus, no physical quantities can be derived from the PLEC best-fit parameters, whereas the PLEC1 best-fit parameters can be used if taken as flux-weighted, average measures of E_{cut} and Γ . Further, in our PLEC fits, the E_{cut} value was often at the minimum boundary (0.1 GeV) with unrealistically small uncertainties. Therefore, we do not report the PLEC fit values in the Tables and instead indicate pulsars with high TS_{bfree} as interesting candidates for phase-resolved spectral analysis, a task beyond the scope of this catalog. For pulsars with $TS_{\text{bfree}} \geq 9$, the PLEC fit results are included in the spectral plots (see the examples amongst Figures A-9 to A-17 in Appendix A) and in the auxiliary files (Appendix B). We encourage comparing predicted spectra from different emission models to the energy sub-band fluxes, which are also included in the auxiliary files. In some pulsars, particularly the Crab, a preference for $b < 1$ may indicate either the presence of a secondary spectral component that dominates above $\gtrsim 10$ GeV or that the curvature radiation assumption is incorrect (as argued by Lyutikov et al. 2012, for example). Such a determination is difficult using the LAT data alone (see Section 8.4).

Table 9. Spectral fitting results for young LAT-detected pulsars

PSR ^a	Photon Flux ($\text{ph cm}^{-2} \text{s}^{-1}$) ($\times 10^{-8}$)	Energy Flux ($\text{erg cm}^{-2} \text{s}^{-1}$) ($\times 10^{-11}$)	Γ	E_{cut} (GeV)	TS	TS_{cut}	$TS_{\text{b, free}}$	Luminosity ($10^{33} \text{ erg s}^{-1}$)	Efficiency ^b (%)
J0007+7303	32.9 ± 0.4	40.1 ± 0.4	1.4 ± 0.1	4.7 ± 0.2	43388	1884	4	$94 \pm 1 \pm 40$	$21.0 \pm 0.2 \pm 8$
J0106+4855	1.7 ± 0.4	1.9 ± 0.2	1.2 ± 0.2	2.7 ± 0.6	544	58	2	$21 \pm 2^{+20}_{-8}$	$71 \pm 7^{+60}_{-30}$
J0205+6449	10.5 ± 0.7	5.4 ± 0.2	1.8 ± 0.1	1.6 ± 0.3	1019	86	7	$24 \pm 1 \pm 1.0$	$0.09 \pm 0.01 \pm 0.01$
J0248+6021	9.9 ± 1.3	5.2 ± 0.4	1.8 ± 0.1	1.6 ± 0.3	578	61	0	$25 \pm 2 \pm 5$	$12 \pm 1 \pm 2$
J0357+3205	9.0 ± 0.4	6.4 ± 0.2	1.0 ± 0.1	0.8 ± 0.1	3468	461	2
J0534+2200	208 ± 1	129.3 ± 0.8	1.9 ± 0.1	4.2 ± 0.2	102653	1461	13	$619 \pm 4 \pm 300$	$0.14 \pm 0.01 \pm 0.1$
J0622+3749	2.0 ± 0.3	1.4 ± 0.1	0.6 ± 0.4	0.6 ± 0.1	302	91	0
J0631+1036	6.4 ± 0.6	4.7 ± 0.3	1.8 ± 0.1	6 ± 1	621	39	1	$5.6 \pm 0.3^{+3}_{-2}$	$3.2 \pm 0.2^{+2}_{-1}$
J0633+0632	9.7 ± 1.1	9.4 ± 0.5	1.4 ± 0.1	2.7 ± 0.3	2448	203	9
J0633+1746	416 ± 1	423.3 ± 1.2	1.2 ± 0.1	2.2 ± 0.1	906994	33861	277	$31.7 \pm 0.1^{+90}_{-20}$	$97.4 \pm 0.3^{+300}_{-50}$
J0659+1414	7.1 ± 0.6	2.5 ± 0.2	1.7 ± 0.5	0.4 ± 0.2	419	33	0	$0.24 \pm 0.02 \pm 0.05$	$0.62 \pm 0.04 \pm 0.1$
J0729-1448 †	54	26	0
J0734-1559	10.8 ± 0.7	5.6 ± 0.2	2.0 ± 0.1	3.2 ± 0.9	916	39	9
J0742-2822	3.2 ± 0.6	1.7 ± 0.2	1.7 ± 0.3	1.6 ± 0.8	112	11	2	$9 \pm 1 \pm 4$	$6.2 \pm 0.7 \pm 3$
J0835-4510	1088 ± 2	906 ± 2	1.5 ± 0.1	3.0 ± 0.1	1659005	43084	916	$89.3 \pm 0.2 \pm 10$	$1.3 \pm 0.1 \pm 0.1$
J0908-4913	7.9 ± 1.3	4.4 ± 0.4	1.0 ± 0.4	0.5 ± 0.2	315	82	0	$35 \pm 3^{+30}_{-20}$	$7.1 \pm 0.7^{+6}_{-4}$
J0940-5428 †	14	13	8
J1016-5857	6.9 ± 2.4	5.4 ± 0.9	1.8 ± 0.2	6 ± 3	290	13	0	$55 \pm 9^{+30}_{-50}$	$2.1 \pm 0.4^{+1}_{-2}$
J1019-5749 †	21	0	0
J1023-5746	30 ± 3	19.5 ± 1.2	1.7 ± 0.1	2.5 ± 0.4	2926	162	20
J1028-5819	31 ± 2	24.3 ± 0.8	1.7 ± 0.1	4.6 ± 0.5	5096	235	28	$158 \pm 5 \pm 40$	$18.9 \pm 0.6 \pm 5$
J1044-5737	26 ± 1	15.6 ± 0.5	1.8 ± 0.1	2.8 ± 0.3	3380	202	19
J1048-5832	25 ± 2	19.6 ± 0.6	1.6 ± 0.1	3.0 ± 0.3	5389	325	30	$176 \pm 5 \pm 40$	$8.8 \pm 0.3 \pm 2$
J1057-5226	32 ± 1	29.5 ± 0.3	1.0 ± 0.1	1.4 ± 0.1	27848	2377	5	$4.3 \pm 0.1^{+5}_{-3}$	$14.4 \pm 0.2 \pm 10$
J1105-6107	7.8 ± 1.6	4.9 ± 0.6	1.5 ± 0.3	1.3 ± 0.6	309	42	8	$150 \pm 20 \pm 50$	$5.9 \pm 0.7 \pm 2$
J1112-6103	1.9 ± 0.9	2.0 ± 0.5	1.6 ± 0.3	6 ± 3	58	6	0	$360 \pm 90^{+600}_{-200}$	$8 \pm 2^{+10}_{-4}$
J1119-6127	11 ± 2	7.1 ± 0.5	1.8 ± 0.1	3.2 ± 0.8	661	37	13	$600 \pm 40 \pm 60$	$26 \pm 2 \pm 2$
J1124-5916	10 ± 1	6.2 ± 0.4	1.8 ± 0.1	2.1 ± 0.4	1058	79	6	$170 \pm 10^{+50}_{-70}$	$1.4 \pm 0.1^{+0.4}_{-0.6}$
J1135-6055	7.4 ± 0.9	4.8 ± 0.3	1.7 ± 0.1	2.4 ± 0.5	498	61	3
J1357-6429	7.8 ± 1.1	3.4 ± 0.3	1.8 ± 0.4	0.9 ± 0.5	187	20	0	$25 \pm 2^{+10}_{-8}$	$0.82 \pm 0.08^{+0.4}_{-0.3}$
J1410-6132 †	3 ± 3	3 ± 1	40	9	0	$800 \pm 300^{+900}_{-400}$	$8 \pm 3^{+9}_{-4}$
J1413-6205	16 ± 2	15.7 ± 0.6	1.5 ± 0.1	4.1 ± 0.5	1795	180	1
J1418-6058	38 ± 3	30.2 ± 1.4	1.8 ± 0.1	5.5 ± 0.5	3487	172	1	$92 \pm 4^{+100}_{-60}$	$1.9 \pm 0.1^{+2}_{-1}$

Table 9—Continued

PSR ^a	Photon Flux ($\text{ph cm}^{-2} \text{s}^{-1}$) ($\times 10^{-8}$)	Energy Flux ($\text{erg cm}^{-2} \text{s}^{-1}$) ($\times 10^{-11}$)	Γ	E_{cut} (GeV)	TS	TS_{cut}	$TS_{\text{b, free}}$	Luminosity ($10^{33} \text{ erg s}^{-1}$)	Efficiency ^b (%)
J1420-6048	26 ± 3	17.0 ± 1.4	1.5 ± 0.1	1.6 ± 0.2	1220	51	2	640 ± 50 ± 200	6.2 ± 0.5 ± 2
J1429-5911	12 ± 1	8.0 ± 0.4	1.6 ± 0.1	2.2 ± 0.3	822	124	0
J1459-6053	24 ± 2	12.9 ± 0.4	2.0 ± 0.1	2.9 ± 0.5	2046	103	15
J1509-5850	20 ± 2	12.7 ± 0.7	1.9 ± 0.1	4.6 ± 0.9	1152	67	0	105 ± 6 ± 40	20 ± 1 ± 7
J1513-5908 †	10 ± 2	3.2 ± 0.7	98	5	1	70 ± 10 ± 20	0.4 ± 0.1 ± 0.1
J1531-5610 †*	0.1 ± 0.05	0.2 ± 0.1	2	2	1	1.0 ± 0.6 ± 0.4	0.11 ± 0.07 ± 0.04
J1620-4927	15 ± 2	15.7 ± 0.8	1.3 ± 0.1	2.5 ± 0.3	1407	199	4
J1648-4611	5.3 ± 1.6	5.4 ± 0.8	1.6 ± 0.3	6 ± 4	176	17	0	160 ± 20 ± 40	80 ± 10 ± 20
J1702-4128 *	4.4 ± 5.2	2.8 ± 2.7	1.1 ± 0.9	0.8 ± 0.5	62	7	0	80 ± 70 ± 20	20 ± 20 ± 5
J1709-4429	160 ± 2	135 ± 1	1.6 ± 0.1	4.2 ± 0.1	96893	3433	132	853 ± 6 ± 200	25.1 ± 0.2 ± 5
J1718-3825	14 ± 1	8.9 ± 0.5	1.5 ± 0.1	1.4 ± 0.1	462	81	1	138 ± 8 ± 30	11.0 ± 0.6 ± 3
J1730-3350 *	3.4 ± 0.1	2.4 ± 0.4	1.5 ± 0.3	1.2 ± 0.3	100	21	1	36 ± 6 ± 9	2.9 ± 0.5 ± 0.7
J1732-3131	17 ± 1	19.4 ± 0.6	1.0 ± 0.1	1.9 ± 0.1	2821	550	0	8.6 ± 0.3 ± 2	5.9 ± 0.2 ± 1
J1741-2054	17 ± 1	11.7 ± 0.4	1.1 ± 0.1	0.9 ± 0.1	3014	464	0	2.1 ± 0.1 ± 0.2	22 ± 1 ± 3
J1746-3239	10 ± 1	7.2 ± 0.5	1.4 ± 0.1	1.5 ± 0.2	654	109	0
J1747-2958	33 ± 3	21.1 ± 1.0	1.6 ± 0.1	1.9 ± 0.1	1689	211	2	570 ± 30 ± 200	23 ± 1 ± 7
J1801-2451 *	1.4 ± 1.0	1.2 ± 0.5	1.5 ± 0.5	3 ± 2	58	10	6	40 ± 10 ⁺⁹ ₋₇	1.5 ± 0.6 ^{+0.4} _{-0.3}
J1803-2149	11 ± 2	9.2 ± 0.8	1.6 ± 0.1	3.6 ± 0.8	410	40	1
J1809-2332	60 ± 2	47.4 ± 0.8	1.6 ± 0.1	3.4 ± 0.2	15781	901	41	164 ± 3 ⁺²⁰⁰ ₋₁₀₀	38 ± 1 ⁺⁶⁰ ₋₃₀
J1813-1246	45 ± 2	25.3 ± 0.6	1.9 ± 0.1	2.6 ± 0.3	4664	272	1
J1826-1256	54 ± 3	38.3 ± 0.9	1.6 ± 0.1	2.2 ± 0.2	5160	533	23
J1833-1034	7.0 ± 1.1	5.9 ± 0.5	0.9 ± 0.2	0.9 ± 0.2	258	78	0	160 ± 10 ± 30	0.46 ± 0.04 ± 0.08
J1835-1106 †*	0.4 ± 0.1	0.6 ± 0.2	30	18	0	6 ± 2 ± 2	3 ± 1 ± 0.9
J1836+5925	63 ± 1	60.6 ± 0.4	1.2 ± 0.1	2.0 ± 0.1	142427	5747	28	20.4 ± 0.1 ⁺³⁰ ₋₂₀	180 ± 1 ⁺²⁰⁰ ₋₁₀₀
J1838-0537	22 ± 2	18.8 ± 0.9	1.6 ± 0.1	4.1 ± 0.4	1325	114	1
J1846+0919	1.4 ± 0.3	2.4 ± 0.2	0.7 ± 0.3	2.2 ± 0.5	428	79	0
J1907+0602	34 ± 2	25.4 ± 0.6	1.6 ± 0.1	2.9 ± 0.3	3773	390	35	314 ± 8 ± 60	11.1 ± 0.3 ± 2
J1952+3252	17 ± 1	13.8 ± 0.3	1.5 ± 0.1	2.5 ± 0.2	4469	365	11	66 ± 2 ⁺⁴⁰ ₋₃₀	1.8 ± 0.1 ^{+1.0} _{-0.8}
J1954+2836	13 ± 1	10.3 ± 0.4	1.6 ± 0.1	3.3 ± 0.4	1592	168	11
J1957+5033	4.0 ± 0.4	2.6 ± 0.1	1.3 ± 0.2	1.0 ± 0.2	846	105	0
J1958+2846	11 ± 1	9.1 ± 0.4	1.4 ± 0.1	2.0 ± 0.3	1519	206	25
J2021+3651	68 ± 2	49.4 ± 0.8	1.7 ± 0.1	3.0 ± 0.2	17821	998	23	5910 ± 90 ⁺³⁰⁰⁰ ₋₄₀₀₀	175 ± 3 ⁺⁸⁰ ₋₁₀₀
J2021+4026	138 ± 2	95.5 ± 0.9	1.6 ± 0.1	2.6 ± 0.1	53955	2343	72	257 ± 2 ⁺²⁰⁰ ₋₁₀₀	225 ± 2 ⁺¹⁰⁰ ₋₁₀₀

Table 9—Continued

PSR ^a	Photon Flux ($\text{ph cm}^{-2} \text{s}^{-1}$) ($\times 10^{-8}$)	Energy Flux ($\text{erg cm}^{-2} \text{s}^{-1}$) ($\times 10^{-11}$)	Γ	E_{cut} (GeV)	TS	TS_{cut}	$TS_{\text{b, free}}$	Luminosity ($10^{33} \text{ erg s}^{-1}$)	Efficiency ^b (%)
J2028+3332	5.8 ± 0.9	5.8 ± 0.4	1.2 ± 0.2	1.9 ± 0.3	1058	161	0
J2030+3641	2.3 ± 0.6	3.1 ± 0.3	0.7 ± 0.4	1.5 ± 0.4	313	91	0	$34 \pm 4^{+30}_{-20}$	$110 \pm 10^{+80}_{-60}$
J2030+4415	9.3 ± 1.0	5.8 ± 0.4	1.6 ± 0.1	1.7 ± 0.3	504	74	0
J2032+4127	7.4 ± 1.1	10.6 ± 0.6	1.1 ± 0.1	3.2 ± 0.5	1383	162	0	$169 \pm 10 \pm 50$	$62 \pm 4 \pm 20$
J2043+2740	1.5 ± 0.4	1.0 ± 0.1	1.4 ± 0.4	1.2 ± 0.6	97	18	2	$3.8 \pm 0.6 \pm 1$	$7 \pm 1 \pm 2$
J2055+2539	6.2 ± 0.4	5.4 ± 0.2	1.0 ± 0.1	1.1 ± 0.1	2751	361	1
J2111+4606	5.3 ± 0.7	4.4 ± 0.3	1.7 ± 0.1	5 ± 1	731	45	0
J2139+4716	3.1 ± 0.5	2.3 ± 0.2	1.3 ± 0.2	1.3 ± 0.3	369	71	2
J2229+6114	38 ± 1	25.3 ± 0.4	1.8 ± 0.1	4.3 ± 0.3	12101	424	54	$19.4 \pm 0.3 \pm 8$	$0.09 \pm 0.01 \pm 0.04$
J2238+5903	9.3 ± 0.9	6.4 ± 0.3	1.6 ± 0.1	2.1 ± 0.3	1165	123	10
J2240+5832	1.1 ± 0.7	1.1 ± 0.3	1.5 ± 0.5	3 ± 2	54	11	0	$80 \pm 20 \pm 10$	$40 \pm 10 \pm 6$

^aA dagger (†) means the spectral fit is unreliable (see text). A star (*) means that the spectrum was calculated using the on-peak data only.

^bOverestimated distances or the assumed beaming factor, $f_{\Omega} = 1$, can result in an efficiency $> 100\%$.

Note. — Unbinned maximum likelihood spectral fit results for the young LAT gamma-ray pulsars, using the PLEC1 model (Eq. 12 in Section 6). Columns 2 and 3 list the phase-averaged integral photon and energy fluxes in the 0.1 to 100 GeV energy band, F_{100} and G_{100} . Columns 4 and 5 list the photon index Γ and cutoff energy E_{cut} . Columns 6, 7, and 8 list the source significance TS , significance TS_{cut} of the exponential cutoff compared to a simple power law, and significance $TS_{\text{b, free}}$ of the PLEC compared to a PLEC1 shape. A value $TS_{\text{cut}} < 9$ indicates that a spectral cutoff is not significantly detected. Columns 9 and 10 give the total gamma-ray luminosity L_{γ} in the 0.1 to 100 GeV energy band, and the gamma-ray conversion efficiency $\eta \equiv L_{\gamma}/\dot{E}$, assuming $f_{\Omega} = 1$ as described in Section 6.3. The first uncertainty in L_{γ} and η comes from the statistical uncertainties in the spectral fit, whereas the second is due to the distance uncertainty. The strong dependence of these quantities on distance (see Table 5) and beaming factor means that these values should be considered with care.

Table 10. Spectral fitting results for LAT-detected millisecond pulsars

PSR ^a	Photon Flux ($\text{ph cm}^{-2} \text{s}^{-1}$) ($\times 10^{-8}$)	Energy Flux ($\text{erg cm}^{-2} \text{s}^{-1}$) ($\times 10^{-11}$)	Γ	E_{cut} (GeV)	TS	TS_{cut}	$TS_{\text{b free}}$	Luminosity ($10^{32} \text{ ergs}^{-1}$)	Efficiency ^b (%)
J0023+0923	1.2 ± 0.4	0.80 ± 0.12	1.4 ± 0.4	1.4 ± 0.6	131	16	7	$4.6 \pm 0.7_{-1}^{+3}$	$3.0 \pm 0.5_{-1}^{+2}$
J0030+0451	6.6 ± 0.3	6.14 ± 0.18	1.2 ± 0.1	1.8 ± 0.2	4788	316	10	$5.8 \pm 0.2_{-2}^{+5}$	$16 \pm 1_{-6}^{+13}$
J0034-0534	2.2 ± 0.3	1.62 ± 0.12	1.4 ± 0.2	1.8 ± 0.4	563	45	0	$5.7 \pm 0.4_{-2}^{+3}$	$3.3 \pm 0.2_{-1.1}^{+1.5}$
J0101-6422	0.75 ± 0.14	1.05 ± 0.09	0.7 ± 0.3	1.5 ± 0.4	491	59	1	$3.8 \pm 0.3 \pm 1$	$3.8 \pm 0.3_{-1.0}^{+1.3}$
J0102+4839	1.3 ± 0.3	1.32 ± 0.16	1.4 ± 0.3	3.2 ± 1.1	251	29	1	$90 \pm 10_{-30}^{+40}$	$49 \pm 6_{-19}^{+23}$
J0218+4232	7.7 ± 0.7	4.56 ± 0.24	2.0 ± 0.1	4.6 ± 1.2	1313	38	1	$380 \pm 20_{-200}^{+300}$	$16 \pm 1_{-7}^{+7}$
J0340+4130	1.5 ± 0.2	2.04 ± 0.15	1.1 ± 0.2	2.6 ± 0.6	553	73	4	$73 \pm 6 \pm 20$	$92.6 \pm 7.0 \pm 29$
J0437-4715	2.7 ± 0.3	1.67 ± 0.11	1.4 ± 0.2	1.1 ± 0.3	687	67	1	$0.49 \pm 0.03 \pm 0.01$	$1.7 \pm 0.1 \pm 0.1$
J0610-2100	0.78 ± 0.25	0.66 ± 0.11	1.2 ± 0.4	1.6 ± 0.8	98	14	1	$100 \pm 20_{-50}^{+500}$	$1180 \pm 190_{-560}^{+6400}$
J0613-0200	2.7 ± 0.4	2.99 ± 0.19	1.2 ± 0.2	2.5 ± 0.5	760	88	0	$29 \pm 2_{-10}^{+30}$	$24.1 \pm 1.5_{-0.5}^{+2.0}$
J0614-3329	8.5 ± 0.3	10.94 ± 0.27	1.3 ± 0.1	3.9 ± 0.3	9408	416	11	$470 \pm 10 \pm 200$	$215 \pm 5_{-72}^{+110}$
J0751+1807	1.1 ± 0.2	1.33 ± 0.12	1.1 ± 0.2	2.6 ± 0.7	427	46	1	$2.5 \pm 0.2_{-1}^{+3}$	$3.5 \pm 0.3_{-1.5}^{+4.4}$
J1024-0719 †	0.2 ± 0.2	0.3 ± 0.1	46	11	0	$0.6 \pm 0.2 \pm 0.1$	$12 \pm 4 \pm 2.3$
J1124-3653	0.94 ± 0.23	1.21 ± 0.13	1.1 ± 0.3	2.5 ± 0.7	293	39	0	$43 \pm 5 \pm 20$	$25.0 \pm 2.7_{-3.4}^{+4.4}$
J1125-5825	1.1 ± 0.5	0.89 ± 0.20	1.7 ± 0.2	4.8 ± 2.4	41	4	1	$70 \pm 20 \pm 20$	$9.1 \pm 2.0_{-2.4}^{+3.0}$
J1231-1411	9.2 ± 0.4	10.28 ± 0.26	1.2 ± 0.1	2.7 ± 0.2	7931	446	3	$24 \pm 0.6 \pm 5$	$45.9 \pm 1.2 \pm 9.9$
J1446-4701	0.73 ± 0.31	0.74 ± 0.14	1.4 ± 0.4	3.0 ± 1.7	71	9	0	$19 \pm 4 \pm 5$	$5.2 \pm 1.0 \pm 1.4$
J1514-4946	4.1 ± 0.6	4.50 ± 0.28	1.5 ± 0.1	5.3 ± 1.1	999	66	0	$48 \pm 3 \pm 10$	$29.7 \pm 1.8 \pm 7.1$
J1600-3053 *	0.22 ± 0.16	0.53 ± 0.28	0.40 ± 0.47	2.0 ± 0.7	147	26	0	$17 \pm 9_{-5}^{+7}$	$23 \pm 12_{-7.0}^{+9.6}$
J1614-2230	2.0 ± 0.4	2.44 ± 0.20	0.96 ± 0.22	1.9 ± 0.4	599	81	1	$12 \pm 1 \pm 2$	$32.6 \pm 2.6 \pm 4.8$
J1658-5324	5.7 ± 0.7	2.89 ± 0.23	1.8 ± 0.2	1.4 ± 0.4	327	34	0	$30 \pm 2 \pm 8$	$9.9 \pm 0.8 \pm 2.6$
J1713+0747	1.3 ± 0.4	1.02 ± 0.14	1.6 ± 0.3	2.7 ± 1.2	126	17	2	$13 \pm 2_{-1}^{+2}$	$39.0 \pm 5.6_{-3.6}^{+4.6}$
J1741+1351 †	0.12 ± 0.04	0.24 ± 0.08	23	7	0	$3 \pm 1 \pm 0.3$	$1.5 \pm 0.5 \pm 0.1$
J1744-1134	4.6 ± 0.7	3.25 ± 0.25	1.3 ± 0.2	1.2 ± 0.3	394	65	0	$6.8 \pm 0.5 \pm 0.5$	$16.5 \pm 1.3 \pm 1.3$
J1747-4036	1.5 ± 0.7	0.99 ± 0.23	1.9 ± 0.3	5.4 ± 3.3	43	4	0	$140 \pm 30_{-50}^{+70}$	$11.7 \pm 2.7_{-4.7}^{+6.3}$
J1810+1744	4.2 ± 0.5	2.34 ± 0.17	1.9 ± 0.2	3.2 ± 1.1	442	23	3	$112 \pm 8_{-30}^{+40}$	$28.2 \pm 2.1_{-7.3}^{+9.4}$
J1823-3021A	1.5 ± 0.4	1.07 ± 0.15	1.6 ± 0.2	2.5 ± 0.6	64	0	5	$700 \pm 100 \pm 80$	$8.9 \pm 1.3 \pm 0.9$
J1858-2216	0.55 ± 0.28	0.72 ± 0.15	0.84 ± 0.74	1.7 ± 1.1	75	17	0	$8 \pm 2_{-2}^{+4}$	$6.8 \pm 1.4_{-1.7}^{+3.2}$
J1902-5105	3.1 ± 0.4	2.16 ± 0.15	1.7 ± 0.2	3.4 ± 1.1	536	31	2	$36 \pm 2 \pm 10$	$5.2 \pm 0.4_{-1.7}^{+2.1}$
J1939+2134 †	1.5 ± 0.8	0.9 ± 0.3	10	4	1	$140 \pm 50 \pm 30$	$1.3 \pm 0.5 \pm 0.2$
J1959+2048	2.4 ± 0.5	1.70 ± 0.19	1.4 ± 0.3	1.4 ± 0.4	185	43	0	$130 \pm 10_{-40}^{+20}$	$16.5 \pm 1.8_{-5.8}^{+2.2}$
J2017+0603	2.0 ± 0.3	3.33 ± 0.21	1.0 ± 0.2	3.4 ± 0.6	1196	100	1	$98 \pm 6 \pm 20$	$75.5 \pm 4.8 \pm 14$
J2043+1711	2.7 ± 0.3	2.70 ± 0.16	1.4 ± 0.1	3.3 ± 0.7	918	65	0	$100 \pm 6_{-30}^{+20}$	$79 \pm 5_{-26}^{+14}$

Table 10—Continued

PSR ^a	Photon Flux ($\text{ph cm}^{-2} \text{s}^{-1}$) ($\times 10^{-8}$)	Energy Flux ($\text{erg cm}^{-2} \text{s}^{-1}$) ($\times 10^{-11}$)	Γ	E_{cut} (GeV)	TS	TS_{cut}	$TS_{\text{b free}}$	Luminosity ($10^{32} \text{ erg s}^{-1}$)	Efficiency ^b (%)
J2047+1053	0.83 ± 0.36	0.62 ± 0.14	1.5 ± 0.5	2.0 ± 1.1	69	11	2	$31 \pm 7_{-8}^{+10}$	$29.4 \pm 6.6_{-7.7}^{+9.9}$
J2051–0827 [*]	0.24 ± 0.13	0.34 ± 0.10	0.50 ± 0.76	1.3 ± 0.7	68	15	0	$4 \pm 1 \pm 1$	$8.1 \pm 2.5 \pm 2.2$
J2124–3358	2.7 ± 0.3	3.68 ± 0.16	0.78 ± 0.13	1.63 ± 0.19	1993	237	3	$4.0 \pm 0.2^{+2}_{-1}$	$10.8 \pm 0.5^{+3.6}_{-3.3}$
J2214+3000	3.0 ± 0.3	3.28 ± 0.15	1.2 ± 0.1	2.2 ± 0.3	1689	146	2	$93 \pm 4 \pm 20$	$48.4 \pm 2.2 \pm 11$
J2215+5135	1.0 ± 0.3	1.18 ± 0.14	1.3 ± 0.3	3.4 ± 1.0	183	25	1	$130 \pm 20 \pm 30$	$24.6 \pm 3.0 \pm 5.7$
J2241–5236	3.0 ± 0.3	3.33 ± 0.16	1.3 ± 0.1	3.0 ± 0.5	2150	115	1	$10.5 \pm 0.5 \pm 3$	$4.0 \pm 0.2 \pm 1.1$
J2302+4442	2.6 ± 0.3	3.67 ± 0.17	0.94 ± 0.12	2.1 ± 0.3	1716	189	2	$62 \pm 3_{-20}^{+10}$	$162.6 \pm 7.5_{-57}^{+26}$

^a A dagger (†) means the spectral fit is unreliable (see text). A star (★) means that the spectrum was calculated using the on-peak data only.

^b Overestimated distances or the assumed beaming factor, $f_{\Omega} = 1$, can result in an efficiency $> 100\%$.

Note. — Unbinned maximum likelihood spectral fit results for the LAT MSPs, using the PLEC1 model (Eq. 12, Section 6). Columns 2 and 3 list the phase-averaged integral photon and energy fluxes in the 0.1 to 100 GeV energy band, F_{100} and G_{100} . Columns 4 and 5 list the photon index Γ and cutoff energy E_{cut} . Columns 6, 7, and 8 list the source significance TS , significance TS_{cut} of the exponential cutoff compared to a simple power law, and significance $TS_{\text{b free}}$ of the PLEC compared to a PLEC1 shape. A value $TS_{\text{cut}} < 9$ indicates that a spectral cutoff is not significantly detected. Column 9 gives the total gamma-ray luminosity L_{γ} in the 0.1 to 100 GeV energy band. The gamma-ray conversion efficiency $\eta \equiv L_{\gamma}/\dot{E}_{\text{int}}$ in Column 10 assumes a beam correction factor $f_{\Omega} = 1$ as described in Section 6.3, and the Shklovskii-corrected \dot{E}_{int} values from Table 6 in Section 4.3. The first uncertainty in L_{γ} and η comes from the statistical uncertainties in the spectral fit, whereas the second is due to the distance uncertainty. The strong dependence of these quantities on distance (see Table 6) and beaming factor means that these values should be considered with care.

For some pulsars with low duty cycles, the phase-averaged analysis returned a low TS value and/or could not constrain the spectral parameters well. We selected the off-peak intervals for these sources and followed the same prescription as for the phase-averaged analyses but without the pulsar in the model. We then fixed the parameters of all sources $> 4^\circ$ from the pulsar, left the normalization parameters of the remaining point sources and diffuse models free, added the pulsar back to the model, and performed an on-peak spectral analysis. These pulsars are indicated by a star (\star) in Tables 9 and 10. The values of F_{100} and G_{100} reported for these pulsars have been corrected to phase-averaged values. Additionally, the spatial residuals for PSR J1702–4128 revealed a large deficit near the pulsar attributed to the Galactic diffuse model. In lieu of a new Galactic diffuse template, we increased the minimum event energy to 300 MeV and performed an on-peak spectral analysis as described previously. Thus, for PSR J1702–4128 the reported values of F_{100} and G_{100} are extrapolations below the energy range of the data, which increases the quoted uncertainties beyond the statistical values.

To estimate systematic uncertainties on the maximum likelihood spectral parameters we selected eight pulsars with representative characteristics (J0218+4232, J0248+6021, J0357+3205, J0614–3329, J0631+1036, J1658–5324, J1833–1034, and J1846+0919) and studied how their spectra changed when perturbing the Galactic diffuse emission and LAT effective area (A_{eff}).

The distribution of Galactic diffuse normalization parameters, from all the fits, has a mean of 1.01 with 1σ deviation of 4%. To estimate possible systematic effects due to an imperfect knowledge of this diffuse component, we repeated the spectral analysis with the normalization of the Galactic diffuse emission fixed to (1 ± 0.06) times the best-fit value, corresponding to $\pm 1.5\sigma$ deviations. The average and largest deviations for Γ , E_{cut} , F_{100} , and G_{100} from this test are listed in the first row of Table 11.

Systematic uncertainties on A_{eff} are estimated to be 10% for $\log_{10} E/1 \text{ MeV} \leq 2$, 5% for $\log_{10} E/1 \text{ MeV} = 2.75$, and 10% for $\log_{10} E/1 \text{ MeV} \geq 4$ with linear extrapolation in between, in log space (Ackermann et al. 2012b). To estimate the effects of these uncertainties

Table 11. Systematic Deviations on Pulsar Spectral Parameters

Systematic	$\langle \Delta \Gamma \rangle$ (%)	$\langle \Delta E_{\text{cut}} \rangle$ (%)	$\langle \Delta F_{100} \rangle$ (%)	$\langle \Delta G_{100} \rangle$ (%)	$\max(\Delta \Gamma)$ (%)	$\max(\Delta E_{\text{cut}})$ (%)	$\max(\Delta F_{100})$ (%)	$\max(\Delta G_{100})$ (%)
Galactic Diffuse	14	4	16	12	80	27	65	46
Bracketing IRFs	5	4	8	6	21	11	13	8

we generated bracketing IRFs in which the usual A_{eff} was replaced by,

$$A_B(E) = A_{\text{eff}}(E)(1 + \text{err}(E)B(E)), \quad (14)$$

where $\text{err}(E)$ represents the A_{eff} uncertainties with $B(E) = \pm 1$ for the normalization factor K and $B(E) = \pm \tanh(\log_{10}(E/E_0)/\kappa)$ for Γ and E_{cut} . Choosing $\kappa = 0.13$ smoothes over twice the energy resolution. When using bracketing IRFs, it is important to isolate changes in the source of interest caused by the modified A_{eff} from changes in the diffuse background spectrum introduced by this perturbation. Since the diffuse background spectra were derived from flight data, we multiply the spectrum of the Galactic diffuse model by $(1 + \text{err}(E)B(E))^{-1}$, ensuring that the predicted counts spectra from the diffuse background remains unchanged during the bracketing studies.

We integrate the fit results for the different bracketing IRFs to obtain F_{100} and G_{100} . The second row of Table 11 lists the average and largest deviations for Γ , E_{cut} , F_{100} , and G_{100} due to bracketing. G_{100} is more robust than F_{100} . The considerably lower average deviations (columns 2 through 5) than maximum values (columns 6 through 9) shows that most deviations are much smaller than the outlying values.

Finally, to plot the spectra, we divided the 0.1 to 100 GeV energy band into 4 (2) bins per decade for pulsars with TS above (below) 250 and fit a power law in each bin, curvature within the bins being negligible. We fixed the spectral parameters of all sources more than 4° from the pulsar of interest and of the diffuse components at their full band fit results. The spectra of the pulsar of interest and the other point sources within the 4° region were modeled as power laws with index fixed at 2. Their flux levels in each energy band were obtained from a two-step fit. In the first step, sources with $TS \leq 0$ were removed, from that energy band only, as leaving them can adversely affect the fit uncertainties. The second step is to re-fit with the modified model. Sample spectra are shown in Figures A-9 to A-17 in Appendix A. For pulsars with $TS_{\text{b free}} \geq 9$, the Figures show both the PLEC1 and PLEC fits. The spectra obtained for all pulsars with $TS \geq 25$ and reliable spectral fits are included in the online material.

6.2. Spectral Results

Table 9 lists the phase-averaged spectral results for the non-recycled pulsars, and Table 10 for the MSPs. For ten pulsars, flagged with a \dagger symbol, the spectral fits were unreliable, for different reasons. Five are undetected as point sources, having $TS < 25$. For two of these the likelihood analysis fails and we report no spectral parameters, as is the case for PSR J0729–1448 in spite of its larger TS value. For the other three the integrated flux is

robust and we report F_{100} and G_{100} . Similarly, we report only the integrated fluxes for four more pulsars, either because the maximum likelihood fit solution favored $\Gamma \approx 0$, or because the parameter uncertainties were of order 100%. For PSRs J1112–6103 and J1410–6132, analysis of the off-peak phase intervals (Section 7) showed significant extended emission, which we added to the phase-averaged source model, improving the spectral results.

Figure 7 shows correlation between Γ and \dot{E} . The Pearson correlation factor is 0.68 for the young, radio-quiet pulsars and 0.58 for the MSPs, with probabilities of occurring by chance (two-sided p-values) of 5×10^{-6} and 1.5×10^{-4} , respectively. For the young radio-loud pulsars, the correlation is smaller (0.40) and marginally significant (two-sided p-value of 0.017). For all pulsars together the correlation factor is 0.57, with high significance (2×10^{-10}), even allowing for trials. We fit the measurements with $\Gamma = \mathcal{A} \log(\dot{E}) + \mathcal{B}$ (dashed lines in the figure). For young pulsars we find similar trends in the radio-loud and radio-quiet populations with $\mathcal{A} \approx 0.2$ and $\mathcal{B} \approx -5$. The exact fit values are sensitive to the outliers with small statistical uncertainties, such as the Crab. The MSPs have more dispersion in Γ and a narrower \dot{E} range, with a steeper slope than for the young population, $\mathcal{A} \approx 0.4$ and $\mathcal{B} \approx -12$.

In 1PC we noted a possible correlation between E_{cut} and the magnetic field strength at the light cylinder ($B_{\text{LC}} = 4\pi^2(1.5I_0\dot{P})^{1/2}(c^3P^5)^{-1/2}$, assuming an orthogonal rotator). For the radio-quiet pulsars, Figure 8 confirms the trend, with a Pearson correlation factor of 0.64 (p-value 4×10^{-5}). The factor is 0.52 for the MSPs, with p-value 0.0007. Here too, the correlation for the young, radio-loud pulsars is small (0.24) and insignificant (p-value 0.17).

6.3. Luminosity

Gamma-ray emission models predict different relations between the spindown power \dot{E} and the gamma-ray luminosity,

$$L_\gamma = 4\pi d^2 f_\Omega G_{100}, \quad (15)$$

making this a discriminating observable when applied to a large sample of gamma-ray pulsars. Following Romani & Watters (2010), we define the beam correction factor f_Ω as

$$f_\Omega(\alpha, \zeta_E) = \frac{\int F_\gamma(\alpha; \zeta, \phi) \sin \zeta d\zeta d\phi}{2 \int F_\gamma(\alpha; \zeta_E, \phi) d\phi}, \quad (16)$$

to extrapolate the observed flux to the full sky for some beam shape model. The angle α between the neutron star’s magnetic and rotation axes is one model parameter. The angle ζ_E between the rotation axis and the Earth line of sight describes the inclination of the system relative to Earth. The numerator integrates emission into all space (all inclinations

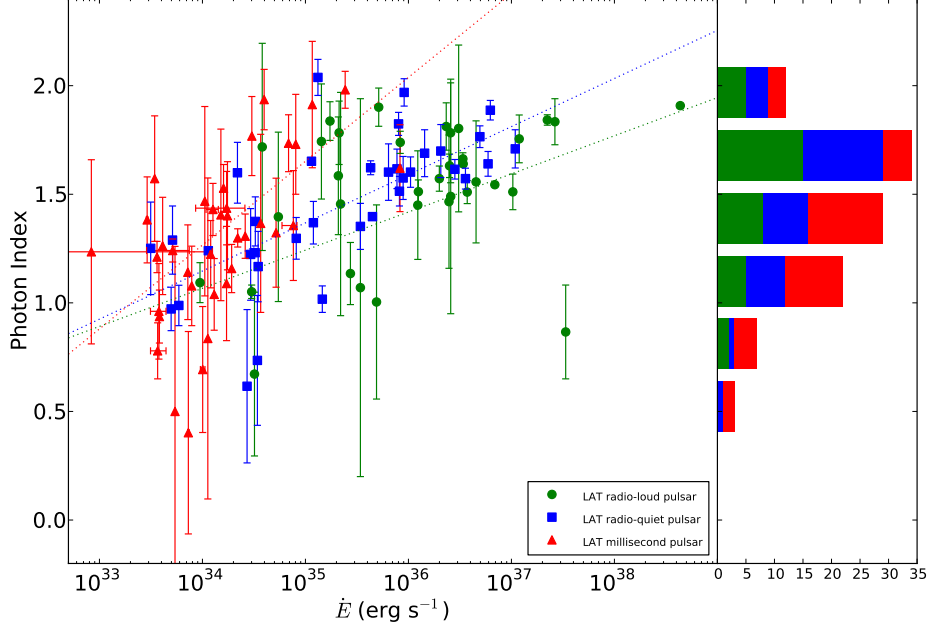


Fig. 7.— Power-law index Γ for the exponentially cutoff gamma-ray spectra versus the Shlovskii-corrected spindown power, for the pulsars bright enough in gamma rays to allow spectral analysis (see text). The straight-line fit results are in Section 6.2. The markers are the same as in Figure 1. The uppermost line is for millisecond gamma-ray pulsars. The middle line fits young, radio-quiet gamma-ray pulsars, while the lowest line is for young, radio-loud gamma-ray pulsars. A histogram of photon index values is projected onto the axis.

ζ) whereas the denominator is the observed flux integrated over a neutron star rotation, with pulsar phase ϕ . In the past, the gamma-ray beam was conventionally assumed to sweep out a 1 sr solid angle, in which case $L_\gamma = d^2 G_{100}$. Such a beam is appropriate to near-surface polar cap emission and corresponds to $f_\Omega = \frac{1}{4\pi} = 0.08$. An outer magnetosphere fan-like beam sweeping the entire sky (4π steradians) gives $f_\Omega \approx 1$, which is the value we adopt for calculating L_γ . However, Pierbattista et al. (2012) found a large spread in f_Ω values for different emission models and for radio-loud versus radio-quiet young pulsars. Values of f_Ω exceeding 1 correspond to beams that are narrow in ϕ , extended in ζ , and/or have average intensity exceeding the value sampled at ζ_E .

Figure 9 shows L_γ versus \dot{E} . Pulsars with poor spectral fits have been excluded. The open field-line voltage is $V \simeq 3.18 \times 10^{-3} \sqrt{\dot{E}}$ volts. Above some threshold voltage, gamma-ray emitting electron-positron cascades occur, and a linear dependence of L_γ on V would

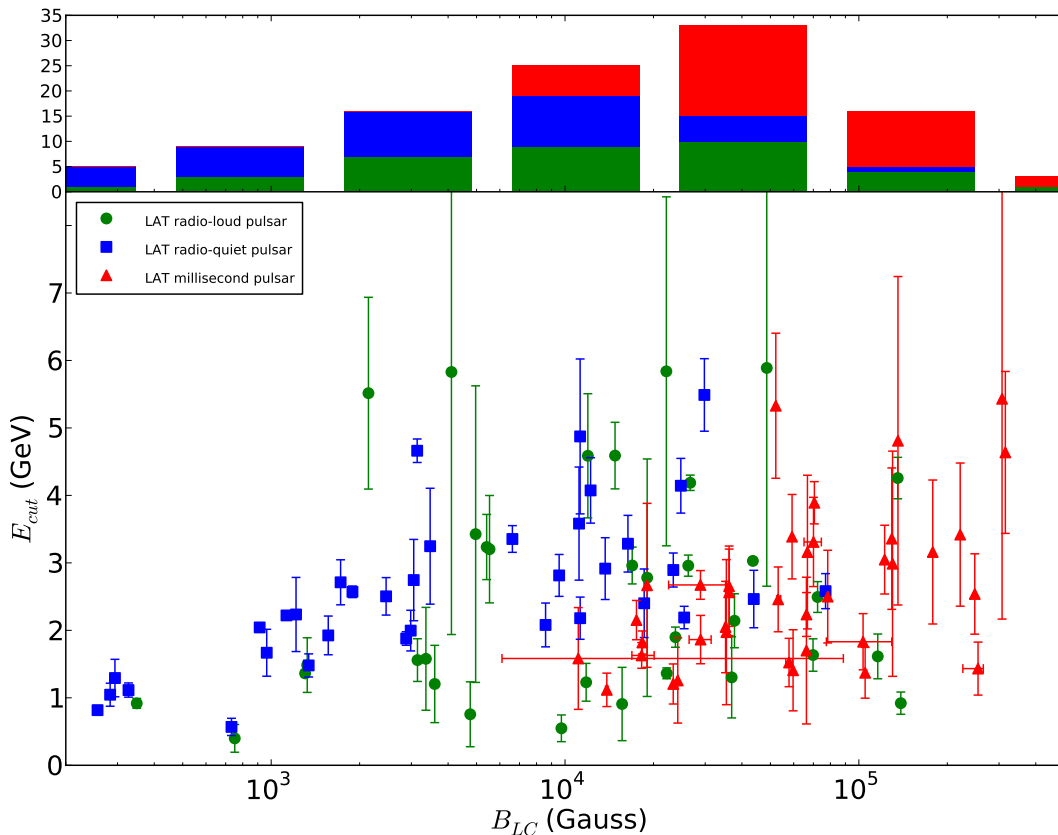


Fig. 8.— The best-fit cutoff energy versus magnetic field at the light cylinder, B_{LC} . The young, radio-quiet gamma-ray pulsars have lower B_{LC} than other gamma-ray pulsars. The markers are the same as in Figure 1. The histogram highlights the different B_{LC} distributions for the three pulsar classes.

give $L_\gamma \propto \sqrt{\dot{E}}$ (Arons 1996), as for the lower diagonal line. With arbitrary normalization, we call this the “heuristic” gamma-ray pulsar luminosity,

$$L_\gamma^h = 10^{33} \sqrt{\dot{E}/10^{33}} = \sqrt{10^{33} \dot{E}} \text{ erg s}^{-1}. \quad (17)$$

At low \dot{E} values L_γ seems to be falling below L_γ^h . The upper diagonal line shows $L_\gamma = \dot{E}$, that is, a 100% efficiency $\eta = L_\gamma/\dot{E}$ for converting spindown power into gamma rays. A few pulsars appear above this line, likely due to over-estimated distances or f_Ω values. Figure 10 plots η versus \dot{E} . Overall, the expected $\eta \propto 1/\sqrt{\dot{E}}$ trend is roughly respected. However, the large dispersion due to the distance uncertainties, as well as beaming effects, limits the extent to which the data constrain the theory.

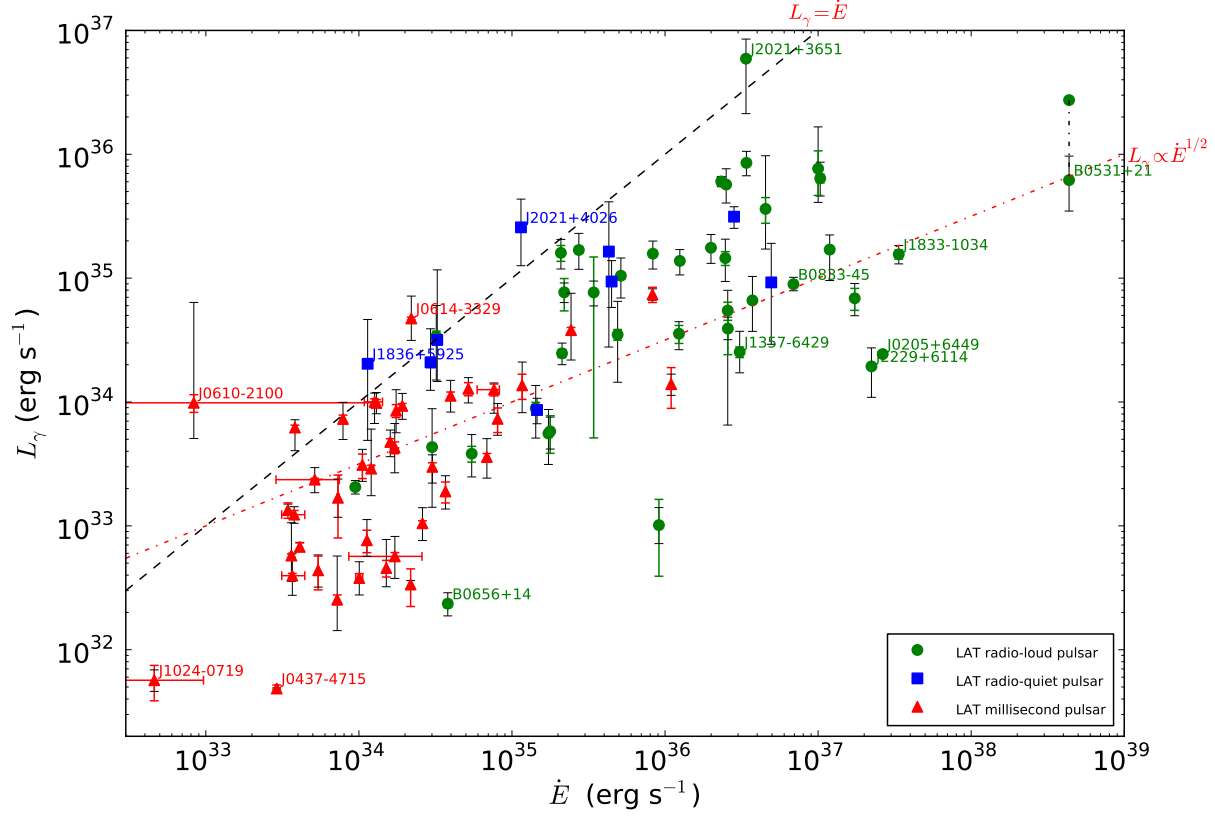


Fig. 9.— Gamma-ray luminosity $L_\gamma = 4\pi f_\Omega d^2 G_{100}$ in the 0.1 to 100 GeV energy band versus spindown power \dot{E} . The vertical error bars from the statistical uncertainty on the energy flux G_{100} are colored in the electronic journal version. The vertical error bars due to the distance uncertainties are black, and generally larger. Doppler corrections (Section 4.3) have been applied to MSPs with known proper motions, leading to visible horizontal error bars in some cases. The upper diagonal line indicates 100% conversion of spindown power into gamma-ray flux: for pulsars above this line, the distance d may be smaller, and/or the assumed beam correction $f_\Omega \equiv 1$ is wrong. The lower diagonal line indicates the heuristic luminosity $L_\gamma^h = \sqrt{10^{33} \dot{E}} \text{ erg s}^{-1}$, to guide the eye. The upper of the two Crab points, at far right, includes the X-ray energy flux (see Section 9.1). The markers are the same as in Figure 1.

The Doppler correction to \dot{E} is small ($|\xi| < 10\%$) for 9 of the 20 pulsars with proper motion corrections (Table 6). For the five with $20\% < \xi < 60\%$, the correction refines their positions in e.g. the L_γ vs \dot{E}^{int} plane. The Doppler correction for PSR J0437–4715, with a

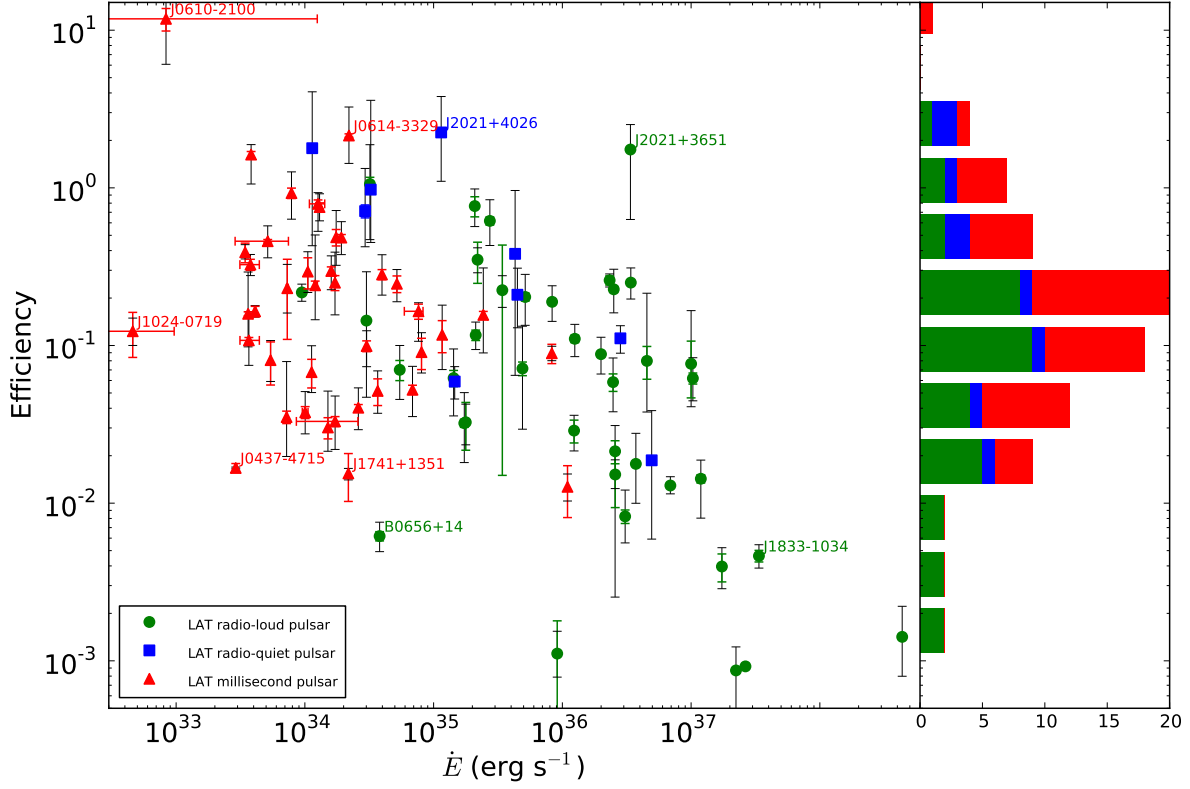


Fig. 10.— Gamma-ray efficiency $\eta = L_\gamma/\dot{E}$ versus spindown power \dot{E} . The error bars are as in Figure 9. The markers and the side histogram use the same color coding as in Figure 1.

large $\xi = 75\%$, produces a qualitative change: observed $\dot{E} = 12 \times 10^{33} \text{ erg s}^{-1}$ decreases to corrected $\dot{E}^{\text{int}} = 3 \times 10^{33} \text{ erg s}^{-1}$, right at the apparent deathline, and the efficiency changes from the lowest outlier amongst MSPs, to a low, but typical, $\eta = 1.7\%$.

The remaining four pulsars with $\xi > 60\%$ bear special discussion. Figure 11 plots lines of constant \dot{E} , L_γ , and transverse velocity v_T in μ vs. d space for different assumptions: $\dot{E} = 0$, $\dot{E} = L_\gamma$, and $\eta\dot{E} = L_\gamma$ with $\eta = 30\%$, at the high end of the observed range. The curve for $v_T = \mu d = 150 \text{ km s}^{-1}$ is the 3σ extremum of the MSP velocity distribution of Lyne et al. (1998). Faster recycled pulsars are possible, but unusual. Allowed (or favored) regions are to the left of the curves. The curve for $\dot{E} = 3 \times 10^{33} \text{ erg s}^{-1}$ shows how an \dot{E} value lower than those seen to date would compare with the other constraints. The shaded zones correspond to the measurements and their uncertainties adopted in this paper along

with previous measurements, for comparison. We recall that here, as throughout the paper, we adopt the moment of inertia $I_0 = 10^{45}$ gm cm², corresponding to a neutron star mass of $1.4 M_\odot$ and radius of 10 km.

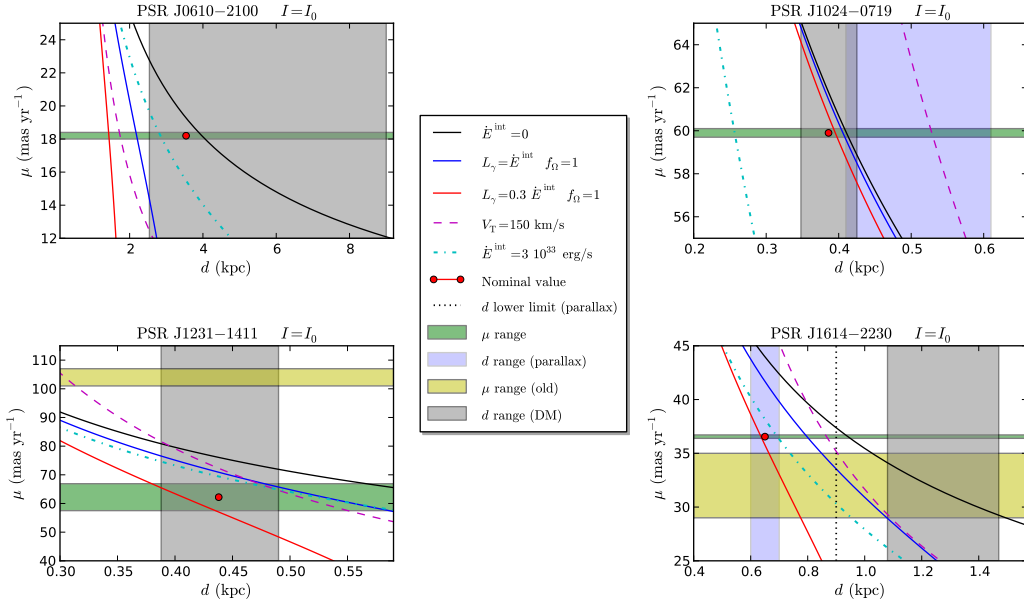


Fig. 11.— Constraints on the proper motion μ and the distance d for the four MSPs with Doppler corrections that lead to unphysical or out-of-family values of the spindown power \dot{E}^{int} (see Section 4.3). Unphysical or uncommon values are to the right of the curves. The uppermost solid curve (black) requires $\dot{E}^{\text{int}} > 0$. The second solid curve (blue) requires $\dot{E}^{\text{int}} > L_\gamma$. The bottom solid curve (red) assumes a large, but common, efficiency $\eta = 30\%$, that is, $\eta \dot{E}^{\text{int}} > L_\gamma$. The violet dashed line corresponds to a pulsar transverse space velocity of 150 km s^{-1} . Finally, the green dot-dashed line traces the empirical gamma-ray deathline value of $\dot{E}^{\text{int}} = 3 \times 10^{33} \text{ erg s}^{-1}$. The red dot is at the adopted values of μ and the distance d . The distances, with uncertainties, from parallax and from DM with the NE2001 model are represented by light blue and gray zones, respectively. The green zone indicates the newer proper motion measurement we use, and the yellow zone shows earlier values. The black dotted line shows the lower limit of parallax distance for J1614–2230 from Demorest et al. (2010).

PSR J0610–2100 was recently discussed by Espinoza et al. (2013): the intrinsic spindown power is well below the empirical deathline, the space velocity is much higher than typical, and the gamma-ray efficiency exceeds 100%. In Figure 11 the nominal (μ, d) point

for this pulsar is to the right of most of the curves. These apparent paradoxes are resolved if the pulsar is closer than 2 kpc. Upon inspection of NE2001, minor changes in the “Local Super Bubble” shape and density for this line of sight could yield a distance less than 2 kpc. Espinoza et al. (2013) also describe material along the line of sight that is unmodeled in NE2001.

They also discussed PSR J1024–0719. We show both the corrected parallax distance from Verbiest et al. (2012) as well as the NE2001 distance, for which the 10% uncertainty is probably underestimated given the small $DM = 6.5 \text{ pc cm}^{-3}$ in that direction. However, any distance compatible with the parallax yields the lowest intrinsic spindown power \dot{E}^{int} of any gamma-ray pulsar, and well below the $3 \times 10^{33} \text{ erg s}^{-1}$ current minimum.

For PSR J1231–1411, discovered in an unassociated LAT source, Ransom et al. (2011) measured a large proper motion, $\mu > 100 \text{ mas yr}^{-1}$, indicated in Figure 11. The Doppler correction using that large μ gave negative \dot{E}^{int} values, which is unphysical. Radio observations since the initial measurement allowed us to update the timing model, resulting in the smaller proper motion value listed in Table 6 and plotted in Figure 11. The Doppler corrections now lead to rather typical parameters for this pulsar.

PSR J1614–2230 is one of the rare pulsars with a neutron star mass measurement. Demorest et al. (2010) measured the Shapiro delay precisely in this binary system to obtain $M = 1.97 \pm 0.04 M_{\odot}$, well above the Chandrasekhar mass of $1.4 M_{\odot}$. The DM distance, 1.27 kpc, along with the proper motion measured in the same paper yields implausible \dot{E}^{int} and η values. A moment of inertia greater than I_0 would improve the situation, favoring “rigid” neutron star equations-of-state. However, the parallax distance and proper motion recently measured by Lassus (2013) brings this pulsar back in line with the rest of the population.

Demorest et al. (2010) conclude that $M > 1.4 M_{\odot}$ is probably true for many or most MSPs. For $I > I_0$, \dot{E}^{int} is larger than the “standard” values in Table 6. This would shrink the spread between the \dot{E}^{int} distributions for the young and MSP gamma-ray pulsars.

7. Unpulsed Magnetospheric Emission

Some pulsars have magnetospheric emission over their full rotation phase with similar spectral characteristics to the emission seen through their peaks. This emission appears in the observed light curves as a low-level, unpulsed component above the estimated background level (i.e., not attributable to diffuse emission or nearby point sources) and can be a powerful discriminator for the emission models.

On the other hand a PWN around the pulsar, or a photon excess due to imprecise knowledge of diffuse emission around the pulsar, would not be modulated at the rotational period and could be confused with a constant magnetospheric signal. Including the discovery of the GeV PWN 3C 58 associated with PSR J0205+6449 described in this section, the LAT sees 17 sources potentially associated with PWNe at GeV energies (Acero et al. 2013b). Some are highlighted in Appendix C. This off-peak emission should be properly modeled when searching for pulsar emission at all rotation phases.

We can discriminate between these two possible signals through spectral and spatial analysis. If the emission is magnetospheric, it is more likely to appear as a non-variable point source with an exponentially cutoff spectrum with a well-known range of cutoff energies. On the other hand, PWNe and diffuse excesses have spectra with a power-law shape and either a hard index continuing up to tens of GeV in the PWN case or present only at lower energies with a very soft index in the diffuse case. In addition, PWNe are often spatially resolvable at GeV energies (e.g., Vela-X has been spatially resolved with the LAT and *AGILE* and HESS J1825–137 with the LAT; Abdo et al. 2010f; Pellizzoni et al. 2010; Grondin et al. 2011, respectively) so an extended source would argue against a magnetospheric origin of the emission. However, given the finite angular resolution of the LAT (see Section 2) not all PWNe will appear spatially extended at GeV energies. The Crab Nebula, for instance, cannot be resolved by the LAT but can be distinguished from the gamma-bright Crab pulsar, in the off-peak interval, by its hard spectrum above ~ 1 GeV (Abdo et al. 2010d). In addition, GeV emission from the Crab Nebula was discovered to be time-variable (e.g., Abdo et al. 2011a) providing another possible way to discern the nature of any observed off-peak signal.

Therefore, to identify pulsars with magnetospheric emission across the entire rotation, we define and search the off-peak intervals of the pulsars in this catalog for significant emission, except PSR J2215+5135 for which the rotation ephemeris covers a short time interval and the profile is noisy. We then evaluate the spectral and spatial characteristics of any off-peak emission to determine if it is likely magnetospheric, related to the pulsar wind, or physically unrelated to the pulsar (e.g., unmodeled diffuse emission).

7.1. Off-peak Phase Selection

We first developed a systematic, model-independent, and computationally-efficient method to define the off-peak interval of a pulsar light curve.

We begin by deconstructing the light curve into simple Bayesian Blocks using the algorithm described in Jackson et al. (2005) and Scargle et al. (2013). We could not apply

the Bayesian Block algorithm to the weighted-counts light curves because they do not follow Poisson statistics, required by the algorithm. We therefore use an unweighted-counts light curve in which the angular radius and minimum energy selection have been varied to maximize the H-test statistic. To produce Bayesian Blocks on a periodic light curve, we extend the data over three rotations, by copying and shifting the observed phases to cover the phase range from -1 to 2 . We do, however, define the final blocks to be between phases 0 and 1 . To avoid potential contamination from the trailing or leading edges of the peaks, we reduce the extent of the block by 10% on either side, referenced to the center of the block.

There is one free parameter in the Bayesian Block algorithm called n_{cprior} which modifies the probability that the algorithm will divide a block into smaller intervals. We found that, in most cases, setting $n_{\text{cprior}} = 8$ protects against the Bayesian Block decomposition containing unphysically small blocks. For a few marginally-detected pulsars, the algorithm failed to find more than one block and we had to decrease n_{cprior} until the algorithm found a variable light curve. Finally, for a few pulsars the Bayesian-block decomposition of the light curves failed to model weak peaks found by the light-curve fitting method presented in Section 5.2 or extended too far into the other peaks. For these pulsars, we conservatively shrink the off-peak region.

For some pulsars, the observed light curve has two well-separated peaks with no significant bridge emission, which leads to two well-defined off-peak intervals. We account for this possibility by finding the second-lowest Bayesian block and accepting it as a second off-peak interval if the emission is consistent with that in the lowest block (at the 99% confidence level) and if the extent of the second block is at least half that of the first block.

Figure 12 shows the energy-and-radius optimized light curves, the Bayesian block decompositions, and the off-peak intervals for six pulsars. Figures A-1 to A-8 overlay off-peak intervals over the weighted light curves of several pulsars. The off-peak intervals for all pulsars in this catalog are given in Tables 7 and 8.

7.2. Off-peak Analysis Method

Characterizing both the spatial and spectral characteristics of any off-peak emission helps discern its origin. We employ a somewhat different analysis procedure here than for the phase-averaged analysis described in Section 6.1. To evaluate the spatial characteristics of any off-peak emission we use the likelihood fitting package `pointlike` (detailed in Lande et al. 2012), and to fit the spectrum we use `gtlike` in binned mode via `pyLikelihood` as was done for the phase-averaged analysis.

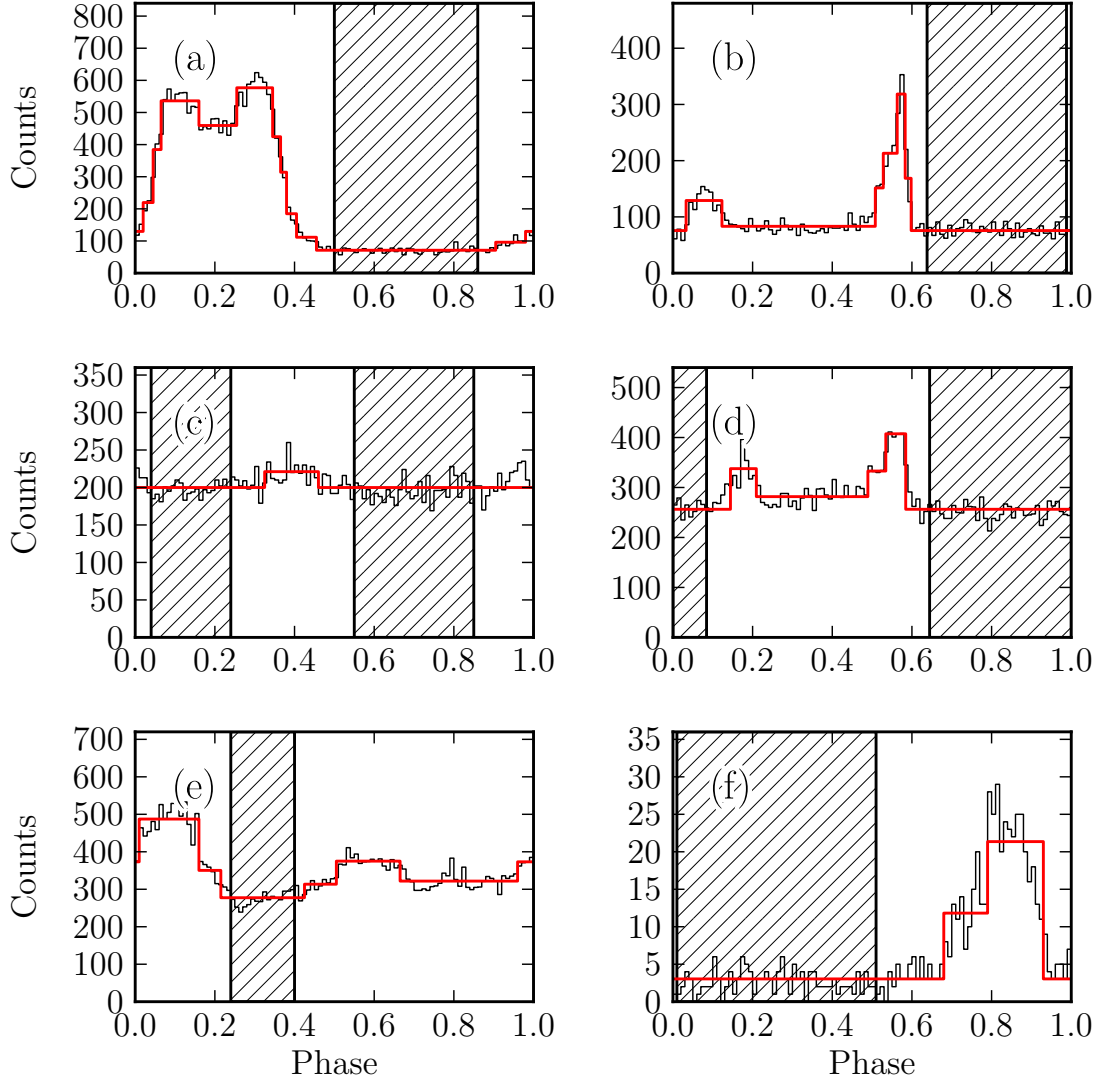


Fig. 12.— The energy-and-radius optimized light curve, Bayesian block decomposition of the light curve, and off-peak interval for (a) PSR J0007+7303, (b) PSR J0205+6449, (c) PSR J1410–6132, (d) PSR J1747–2958, (e) PSR J2021+4026, and (f) PSR J2124–3358. The black histograms represent the light curves, the gray lines (colored red in the electronic version) represent the Bayesian block decompositions of the pulsar light curves, and the hatched areas represent the off-peak intervals selected by this method.

For each pulsar we start from the same temporal and spatial event selections described in Section 2 but we increase the maximum energy to 400 GeV (the highest event energy for any ROI under this selection is ~ 316 GeV). For the `pointlike` analysis we further select a 10° radius ROI and for `gtlike` a $14^\circ \times 14^\circ$ square ROI, both centered on the pulsar position. Finally, we only consider photons with pulse phases within the corresponding off-peak interval.

We search for off-peak emission assuming a point source and (except for the Crab Nebula and Vela-X, described below) a power-law spectral model. We fit the position of this putative off-peak source using `pointlike` as described by Nolan et al. (2012) and then use the best-fit position in a spectral analysis with `gtlike`. From the spectral analysis we require $TS \geq 25$ (just over 4σ) to claim a detection. If $TS < 25$, we compute upper limits on the flux in the energy range from 100 MeV to 316 GeV assuming a power law with photon index fixed to 2.0 and a PLEC1 model with $\Gamma = 1.7$ and $E_{\text{cut}} = 3$ GeV.

The spectrum of the Crab Nebula (associated with PSR J0534+2200) is uniquely challenging because the GeV spectrum contains both a falling synchrotron and a rising inverse Compton component (Abdo et al. 2010d). For this particular source we used the best-fit two-component spectral model from Buehler et al. (2012) and fit only the overall normalization of the source. In addition, for Vela-X (associated with PSR J0835-4510) we took the best-fit spectral model from Grondin et al. (2013) and fit only the overall normalization of this source. This spectrum has a smoothly broken power law spectral model and was fit assuming Vela-X to have an elliptical disk spatial model.

If the off-peak source is significant, we test whether the spectrum shows evidence for a cutoff, as described in Section 6.1 and by Ackermann et al. (2011), assuming the source is at the pulsar position. We say that the off-peak emission shows evidence for a cutoff if $TS_{\text{cut}} \geq 9$, corresponding to a 3σ detection.

For a significant off-peak point source, we use `pointlike` to test if the emission is significantly extended. We assume a radially-symmetric Gaussian source and fit the position and extension parameter (σ) as described in Lande et al. (2012). The best-fit extended source parameters are then given to `gtlike`, which is used to fit the spectral parameters and the significance of the extension over a point source, TS_{ext} , evaluated as described in Lande et al. (2012). That paper established that $TS_{\text{ext}} \geq 16$ means highly probable source extension. In the present work we aim only to flag possible extension, and use $TS_{\text{ext}} \geq 9$.

To test for variability, even without significant emission over the 3-year time range, we divide the dataset into 36 intervals and fit the point-source flux independently in each interval, computing TS_{var} as in 2FGL. For sources with potential magnetospheric off-peak

emission and for regions with no detection, we performed the test at the pulsar’s position. Otherwise, we test at the best-fit position. The off-peak emission is said to show evidence for variability if $TS_{\text{var}} \geq 91.7$, corresponding to a 4σ significance. As noted in Section 2, our timing solutions for PSRs J0205+6449 and J1838–0537 are not coherent across all three years. For these two pulsars, we excluded the time ranges without ephemerides and only tested for variability during months that were completely covered. For J1838–0537 only one month is lost, whereas for J0205+6449 the 7% data loss is spread across three separate months. As a result, TS_{var} for these pulsars is a conservative estimate of variability significance.

The procedure described above, especially the extension analysis, is particularly sensitive to sources not included in 2FGL that are near the pulsar of interest, for two reasons. First, we are using an additional year of data and second, when “turning off” a bright pulsar nearby, faint sources become more important to the global fit. Therefore, in many situations we had to run the analysis several times, iteratively improving the model by including new sources, until we removed all $TS > 25$ residuals. The final `gtlike`-formatted XML source model for each off-peak region is included in the auxiliary material.

There are still, however, pulsars for which we were unable to obtain an unbiased fit of the off-peak emission, most likely due to inaccuracies in the model of the Galactic diffuse emission and incorrectly modeled nearby sources. The most common symptom of a biased fit is an unphysically large extension. In these cases, the extended source attempts to account for multiple point sources or incorrectly-modeled diffuse emission, not just the putative off-peak emission. Systematics associated with modeling extended sources are discussed more thoroughly in Lande et al. (2012). For the purposes of this catalog, we have flagged the pulsars where off-peak analysis suffered from these issues and do not attempt a complete understanding of the emission.

7.3. Off-peak Results

The off-peak intervals of 54 LAT-detected pulsars have been evaluated by Ackermann et al. (2011) using 16 months of sky survey observations. This led to the discovery of PWN-like emission in the off-peak interval of PSR J1023–5746, coincident with HESS J1023–575, and identification of 5 pulsars that appear to have near 100% duty cycles. Our results, summarized in Table 12, extend the analysis to 116 pulsars over 3 years of data. Sample off-peak spectra are shown in Figure 13. Using the procedures outlined in Sections 7.1 and 7.2, we have identified 34 pulsars that have significant emission ($TS \geq 25$) in their off-peak intervals. We classify the likely nature of the emission as follows.

If the spectrum cuts off ($TS_{\text{cut}} \geq 9$), the emission could be magnetospheric (‘M’). An indication of spatial extension ($TS_{\text{ext}} \geq 9$) flags sources where the emission may instead be an artifact of defects in the Galactic diffuse emission model and we list such sources as type ‘U’, for ‘unidentified’. Similarly, emission from sources without evidence for a spectral cutoff could originate in the pulsar wind, type ‘W’. Spatial extension alone is not a sufficient indicator, since the LAT PSF is larger than many PWNe. A hard spectral index also suggests a PWN contribution. The table lists the four solid PWN detections. PSR J0205+6449 in 3C58 is a new detection at GeV energies. Only one of the four, the previously identified Vela-X PWN (Abdo et al. 2010f), is spatially extended for the LAT.

We identify 19 type ‘U’ regions, and 11 type ‘M’ sources, significantly expanding the number of pulsars that perhaps have detectable magnetospheric emission across all rotational phases. One caution is that many of these ‘M’ pulsars, especially the young objects, are in regions of particularly bright diffuse gamma-ray emission, where small fractional uncertainties in the level of diffuse emission can account for much of the apparent unpulsed emission. However, if established as true magnetospheric components, these will be important test cases for pulsar emission models. For type ‘M’ and ‘U’ sources, we present the best-fit spectral parameters using a point source at the pulsar’s position with a PLEC1 spectral model in Table 12. For all other sources (except the Crab Nebula described in Section 7.2), we present the spectral results using a power-law spectral model and the best-fit spatial representation.

For a few sources, the spectral analysis performed here disagrees with that in Ackermann et al. (2011). For soft and faint sources (including J1044–5737 and J1809–2332), the spectral discrepancy is mainly caused by our use of a newer Galactic diffuse model. At lower energies, small changes in the diffuse model can have a significant impact on the analysis of a region. For bright magnetospheric sources (including J0633+1746 and J2021+4026), the spectral discrepancy is mainly due to using different phase ranges (see Section 7.1).

Figure 14 shows that only a small fraction of the spindown power goes into the gamma-ray emission from LAT-detected PWNe. Similarly, Figure 15 shows that the LAT only detects PWNe from the youngest pulsars with the highest spindown power. GeV emission from the Crab Nebula is highly time variable (Section 7.2). Indeed, we find $TS_{\text{var}} = 373$ for the Crab Nebula; however no other source demonstrated flux variability (all have $16 < TS_{\text{var}} < 65$). Other GeV PWNe may be variable, but the combination of lower fluxes and less-extreme variations limits our ability to identify them as such.

The off-peak results for several interesting sources are presented in Appendix C. The complete off-peak search results can be found in the auxiliary information described in Appendix B. For regions where we find $TS < 25$, the auxiliary information contains upper limits computed for both a power-law spectral model and a PLEC1 model with $E_{\text{cut}} = 3$

GeV and $\Gamma = 1.7$. The auxiliary information also contains TS_{var} for each off-peak interval.

Table 12. Off-Peak Spatial and Spectral Results

PSR	Type	TS	TS_{ext}	TS_{cutoff}	Energy Flux ($10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$)	Γ	E_{cutoff} (GeV)
Young Pulsars							
J0007+7303	U	71.4	10.8	0.0	1.98 ± 0.26	2.61 ± 0.14	
J0205+6449	W	33.7	0.5	0.0	1.75 ± 0.68	1.61 ± 0.21	
J0534+2200	W	5247.	0.0	0.3	67.2 ± 1.6	^a	
J0631+1036	U	33.1	0.0	5.4	1.70 ± 0.33	2.38 ± 0.14	
J0633+1746	M	3666.	2.3	239.	41.4 ± 1.1	1.37 ± 0.09	0.93 ± 0.10
J0734–1559	U	28.3	12.4	30.8	1.61 ± 0.24	0.01 ± 0.08	0.17 ± 0.03
J0835–4510	W	473.	283.	22.8	30.3 ± 1.2	^b	
J0908–4913	U	65.1	41.4	60.4	3.04 ± 1.07	0.15 ± 0.59	0.30 ± 0.01
J1023–5746	U	59.7	30.0	10.9	5.35 ± 1.17	0.57 ± 0.80	0.49 ± 0.21
J1044–5737	U	42.0	98.1	22.4	3.12 ± 0.75	0.80 ± 0.93	0.40 ± 0.18
J1105–6107	U	33.3	37.5	21.7	3.81 ± 0.77	0.92 ± 0.56	0.48 ± 0.22
J1112–6103	U	65.0	71.1	0.9	5.10 ± 0.74	2.17 ± 0.09	
J1119–6127	U	61.3	1.0	0.9	4.11 ± 0.63	2.22 ± 0.09	
J1124–5916	M	95.9	0.0	18.2	2.87 ± 0.71	1.31 ± 0.91	1.43 ± 1.42
J1410–6132	U	27.5	71.2	0.4	4.29 ± 1.05	1.90 ± 0.15	
J1513–5908	W	102.	3.5	0.0	4.95 ± 0.83	1.78 ± 0.12	
J1620–4927	M	28.9	0.5	35.2	5.25 ± 0.96	0.35 ± 0.94	0.57 ± 0.29
J1746–3239	U	53.3	34.3	34.2	3.65 ± 0.59	0.94 ± 0.31	0.60 ± 0.10
J1747–2958	M	45.5	5.4	49.8	8.41 ± 2.84	0.02 ± 0.32	0.28 ± 0.01
J1809–2332	U	32.5	13.6	21.9	4.10 ± 0.80	0.24 ± 0.83	0.31 ± 0.11
J1813–1246	M	62.8	0.0	9.0	6.31 ± 1.40	1.60 ± 0.73	0.99 ± 0.95
J1836+5925	M	10407.	0.0	365.	36.9 ± 0.7	1.47 ± 0.03	1.98 ± 0.09
J1838–0537	U	51.3	32.9	21.9	8.35 ± 1.31	1.39 ± 0.54	2.55 ± 2.48
J2021+4026	M	1717.	8.7	244.	64.0 ± 1.4	1.64 ± 0.02	1.82 ± 0.04
J2032+4127	U	53.6	76.1	1.5	4.36 ± 0.77	2.07 ± 0.12	
J2055+2539	M	123.	0.0	30.0	1.63 ± 0.19	1.05 ± 0.28	0.64 ± 0.12
Millisecond Pulsars							
J0034–0534	U	41.0	0.0	6.0	0.82 ± 0.16	2.40 ± 0.19	
J0102+4839	U	49.7	0.0	7.4	1.29 ± 0.20	2.51 ± 0.14	
J0218+4232	U	50.1	0.0	6.8	2.13 ± 0.33	2.72 ± 0.26	
J0340+4130	M	26.9	0.1	16.3	0.53 ± 0.11	0.02 ± 0.22	0.94 ± 0.28
J1658–5324	U	42.3	0.0	1.9	1.69 ± 0.29	2.52 ± 0.76	
J2043+1711	U	52.5	0.0	8.8	1.46 ± 0.27	2.29 ± 0.14	
J2124–3358	M	129.	0.0	19.8	1.08 ± 0.15	0.70 ± 0.51	1.21 ± 0.49
J2302+4442	M	114.	0.0	9.8	1.45 ± 0.20	1.54 ± 0.40	1.61 ± 0.82

^aThe spectral shape of the Crab Nebula was taken from Buehler et al. (2012).

^bThe spectral shape of Vela-X was taken from Grondin et al. (2013).

Note. — Off-peak regions with a significant detection of emission. The source classification is ‘M’ for likely magnetospheric, ‘W’ for likely pulsar wind, and ‘U’ for unidentified. The table includes the significance of the source (TS), of the source extension (TS_{ext}), and of a spectral cutoff (TS_{cut}). The best-fit energy flux and photon index are computed in the energy range from 100 MeV to 316 GeV. Exponential cutoff energies

are listed for sources with large TS_{cut} . The quoted errors are statistical only. A few sources are discussed in Appendix C.

8. The Pulsars Not Seen

This catalog is a milestone in the progress toward the long-term goal of acquiring the most uniform sample of neutron stars possible, so that comparisons with model predictions (e.g., Gonthier et al. 2007; Watters & Romani 2011; Pierbattista et al. 2012) will allow improved tests of emission models and of their links with their parent population of massive stars or with diffuse Galactic emission. Selection biases can be subtle and the advantage of pulsar searches in the coming years is not so much to increase the absolute numbers, but to be sure to have explored the dark corners of parameter space. Continued support from pulsar radio astronomers is crucial to maintain sensitivity to the more unusual gamma-ray pulsars in the coming years of the *Fermi* mission. Here, we consider pulsars that might have been expected to be seen with the LAT, but were not, to highlight “gamma-quiet” or “sub-luminous” pulsars.

8.1. High Spindown Power Pulsars

Of the 64 known RPPs with $\dot{E} > 10^{36}$ erg s⁻¹, Table 13 lists the 28 for which we did *not* see gamma-ray pulsations when the data set for this paper was frozen. When no steady LAT point source lies within 0°2 we provide a 95% confidence level upper limit (UL) on G_{100} (Section 8.2). The Galactic latitude b roughly indicates the diffuse background level. The last column compares the UL with the “heuristic” energy flux, $\sqrt{\dot{E}}/(4\pi d^2)$. Figure 15 plots $\sqrt{\dot{E}}/d^2$ for all pulsars. Note the absence of the factor 4π in the latter case. To convert the plot’s scale, in units of erg^{1/2} s^{-1/2} kpc⁻², to the units used for G_{100} and the flux ULs, using L_{γ}^h , (Eq. 17), we multiply the scale of Figure 15 by

$$\frac{\sqrt{10^{33} \text{ erg s}^{-1}}}{4\pi(3.08 \times 10^{21} \text{ cm kpc}^{-1})^2} = 2.65 \times 10^{-28} \text{ erg}^{1/2} \text{ s}^{-1/2} \text{ kpc}^2 \text{ cm}^{-2}. \quad (18)$$

For several of the pulsars, the predicted flux is less than twice the UL. These pulsars seem to be near our current sensitivity limit, sensitive to distance, beaming, or local background uncertainties. We highlight a few pulsars from Table 13, in order of decreasing spindown power. The large distance (54 kpc) to the Large Magellanic Cloud (LMC) is the simplest explanation for the non-detections of PSRs J0537–6910 and J0540–6919. Using highly accurate rotation ephemerides based on RXTE X-ray observations and varied data selection cuts, we confirm the non-detection reported by Abdo et al. (2010l), with over three times as much data.

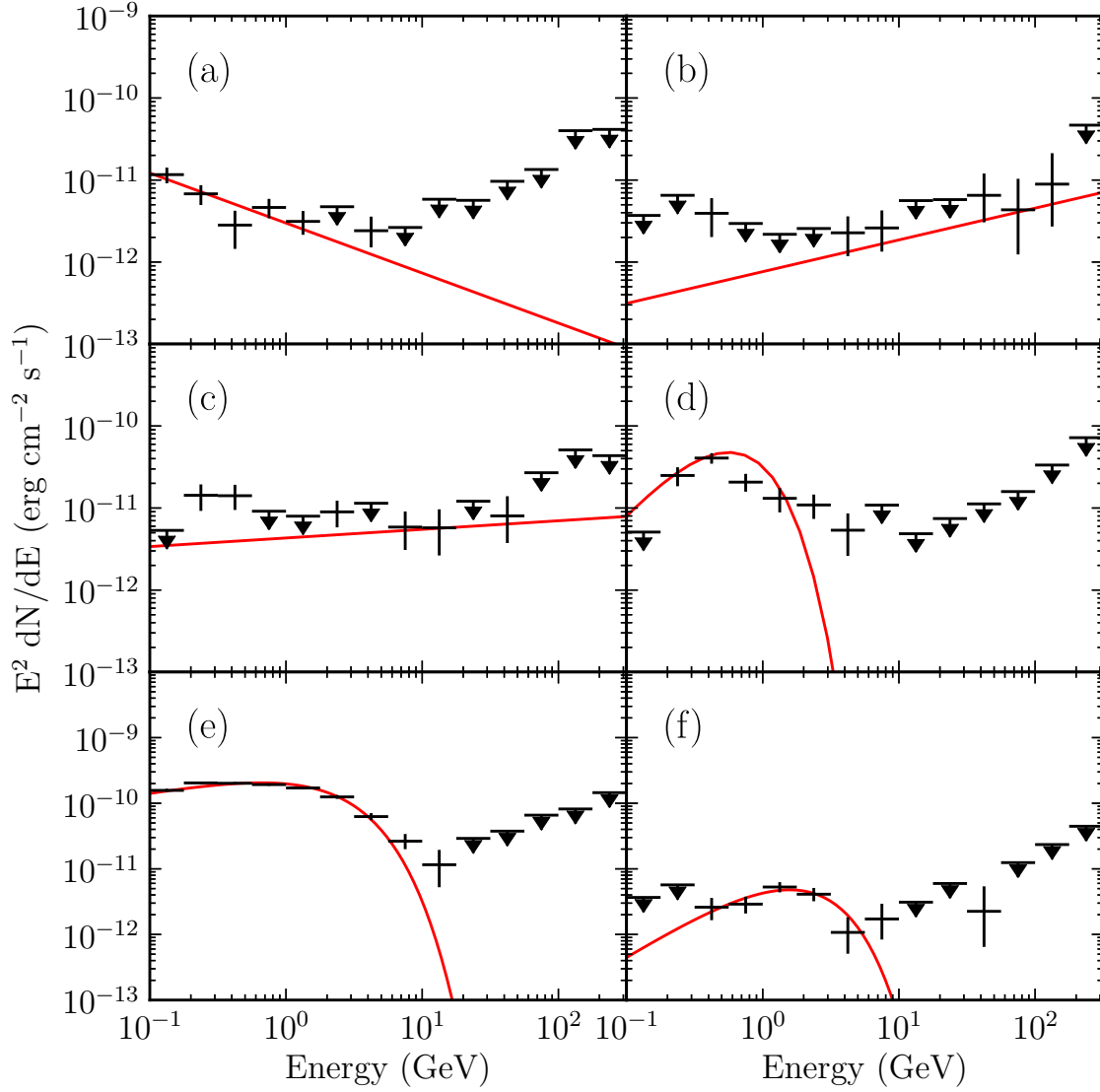


Fig. 13.— Spectral energy distributions for the off-peak phase intervals around (a) PSR J0007+7303 (b) PSR J0205+6449, (c) PSR J1410–6132, (d) PSR J1747–2958, (e) PSR J2021+4026, and (f) PSR J2124–3358. We plot a detection in those energy bands in which the source is found with $TS \geq 4$ (a 2σ detection) and report a Bayesian 95% confidence-level upper limit otherwise. The best-fit spectral model, using the full energy range, is also shown for comparison.

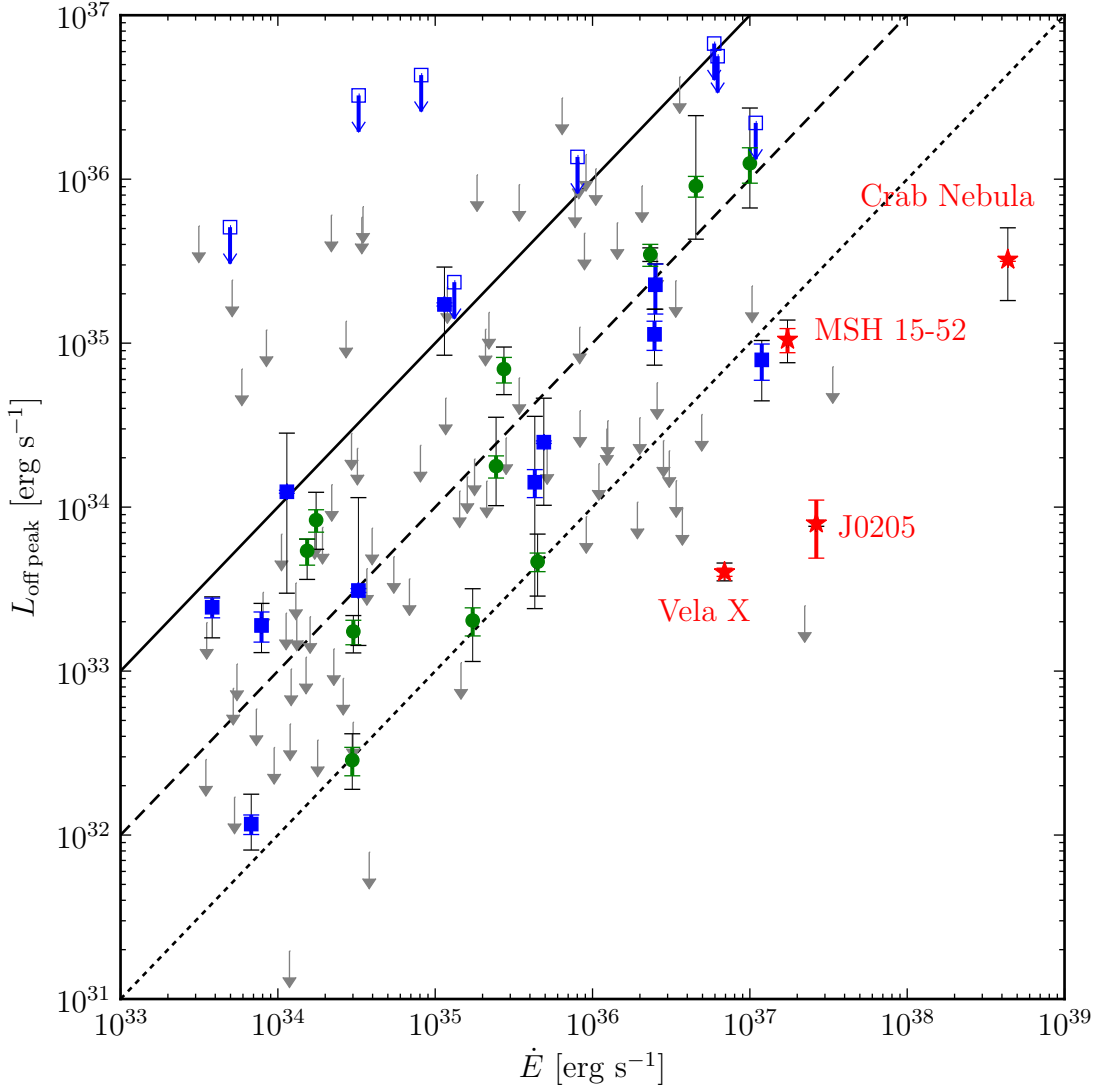


Fig. 14.— The off-peak luminosity compared to the observed pulsar spindown power. The luminosity is computed and plotted with the same convention as Figure 9. A luminosity upper limit is plotted when there is no significant off-peak emission or when there is only a distance upper limit. The star-shaped markers (colored red in the online version) represent type ‘W’ sources, the square-shaped markers (colored blue) represent type ‘M’ sources, circular markers (colored green) represent type ‘U’ sources, and the gray arrows represent non-detections. The filled blue square-shaped markers represent ‘M’ sources with a detected luminosity and the unfilled markers represent luminosity upper limits where there is only a distance upper limit. The solid, dashed, and dotted diagonals show 100%, 10%, and 1% efficiency (respectively).

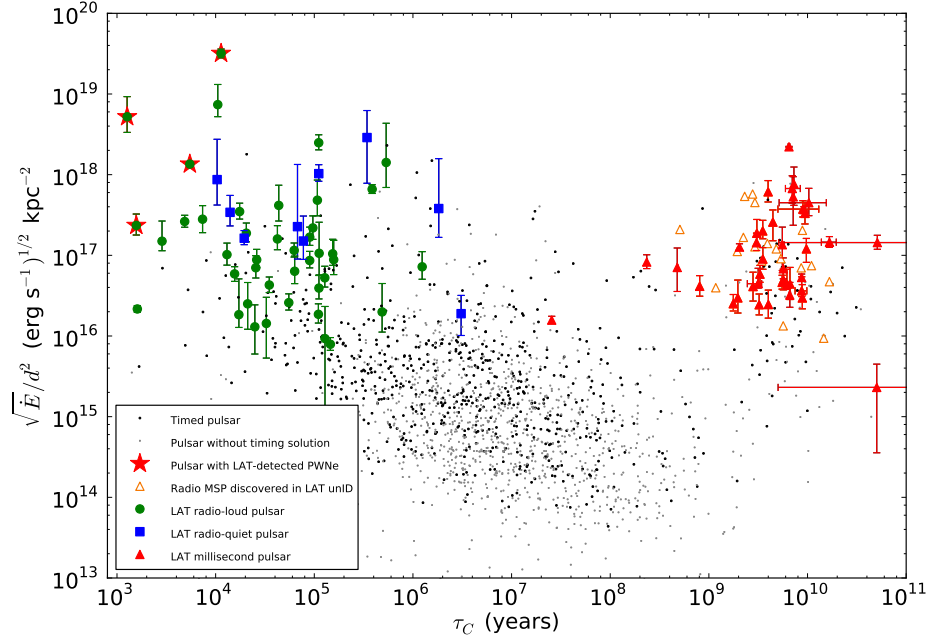


Fig. 15.— Figure-of-merit for the gamma-ray flux from a given pulsar, $\sqrt{\dot{E}/d^2}$, versus the pulsar characteristic age τ_C . For comparison with the integral energy flux G_{100} , a scale factor 2.65×10^{-28} can be applied to the y-axis (see Section 8.1). The markers are the same as in Figure 1, with red stars added for the four pulsars associated with GeV PWNe (see Section 7). Black and grey dots include all \dot{E} values, even though only high \dot{E} pulsars have been seen in gamma rays to date.

PSR J2022+3842 was thought to have the second-highest spindown power, after the Crab, of any known pulsar in the Milky Way (Arzoumanian et al. 2011). However, recent XMM data revealed that the period is 48 ms, twice the period previously seen with RXTE and GBT, with a $4\times$ smaller \dot{E} value (Arumugasamy et al. 2013). Its DM distance is 10 kpc but, situated in the heart of the Cygnus region, it (or PSR J2021+3651) may be as close as 2 kpc. For $d = 10$ kpc, $\sqrt{\dot{E}/d^2}$ is $\approx 5 \times 10^{16}$ erg $^{1/2}$ s $^{-1/2}$ kpc $^{-2}$. The LAT detects a source 0 $^{\circ}$ 06 away, with a spectrum adequately modeled by a power law with or without an exponential cutoff. The pulsar is difficult to detect in radio and suffers large timing noise; consequently, phase-folding the LAT data is difficult and blind period searches of the LAT data are hampered by the intense background. If unpulsed emission is confirmed, this could be an example of magnetospheric emission with low modulation and 100% duty cycle. The second-highest spindown power now belongs to the radio-quiet, X-ray PSR J1813–1749. It has a highly uncertain distance, probably greater than 5 kpc (Halpern et al. 2012), and is undetected with the LAT.

PSRs J1400–6325 and J1747–2809 are also distant, with slightly smaller $\sqrt{\dot{E}/d^2}$. The latter lies toward the Galactic center where the diffuse background level is intense and source density is high. The former is undetected in radio and was timed in X-rays; however, phase-connected ephemerides cover the *Fermi* mission epoch only partially. The nearest source with $TS \geq 25$ in an internal 3-year source list, constructed in a similar fashion to 2FGL, is more than 1° away. Curiously, Geminga and PSR J0007+7303 are the only X-ray loud, radio-quiet pulsars detected by the LAT.

Both J1617–5055 and J1930+1852 are roughly in the same $\sqrt{\dot{E}/d^2} \approx 10^{17}$ erg $^{1/2}$ s $^{-1/2}$ kpc $^{-2}$ range as the previous pulsars. J1617–5055 has particularly strong timing noise and the weekly to monthly Parkes observations did not allow a phase-connected timing solution covering the mission epoch. The rotation ephemeris was finally obtained in 2011 but gamma-ray phase-folding reveals no hint of pulsations⁸. The nearest LAT source (0 $^{\circ}$ 2 away) is the PWN HESS J1616–508 (Lande et al. 2012). PSR J1930+1852 has an accurate radio ephemeris but no hint of gamma-ray pulsations. The nearest 2FGL source with $TS \geq 25$ is 0 $^{\circ}$ 55 away, nine times the 95% error radius for that source. Of the 14 pulsars for which Ng & Romani (2008) fit the nebular torus seen in X-rays, PSR J1930+1852, in PWN G54.1 + 03 (Acero et al. 2013b), has the smallest angle between the spin axis and the line of sight. The radio beam intersects the Earth but the equatorial gamma-ray emission may not.

Skipping down the list, PSR J1928+1746 is within twice the 95% error radius of a source in an internal 3-year source list, but phase-folding provides no significant pulsations.

⁸R. Shannon, private communication

Table 13. Undetected RPPs with spindown power $\dot{E} > 10^{36}$ erg s $^{-1}$

PSR	\dot{E} (10^{36} erg s $^{-1}$)	Distance (kpc)	b ($^{\circ}$)	Flux UL ^a (10^{-12} erg cm $^{-2}$ s $^{-1}$)	$\sqrt{10^{33}\dot{E}}/(4\pi d^2)$ (10^{-12} erg cm $^{-2}$ s $^{-1}$)
J0537–6910	488	53.7	–31.7	LMC	2
J0540–6919	146	53.7	–31.5	LMC	1
J1813–1749	56	> 5 ^b	–0.02	< 26	< 80
J1400–6325	51	11.3	–1.6	< 6	15
J1747–2809	44	13.3	0.08	< 15	10
J2022+3842	30	10.0	0.96	SNR G76.9 + 1.0	14
J1617–5055	16	6.8	–0.28	PWN ^c	23
J1930+1852	12	7.0	0.27	< 16	18
J1849–0001	9.8	7.0	0.53	< 15	17
J1846–0258	8.1	5.8	–0.24	< 26	22
J1811–1925	6.5	5 ^b	–0.35	< 24	27
J1838–0655	5.6	...	–0.20	UnId ^c	...
J1856+0245	4.6	9.0	0.06	< 27	7
J1935+2025	4.6	6.2	–0.20	< 18	15
J1524–5625	3.2	2.8	0.35	< 6	61
J1913+1011	2.9	4.8	–0.17	< 23	20
J1826–1334	2.8	3.9	–0.69	PWN ^c	29
J1803–2137	2.2	4.4	0.15	J1803.3–2148 (110)	20
J1837–0604	2.0	6.4	0.27	< 31	9
J1809–1917	1.8	3.5	0.08	< 26	28
J1301–6305	1.7	6.7	–0.24	< 14	8
J1614–5048	1.6	7.9	0.17	< 31	5
J1828–1101	1.6	6.6	0.04	< 23	8
J1928+1746	1.6	5.8	0.11	J1928.8+1740 (33)	10
J1341–6220	1.4	11.1	–0.04	< 18	3
J1437–5959	1.4	8.1	0.23	< 19	5
J0855–4644	1.1	< 0.9 ^b	–1.0	< 10	> 340
J1831–0952	1.1	4.0	–0.13	< 17	17

^aAn energy flux upper limit is given if the pulsar is not within $0^{\circ}2$ of a LAT 2FGL source; otherwise the source name is given, with the integral energy flux above 100 MeV in parentheses. LMC is the Large Magellanic Cloud (Abdo et al. 2010).

^bDistance constraint references: PSR J0855–4644 from Acero et al. (2013a); PSR J1811–1925 by association with SNR G11.2–0.3 (Tam & Roberts 2003); PSR J1813–1749 from Halpern et al. (2012).

^cThe LAT detects a (possible) PWN extending to the pulsar position: HESS J1616–508, centered $\approx 0^{\circ}2$ away (Lande et al. 2012); HESS J1825–137, centered $\approx 0^{\circ}5$ away (Grondin et al. 2011); HESS J1837–069, centered $\approx 0^{\circ}4$ away, may be a PWN powered by this pulsar (Lande et al. 2012).

Note. — Of the 64 known RPPs with spindown power $\dot{E} > 10^{36}$ erg s $^{-1}$, the above 28 were unpulsed in GeV gamma rays as this catalog was being prepared. Column 3 is distance (the ATNF database DIST1 parameter unless noted otherwise) and Column 4 is Galactic latitude. The upper limit in Column 5 is calculated using the all-sky model as described in Section 8.2. The heuristic spindown luminosity in Column 6 corresponds to the lower diagonal in Figure 9.

Similar explorations for $10^{35} < \dot{E} < 10^{36}$ erg s⁻¹ led to the discovery of PSR J1913+0904 as a gamma-ray pulsar, listed in Table 4. Three pulsars for which the nearby high- TS LAT source is a gamma-ray pulsar are PSRs J1524–5625, J1803–2137, and J1831–0952. Searching the off-peak phase interval of the nearby pulsar reduces background but did not allow us to detect pulsations. PSR J1524–5625 bears special mention because of the large spread between the upper limit and the heuristic energy flux. If the NE2001 DM distance is accurate this could be a gamma-quiet pulsar candidate. The same is also true for PSR J0855–4644, again with a distance caveat.

In 2FGL, 83 sources have a ‘PSR’ identification and are in this catalog. An additional 27 sources have a ‘psr’ association, meaning that the pulsar lies within the error ellipse of the source but 5σ pulsations were not seen when 2FGL was written. Since then, pulsations for 12 of the 27 have allowed firm identification, included in this catalog. Of the remaining 15, 12 are radio MSPs discovered at LAT source positions (Section 3.3). The remaining three spatial associations of young RPPs with 2FGL sources are PSRs J1632–4818, J1717–5800, and J1928+1746. The last was discussed above. PSR J1632–4818 is a typical gamma-ray candidate ($\dot{E} = 4.8 \times 10^{34}$ erg s⁻¹, $d \approx 8$ kpc, $b = -0^\circ 21$) except that it has one of the highest surface B-fields of any RPP. It shows no hint of pulsations and the 2FGL association could be a chance spatial coincidence, or PSR J1632–4818 may be a candidate for unpulsed magnetospheric emission. PSR J1717–5800 is almost certainly a chance spatial coincidence: it is well below the empirical deathline ($\dot{E} = 2.3 \times 10^{32}$ erg s⁻¹) and not nearby ($d \approx 3.5$ kpc), although with low background so far from the Galactic plane ($b = -11^\circ 5$). Folding with an archival ephemeris showed no hint of pulsations.

Figure 15 illustrates the utility of the commonly used detectability metric $\sqrt{\dot{E}}/d^2$, showing an approximate LAT threshold of $\sim 10^{16}$ erg^{1/2}s^{-1/2}kpc⁻². While the difficulty of establishing reliable distances for many pulsars (Section 4.2) makes quantitative comparison challenging, it is clear that this is a good predictor of pulsar detectability. However, as emphasized by Romani et al. (2011) there are a number of pulsars with a high detectability metric not seen by the LAT (Section 8), indicated by black dots above the LAT threshold in Figure 15. In some cases, underestimated distances may explain the non-detections; however, in other cases the pulsars have accurate parallax measurements and the gamma-ray beam must either miss Earth (e.g., a pulsar viewed at small ζ for outer-magnetosphere models) or have light curves with very small modulation amplitude (e.g., a large, unpulsed component for SG or aligned polar cap models). We note that in many cases no LAT flux is detected in the pulsar direction, so beaming presents the most likely explanation.

In summary: many non-detections are due to large distances and/or background. In a few rare cases, our rotation ephemeris allows inadequate phase-folding. However, we are also

accumulating a sample of “gamma-quiet” pulsars as well as a sample of possibly unpulsed gamma-ray pulsars, that is, for which the emission from the magnetosphere is unmodulated.

8.2. Flux Upper Limits and Sensitivity

Figure 16 maps the LAT sensitivity on the sky for the phase-averaged detection of a point source with a pulsar-like spectrum, for the 3-year data set. To build the map, we started from the all-sky source model. For each point on a $0^\circ.15$ grid we added an additional pulsar-like point source with fixed parameters of $\Gamma = 1.8$ and $E_{\text{cut}} = 2$ GeV for the PLEC1 spectral shape (see Section 6) and re-fit the data. We then use the corresponding likelihood, as a function of the flux, to determine 95% confidence-level upper limits. This underestimates the actual sensitivity for two reasons. First, for weak sources near the threshold, leaving Γ and E_{cut} as free parameters during the data fitting increases the likelihood function peak width by a factor of two. Second, we claim a point-source detection only if the signal can be localized. To account for these effects, we increased the derived flux limit by a factor of two. The result is consistent with the measured fluxes of detected sources. The apparent fluctuations along the Galactic plane are a consequence of the limitations of the interstellar emission model used to represent the diffuse background. Discontinuities result from different optimized normalizations for the regions of interest selected for the all-sky analysis that were used to calculate the sensitivity limits. The upper limits in Table 13 correspond to the values nearest the positions of those pulsars.

Figure 17 shows the LAT pulsars’ integral energy flux from 0.1 to 100 GeV (G_{100} , see Section 6), versus their Galactic latitude b . Also shown is the latitude dependence of the sensitivity from Figure 16, averaged over longitude, with the 10% and 90% percentile limits. The minimum detectable flux increases with the background level, causing a selection bias against low-latitude pulsars seen against the bright Galactic background. Below 20×10^{-12} erg cm $^{-2}$ s $^{-1}$ all but one pulsar was discovered in gamma rays by phase-folding using rotation ephemerides obtained from radio or X-ray data. The gamma-selected pulsars, those discovered in a blind period search of LAT data, are brighter in gamma rays. Dormody et al. (2011) found that the blind-search sensitivity is a factor of 2.5 worse than for searches using ephemerides. The nine pulsars directly on the plane, and/or with very low fluxes, having phase-averaged significances below the formal detection threshold ($TS < 25$) were all found with ephemerides. The observed minimum energy fluxes for LAT pulsars are below 5×10^{-12} erg cm $^{-2}$ s $^{-1}$ far from the Galactic plane. PSR J2240+5832 ($l, b = 106^\circ.6, -0^\circ.11$) is an example of a pulsar with measured flux right at the sensitivity threshold. Figure A-8 shows that it has particularly narrow peaks, facilitating its pulsed detection without

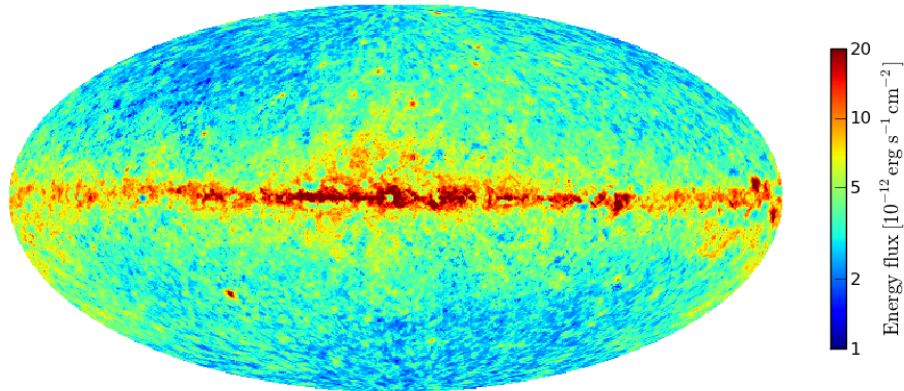


Fig. 16.— Hammer-Aitoff projection of the LAT 3-year sky-survey energy flux sensitivity above 100 MeV, assuming a pulsar-like exponentially cutoff power law energy spectrum.

improving its phase-averaged significance. Pulsars with spectral parameters far from the average values assumed for the sensitivity sky map also lead to outlying points.

8.3. SNRs and PWNe without Detected Gamma-ray Pulsars

Table 14 compiles pulsars explored in *Fermi*-LAT studies of 19 SNRs. Eleven of the SNRs are spatially extended at GeV energies. Of these, seven have known CCOs: IC 443, Puppis A, RX J1713.7–3946, S147, W30, W44, and W28. The remaining four (Cygnus Loop, HB21, RX J0852.0–4622 and W51C) have associated CCOs, or candidate PWNe indicating the likely presence of a CCO. Their gamma-ray emission is more consistent with a single extended source than a composite system of a pulsar and a remnant. Three of the remaining GeV SNRs (Cassiopeia A, Tycho and W49B) show no evidence for extended emission at GeV or TeV energies. The final five SNRs (3C 58, Crab, Vela, MSH 15-52, and MSH 11-62) contain gamma-ray pulsars. MSH 11-62 has no off-peak detection, meaning the gamma-ray emission is completely due to nearby PSR J1105–6107 with no detected SNR contribution.

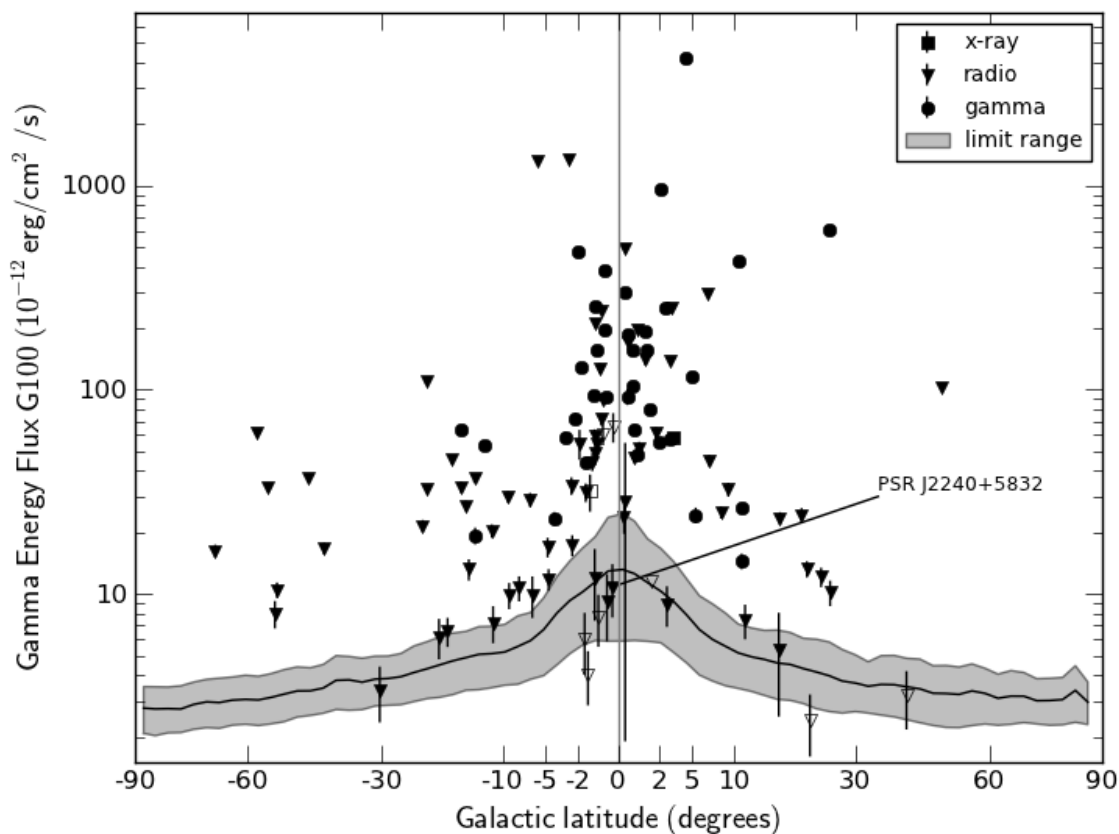


Fig. 17.— Integral energy flux from 0.1 to 100 GeV, G_{100} , versus Galactic latitude b (scaled as $b^{0.65}$ for clarity). Circles indicate gamma-selected pulsars discovered in blind period searches, while triangles and squares indicate previously known pulsars discovered in gamma rays by phase-folding with rotation ephemerides obtained from radio or X-ray data. Open symbols indicate $TS < 25$. The solid curve shows the 3-year mean sensitivity for point-source detection averaged over longitude, and the gray bands show the 10% to 90% percentile sensitivity range.

Table 14. GeV-detected SNRs with dedicated studies

SNR name	Common name(s)	Objects within Region of Interest	Gamma-ray pulsar?	Emission source	Method	Ref.
SNR 120.1+01.4	Tycho, 3C 10, SN 1572	None	...	remnant only	spectral	15
SNR 130.7+03.1	3C 58, SN 1181	PSR J0205+6449 <i>a</i>	Y	composite	off-pulse	1
SNR 184.6-05.8	Crab	PSR J0534+2200 <i>a</i>	Y	composite	off-pulse	3
SNR 180.0-01.7	S147	PSR J0538+2817	N	remnant only	spatial	18
SNR 189.1+03.0	IC443	PWN CXOU J061705.3+222127	...	remnant only	spectral	5
SNR 260.4-03.4	Puppis A	PSR J0821-4300 <i>a</i>	...	remnant only	spatial	20
SNR 263.9-03.3	Vela	PSR J0831-4406, PSR J0835-4510 <i>a</i> , PWN Vela-X	Y	composite	off-pulse	6
SNR 266.2-01.2	Vela Jr.	PSR J0855-4644, PSR J0855-4658	N	remnant only	spectral	12
SNR 291.0-00.1	MSH 11-62	PSR J1105-6107	Y	pulsar only	off-pulse	17
SNR 320.4-01.2	MSH 15-52	PSR J1513-5908 <i>a</i>	Y	composite	off-pulse	7
SNR 347.3-00.5	RX J1713.7-3946	PSR J1712-391, J1712-392, J1713-3949	...	remnant only	spectral	11
SNR 006.4-00.1	W28	PSR J1759-2307, PSR J1800-2343	N	remnant only	spectral	8
SNR 008.7-00.1	W30	PSR J1803-2137 <i>a</i> , PSR J1806-2125	N	remnant only	spatial	14
SNR 034.7-00.4	W44	PSR J1856+0113 <i>a</i>	N	remnant only	spatial	10, 16
SNR 043.3-00.2	W49B	None	N	remnant only	spatial	9
SNR 049.2-00.7	W51c	PWN CXOU J192318.5+140335	N	remnant only	spectral	2
SNR 074.0-08.5	Cygnus Loop	PWN 2XMM J204920.2+290106	...	remnant only	spatial	13, 19
SNR 089.0+04.7	HB21	None	...	remnant only	spatial	21
SNR 111.7-02.1	Cassiopeia A	CCO	N	remnant only	spectral	4

Note. — Nineteen supernova remnants with dedicated investigations of their GeV emission mechanisms. An “*a*” after a pulsar name indicates physical association with the SNR. Four of the five gamma-ray pulsars (‘Y’ in Column 4) are associated with SNRs, while PSR J1105-6107 merely overlaps.

References. — (1) Abdo et al. (2009b); (2) Abdo et al. (2009c); (3) Abdo et al. (2010d); (4) Abdo et al. (2010g); (5) Abdo et al. (2010k); (6) Abdo et al. (2010f); (7) Abdo et al. (2010b); (8) Abdo et al. (2010e); (9) Abdo et al. (2010i); (10) Abdo et al. (2010j); (11) Abdo et al. (2011b); (12) Tanaka et al. (2011); (13) Katagiri et al. (2011); (14) Ajello et al. (2012); (15) Giordano et al. (2012); (16) Uchiyama et al. (2012); (17) Slane et al. (2012); (18) Katsuta et al. (2012); (19) Katsuda et al. (2012); (20) Hewitt et al. (2012); (21) Reichardt et al. (2012);

8.4. Towards TeV Energies

Both the MAGIC and VERITAS atmospheric Cherenkov telescopes (ACTs) detected pulsations from the Crab, with an integral photon flux above 100 GeV of $\sim 6 \times 10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$ (Aliu et al. 2008, 2011; Aleksić et al. 2012). Extrapolating the LAT Crab pulsar spectrum (Table 9) predicts an integral flux above 100 GeV of less than $10^{-19} \text{ cm}^{-2} \text{ s}^{-1}$ for a power law with a pure exponential cutoff ($b = 1$ in Eq. 12). Fitting the cutoff shape yields $b = 0.44$ and the extrapolation again yields a value far below their measurement. In consequence these authors fit the joint LAT-ACT data with a broken power law, bridging any dip that may exist between the LAT and ACT energy ranges. We have examined whether the high-energy tails of the LAT data point to other pulsars that may be detectable by ground-based instruments, or that may help distinguish between the various emission models.

For our brightest and hardest pulsars (some LAT events aligned in position and phase with Vela’s second gamma-ray peak exceed 50 GeV) we fit the LAT data to a broken power law. We also simulated 3-year data sets using the standard LAT tool `gtobssim` and fit the simulated data in the same way. We find that the extrapolations to ACT energies are unreliable. The main problem is that the few gamma rays with the highest energies greatly influence the spectral parameters. The fits are also sensitive to the low-energy bound of the data set. Varying the functional form further broadens the range of extrapolated fluxes, predicting anything from quick to extremely difficult detections by ground-based instruments. We chose to make no such predictions here and advise caution in extrapolating the GeV data. The LAT collaboration is currently preparing a catalog of sources detected above 10 GeV (The Fermi-LAT Collaboration 2013). Of the 27 sources associated with known pulsars, 20 (11) have significant pulsations in the range $> 10 \text{ GeV}$ ($> 25 \text{ GeV}$).

9. Multiwavelength Counterparts

9.1. X-ray Properties

Gamma-ray pulsars are usually observed to release most of their pulsed energy in the GeV range, but they are inherently multiwavelength objects. X-ray emission associated with individual pulsars is often detected with a significance higher than 5σ . Extensive work on X-ray pulsars was enabled by *ROSAT* and *ASCA* (Becker & Truemper 1997). The X-ray flux can be pulsed non-thermal emission from the magnetosphere; blackbody thermal emission from the neutron star surface, either pulsed or unpulsed; or extended emission from a PWN energized by particles accelerated by the pulsar. Information from X-ray observations should be included in any audit of the rotational energy loss of a given pulsar.

To characterize the X-ray spectra of LAT-detected pulsars, we use only photons with energies from 0.3-10 keV collected by any of the major contemporary observatories operating in the soft X-ray band: *Chandra/ACIS* (Garmire et al. 2003), *XMM-Newton* (Strüder et al. 2001; Turner et al. 2001), *SWIFT/XRT* (Burrows et al. 2005), and *Suzaku* (Mitsuda et al. 2007). Unlike in the gamma-ray band, the X-ray coverage of LAT pulsars is uneven since the majority of the newly discovered pulsars (which account for half of the entries in the present catalog) have never been the targets of deep X-ray observations, while for other well-known gamma-ray pulsars, such as Crab, Vela, and Geminga extensive observations have been carried out. However, all LAT pulsars do have some degree of X-ray coverage, ranging from few-ksec shallow snapshots with *SWIFT/XRT* to orbit-long, deep observations by *Chandra*, *XMM-Newton* or *Suzaku*.

Tables 15 and 16 compile X-ray spectral results for all pulsars in this catalog. Given the complex phenomenology of pulsar X-ray emission, we have attempted to categorize the fluxes in a manner that will support comparisons and statistical studies focused on the system energy audit.

The status and quality for X-ray detections are indicated as follows: ‘0’ indicates no confirmed X-ray counterpart (or a purely thermal emission without a non-thermal spectral component), ‘1’ indicates that a counterpart has been identified but with too few counts for further characterization, and ‘2’ indicates sufficient information for spectral characterization (e.g., Ray et al. 2011; Abdo et al. 2012). An ad hoc analysis was performed for some pulsars for which the standard analysis could not be applied (e.g., owing to the very intense thermal component of the spectrum of Vela or to an active galaxy near PSR J1418–6058); we designate them as type ‘2*’ pulsars.

We consider an X-ray counterpart to be established if (i) X-ray pulsations have been detected, (ii) X-ray and radio pulsar positions coincide, or (iii) LAT timing (Ray et al. 2011) yields a position good enough to claim a high confidence identification with an X-ray source. The probability of finding a serendipitous source located inside a typical *Chandra* error circle is less than 0.0005 (Ebisawa et al. 2005; Novara et al. 2006); however, the probability increases by a factor of ~ 50 for *Suzaku* observations owing to a more-limited spatial resolution. Thus, we label all the objects found by *Suzaku* as ‘1*’ to indicate that there is a non-negligible possibility of a chance coincidence. All the pulsar and nebular spectra have been modeled as absorbed power laws. Blackbody components have been added to the spectra when statistically needed. Absorption along the line of sight has been evaluated through the fitting procedure. However, for pulsars with very low statistics we used values derived from observations taken in different bands, when available. We note that six MSPs can be fitted only with a thermal model: thus, owing to the lack of any non-thermal component, we

designate them as ‘type 0’. According to our classification scheme, we have 50 type 0, 11 type 1 and 56 type 2 pulsars. In total, 67 gamma-ray pulsars (30 radio-loud, 19 radio-quiet, and 18 millisecond) have an X-ray counterpart with a non-thermal spectral component.

For each type 2 pulsar, we checked for a possible PWN contribution. We analyzed all the data to search for extended emission through a radial brilliance study. When a PWN was found, its contribution was evaluated as follows: we extracted photons from an inner circular region containing $\sim 95\%$ of the point-like source counts, following the prescriptions suggested for each telescope. Such a region contains both the PSR and the PWN so that its X-ray flux must be fitted with two absorbed (PWN and PSR) power laws (plus a blackbody, if needed). We also selected an ad hoc outer region containing the brightest part of the nebula and fitted it with a single (PWN) power law, forcing the N_{H} and the PWN photon index values to be identical for the two (inner and outer) spectra. The PWN fluxes listed in Tables 15 and 16 are the spatially integrated fluxes for the two regions. Details on data analysis and fitting procedures for each telescope can be found in Marelli et al. (2011) and Marelli (2012).

For pulsars with a confirmed counterpart but too few photons to distinguish the spectral shape (type ‘1’) the unabsorbed flux is estimated assuming a single power-law spectrum with photon index 2 to characterize all components, an absorbing column obtained by rescaling the Galactic column density estimated from the Leiden-Argentine-Bonn Survey of Galactic H I (Kalberla et al. 2005)⁹, and the distances from Tables 5 and 6. For type 1 pulsars we assume that the combined PWN and PSR thermal contributions account for 30% of the total source flux, a value comparable to the mean value obtained for type 2 objects. For pulsars without a confirmed counterpart (type ‘0’) an upper limit is shown, again derived assuming a photon index of 2 and a signal-to-noise ratio of 3.

A plot G_{100}/F_X as a function of \dot{E} (see Figure 18) for type 2 pulsars shows a three decade spread in the G_{100}/F_X values for a given value of \dot{E} . This lack of correlation between gamma-ray and non-thermal X-ray fluxes may point to important (yet poorly-understood) differences in the geometry and/or height of the X- and gamma-ray emitting regions within the magnetosphere.

In general, radio-quiet pulsars are characterized by fainter X-ray counterparts than radio-loud pulsars. Indeed, the X-ray fluxes of LAT-discovered radio-quiet population have less scatter than the radio-loud pulsars. MSPs have the lowest gamma-to-X flux ratio with less apparent scatter than that observed in young pulsars. These results confirm and expand those obtained by Marelli et al. (2011) and Marelli (2012) for smaller samples of LAT pulsars.

⁹<http://heasrc.gsfc.nasa.gov/cgi-bin/Tools/w3nh/w3nh.pl>

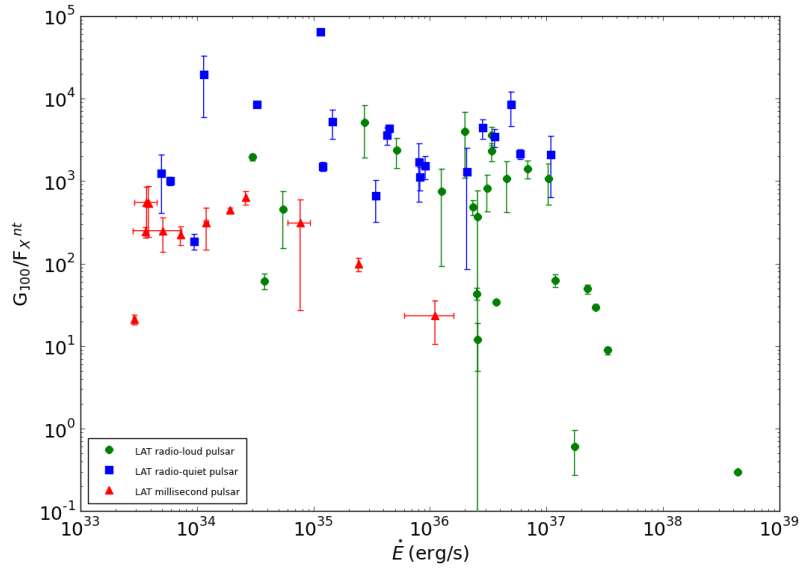


Fig. 18.— The ratio of G_{100} to the unabsorbed non-thermal X-ray flux versus \dot{E} for pulsars with good X-ray spectral measurements (‘type 2’). All G_{100} values are included, even when flagged with a † in Tables 9 and 10. Correlation with \dot{E} , if any, is weak. The markers are the same as in Figure 1. The young radio-loud pulsars have $\langle \log(G_{100}/F_X) \rangle = 2.37$ (standard deviation of 1.10), the radio-quiet population has $\langle \log(G_{100}/F_X) \rangle = 3.48$ (s.d. of 0.49) while the MSPs have $\langle \log(G_{100}/F_X) \rangle = 2.31$ (s.d. of 0.48).

Table 15. X-ray spectral parameters of young LAT-detected pulsars and their nebulae

PSR ^a	Inst ^b	N_{H} 10^{20} cm^{-2}	F_X^{nt} $10^{-13} \text{ erg cm}^{-2} \text{ s}^{-1}$	Sp. Type, pulsed ^d	G_{100}/F_X^{nt}	F_X^{pwn} $10^{-13} \text{ erg cm}^{-2} \text{ s}^{-1}$
J0007+7303 ²	X+C	$16.6^{+8.9}_{-7.6}$	0.98 ± 0.01	BB+Pow, P	4320 ± 70	$21.40^{+0.14}_{-0.17}$
J0106+4855 ⁰	Su	10^c	< 0.84	...	> 229	...
J0205+6449 ²	C	$45.0^{+1.3}_{-1.1}$	19.70 ± 0.70	BB+Pow, P	29.7 ± 2.1	24.0 ± 0.5
J0248+6021 ⁰	S	80^c	< 9.00	...	> 57.4	...
J0357+3205 ²	C	8.0 ± 4.0	$0.64^{+0.09}_{-0.06}$	Pow, P	1000^{+150}_{-100}	$3.72^{+0.62}_{-1.36}$
J0534+2200 ²	L	34.5 ± 0.2	44300 ± 1000	Pow, P	0.296 ± 0.007	396000 ± 1000
J0622+3749 ⁰	C	10^c	< 2.58	...	> 56.1	...
J0631+1036 ⁰	X	20^c	< 0.23	...	> 2070	...
J0633+0632 ²	C	$6.08^{+21.91}_{-6.08}$	0.63 ± 0.05	BB+Pow	1510 ± 170	$2.92^{+0.79}_{-0.81}$
J0633+1746 ²	L	1.07^c	$4.97^{+0.09}_{-0.27}$	BB+Pow, P	8520^{+160}_{-460}	0.172 ± 0.001
J0659+1414 ²	L	4.3 ± 0.2	$4.06^{+0.03}_{-0.59}$	BB+Pow, P	$61.8^{+6.3}_{-10.9}$	N
J0729-1448 ⁰	C	50^c	< 0.37	...	> 318	...
J0734-1559 ⁰	S	20^c	< 2.36	...	> 236	...
J0742-2822 ⁰	X	20^c	< 0.23	...	> 771	...
J0835-4510 ^{2*}	L	2.2 ± 0.5	65.1 ± 15.7	BB+Pow, P	1410 ± 340	128 ± 1
J0908-4913 ⁰	C+X	80^c	< 0.39	...	> 1130	...
J0940-5428 ⁰	C	50^c	< 0.13	...	> 314	...
J1016-5857 ²	C	$57.5^{+23.5}_{-19.5}$	$1.47^{+0.40}_{-1.31}$	Pow	370^{+137}_{-343}	$3.53^{+0.26}_{-2.77}$
J1019-5749 ⁰	S	150^c	< 1.50	...	> 51.4	...
J1023-5746 ^{2*}	C	117^{+37}_{-33}	$0.94^{+0.19}_{-0.60}$	Pow	2070^{+460}_{-1320}	$0.853^{+0.193}_{-0.593}$
J1028-5819 ¹	C+Su	< 15.0	0.45 ± 0.13	Pow	5390 ± 1660	...
J1044-5737 ^{1*}	Su	< 12.9	$0.92^{+0.26}_{-0.59}$	Pow	1700^{+490}_{-1090}	...
J1048-5832 ^{2*}	C+X	46.0 ± 2.3	$0.49^{+0.18}_{-0.34}$	Pow	4000^{+1490}_{-2800}	$0.608^{+0.224}_{-0.426}$
J1057-5226 ²	C+X	2.7 ± 0.2	$1.51^{+0.02}_{-0.13}$	BB+Pow, P	1950^{+40}_{-170}	N
J1105-6107 ⁰	C	50	< 0.08	...	> 6130	...
J1112-6103 ²	C	121^{+76}_{-50}	0.57 ± 0.28	Pow	1070 ± 560	0.63 ± 0.24
J1119-6127 ²	C+X	185^{+42}_{-38}	1.48 ± 0.21	BB+Pow, P	483 ± 84	0.601 ± 0.194
J1124-5916 ²	C	$30.0^{+2.8}_{-4.8}$	$9.78^{+1.18}_{-1.08}$	BB+Pow, P	$63.1^{+9.5}_{-9.0}$	$5.17^{+0.24}_{-0.30}$
J1135-6055 ²	C	$41.9^{+18.9}_{-15.2}$	$0.37^{+0.15}_{-0.32}$	Pow	1290^{+520}_{-1130}	$1.87^{+0.39}_{-1.05}$
J1357-6429 ²	X	$18.9^{+4.8}_{-4.5}$	0.42 ± 0.16	BB+Pow, P	809 ± 324	$3.81^{+0.38}_{-0.52}$
J1410-6132 ⁰	S	200^c	< 1.80	...	> 366	...
J1413-6205 ^{1*}	C+Su	40^c	1.41 ± 0.38	Pow	1120 ± 310	...
J1418-6058 ²	C+X	225^{+52}_{-45}	0.36 ± 0.14	Pow	8400 ± 3420	Y
J1420-6048 ^{2*}	X	202^{+161}_{-106}	1.60 ± 0.70	Pow	1060 ± 480	...
J1429-5911 ⁰	Su	80^c	< 0.73	...	> 1100	...
J1459-6053 ²	X	$42.0^{+47.0}_{-18.0}$	0.85 ± 0.23	Pow	1520 ± 420	N
J1509-5850 ²	C+X	$79.5^{+22.1}_{-16.5}$	$0.53^{+0.20}_{-0.18}$	Pow	2380^{+900}_{-830}	$2.47^{+0.32}_{-0.54}$
J1513-5908 ^{2*}	C+L ^f	91.8 ± 0.2	52 ± 18	Pow	0.612 ± 0.284	$1459.5^{+12.9}_{-12.7}$
J1531-5610 ¹	C	40^c	1.62 ± 0.53	Pow
J1620-4927 ⁰	X	40^c	< 0.67	...	> 2330	...
J1648-4611 ⁰	C	100^c	< 0.22	...	> 2520	...
J1702-4128 ¹	C+X	100^c	0.23 ± 0.07	Pow	3150^{+4500}_{-3150}	...
J1709-4429 ²	C+X	$45.6^{+4.4}_{-2.9}$	$3.78^{+0.37}_{-0.94}$	BB+Pow, P	3560^{+350}_{-890}	$8.36^{+0.52}_{-0.67}$
J1718-3825 ²	X	$40.7^{+14.6}_{-15.5}$	$1.18^{+0.58}_{-0.97}$	Pow	753^{+375}_{-622}	$1.33^{+0.55}_{-0.95}$
J1730-3350 ⁰	C+X	100^c	< 0.26	...	> 3280	...

Table 15—Continued

PSR ^a	Inst ^b	N_{H} 10 ²⁰ cm ⁻²	F_X^{nt} 10 ⁻¹³ erg cm ⁻² s ⁻¹	Sp. Type, pulsed ^d	G_{100}/F_X^{nt}	F_X^{pwn} 10 ⁻¹³ erg cm ⁻² s ⁻¹
J1732–3131 ²	C	9.39 ^{+28.58} _{-9.39}	0.37±0.13	Pow	5260±1870	...
J1741–2054 ²	C	15.3 ^{+5.1} _{-3.6}	6.24 ^{+0.34} _{-1.14}	BB+Pow	187 ⁺¹³ ₋₃₅	1.68 ^{+0.28} _{-0.34}
J1746–3239 ⁰	S	10 ^c	<1.74	...	>416	...
J1747–2958 ^{2*}	C+X	256 ⁺⁹ ₋₆	48.70 ^{+21.30} _{-6.00}	Pow	43.3 ^{+19.2} _{-6.1}	84.5 ^{+10.2} _{-4.0}
J1801–2451 ²	C+X	374 ⁺¹²⁰ ₋₁₀₈	9.97±2.02	Pow	75.3±45.9	3.27±1.24
J1803–2149 ⁰	Su	50 ^c	<0.46	...	>2030	...
J1809–2332 ²	C+X	49.2 ^{+6.8} _{-5.5}	1.32±0.30	BB+Pow	3590±820	14.4 ^{+1.6} _{-2.2}
J1813–1246 ¹	Su	153 ⁺⁶¹ ₋₅₀	1.37 ^{+0.24} _{-0.45}	Pow	1840 ⁺³³⁰ ₋₆₁₀	Y
J1826–1256 ²	C	126 ⁺⁵³ ₋₄₆	1.12±0.25	Pow	3420±770	1.52±0.33
J1833–1034 ²	X+C	210±1	66.30±1.50	Pow	8.89±1.15	721±5
J1835–1106 ⁰	C	90 ^c	<0.28
J1836+5925 ²	X+C	0.7 ^{+10.6} _{-0.7}	0.31 ^{+0.04} _{-0.21}	BB+Pow	19500 ⁺²³⁰⁰ ₋₁₃₄₀₀	N
J1838–0537 ^{1*}	Su	100 ^c	0.88±0.07	Pow	2130±230	...
J1846+0919 ⁰	S	20 ^c	<2.92	...	>83.3	...
J1907+0602 ²	X+C+Su	41.1 ^{+3.5} _{-3.0}	0.58±0.14	Pow	4410±1050	N
J1952+3252 ²	X+C	33.3±0.9	40.70±1.50	BB+Pow	33.9±1.8	77.7±1.5
J1954+2836 ⁰	Su	50 ^c	<0.75	...	>1370	...
J1957+5033 ⁰	Su	10 ^c	<0.33	...	>810	...
J1958+2841 ²	C+Su	122 ⁺⁷¹ ₋₅₄	1.37±0.66	Pow	667±325	N
J2021+3651 ²	C+X	63.8 ^{+0.50} _{-0.39}	2.15 ^{+0.24} _{-0.49}	BB+Pow	2300 ⁺²⁶⁰ ₋₅₃₀	10.4±0.6
J2021+4026 ²	C	65.2 ^{+30.5} _{-37.3}	0.15±0.01	BB+Pow	64600±4000	N
J2028+3332 ⁰	S	10 ^c	<1.57	...	>370	...
J2030+3641 ⁰	S	80 ^c	<4.52	...	>69.5	...
J2030+4415 ⁰	S	40 ^c	<2.53	...	>228	...
J2032+4127 ²	C+X	47.8 ^{+13.1} _{-14.9}	0.27 ^{+0.14} _{-0.16}	Pow	5110 ⁺²⁶³⁰ ₋₂₉₅₀	...
J2043+2740 ²	X	<3.62	0.22 ^{+0.03} _{-0.11}	Pow	453 ⁺¹¹⁷ ₋₂₅₅	...
J2055+2539 ²	X	15.1 ^{+14.2} _{-11.7}	0.43 ^{+0.12} _{-0.28}	Pow	1240 ⁺³⁵⁰ ₋₈₀₀	2.61±0.84
J2111+4606 ⁰	S	30 ^c	<2.25	...	>196	...
J2139+4716 ⁰	S	10 ^c	<3.20	...	>73.1	...
J2229+6114 ²	C+X	30 ⁺⁹ ₋₄	51.30 ^{+9.30} _{-5.80}	Pow, P	49.4 ^{+9.0} _{-5.7}	11.4 ^{+0.8} _{-1.0}
J2238+5903 ⁰	S	70 ^c	<4.49	...	>143	...
J2240+5832 ⁰	S	70 ^c	<4.60	...	>23.5	...

^aSuperscripts: 0: no X-ray detection, or no non-thermal component to the X-ray spectrum; 1: X-ray spectrum is poorly constrained; 2: Non-thermal X-ray source (see Section 9.1). An asterisk means an ad hoc analysis was necessary. The 1* indicates a Suzaku detection positionally consistent with a *Fermi* source, see Marelli et al. (2011).

^bC = *Chandra*/ACIS, X = *XMM-Newton*/PN+MOS, S = *SWIFT*/XRT, Su = *Suzaku*/XIS. For L, the results were taken from Kargaltsev & Pavlov (2008) for the Crab, from De Luca et al. (2005) for Geminga and J0659+1414, and from Mori et al. (2004) for Vela.

^cThe column density N_{H} was set to the Galactic value for the pulsar direction obtained with Webtools (<http://heasarc.gsfc.nasa.gov/docs/tools.html>), scaled for the distance.

^d“P” indicates observation of pulsed X-rays.

Note. — X-ray characteristics of young LAT pulsars. The listed fluxes are unabsorbed and non-thermal in the 0.3-10 keV energy band. The model used is an absorbed power law, plus a blackbody (BB) when statistically necessary. The

exceptions are PSRs J0633+1746 and J0659+1414 (double BB plus power law). The errors are at the 90% confidence level. For type 1 and 1* pulsars we assumed that the PWN and pulsar thermal contributions are 30% of the entire source flux. X-ray nebulae have been detected (or excluded) through brilliance profile analyses; when spectral analysis was possible the flux is in the last column, otherwise the confirmed (or not) presence of a PWN is noted by ‘Y’ (‘N’).

Table 16. X-ray spectral parameters of LAT-detected MSPs

PSR ^a	Inst ^b	N_{H} 10^{20} cm^{-2}	F_X^{nt} $10^{-13} \text{ erg cm}^{-2} \text{ s}^{-1}$	Sp.Type, pulsed ^d	G_{100}/F_X^{nt}	F_X^{pwn} $10^{-13} \text{ erg cm}^{-2} \text{ s}^{-1}$
J0023+0923 ¹	C	5 ^c	0.21 ^{+0.20} _{-0.17}	Pow	381 ⁺³⁷⁴ ₋₃₂₁	...
J0030+0451 ²	X	6.4 ^{+3.4} _{-2.4}	2.55±0.29	BB+Pow, P	241±29	N
J0034–0534 ⁰	X	<56.3	<0.06	BB	>2800	N
J0101–6422 ⁰	S	1 ^c	<2.31	...	>45.3	...
J0102+4839 ⁰	Su	5 ^c	<0.17	...	>777	...
J0218+4232 ²	X	2.70 ^{+3.76} _{-2.70}	4.62 ^{+0.43} _{-0.63}	Pow, P	98.7 ^{+12.0} _{-15.5}	N
J0340+4130 ⁰	X	5 ^c	<0.20	...	>1020	...
J0437–4715 ²	X+C	1.58 ^{+0.93} _{-1.09}	7.91 ^{+0.50} _{-0.60}	BB+Pow, P	21.1 ^{+2.5} _{-2.6}	N
J0610–2100 ⁰	S	8 ^c	<1.15	...	>57.1	...
J0613–0200 ^{2*}	X	<3.30	0.96 ^{+0.69} _{-0.44}	Pow	312 ⁺²²⁵ ₋₁₄₇	N
J0614–3329 ¹	X+Su	6.44 ^{+6.32} _{-2.01}	1.41 ^{+0.48} _{-0.58}	Pow	776 ⁺²⁶⁶ ₋₃₂₀	Y
J0751+1807 ²	X	8.74 ^{+2.10} _{-8.74}	0.59±0.09	BB+Pow	224±47	N
J1024–0719 ⁰	X	<3.58	<0.11	BB	>286	...
J1124–3653 ¹	S	5 ^c	0.45±0.25	Pow	269±156	...
J1125–5825 ⁰
J1231–1411 ²	X	11.3±5.1	4.12 ^{+0.88} _{-1.74}	BB+Pow	250 ⁺⁵⁴ ₋₁₀₆	N
J1446–4701 ⁰	S	10 ^c	<1.50	...	>49.5	...
J1514–4946 ⁰	C	50 ^c	<0.16	...	>2760	...
J1600–3053 ⁰	X	10 ^c	<0.07	BB	>2500	...
J1614–2230 ⁰	C+X	2.9 ^{+4.3} _{-2.9}	<0.29	BB	>852	N
J1658–5324 ¹	C	<14.9	1.24 ^{+0.39} _{-0.80}	Pow	233 ⁺⁷⁸ ₋₁₅₃	...
J1713+0747 ⁰	S	5 ^c	<1.80	...	>56.5	...
J1741+1351 ⁰	S	5 ^c	<2.30	Pow	>10.5	...
J1744–1134 ⁰	C	9.40 ^{+11.50} _{-9.40}	<0.26	BB	>1270	N
J1747–4036 ⁰	S	20 ^c	<1.80	...	>55.1	...
J1810+1744 ¹	C	6 ⁺²⁰ ₋₆	0.18±0.07	Pow	1290±520	...
J1823–3021A ⁰	X+C
J1858–2216 ⁰	S	10 ^c	<1.95	...	>37.1	...
J1902–5105 ⁰	Su	3 ^c	<0.56	...	>382	...
J1939+2134 ²	C	109 ⁺⁶³ ₋₄₄	3.95±0.71	Pow, P	23.2±13.3	N
J1959+2048 ²	X+C	3.72 ^{+3.79} _{-3.72}	0.55 ^{+0.10} _{-0.44}	BB+Pow	309 ⁺⁷⁶ ₋₂₅₃	0.168 ^{+0.061} _{-0.071}
J2017+0603 ¹	C	10 ^c	0.11±0.02	Pow	3030±620	...
J2043+1711 ⁰	S	6 ^c	<0.98	...	>276	...
J2047+1053 ⁰	S	7 ^c	<1.90	...	>32.5	...
J2051–0827 ⁰	C+X	<17.1	<0.04	BB	>1530	...
J2124–3358 ²	X	2.76 ^{+4.87} _{-2.76}	0.67 ^{+0.15} _{-0.34}	BB+Pow, P	550 ⁺¹²⁹ ₋₂₈₆	0.140 ^{+0.094} _{-0.069}
J2214+3000 ²	C	<21.3	0.74±0.03	Pow	441±34	...
J2215+5135 ¹	C	10 ^c	0.77±0.35	Pow	153±75	...
J2241–5236 ²	C	<24.8	0.52±0.07	Pow	638±99	...
J2302+4442 ²	X	13.0 ^{+9.1} _{-5.2}	0.68 ^{+0.14} _{-0.38}	Pow	539 ⁺¹¹⁷ ₋₃₀₅	N

^aSuperscripts: 0: no X-ray detection, or no non-thermal component to the X-ray spectrum; 1: X-ray spectrum is poorly constrained; 2: Non-thermal X-ray source (see Section 9.1). An asterisk means an ad hoc analysis was necessary.

^bC = *Chandra*/ACIS, X = *XMM-Newton*/PN+MOS, S = *SWIFT*/XRT, Su = *Suzaku*/XIS.

^cThe column density N_{H} was set to the Galactic value for the pulsar direction obtained with Webtools, scaled for

the distance.

^d“P” indicates observation of pulsed X-rays.

Note. — X-ray characteristics of LAT MSPs. The listed fluxes are unabsorbed and non-thermal in the 0.3-10 keV energy band. The model used is an absorbed power law, plus a blackbody (BB) component when statistically necessary. The exceptions is PSR J0437–4715 (double BB plus power law). The errors are at the 90% confidence level. For type 1 and 1* pulsars we assumed that the PWN and PSR thermal contributions are 30% of the entire source flux. X-ray nebulae have been detected (or excluded) through brilliance profile analyses; when spectral analysis was possible the flux is in the last column of the table, otherwise the confirmed (or not) presence of a PWN is noted by ‘Y’ (‘N’).

9.2. Optical Properties

Only seven of the *Fermi* pulsars are firmly identified at ultraviolet (UV), optical and/or infrared (IR) wavelengths. Six are solitary young-to-middle-aged pulsars (Crab, Vela, PSR B1509–58, PSR B0656+14, Geminga, PSR B1055–52), all detected in the optical and some also in the IR and UV. One is an MSP in a binary system, only detected in the near-UV (PSR J0437–4715). All were identified prior to the launch of *Fermi*, mostly in the 1990s (Mignani 2011). In the last decade possible counterparts were found for PSR B1951+32 (Butler et al. 2002) and the solitary MSP PSR J1024–0719 (Sutaria et al. 2003), prior to their detection as gamma-ray pulsars. Furthermore, PSR J1124–5916 has been associated with a bright, optical PWN, although no point source has been identified as a potential counterpart (Zharikov et al. 2008). While companion stars have been identified for eight of the binary *Fermi* pulsars, with four more obtained recently (Breton et al. 2013), our discussion is focused on the optical emission properties of the pulsars only, and not of their binary companions. Table 17 summarizes these results, and includes 49 *Fermi* pulsars, both solitary and binary.

Since the launch of *Fermi* there have been no deep systematic optical observations of gamma-ray pulsars. A quick-look survey carried out with 2–4 m class telescopes did not discover any potential counterparts (Collins et al. 2011). Dedicated follow-up observations with either the *HST* or 8 m-class telescopes have been made in a few cases, e.g. for PSR B1055–52 (Mignani et al. 2010b,a), PSR J1357–6429 (Mignani et al. 2011; Danilenko et al. 2012), PSR J1048–5832 (Mignani et al. 2011; Razzano et al. 2013), PSR J1028–5819 (Mignani et al. 2012b), PSR J0205+6449 (Shearer et al. 2013), and PSR J0007+7303 (Mignani et al. 2013b). Apart from PSR B1055–52 for which an optical counterpart had been previously identified, new counterparts were detected only for PSR J1357–6429 and PSR J0205+6449. A bright, near-IR PWN was also found coincident with PSR J1833–1034 (Zajczyk et al. 2012).

We derived optical upper limits for gamma-ray pulsars in two ways: by compiling information from previous publications and by searching public optical archives for unpublished or serendipitous observations of *Fermi* pulsars. We did not include observations from optical/IR sky surveys, as the limiting fluxes of these surveys are usually too shallow to derive constraining upper limits on the pulsar optical/IR emission. We considered both solitary pulsars and pulsars in binary systems since, in the latter case, the upper limit applies to both the pulsar and its companion. For pulsars associated with PWNe or with a detected binary companion, we assumed the nebula or the companion flux as a very conservative upper limit on the pulsar emission. In all, upper limits exist for 38 *Fermi* pulsars, though with different sensitivity limits.

Pulsars are located at different distances and are affected by different amounts of in-

terstellar extinction. To investigate their emission properties at optical energies, we first computed the extinction-corrected energy fluxes in the different pass bands. To avoid bias in comparing with X-ray energy flux densities, we used the hydrogen column density N_{H} derived from the X-ray spectral fits (Marelli et al. 2011) for this calculation. We derived the interstellar reddening $E(B - V)$ from N_{H} , using the relation of Predehl & Schmitt (1995) with $R_V \equiv A_V/E(B - V) = 3.1$, and computed the extinction in the different pass bands according to the differential extinction coefficients of Fitzpatrick (1999). We note that the uncertainties of the extinction values derived in this way are dominated by the accuracy on the N_{H} determination from the X-ray spectral fits and the uncertainty on the $N_{\text{H}}/E(B - V)$ ratio that depends on the dust-to-gas ratio along the line of sight and the grain properties. Ideally, to mitigate these uncertainties, one should directly measure the $E(B - V)$ using color-magnitude diagram comparison techniques, as done by Mignani et al. (2013a), for example. However, this requires suitable multiband optical/IR data over a sufficiently large angular scale for all pulsars listed in Table 17 and photometric analysis for each field, which is beyond the goal of this work. For consistency, we use V-band measurements where available. For the other cases, we extrapolated the measured flux or upper limit in the pass band closest to the V band assuming either the measured spectrum for the identified pulsars or, as a first approximation, a power law with spectral index $\alpha_{\text{O}} = 0$ for the unidentified ones. While the true spectral index of a given pulsar may differ from that we have assumed, we note that in most cases the uncertainty of the extinction-corrected flux upper limits due to this assumption is negligible compared to the uncertainty of the interstellar extinction correction.

Figure 19 shows the extinction-corrected optical energy flux plotted against the gamma-ray energy flux for 38 pulsars from Table 17 (binary pulsars where the companion star is detected are excluded). No correlation is apparent but it is clear that the gamma-ray energy flux is much greater than the optical energy flux for all of these pulsars. This is due, in part, to the fact that only seven pulsars are firmly detected in the optical. In addition, the different sensitivities of the observations produce a rather inhomogeneous set of upper limits. If one excludes the Crab and considers only the faintest detected pulsars (J0437–4715, J1057–5226, and Geminga) the optical energy flux appears to be independent of the gamma-flux across about three orders of magnitude. This is surprising as the optical luminosity ($L_{\text{opt,IR}} \propto \dot{E}^{1.70 \pm 0.03}$, Mignani et al. 2012a) and the gamma-ray luminosity ($L_{\gamma} \propto \dot{E}^{1/2}$, see Figure 9) both scale as a power of the rotational energy loss \dot{E} , albeit with different slopes. Because of this mutual dependence, one would expect the optical luminosity, and hence the unabsorbed optical energy flux, to scale with L_{γ} . However, the luminosity- \dot{E} relation in the optical is computed from a very limited sample of objects and is sensitive to possible outliers. Expanding the sample of gamma-ray pulsars detected in the optical is therefore crucial to

establishing possible correlations between the emission in the two energy bands.

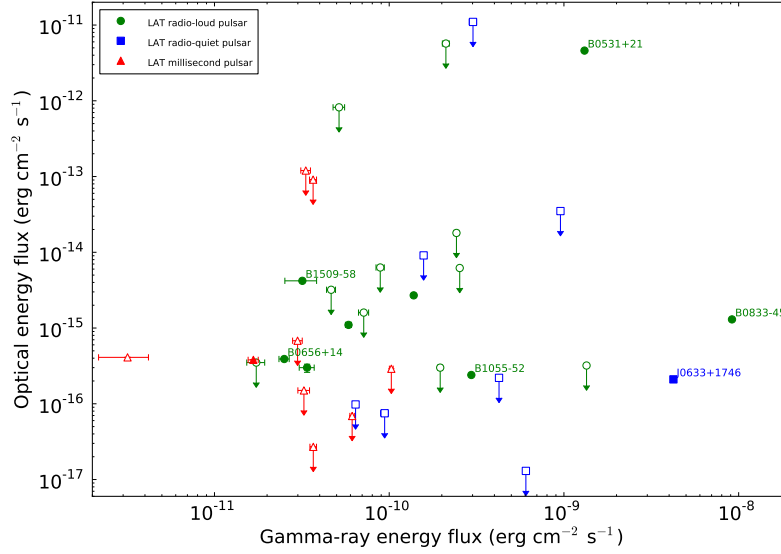


Fig. 19.— Optical energy fluxes and upper limits when available (see Table 17) versus the 0.1 to 100 GeV gamma-ray energy fluxes. All G_{100} values are included, even when flagged with a † in Tables 9 and 10. The markers are the same as in Figure 1.

Table 17. Optical observations of LAT pulsars

PSR	Magnitude	Filter	Extinction	Corrected Flux (10^{-15} erg cm^{-2} s^{-1})	Type ^a	References
J0007+7303	>27.6	R	$0.92^{+0.49}_{-0.42}$	< 0.076	U	Mignani et al. (2013b)
J0023+0923	25	g	0.37	<0.86	BC	(Breton et al. 2013)
J0030+0451	>27.0	V	$0.36^{+0.19}_{-0.13}$	< 0.08	U	Koptsevich et al. (2003)
J0034–0534	24.80	I	3.13	< 0.66	BC	Lundgren et al. (1996)
J0205+6449	>25.50	R	$2.50^{+0.07}_{-0.06}$	< 1.10	U	Collins et al. (in prep.)
J0218+4232	24.20	V	$0.15^{+0.21}_{-0.15}$	< 0.75	BC	Bassa et al. (2003)
J0248+6021	>20.9	U	4.44	< 3400	U	Theureau et al. (2011)
J0357+3205	>27.3	V	0.44 ± 0.22	< 0.07	U	De Luca et al. (2011)
J0437–4715	24.80	V	0.01	0.38 ± 0.001	BP	Kargaltsev et al. (2004)
J0534+2200	16.50	V	1.92 ± 0.01	4600 ± 47	P*	Cocke et al. (1969)
J0610–2100	26.70	V	0.44	< 0.098	BU	Pallanca et al. (2012)
J0613–0200	>26.0	B	0.18	< 0.28	BU	Collins et al. (in prep.)
J0631+1036	>24.2	V	1.11	< 1.90	U	Collins et al. (in prep.)
J0633+0632	>27.4	R	$0.34^{+1.22}_{-0.34}$	< 0.10	U	Collins et al. (in prep.)
J0633+1746	25.50	V	0.06	0.21	P*	Bignami et al. (1993)
J0659+1414	25.00	V	0.24 ± 0.01	0.39 ± 0.004	P*	Caraveo et al. (1994)
J0742–2822	>26.0	V	1.11	< 0.35	U	...
J0751+1807	25.08	R	$0.49^{+0.12}_{-0.49}$	< 0.41	BC	Bassa et al. (2006)
J0835–4510	23.60	V	0.12 ± 0.03	1.3 ± 0.03	P*	Lasker (1976)
J1024–0719	24.90	V	0.20	0.41	P+	Sutaria et al. (2003)
J1028–5819	>25.4	B	2.78	< 12	U	Mignani et al. (2012b)
J1048–5832	>27.6	V	2.56 ± 0.13	< 0.34	U	Mignani et al. (2011)
J1057–5226	25.43	V	0.15 ± 0.01	0.24 ± 0.003	P	Mignani et al. (2010b)
J1119–6127	>24.0	J	10.3 ± 2.2	< 2.90	U	Mignani et al. (2007)
J1124–5916	24.93	V	$1.67^{+0.16}_{-0.27}$	1.5	N	Zharikov et al. (2008)
J1231–1411	>26.3	V	0.63 ± 0.28	< 0.23	BU	Collins et al. (in prep.)
J1357–6429	24.60	I	$1.05^{+0.27}_{-0.25}$	0.3	P+	Mignani et al. (2011)
J1413–6205	>23.0	R	2.22	< 9.1	U	...
J1418–6058	>23.0	R	$12.5^{+2.9}_{-2.5}$	< 80000	U	...
J1513–5908	26.00	R	5.10 ± 0.01	4.2 ± 0.032	P	Wagner & Seifert (2000)
J1614–2230	24.30	R	$0.16^{+0.24}_{-0.16}$	< 0.67	BC	Bhalerao & Kulkarni (2011)
J1709–4429	>27.5	V	$2.53^{+0.24}_{-0.16}$	< 0.40	U	Mignani et al. (1999)
J1713+0747	26.00	V	0.00	< 0.12	BC	Lundgren et al. (1996)
J1718–3825	>24.0	V	$2.3^{+0.8}_{-0.9}$	< 13.3	U	...
J1744–1134	>26.3	V	$0.52^{+0.64}_{-0.52}$	< 0.27	U	Sutaria et al. (2003)
J1747–2958	>25.0	R	$14.22^{+0.33}_{-0.50}$	< 7200	U	...
J1833–1034	15.86	K	11.67 ± 0.06	64	N	Zajczyk et al. (2012)
J1810+1744	20.20	g	0.43	< 76	BC	Breton et al. (2013)
J1836+5925	>28.5	V	$0.04^{+0.59}_{-0.04}$	< 0.02	U	Halpern et al. (2002)
J1952+3252	24.50	V	1.85 ± 0.05	2.70	P+	Butler et al. (2002)
J1959+2048	20.00	V	0.21 ± 0.21	< 38	BC	Kulkarni et al. (1988)
J2017+0603	>19.1	V	0.56	< 120	BU	Cognard et al. (2011)
J2021+4026	>25.2	R	$3.62^{+1.69}_{-2.07}$	< 10.20	U	Weisskopf et al. (2011)
J2051–0827	22.3	R	0.95	< 7.2	BC	Stappers et al. (1996)
J2124–3358	>27.8	V	$0.15^{+0.27}_{-0.15}$	< 0.03	U	Mignani & Becker (2004)

Table 17—Continued

PSR	Magnitude	Filter	Extinction	Corrected Flux (10^{-15} erg cm $^{-2}$ s $^{-1}$)	Type ^a	References
J2215+5135	18.70	g	1.15	< 600	BC	Breton et al. (2013)
J2229+6114	>23.0	R	$1.67^{+0.50}_{-0.22}$	< 8.80	U	Halpern et al. (2001)
J2256-1024	26.80	g	0.14	< 0.13	BC	Breton et al. (2013)
J2302+4442	>19.6	V	$0.72^{+0.51}_{-0.29}$	< 145	BU	Cognard et al. (2011)

^aP indicates that the optically detected object is the pulsar, and P* means the detection is pulsed. P+ means that the pulsar may have been detected (candidate). B indicates a binary system, and C means that the optical detection is of the companion star. N is for an optical detection of the host nebula. U means that no optical detection was achieved, and magnitude lower limits are provided. Upper limits for binary systems apply to both the pulsar and the companion. For undetected pulsars, a measured PWN or companion flux is taken as a conservative upper limit on the pulsar flux.

Note. — LAT pulsars with optical or infrared detections or upper limits. Sensitivities vary between magnitude 22 and 27, depending on the telescope used, the pass band, the instrument, the exposure time, and the observing conditions. The Table reports the observed magnitudes (or limits) in a given band, either for the pulsar or its binary companion, the computed interstellar reddening based upon N_{H} , and the unabsorbed optical flux of the pulsar (or upper limit) in the V-band (peak wavelength $\lambda = 5500\text{\AA}$; bandwidth $\Delta\lambda = 890\text{\AA}$).

10. Discussion

This catalog expands on the results of 1PC, with the uniformly analyzed pulsar sample growing from 46 (six months) to 117 (36 months). Nearly half these pulsars were unknown before *Fermi*, having either been discovered in blind gamma-ray searches or in LAT-directed radio searches. These pulsars fall nearly equally into three main classes: young radio-loud, young radio-quiet, and millisecond. Compared to 1PC, the larger sample and detailed, uniform analysis of the 117 gamma-ray pulsars in this catalog enables more extensive population studies and evaluations of pulsar models. Although the present work is not an exhaustive review of such work, the following sections outline some of the implications of the catalog. A striking change in the pulsar sample is the dramatic increase in the MSP fraction, 34% (40/117). This doubles the 1PC fraction and is a testimony to the remarkable success of radio follow-up studies (Ray et al. 2012).

10.1. Radio and Gamma-ray Detectability

In our sample, 53% (41/77) of the young pulsars are radio loud (see Table 1), close to the 55% (21/38) fraction in 1PC. Because the sensitivity for blind searches is lower than for simple folding (see Section 8.2), the parent population must contain substantially more radio-quiet than radio-loud pulsars (to a given gamma-ray flux limit). This again is quite similar to the inference 1PC. By contrast, all known gamma-ray MSPs are radio loud (MSP J1311–3430 was detected in a blind search of LAT data, but radio pulsations were subsequently detected, Pletsch et al. 2012a; Ray et al. 2013). Although this circumstance might seem to be just a result of the difficulty of blind searches for MSPs, an analysis first done by Romani (2012) showed that the lack of gamma-ray MSPs without detectable radio emission is not an artifact. Among the 250 brightest 2FGL sources, for which the counterpart identifications are nearly complete and all sources have sensitive blind searches, only four remain unassociated with objects seen at longer wavelengths. Within this sample, only 41% (17/41) of the young pulsars are radio loud, while all 12 of the MSPs are radio loud. Even if all four of the remaining unassociated sources in this sample are gamma-ray MSPs with radio beams that do not cross our line of sight (a highly unlikely scenario) the radio-loud MSP fraction can be no smaller than 75%, a much larger fraction than seen for young pulsars.

The detected gamma-ray pulsars are clearly highly energetic, with no young pulsars and only a few MSPs detected below $\dot{E} \approx 3 \times 10^{33} \text{ erg s}^{-1}$ (Figure 1). There are apparent differences in detectability between the pulsar classes, which likely reflect differences in the radio and gamma-ray beaming. Figure 20 shows that the fraction of radio-loud young gamma-ray pulsars increases with \dot{E} (a feature first noted by Ravi et al. 2010). For $\dot{E} > 1 \times 10^{37} \text{ erg}$

s^{-1} , only one of the nine pulsars is radio-quiet. This may be an effect of the size of the magnetosphere rather than the spindown power, supported by the observation that there is only one radio-quiet pulsar with $P < 70$ ms. Thus, for gamma-ray pulsars with light-cylinder radius $R_{LC} < 200R_{NS}$, one nearly always detects the radio beam.

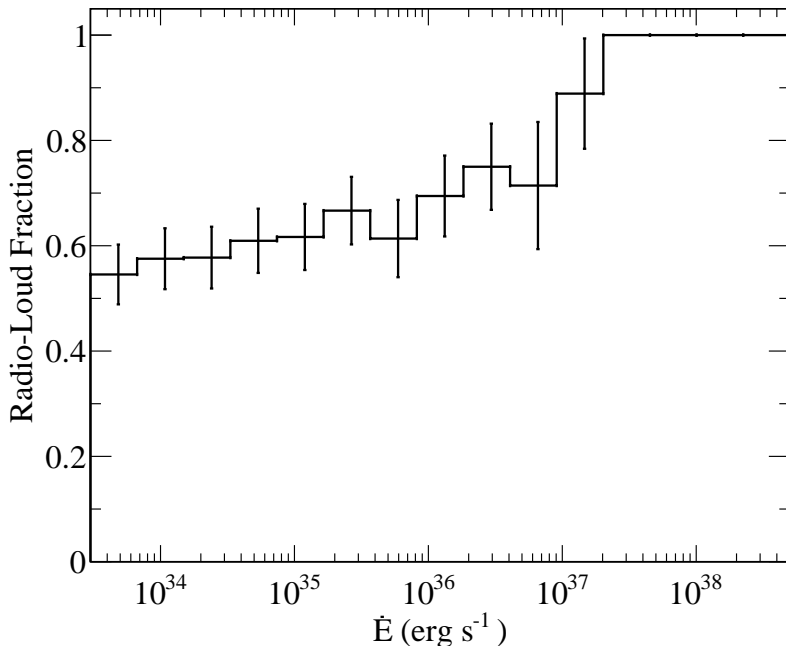


Fig. 20.— Fraction of young gamma-ray pulsars above a given \dot{E} which are radio-loud, ranging from ~ 0.5 for the least energetic to 1.0 for the most energetic.

Watters & Romani (2011) showed that these trends can indeed be explained by beaming evolution. They found that if radio beams move to high altitudes for young, short period pulsars, as argued by Karastergiou & Johnston (2007), then nearly all such gamma-ray pulsars are also radio detectable. Also, the increased fraction of low- \dot{E} radio-quiet pulsars was shown to be a natural consequence of outer-magnetosphere models with emission dominated by regions above the null charge surface, especially if the pulsars evolve toward alignment (decreasing α) on Myr timescales (Young et al. 2010). This is because the radio beam, increasingly aligned to the spin pole, is seldom visible when one views the gamma-ray beam, which is concentrated to the spin equator. Pierbattista et al. (2012) also discuss the lack of high- \dot{E} radio-quiet pulsars and the increase in the radio-quiet fraction at low- \dot{E} , plausibly addressed by an evolution toward spin alignment. However, these authors note that such alignment on Myr timescales cannot address the discrepancy seen at high- \dot{E} because of the young age of those objects. Also, while Watters & Romani (2011) find that the outer

gap (OG, see Cheng et al. 1986; Romani & Yadigaroglu 1995b) model is preferred over the lower-altitude two-pole caustic (TPC, see Dyks & Rudak 2003) model with very large statistical significance, Pierbattista et al. (2012) argue that a slot gap (SG, see Muslimov & Harding 2004) model extending to the light cylinder can provide acceptable numbers of pulsar detections and radio-loud/radio-quiet ratios. They also stress that, unless the radio beam is broader than currently modeled, both the SG and OG geometries fail to reproduce the 100% radio-loud fraction found at high \dot{E} . Further work is needed to determine whether any of these models can reproduce the detailed ratios and their evolution with spin period and spindown power.

Compared with the young, radio-loud pulsars, we see in Figures 2 and 4 that MSPs are often detected at smaller distances because of their lower luminosities and at larger Galactic latitude $|b|$ because of the much greater age of this population. The radio-quiet young pulsars are intermediate in this respect, reflecting the increased tendency noted above for gamma-only detection as the pulsars spin down. In general MSPs lie within 2 kpc, although the detection of the MSP J1823–3021A in the globular cluster NGC 6624 (Freire et al. 2011) illustrates that very energetic (young) MSPs in low background regions can be detectable by the LAT across much of the Galaxy.

In Figure 15, it is apparent that a large fraction of the MSPs above the LAT sensitivity threshold are detected as gamma-ray pulsars. Thus, the absence of gamma-ray only MSPs and the paucity of above-threshold radio MSPs undetected by the LAT suggests that MSP radio and gamma-ray beams cover a comparably-sized, and nearly-coincident, fraction of the sky.

10.2. Light Curve Trends

Additional clues to the pulsar beaming and detectability can be extracted from patterns in the radio and gamma-ray light curves. For young pulsars, the trends visible in 1PC are strengthened and extended in this catalog. Most (58/77) of these pulsars show two strong, caustic peaks significantly separated, often with significant bridge emission. The most prominent light curve trend is the anti-correlation (Figure 5) between Δ and δ , shown by Romani & Yadigaroglu (1995a) to be a general property of outer-magnetosphere models with caustic pulses. However, the sample in this catalog makes it clear that the trend is not universally followed. Certainly, MSPs show less correlation, and a significant number of young pulsars have only one strong gamma-ray peak. Watters & Romani (2011) argued that many single-peak, young pulsars fall in the $\delta \approx 0.3 - 0.6$ range, which can be attributed to either missing the first peak (P1) or a blended combination of P1 and the second peak

(P2). However several objects, especially MSPs, depart strongly from this pattern. The LAT MSPs show, on average, larger radio lags than young pulsars, as expected since MSPs have smaller magnetospheres and hence stronger aberration of the radio pulses at typical radio emission heights. The distribution in Δ for MSPs, shown in the side histogram of Figure 5, also shows a larger fraction with $\Delta > 0.5$ than for young pulsars.

One feature noted in 1PC that persists in the current sample is that for pulsars with two strong caustic peaks, the P2/P1 ratio increases with energy. This often helps us identify the harder P2 component. However, when $\Delta \sim 0.5$ the spectral evolution of the peak ratio is often weak, making peak assignment more difficult (e.g., PSR J0908–4913). For a few single-peak pulsars with sharp trailing edges, we can see the peak strength grow with photon energy – for these objects we suspect that the observed peak is a P2 component and P1 may be detectable only below the LAT energy band.

The uniformity of the light curve fitting in this catalog (Section 5.2) should greatly facilitate use of the $\delta - \Delta$ distribution to test magnetospheric models. For example, in such a treatment of the 1PC sample, Watters & Romani (2011) found that the data gave a strong statistical preference for an OG geometry to that of the TPC model. However, accurate δ values need the true phase of the magnetic axis, likely requiring careful modeling of the radio light curve and polarization. In addition, magnetospheric currents (see Kalapotharakos et al. 2012a) can provide systematic δ shifts with respect to the vacuum approximation. Even when radio emission is not detected the Δ distribution can provide a statistical test of the beaming model. For example, the preponderance of $\Delta \approx 0.5$ is very natural in SG or TPC models. Figure 6 does not show any strong correlation of Δ with \dot{E} , although some increase in the incidence of smaller Δ is expected in OG models for lower \dot{E} .

The majority of pulsars have two peaks: three-quarters for the young pulsars, and 60% of the MSPs. Figure 21 shows a sample of four recurring profile shapes, which we classify using the ratio $R_{r,f}^i$ of the half-widths of the rising and falling peak edges, included in the auxiliary files. Here, the index $i = 1, 2$ indicates the first and second gamma-ray peaks following phase $\phi = 0$. In panel (a) of the figure, the sharpest edges are the trailing edge of the first peak, $R_{r,f}^1 > 1$, and the leading edge of the second peak, $R_{r,f}^2 < 1$. In panel (d), it is the opposite: the “outer” peak edges are the most abrupt, $R_{r,f}^1 < 1$ and $R_{r,f}^2 > 1$. In panel (b), the leading edges are sharpest, $R_{r,f}^1 < 1$ while for panel (c), the trailing edges are steep, $R_{r,f}^{1,2} > 1$. About half of the profiles are classifiable with good statistical significance, and nearly all of these have sharp outer edges as in panel d. This is the expected pattern for caustics from a hollow cone. A few pulsars depart significantly from this pattern, most prominently those for which both peaks fall sharply (panel c). These tend to have $\Delta \approx 0.5$, suggesting emission from both poles. Many MSPs do not fit this simple scheme and cannot

be classified in this way with any confidence. More detailed analysis of the pulse width statistics may uncover other trends.

For two of our highest-statistics young pulsars, Vela and J1057–5226, the bridge emission shows persistent structure that can be identified as a third peak. In addition, two MSPs (J1231–1411 and J0751+1807) show significant third peaks. For Vela we have sufficient statistics to see that the phase of the third peak shifts with energy (Abdo et al. 2010o); this is not expected for a simple caustic-induced peak. To explain such structure we must go beyond simple geometrical approximation and employ models with a full treatment of the radiation spectrum and its variation through the magnetosphere (e.g., Du et al. 2011; Wang et al. 2011). As LAT statistics improve, we can expect further such ‘P3’ detections and better constraints on the origin of such pulse components.

Table 12 lists 11 pulsars with significant ‘M’-type off-peak emission, suggesting the possibility of nearly constant magnetospheric emission. However, several of these pulsars are faint or reside in regions of high and/or complex diffuse background, which complicates the spectral analysis. Therefore, we identify only four young pulsars (PSRs J0633+1746 ‘Geminga’, J1836+5925, J2021+4026, and J2055+2539) and two MSPs (PSRs J2124–3358 and J2302+4442) that have strong evidence for exponentially cutoff off-peak emission well in excess of the inferred diffuse background (including estimated systematic uncertainty). For the young pulsars this is a serious challenge to outer-magnetosphere models radiating only above the null-charge surface; such weakly pulsed emission should be rare, being expected only for nearly aligned pulsars with $\zeta \approx \pi/2$. In contrast, lower altitude emission (such as from SG or extended polar cap models, Dyks et al. 2004; Venter et al. 2009) provides a natural explanation for off-peak, non-caustic emission. The remaining five pulsars with ‘M’-type off-peak emission (J0340+4130, J1124–5916, J1620–4927, J1747–2958, and J1813–1246) should be treated with caution as the uncertainty in the flux from systematic error in distinguishing the diffuse background from magnetospheric emission for these sources makes it difficult to probe faint off-peak emission. However, lower-altitude emission may also be present in a number of these cases. For MSPs, and especially for radio-quiet young pulsars, we expect that the magnetic impact angle $|\beta| \equiv |\alpha - \zeta|$ is relatively large, i.e. the Earth line of sight passes far from the radio pole. Thus, off-peak emission for these objects may be associated with large $|\beta|$. The separatrix layer in the wind zone just outside the magnetosphere has also been suggested (Bai & Spitkovsky 2010) as a site for such gamma-ray flux.

The majority of MSPs (27) have profiles that are very similar to those of young pulsars, with a variety of double and single-peaked profiles with radio lags following the $\delta - \Delta$ trend seen in the normal radio-loud pulsars. A standard SG or OG geometry with narrow gaps can

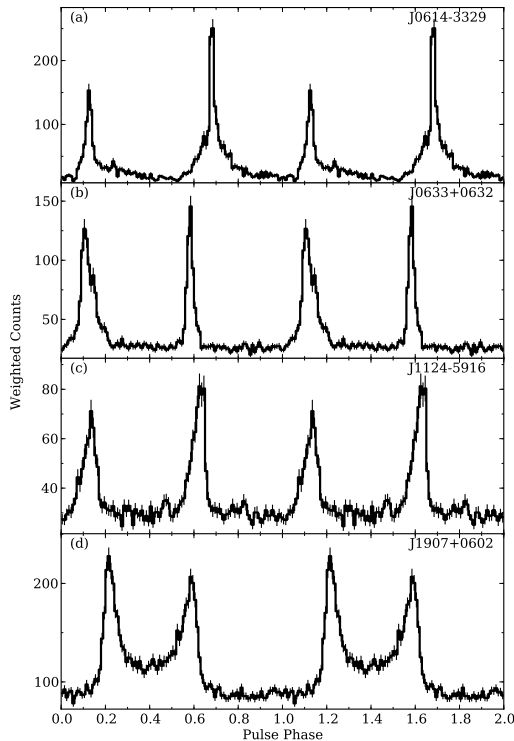


Fig. 21.— Examples of four types of two-peak gamma-ray pulse profiles: (a) PSR J0614–3329: sharpest peak edge between the two peaks, (b) PSR J0633+0632: sharp leading edge for both peaks, (c) PSR J1124–5916: sharp trailing edge for both peaks, and (d) PSR J1907+0602: “outside” peak edges are sharpest.

fit these (Venter et al. 2009), requiring surprisingly high-multiplicity pair cascades in these old pulsars. The rest have very different profile types. In particular, a relatively large fraction (6) have phase-aligned radio and gamma-ray pulses, four have $\delta \approx 0.99$ while two (PSRs J0030+0451 and J1810+1744) have $\delta \approx 0.85$ due to the particular method of defining the fiducial radio phase. The peak alignment suggests that the radio as well as the gamma-ray pulses result from caustic emission at altitudes that are a large fraction of the light cylinder in the relatively small MSP magnetospheres (Venter et al. 2012). The Crab pulsar, with sharp (caustic) main and inter-pulse radio components composed of ‘giant pulse’ emission and aligned with the two gamma-ray pulses, likely represents a similar case. If the radio emission is caustic, a very low level of polarization may result from depolarization of emission over a large range of altitudes (Dyks et al. 2004). In fact, at least four of the six MSPs that have aligned profiles do show little or no linear polarization, as well as unusually steep radio spectra and high B_{LC} (Venter et al. 2012; Espinoza et al. 2013). There are also MSP profiles

(7) where the gamma-ray peaks lead the main radio peak by ~ 0.2 – 0.3 in phase, a profile class that is so far not seen in the young gamma-ray pulsar population. A ‘pair-starved polar cap’ (PSPC) model provides good fits for this class (Venter et al. 2009), implying emission from nearly the full open zone at all altitudes and the same pole that dominates the radio pulse. Three of these MSPs, PSRs J1124–3653, J1744–1134 and J2214+3000, do not have the expected value of δ in Table 8; this is because we chose to use the opposite hemisphere when defining the fiducial radio phase. The best-fit magnetic inclination angles for all of these profile classes show a very wide range, from nearly aligned to nearly orthogonal (Johnson et al., in prep.). The $|\beta|$ values show a large range since MSPs have relatively larger radio beams than young, shorter-period pulsars, allowing the radio pulse to be visible at larger ζ .

The gamma-ray light curves in this catalog strongly suggest that the gamma-ray emission is distributed in a narrow gap bordering the closed field line boundary. Although the light curves show a large degree of regularity, none of the currently proposed emission models alone is able to account for all of the observed properties. While many light curves follow the OG pattern, some, particularly MSPs and some radio-quiet pulsars, do not fit this model and may require additional emission from other zones, such as below the null-charge surface.

10.3. Luminosity and Spectral Trends

As in 1PC, the dependence of gamma-ray luminosity on spindown power is one of our most important results. Figures 9 and 10 confirm the general trend of 1PC: young pulsars show increased efficiency for producing gamma rays as the spindown luminosity decreases toward $\dot{E} \approx 5 \times 10^{35} \text{erg s}^{-1}$, an observation which is in conflict with the two-layer OG model of Takata et al. (2010) that predicts a much flatter trend at high \dot{E} . Above $\approx 5 \times 10^{35} \text{erg s}^{-1}$ MSPs show a similar trend but there is a wide range of efficiencies below this \dot{E} . Figure 10 shows this more directly, and emphasizes the point that for most gamma-ray pulsars the apparent efficiency exceeds 10% and for the lowest \dot{E} efficiencies approach unity. This tension may be slightly mitigated if MSP moments of inertia, and thus spindown luminosities, are a few times larger than the standard $I_0 = 10^{45} \text{g cm}^2$ assumed here (Demorest et al. 2010). Nevertheless, the high efficiencies are remarkable, meaning that gamma rays trace the bulk energetics of the pulsar machine and implying that studies correlating this output with other observables provide excellent prospects for understanding pulsar magnetosphere physics.

Two factors complicate such studies. The first is the difficulty of obtaining accurate pulsar distances, which we have discussed in Section 4.2. Improvements to the distance determinations (e.g. from additional VLBI parallaxes for the radio-bright pulsars, Deller et al. 2011) provides the best hope for eliminating this large uncertainty. An additional

challenge is the uncertainty in the geometry-dependent beaming correction f_Ω (Eq. 16). While for most models and viewing geometries this is within a factor of a few of unity, for some situations (e.g., OG geometries at $\dot{E} \approx 10^{34} \text{ergs}^{-1}$ where $f_\Omega \sim 0.1$, Pierbattista et al. 2012) the correction can be quite substantial. Further, physical effects such as α -dependent variation of the gap width, changes in the gamma-ray emissivity along the gap, or differences in the detected pulse spectrum can increase the variation of f_Ω beyond simple geometrical factors.

The large scatter in Figures 9 and 10 may well be due to these two factors. Attempts to quantitatively constrain pulsar physics by fitting the luminosity evolution should certainly marginalize over these uncertainties. However, the current sample does provide a greatly improved testbed for comparing predictions of the heuristic $L_\gamma \propto \dot{E}^{1/2}$ law with more detailed predictions (e.g., Takata et al. 2010).

We can also check the \dot{E} dependence of the spectral fit parameters Γ (Figure 7) and E_{cut} . As for 1PC we see that the lowest measured Γ are near the limit of $\Gamma = 2/3$ for single-particle curvature or synchrotron radiation. In addition, there is a trend toward a softer spectrum (larger Γ) at high spindown luminosity with $\Gamma \propto \dot{E}^{0.2}$. One explanation for this trend is increased pair formation activity in high- \dot{E} pulsars, leading to pair cascades and steep radiating particle spectra. With this catalog we now have the statistics to separately probe the trend in MSPs, which appears steeper with $\Gamma \propto \dot{E}^{0.4}$.

There is no apparent trend of E_{cut} with \dot{E} . If this cutoff follows the radiation-reaction limited energy, and $L_\gamma \propto \sqrt{\dot{E}}$, then we expect $E_{\text{cut}} \propto P^{-1/4} \dot{E}^{1/8}$ so the dependence on spindown power should indeed be weak. One might also expect that MSPs with small P would have larger E_{cut} . However, the small radius of curvature in their compact magnetospheres tends to ensure that the primaries reach lower maximum energy, if radiation-reaction limited.

We note two important caveats to those wishing to use these phase-average spectral fits to constrain models. First, for the fainter pulsars (especially the MSPs) the covariance between Γ and E_{cut} is substantial and the apparent trends may be affected. Second, it is clear from our phase-resolved studies of the brighter pulsars (Abdo et al. 2010d,o,h) that phase-dependent variations in the spectral parameters can be nearly as large as the variations in the full population. Additional evidence for varying E_{cut} appears when we allow the exponent b to vary in the spectral fits, as discussed in Section 6.1. Nevertheless, non-exponential cutoffs may be present in some cases as discussed in Section 8.4.

10.4. Pulsar Population and the Millisecond Pulsar Revolution

As the LAT pulsar sample grows, our ability to make statistically powerful statements about the Galactic neutron stars and their evolution increases. Compared to radio surveys, the LAT provides a new and differently biased sample of the energetic pulsars, so many conclusions drawn from the classic radio samples need to be revisited.

Perhaps the most dramatic progress presented in this catalog is the major increase in the MSP sample from 8 objects to 40, with discoveries continuing (Table 4). In fact, the LAT has proved such an excellent signpost to nearby energetic MSPs that the LAT-guided discoveries represent a large and increasing fraction of the known energetic Galactic MSPs. For example, 70 Galactic (non-globular cluster) MSPs were known before *Fermi*; there are now 120 such MSPs, 39 of which are in the present catalog. This dominance is especially obvious for $P < 3$ ms; as noted by Ray et al. (2012) the LAT-detected MSPs are a shorter-period, more-energetic population than radio-selected MSPs.

Also, Roberts (2013) has noted that *Fermi*-guided pulsar searches have resulted in a dramatic ten-fold increase in the number of Galactic MSPs in tight binaries with pulsar-driven companion winds, the so-called ‘black widows’ and ‘redbacks’. This is because the gamma-ray signal penetrates the companion wind, flagging the source as a possible MSP and guiding repeated radio searches for the intermittently visible radio pulsations.

In addition to the Galactic MSPs, there is a population of some 120 pulsars in the globular cluster system, discovered through radio searches. As these clusters are relatively distant it is not surprising that, to date, the LAT has detected pulsed signals from only two of the youngest, most energetic cluster MSPs (Freire et al. 2011; Johnson et al. 2013). However, the detection of gamma-ray sources coincident with a dozen globular clusters (Abdo et al. 2010a; Kong et al. 2010; Tam et al. 2011) indicates a substantial MSP population (first predicted by Chen 1991). Indeed, the cluster gamma-ray flux appears to correlate with the expected MSP formation rate (Abdo et al. 2010a). It had been noticed that a large fraction of the radio MSPs in globular clusters are tight, often evaporating, black-widow-type binaries and it was suggested that this was a true difference from the Galactic field population (King et al. 2003). However, it now seems that this was largely an artifact of the very long radio dwell times used for cluster searches that allowed discovery of radio-intermittent MSPs generally undetectable to field surveys. With the LAT unassociated sources providing a signpost for deep targeted searches, a similar field black widow population has now been discovered. Thus, the LAT has uncovered a new sample of MSPs with less (or at least different) bias than the classical radio population, demanding a re-assessment of the MSP population and its evolution.

For young pulsars, there is good hope that the LAT can similarly provide a much more complete census of massive star remnants in the nearby Galaxy. For example, Watters & Romani (2011) found that the 1PC sample implied an energetic young pulsar birthrate of 1.69 ± 0.24 pulsars/100 yr, a substantial fraction of the 2.4/100 yr OB star birthrate and the 2.30 ± 0.48 SNe/100 yr rate (Li et al. 2011). This catalog will allow further refinement of this comparison.

Finally, we conclude by noting at least one expected pulsar population remains missing from the present catalog. As in 1PC, no pulsed detection of a young spin-powered object in a massive binary has yet been made. Several percent of the pulsar population is expected to pass through this channel (the progenitor of the double neutron star binaries). A few such pulsars are known from the radio, but like PSR B1259–63 are detected only at large orbital separation. For short period systems, plasma from the companion wind presumably prevents radio pulsar detection. However, gamma rays provide an excellent signpost to such wind-absorbed pulsars as witnessed by recent success in detecting short-period black widows and redbacks in LAT sources. It is probable that some of the gamma-ray detected massive binaries host spin-powered pulsars, a possibility discussed by Dubus (2006). With improved search techniques, such objects may appear in future LAT pulsar catalogs.

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A. Appendix: Sample light curves and spectra.

Figures A-1 to A-8 show sample multiband light curves. Figures A-9 to A-17 show sample gamma-ray spectra. Plots for all pulsars in this catalog are available in the on-line auxiliary material, detailed in Appendix B.

For each pulsar, the top frame of the light curve figure shows the 0.1 to 100 GeV gamma-ray light curve, with the same data repeated over two rotations, to clarify structures near phases 0 and 1. The pulsar name follows the same color-code as the markers in the plots in the body of the text: red for MSPs ; blue for young radio-quiet, and green for young loud. The letter P gives the rotation period. The letter H gives the H-test value for the gamma-ray light curve in the top frame. The letters d and D give the lag δ of the first gamma peak relative to the radio fiducial phase, and the separation Δ of the outermost gamma peaks, respectively. The dark (blue online) curve over the phase range $\phi \in [0.0, 1.0)$ in the top frame is the 0.1 to 100 GeV profile fit described in Section 5. The radio profile, when it exists, is drawn (red curve online) in the top frame, also repeated over 2 rotations, with the observing frequency as indicated on the figure. The radio telescopes that provided the profiles shown are indicated by NAN (Nançay), PKS (Parkes), JBO (Jodrell Bank), AO (Arecibo), GBT (Green Bank), and WSRT (Westerbork). The shaded gray region in the top frame represents the off-peak interval as defined in Section 7.1. The lower frames show the gamma-ray light curves in the indicated energy bands. The horizontal, dashed lines indicate the estimated background levels, for the gamma-ray light curves, and the associated uncertainties as described in Section 5.

For each pulsar, the gamma-ray spectral points are from individual energy-band fits in which the pulsar spectrum is approximated as a pure power law. For energy bands in which the pulsar is detected with $TS < 9$ we report 95% confidence level upper limits. The solid black line corresponds to the best-fit PLEC1 model from the full energy range fit, while the red (online) dashed lines represent the 1σ confidence region on the best-fit model. For

pulsars where a pure exponential cutoff ($b = 1$) is disfavored ($TS_{b\text{ free}} \geq 9$) we also show the PLEC fit (blue online).

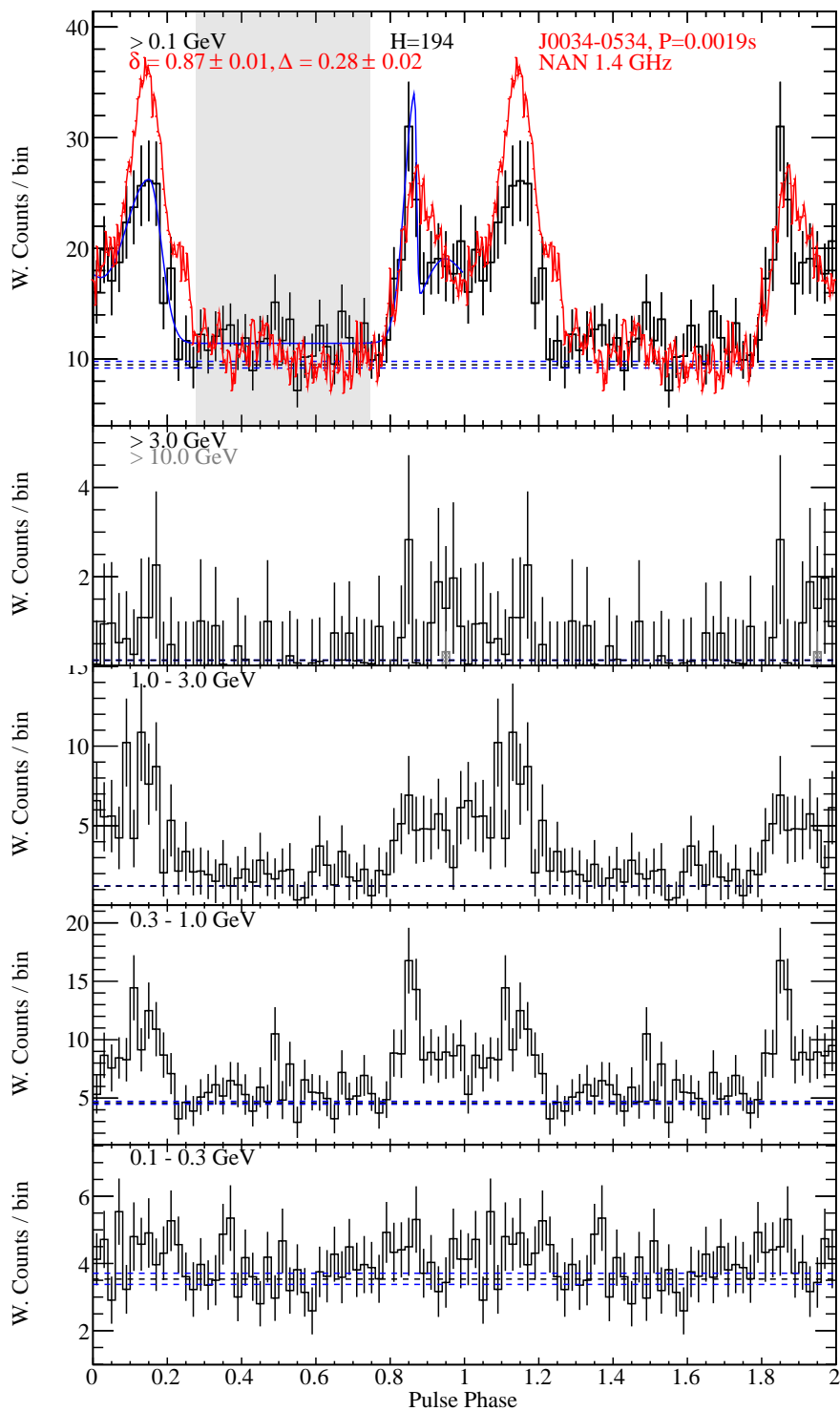


Fig. A-1.— Light curves of PSR J0034–0534. Appendix A text describes the different frames. This MSP has radio and gamma-ray peaks occurring at the same phase and is a strong case for near-constant magnetospheric emission. Phase 0 (the “fiducial point”, see Section 5) for this pulsar was set to the median point between the two radio peaks.

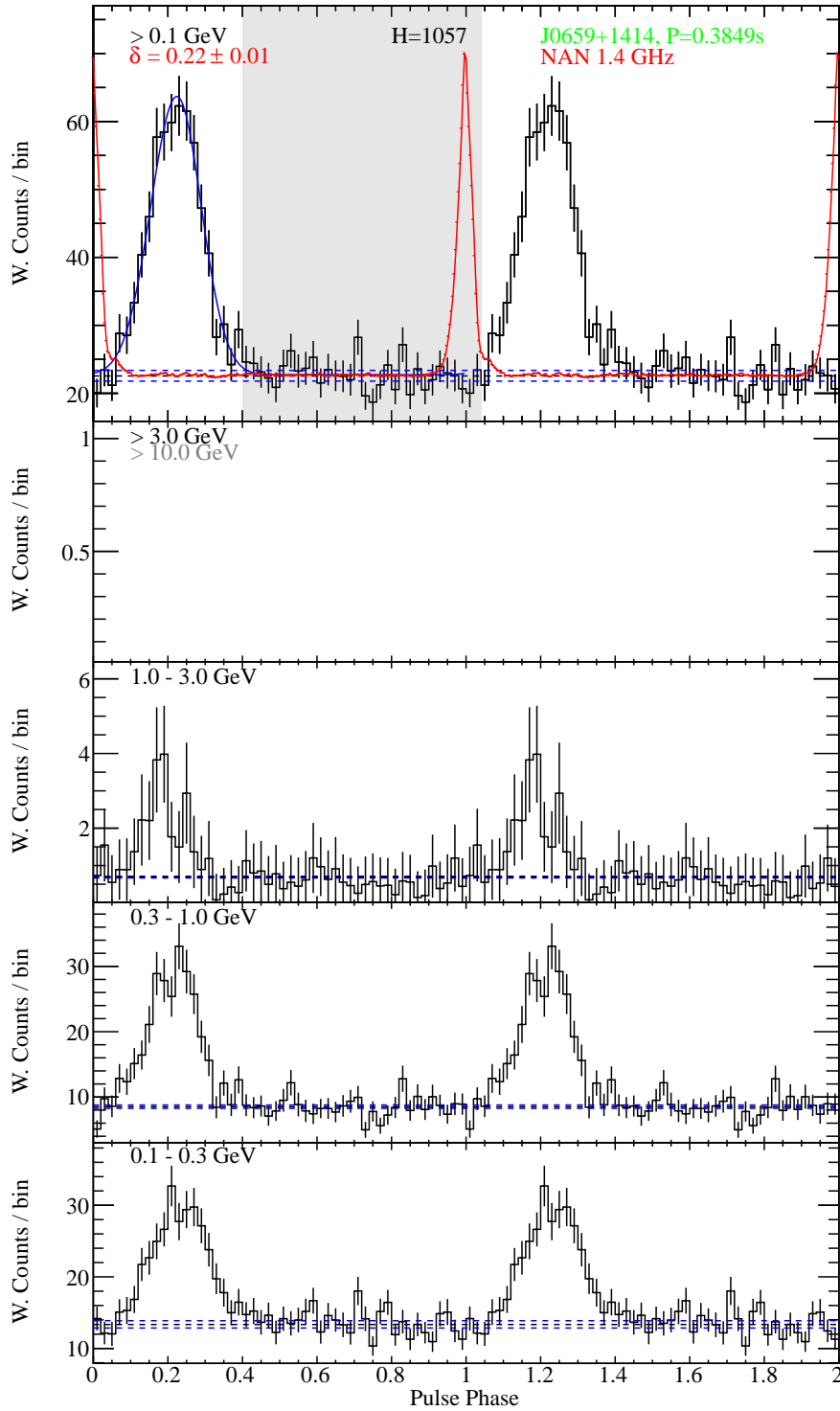


Fig. A-2.— Light curves of PSR J0659+1414. Appendix A text describes the different frames. This is an example of a single-peaked gamma-ray light curve.

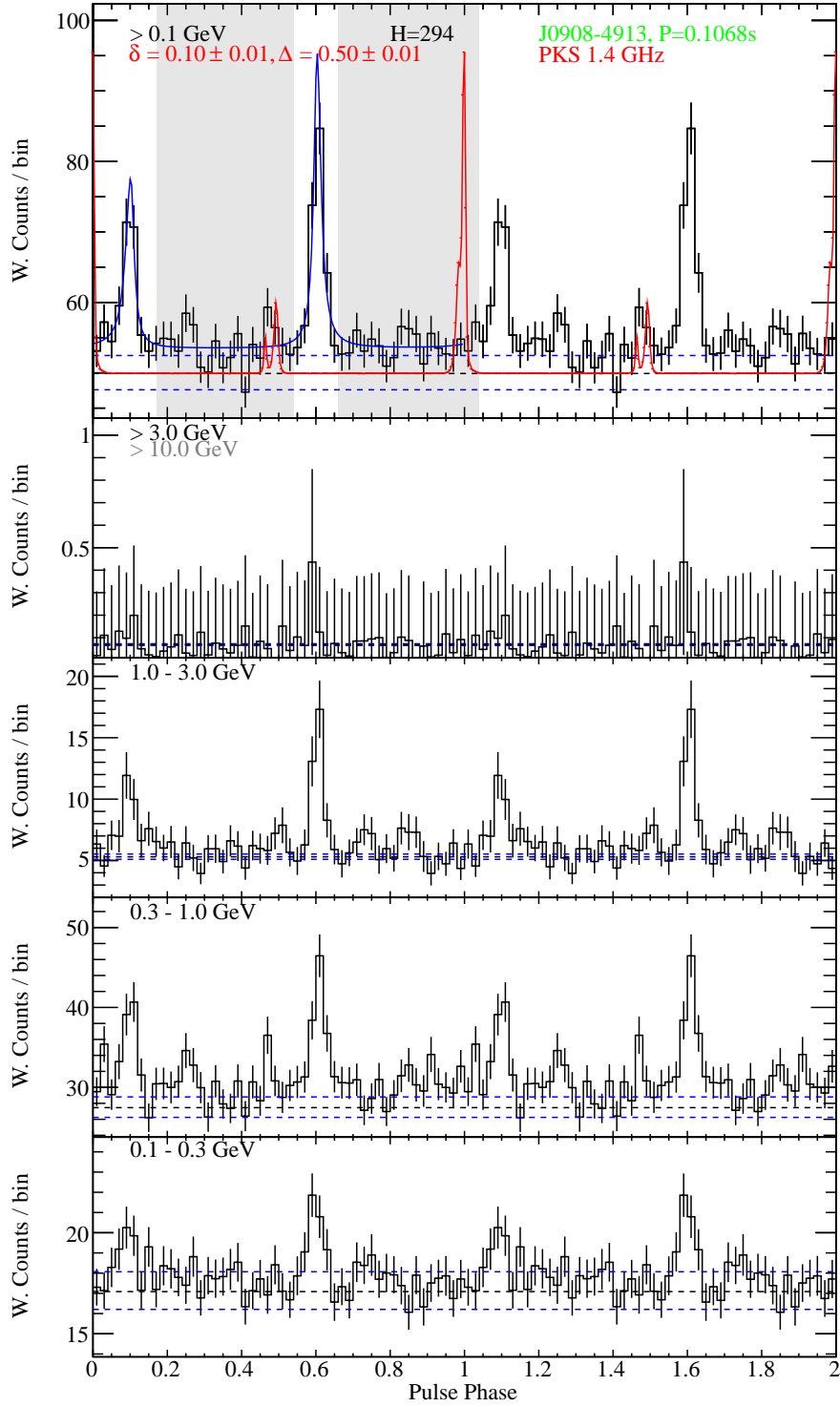


Fig. A-3.— Light curves of PSR J0908–4913. Appendix A text describes the different frames. This is an example of a two-peaked gamma-ray light curve with no apparent emission between the peaks, leading to a disconnected off-peak region as evident from the shaded regions in the figure.

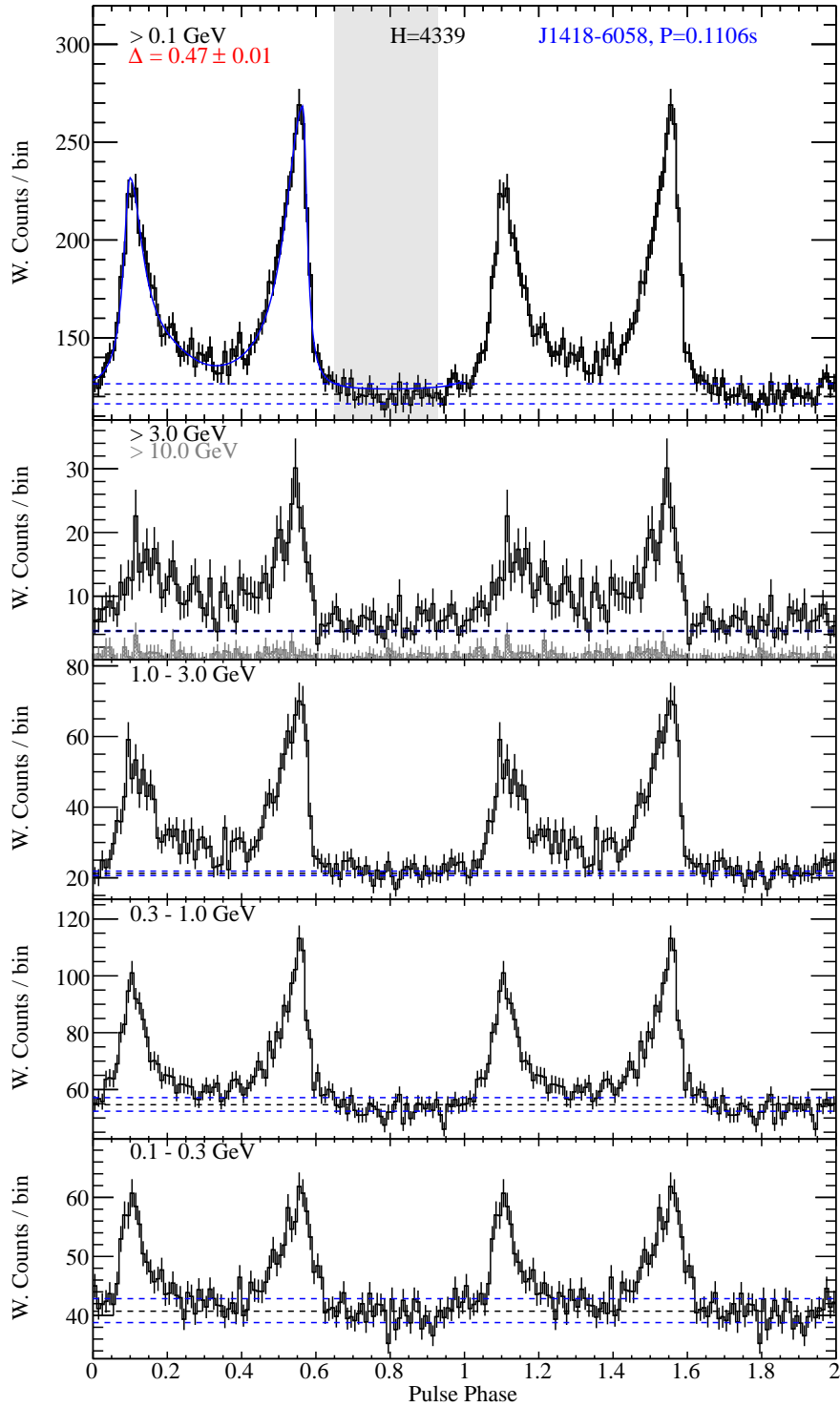


Fig. A-4.— Light curves of PSR J1418–6058. Appendix A text describes the different frames. This is an example of a radio-quiet pulsar with no radio detection and demonstrates the common two-peaks with sharp, asymmetric structure.

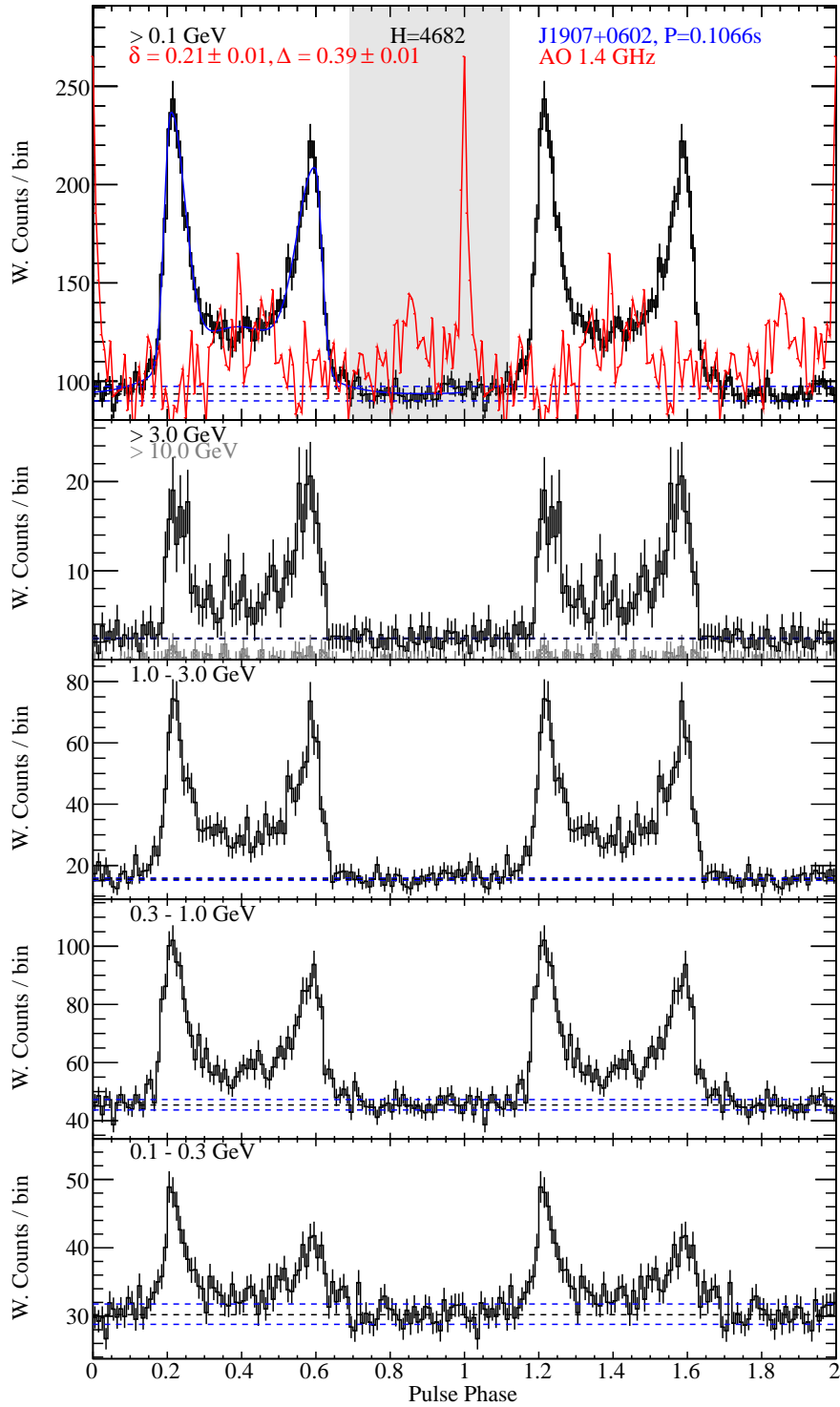


Fig. A-5.— Light curves of PSR J1907+0602s. Appendix A text describes the different frames. This gamma-ray light curve illustrates the sharp, asymmetric peak structure described in Section 10.2.

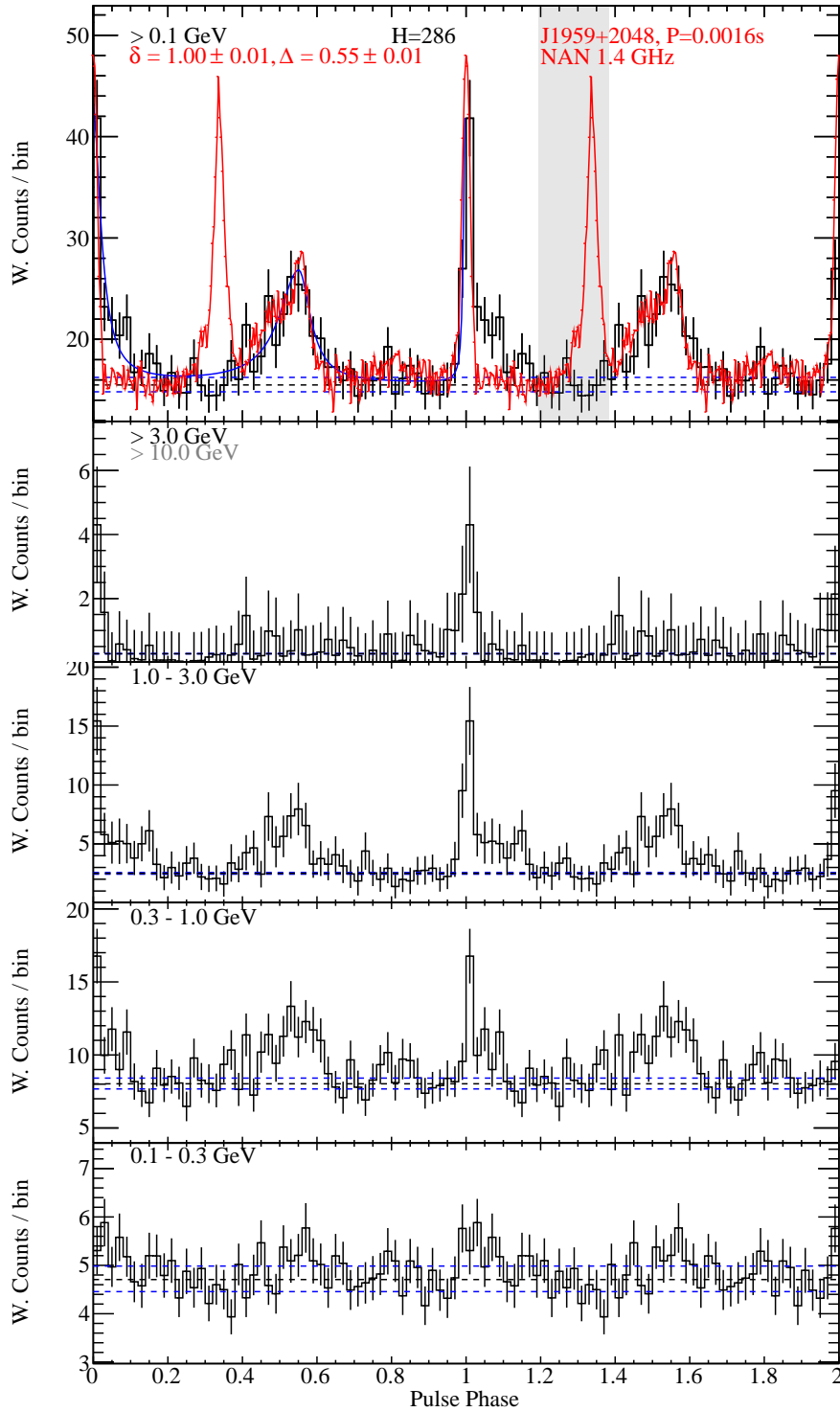


Fig. A-6.— Light curves of PSR J1959+2048. Appendix A text describes the different frames. This MSP has gamma-ray and radio peaks occurring at the same phase but with one unmatched radio peak that is not present at lower radio frequencies.

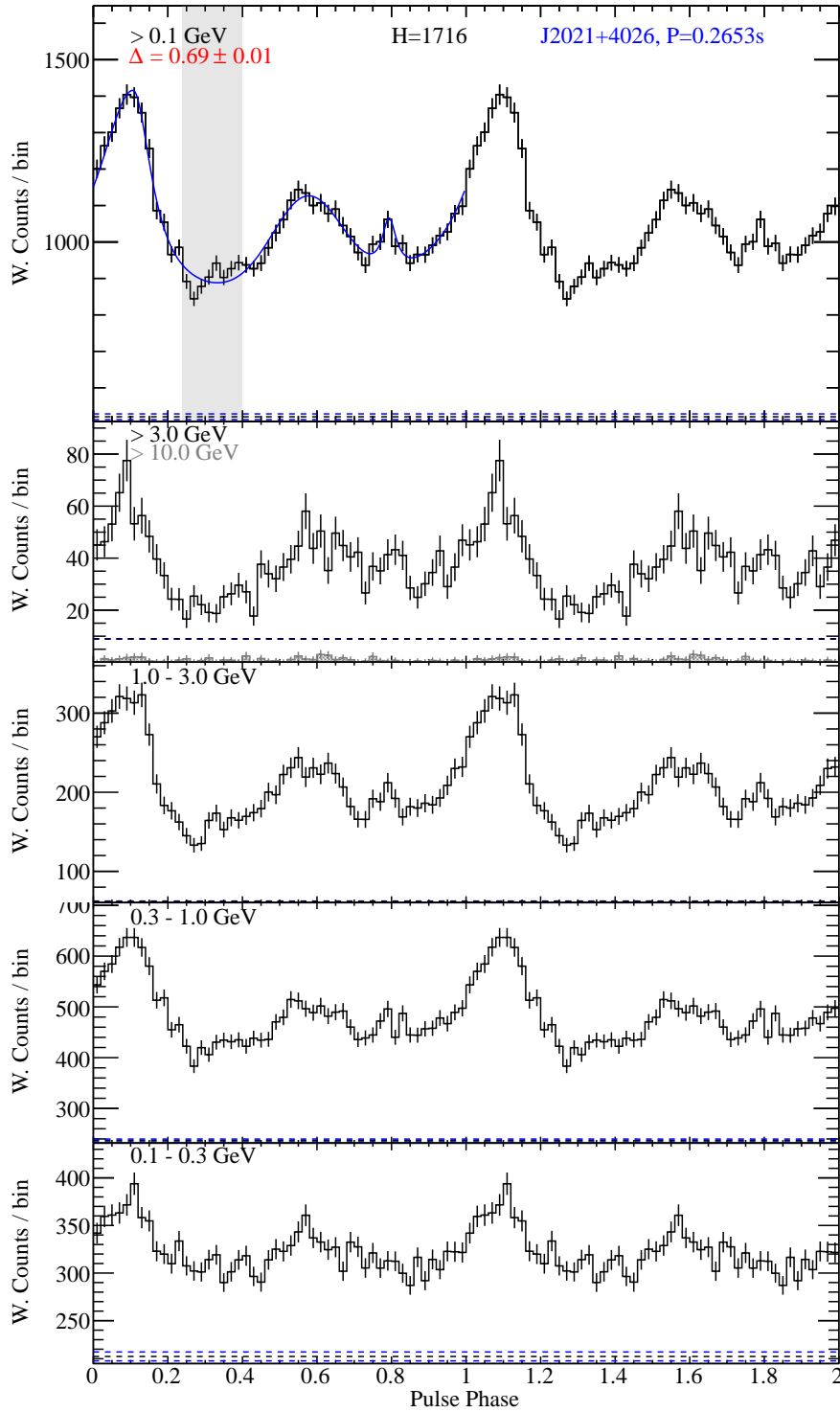


Fig. A-7.— Light curves of PSR J2021+4026. Appendix A text describes the different frames. This is an example of a radio-quiet pulsar with no radio detection and also a strong case for magnetospheric emission across the entire pulse.

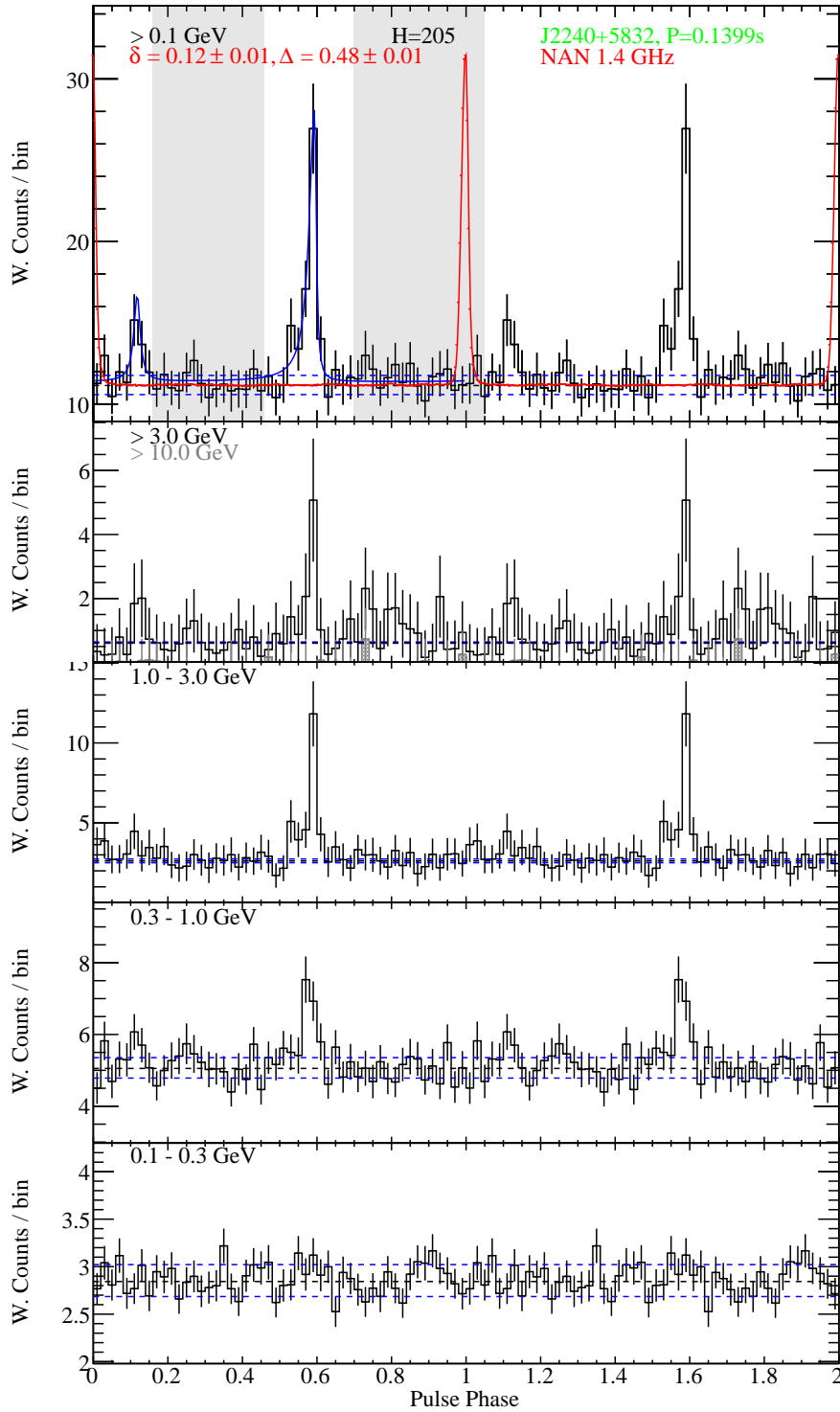


Fig. A-8.— Light curves of PSR J2240+5832. Appendix A text describes the different frames. This is an example of a two-peaked gamma-ray profile with clear evolution of the P2/P1 ratio with energy.

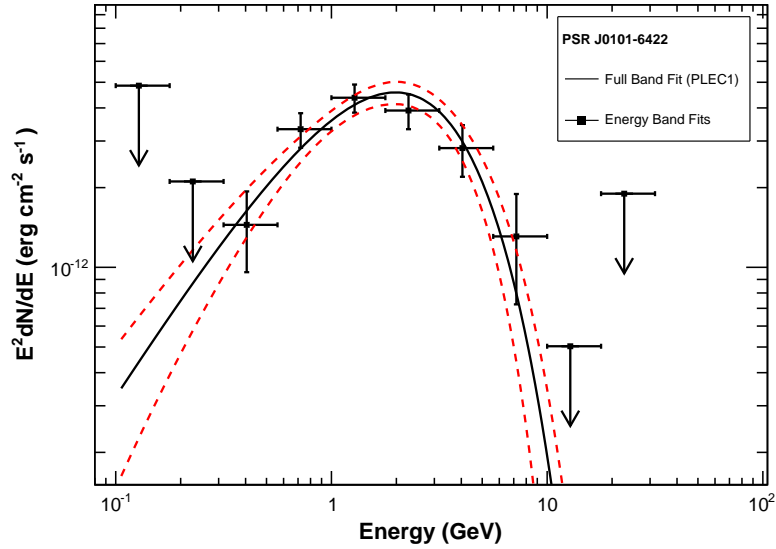


Fig. A-9.— Gamma-ray spectrum of PSR J0101–6422. Appendix A text describes the figure components. This spectrum demonstrates that pulsars are typically most significant near 1 GeV and often cannot be detected as point sources near 0.1 GeV.

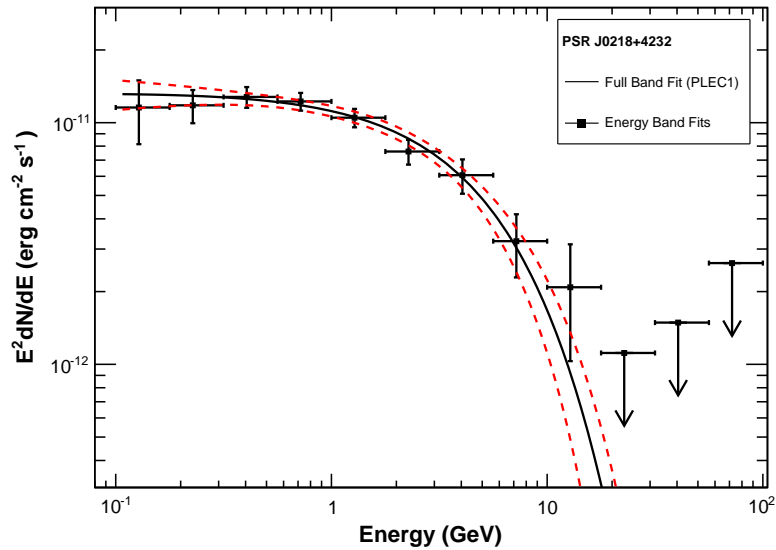


Fig. A-10.— Gamma-ray spectrum of PSR J0218+4232. Appendix A text describes the figure components. This spectrum has a cutoff on the high-energy tail of the E_{cut} distribution.

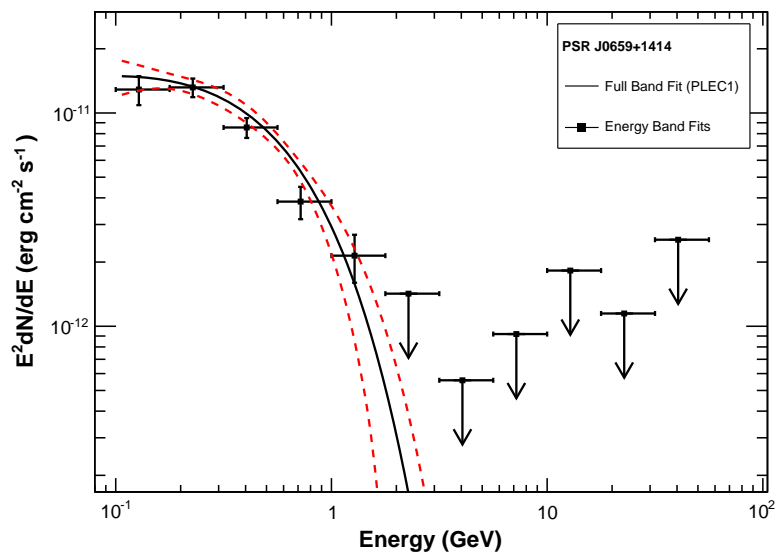


Fig. A-11.— Gamma-ray spectrum of PSR J0659+1414. Appendix A text describes the figure components. This spectrum has a cutoff on the low-energy tail of the E_{cut} distribution.

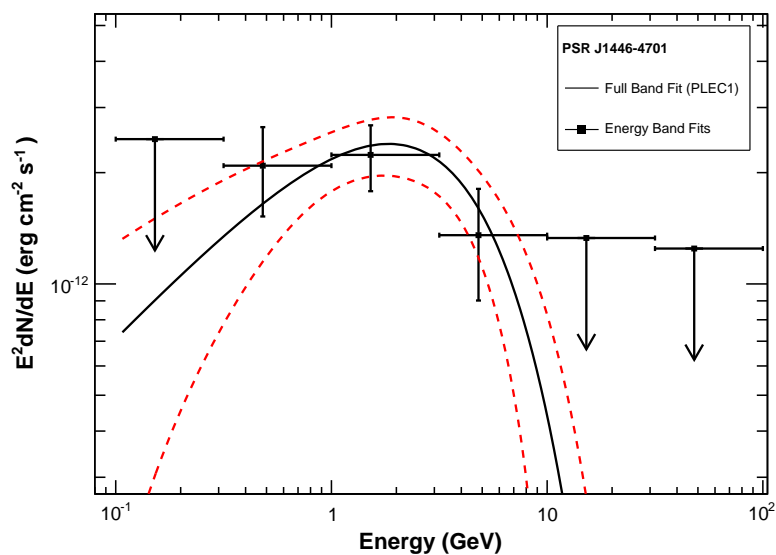


Fig. A-12.— Gamma-ray spectrum of PSR J1446–4701. Appendix A text describes the figure components. This is a faint pulsar with only a few points in the spectrum.

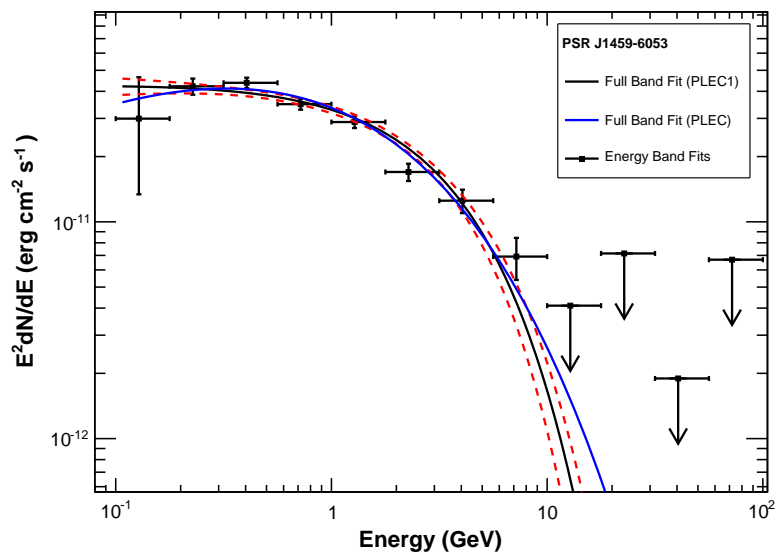


Fig. A-13.— Gamma-ray spectrum of PSR J1459–6053. Appendix A text describes the figure components. For this pulsar, $TS_{\text{bfree}} \geq 9$ and the PLEC fit (exponential cutoff parameter not fixed to $b = 1$) is shown (blue in the online version). The full energy range maximum likelihood fits are, typically, a better measure of the spectrum than fits to the points.

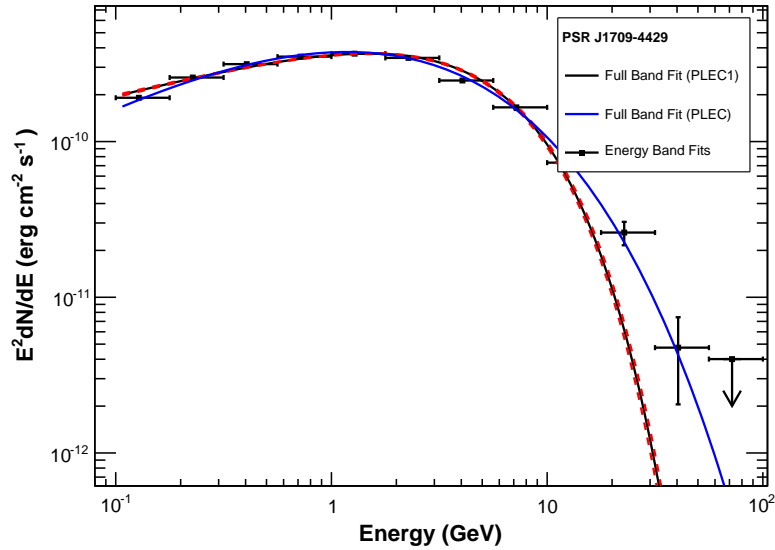


Fig. A-14.— Gamma-ray spectrum of PSR J1709–4429. Appendix A text describes the figure components. This spectrum illustrates deviations from the PLEC1 model above 1 GeV (PLEC fit shown in blue online).

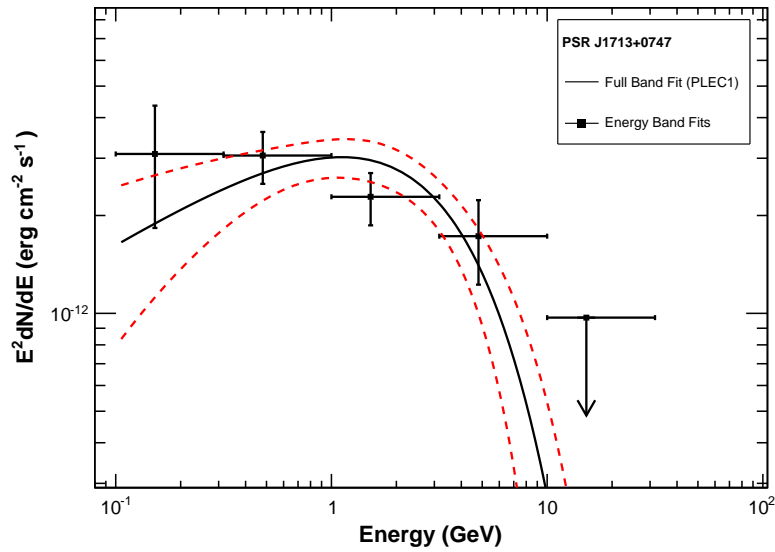


Fig. A-15.— Gamma-ray spectrum of PSR J1713+0747, Appendix A text describes the figure components. This is a faint pulsar with a reliable spectral fit and significant emission out to at least a few GeV.

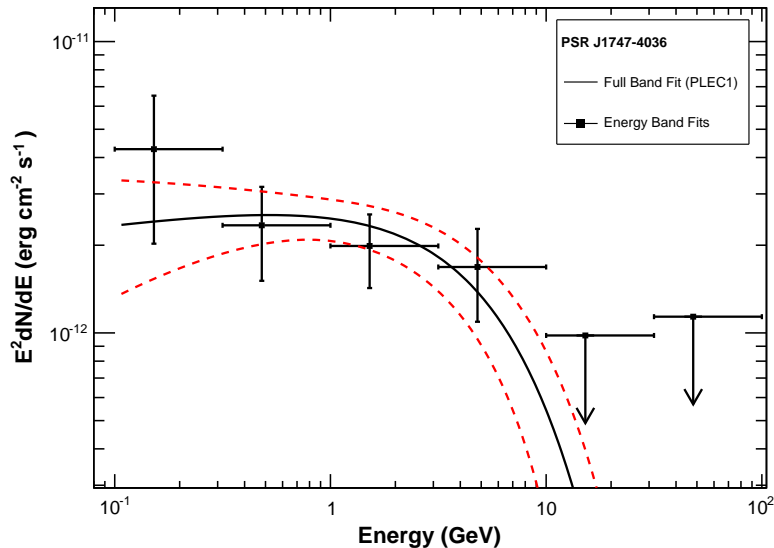


Fig. A-16.— Gamma-ray spectrum of PSR J1747–4036. Appendix A text describes the figure components. This spectrum is an example of a faint pulsar for which significant spectral curvature cannot be seen with 3 years of data.

B. Appendix: Description of the online catalog files

The complete results of the analyses described in this paper are reported in the on-line supplemental material. This supplement is an electronic archive, provided as a tarred and gzipped file (2PC_auxiliary_files_v###.tgz). Inside the archive is a directory structure containing FITS files with tables reporting the analysis results, images of the light curves and spectral results for each pulsar, a text file containing the rotation ephemeris used in the analysis of each pulsar, and individual FITS files for each pulsar with the light curves and spectra in numerical form. This structure is described in Table 18. The online material is available at http://fermi.gsfc.nasa.gov/ssc/data/access/lat/2nd_PSR_catalog/.

Additional information, such as for example the half-widths at half-maximum of the leading and trailing edges of the peak fits described in Section 5, and their uncertainties, is also provided. Another example is that in addition to the exponentially cut-off power law spectral parameters listed in Tables 9 and 10, we provide the results of the pure power-law fits, and of the fits with the b parameter free for pulsars with $TS_{b\text{ free}} \geq 9$ (see Section 6). Detailed column descriptions for the main FITS tables are in Section B.1, and detailed column descriptions for the individual pulsar FITS tables are in Section B.2.

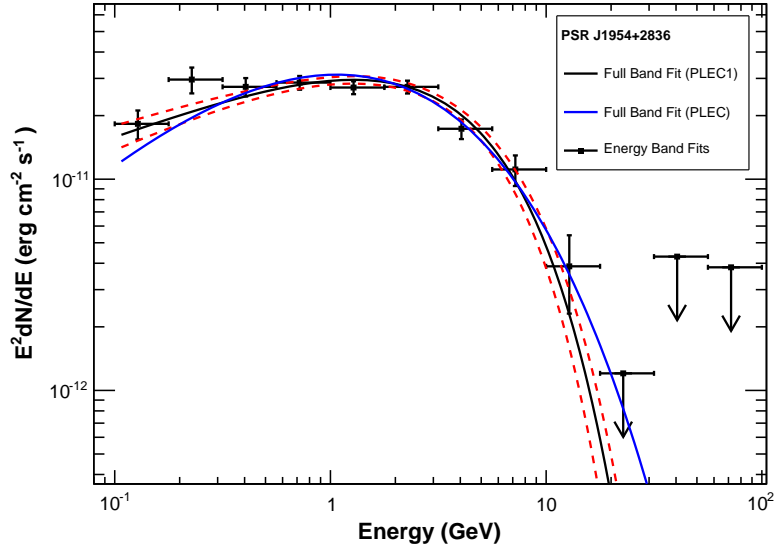


Fig. A-17.— Gamma-ray spectrum of PSR J1954+2835. Appendix A text describes the figure components. This spectrum has a photon index on the soft (larger) end of the Γ distribution. For this pulsar, $TS_{\text{bfree}} \geq 9$ and the PLEC fit (exponential cutoff parameter not fixed to $b = 1$) is shown (blue in the online version).

B.1. Detailed column descriptions of Main catalog FITS tables

The main catalog file, `2PC_catalog_v###.fits` contains summary results for all 117 pulsars. The file contains four FITS table extensions: `PULSAR_CATALOG`, `SPECTRAL`, `OFF_PEAK`, and `REFERENCES`. The primary extension is empty. Table 19 details the contents of the `PULSAR_CATALOG` extension. This file is also duplicated in ascii format (`2PC_catalog_v###.asc`) with the four tables appended in the order listed above.

The `PULSAR_CATALOG` extension contains one row for each pulsar with most of the information provided in the Tables in this paper, as well as some additional quantities that can be computed from these results. Exceptions are the complete spectral analysis (reported in the `SPECTRAL` extension, Table 20), and the results of the off-peak analysis (reported in the `OFF_PEAK` extension, Table 21).

Table 18. LAT 2nd Pulsar Catalog Auxiliary Files Description

Description	Filename(s) and path	Description of file contents
Main catalog FITS table	2PC_catalog.v###.fits	This file contains summary results for all 117 pulsars. The file contains four FITS table extensions: PULSAR_CATALOG, SPECTRAL, OFF_PEAK, and REFERENCES. The primary extension is empty. Section B.1 provides descriptions of these extensions.
Main catalog ascii file	2PC_catalog.v###.asc	This file contains the same summary results for all 117 pulsars as in the previous file, but in ascii format. The four tables are appended in the ascii file in the order listed above.
Light curve images	/images/lightcurves/PNG & /images/lightcurves/PDF	These directories contain plots in both .pdf and .png format for each pulsar of the 100 MeV - 100 GeV and best-fit gamma-ray light curves, light curves in the 100 - 300 MeV, 300 MeV - 1 GeV, 1 - 3 GeV, greater than 3 GeV, and greater than 10 GeV bands, as well as the radio profile for the pulsar, when one exists.
SED images	/images/SED/PNG & /images/SED/PDF	These directories contain plots in both .pdf and .png format for each pulsar of the band-by-band energy fits plus the best-fit spectrum using the <code>PLSuperExpCutoff</code> model (a power law with exponential cutoff model where the exponential index is fixed at a value of 1).
Pulsar timing files	/par_files	This directory contains text files with the ephemerides and other timing parameters used in this work. The files are in TEMPO2 format.
Individual pulsar data FITS files	/PSR_data/FITS	This directory contains FITS format tables for each pulsar with the spectral, light curve, and model data reported in this work. Each data file contains a number of FITS table extensions: PULSAR_SED, MODEL_SED, GAMMA_LC, BEST_FIT_LC, and RADIO_PROFILE (where a radio profile exists). The primary extension is empty. Detailed descriptions of these extensions are provided in Section B.2
Individual pulsar data ascii files	/PSR_data/ascii	This directory contains the same results for each pulsar as in the previous directory, but in ascii format. The relevant extensions are appended in the ascii file in the order listed above.

Table 19. LAT 2nd Pulsar Catalog FITS format: PULSAR_CATALOG Extension

Name	Units	Description
PSR_Name	...	Pulsar name
RAJ2000, DEJ2000	deg	The pulsar position in celestial coordinates (J2000).
GLON, GLAT	deg	The pulsar position in galactic coordinates.
Period	ms	The pulsar rotation period.
P_Dot	s s ⁻¹	The period first derivative, uncorrected for Shklovskii effect or Galactic acceleration, see Section 4.3.
E.Dot	erg s ⁻¹	The pulsar spin-down luminosity, uncorrected for Shklovskii effect or Galactic acceleration, see Section 4.3.
F100, Unc_F100	ph cm ⁻² s ⁻¹	The best-fit photon flux and statistical error, integrated from 100 MeV to 100 GeV. NULL values indicate unreliable spectral fits.
G100, Unc_G100	erg cm ⁻² s ⁻¹	The best-fit energy flux and statistical error, integrated from 100 MeV to 100 GeV. NULL values indicate unreliable spectral fits.
TS_DC	...	The test statistic obtained at the position of the pulsar, assuming a <code>PLSuperExpCutoff</code> spectral model with the exponential index fixed to 1. The fit uses data from 100 MeV to 100 GeV, and includes all pulse phases except as noted in Tables 9 and 10.
TS_Cutoff	...	The significance of the spectral cutoff, obtained from the improvement in log(Likelihood) from the <code>PLSuperExpCutoff</code> spectral model fit over the <code>PowerLaw</code> spectral fit.
TS.bfree	...	The improvement in the test statistic when the exponential index is left free in the <code>PLSuperExpCutoff</code> spectral fit. If there is no improvement, or the fit is worse, this value is zero.
Photon_Index, Unc_Photon_Index	...	The best-fit photon index and statistical error for the <code>PLSuperExpCutoff</code> spectral model. NULL values indicate unreliable spectral fits.
Cutoff, Unc_Cutoff	MeV	The best-fit cutoff energy and the statistical error for the <code>PLSuperExpCutoff</code> spectral model. NULL values indicate unreliable spectral fits.
Distance, Neg_Unc_Distance, Pos_Unc_Distance	pc	The pulsar's distance measurement and its uncertainties. NULL values indicate that only an upper limit has been determined.
UL_Distance	pc	Upper limit on pulsar distance when no measured value has been determined, NULL values indicate that a distance measurement has been determined.
Distance_Method,	...	The method used to determine the pulsar's distance. Methods are: <code>K</code> for the kinematic model, <code>DM</code> for the dispersion measure using the NE2001 model of Cordes & Lazio (2002), <code>O</code> for optical measurements, and <code>X</code> for X-ray measurements. <code>DMM</code> means that the distance to the Galaxy's edge, as determined by the maximum DM value provided by the NE2001 model for that line of sight, is taken as an upper limit.
Distance_Ref	...	Numerical reference for the distance measurement. The full reference is in the REFERENCES extension of this FITS file.
Prop_Motion, Unc_Prop_Motion	mas yr ⁻¹	The proper motion and errors for the pulsar when available.
Prop_Motion_Ref	...	Numerical reference for the proper motion measurement. The full reference is in the REFERENCES extension of this FITS file.
P_Dot_Int, Neg_Unc_P_Dot_Int Pos_Unc_P_Dot_Int	s s ⁻¹	The intrinsic P-dot and associated errors, after contributions from the Shklovskii effect and Galactic acceleration have been removed, see Section 4.3.
E.Dot_Int, Neg_Unc_E.Dot_Int, Pos_Unc_E.Dot_Int	erg s ⁻¹	The intrinsic spin down power and associated errors.

Table 19—Continued

Name	Units	Description
Luminosity, Unc_Luminosity, Neg_Sys_Luminosity, Pos_Sys_Luminosity	erg s ⁻¹	The pulsar luminosity, statistical error, and systematic errors. Systematic errors are derived from the distance uncertainty. Values are NULL when only an upper limit exists.
UL_Luminosity	erg s ⁻¹	Upper limit on the luminosity when no value has been determined. Entries are NULL when a value has been determined. NULL values in all Luminosity columns indicate unreliable spectral fits.
Efficiency, Unc_Efficiency, Neg_Sys_Efficiency, Pos_Sys_Efficiency	percent	The pulsar efficiency, statistical error, and systematic errors from the distance measurement. Values are NULL when only an upper limit has been determined.
UL_Efficiency	percent	Upper limit on the pulsar efficiency. Entries are NULL when a value has been determined. NULL values in all Efficiency columns indicate unreliable spectral fits.
S1400	mJy	Radio flux density at 1400 MHz. In some cases, documented in Section 4.1, this value is extrapolated from measurements at other frequencies. Entries are NULL when only an upper limit has been reported.
UL_S1400	...	Upper limit on the radio flux density at 1400 MHz when no measurement has been reported. Entries are NULL when a value has been reported.
S1400_Ref	...	Numerical reference for the radio flux density measurement. The full reference is in the REFERENCES extension of this FITS file.
Num_Peaks	...	Number of peaks in the gamma-ray profile.
Shift_Method	...	Method used to choose the radio fiducial phase. Methods are: p for the peak radio intensity, h for an opposite hemisphere shift (0.5 phase shift from the peak intensity), s for the point of symmetry in the radio profile, and o for some other method as described in the text (used only for PSR J0534+2200).
Radio_Lag, Unc_Radio_Lag	...	Phase separation (δ) between the first gamma-ray peak and the radio peak, and the associated error on that separation.
Peak_Sep, Unc_Peak_Sep	...	Phase separation (Δ) between the first and last gamma-ray peaks, and the associated error on that separation. This value is NULL for pulsars with only a single gamma-ray peak.
HWHM_P1_L, Unc_HWHM_P1_L	...	Half-width half-max and corresponding uncertainty of the leading (left) first peak edge, as fitted. The best-fit light curve is in the BEST_FIT_LC extension in the individual FITS file for each pulsar.
HWHM_P1_R, Unc_HWHM_P1_R	...	Half-width half-max and corresponding uncertainty of the trailing (right) first peak edge, as fitted. The best-fit light curve is in the BEST_FIT_LC extension in the individual FITS file for each pulsar.
HWHM_P2_L, Unc_HWHM_P2_L	...	Half-width half-max and corresponding uncertainty of the leading (left) second peak edge, as fitted. The best-fit light curve is in the BEST_FIT_LC extension in the individual FITS file for each pulsar.
HWHM_P2_R, Unc_HWHM_P2_R	...	Half-width half-max and corresponding uncertainty of the trailing (right) second peak edge, as fitted. The best-fit light curve is in the BEST_FIT_LC extension in the individual FITS file for each pulsar.

Table 19—Continued

Name	Units	Description
H_ColDensity, Neg_Unc_H_ColDensity, Pos_Unc_H_ColDensity UL_H_ColDensity	cm^{-2} cm^{-2}	Hydrogen column density and associated systematic errors from the distance measurement. The values are NULL when only an upper limit for the hydrogen column density exists. Upper limit on the hydrogen column density. Entries are NULL when a value has been reported.
XFlux_NonTherm, Unc_XFlux_NonTherm	$\text{erg cm}^{-2}\text{s}^{-1}$	Non-thermal unabsorbed X-ray energy flux and 90% CL statistical errors, in the 0.3-10 keV energy band. Spectrum is an absorbed power law, plus black body model when needed. Exceptions are PSRs J0633+1746 and J0659+1414 where a double black body plus power law model was used. Entries are NULL when only an upper limit has been reported.
UL_XFlux_NonTherm	$\text{erg cm}^{-2}\text{s}^{-1}$	Non-thermal X-ray energy flux upper limit. Entries are NULL when a value has been reported.
XFlux_PWN, Unc_XFlux_PWN	$\text{erg cm}^{-2}\text{s}^{-1}$	Estimated non-thermal X-ray flux and 90% CL statistical errors, from the brightest part of the associated plerion, in the 0.3-10 keV energy band.
X_Qual	...	Quality of X-ray detections: ‘0’ indicates no confirmed counterpart, ‘1’ indicates that a counterpart has been identified but with too few counts for further characterization, and ‘2’ indicates that a counterpart has been identified with sufficient counts for spectral characterization.
Opt_Mag	...	Optical magnitude of the optical counterpart for the pulsar or pulsar system, where a counterpart is detected. NULL if no observation available.
LL_Opt_Mag	...	Y indicates that Opt_Mag is a lower limit on the optical magnitude.
Opt_Band	...	The filter used for the optical observation.
Opt_Object	...	Object to which the measured optical flux pertains. The codes are B for binary system; U for upper limit; P = neutron star detected; P* = pulsed optical detection; P+ = pulsar candidate (possible unpulsed pulsar detection); C = companion detected ; N = nebula (PWN) detected.
Extinction, Neg_Unc_Extinction, Pos_Unc_Extinction UL_Extinction	Optical extinction and associated errors derived from the hydrogen column density and using the relation of Fitzpatrick (1999). Entries are NULL when only an upper limit is reported. Optical extinction upper limit when no value has been reported. Entries are NULL when a value has been reported.
Corr_OptFlux, Neg_Unc_OptFlux, Pos_Unc_OptFlux	$\text{erg cm}^{-2}\text{s}^{-1}$	Corrected (unabsorbed) optical energy flux in the V-band, and associated errors. The optical flux has been corrected for interstellar reddening, and scaled to the V-band (peak wavelength $\lambda = 5500\text{\AA}$, bandwidth $\Delta\lambda = 890\text{\AA}$) where necessary. Entries are NULL when only an upper limit has been reported.
UL_OptFlux	...	Upper limit on the corrected (unabsorbed) optical energy flux in the V-band when no measurement has been reported. Entries are NULL when a value has been reported.
Type	...	Indicates whether the pulsar is a young radio loud (YRL), young radio quiet (YRQ, $S_{1400} > 30 \mu\text{Jy}$), or millisecond (MSP) pulsar.
Binary	...	Y indicates the pulsar is in a binary system.
History	...	Indicates whether the pulsar was discovered in radio (Radio), X-rays (X-rays), or gamma rays (Gamma).

The SPECTRAL extension (Table 20) contains the results of the spectral analysis, one row for each pulsar. Models used in the spectral analysis are PLEC1, PLEC, and PL. The spectral analysis is described in Section 6.1.

The Prefactor, Scale, Photon_Index, and Cutoff values for each pulsar using the PLEC1 model are provided. Results from the fit using the PLEC spectral model are only reported when $TS_{b_{free}} \geq 9$. The differential spectrum of the PLEC spectral model (PLSuperExpCutoff) is defined as:

$$\frac{dN}{dE} = \text{Prefactor} \left(\frac{E}{\text{Scale}} \right)^{-\text{Photon_Index}} \exp \left\{ - \left(\frac{E}{\text{Cutoff}} \right)^{\text{Exponential_Index}} \right\}. \quad (\text{B1})$$

while the differential spectrum of the PLEC1 model is Eq. B1 with Exponential_Index fixed to 1. When the PLEC fit is reported in the SPECTRAL extension, it includes the value for the Exponential_Index

The differential spectrum of the PL (PowerLaw) spectral model is defined as:

$$\frac{dN}{dE} = \text{Prefactor} \left(\frac{E}{\text{Scale}} \right)^{-\text{Photon_Index}} \quad (\text{B2})$$

with the Prefactor, Scale, and Photon_Index for each pulsar using the PL model given in the SPECTRAL extension.

Table 20. LAT 2nd Pulsar Catalog FITS format: SPECTRAL Extension

Name	Units	Description
PSR_Name	...	Pulsar name
On_Peak	...	Y indicates the spectral fit used only on-peak events.
TS_DC	...	The test statistic obtained at the position of the pulsar, assuming a <code>PLSuperExpCutoff</code> spectral model with the exponential index fixed to 1. The fit uses data from from 100 MeV to 100 GeV, and includes all pulse phases.
TS_Cutoff	...	The significance of the spectral cutoff, obtained from the improvement in $\log(\text{Likelihood})$ from the <code>PLSuperExpCutoff</code> spectral model fit over the <code>PowerLaw</code> spectral fit.
TS_bfree	...	The improvement in the test statistic when the photon index is left free in the <code>PLSuperExpCutoff</code> spectral fit. If there is no improvement, or the fit is worse, this value is zero.
PLEC1_Prefactor, Unc_ECPL1_Prefactor	$\text{ph cm}^{-2}\text{s}^{-1}\text{MeV}^{-1}$	The best-fit prefactor and associated error for the spectral fit using a power law with exponential cutoff model where the exponential index is fixed at a value of 1.
PLEC1_Photon_Index, Unc_ECPL1_Photon_Index	...	The best-fit photon index and associated error for the spectral fit using a <code>PLSuperExpCutoff</code> model where the exponential index is fixed at a value of 1.
PLEC1_Scale	MeV	The scaling energy for the spectral fit using a <code>PLSuperExpCutoff</code> model where the exponential index is fixed at a value of 1.
PLEC1_Cutoff, Unc_PLEC1_Cutoff	MeV	The best-fit cutoff energy and associated error for the spectral fit using a <code>PLSuperExpCutoff</code> model where the exponential index is fixed at a value of 1.
PLEC1_Flux, Unc_PLEC1_Flux	$\text{ph cm}^{-2}\text{s}^{-1}$	The photon flux integrated from 100 MeV to 100 GeV and associated error for the spectral fit using a <code>PLSuperExpCutoff</code> model where the exponential index is fixed at a value of 1.
PLEC1_EFlux, Unc_PLEC1_EFlux	$\text{erg cm}^{-2}\text{s}^{-1}$	The energy flux integrated from 100 MeV to 100 GeV and associated error for the spectral fit using a <code>PLSuperExpCutoff</code> model where the exponential index is fixed at a value of 1.
PLEC_Prefactor, Unc_PLEC_Prefactor	$\text{ph cm}^{-2}\text{s}^{-1}\text{MeV}^{-1}$	The best-fit prefactor and associated error for the spectral fit using a <code>PLSuperExpCutoff</code> model where the exponential index is left free.
PLEC_Photon_Index, Unc_PLEC_Photon_Index	...	The best-fit photon index and associated error for the spectral fit using a <code>PLSuperExpCutoff</code> model where the exponential index is left free.
PLEC_Scale	MeV	The scaling energy for the spectral fit using a <code>PLSuperExpCutoff</code> model where the exponential index is left free.
PLEC_Cutoff Unc_PLEC_Cutoff	MeV	The best-fit cutoff energy and associated error for the spectral fit using a <code>PLSuperExpCutoff</code> model where the exponential index is left free.
PLEC_Exponential_Index, Unc_PLEC_Exponential_Index	...	The best-fit value and associated error for the spectral fit using a <code>PLSuperExpCutoff</code> model where the exponential index is left free.
PLEC_Flux, Unc_PLEC_Flux	$\text{ph cm}^{-2}\text{s}^{-1}$	The photon flux integrated from 100 MeV to 100 GeV and associated error for the spectral fit using a <code>PLSuperExpCutoff</code> model where the exponential index is left free.
PLEC_EFlux, Unc_PLEC_EFlux	$\text{erg cm}^{-2}\text{s}^{-1}$	The energy flux integrated from 100 MeV to 100 GeV and associated error for the spectral fit using a <code>PLSuperExpCutoff</code> model where the exponential index is left free.
PL_Prefactor, Unc_PL_Prefactor	$\text{ph cm}^{-2}\text{s}^{-1}\text{MeV}^{-1}$	The best-fit prefactor and associated error for the spectral fit using a <code>PowerLaw</code> model.
PL_Photon_Index, Unc_PL_Photon_Index	...	The best-fit photon index and associated error for the spectral fit using a <code>PowerLaw</code> model.
PL_Scale	MeV	The scaling energy for the spectral fit using a <code>PowerLaw</code> model.
PL_Flux, Unc_PL_Flux	$\text{ph cm}^{-2}\text{s}^{-1}$	The photon flux integrated from 100 MeV to 100 GeV and associated error for the spectral fit using a <code>PowerLaw</code> model.
PL_EFlux, Unc_PL_EFlux	$\text{erg cm}^{-2}\text{s}^{-1}$	The energy flux integrated from 100 MeV to 100 GeV and associated error for the spectral fit using a <code>PowerLaw</code> model.

The OFF_PEAK extension (Table 21) contains the spatial and spectral results of the search for off-peak emission. The table contains one row for each pulsar. Details of this analysis are given in Section 7.2.

For sources reported with a PL spectral model, the differential spectrum is defined as:

$$\frac{dN}{dE} = \text{Prefactor_OP} \left(\frac{E}{\text{Scale_OP}} \right)^{-\text{Index_OP}}. \quad (\text{B3})$$

The Prefactor_OP, Index_OP, and Scale_OP are given in the OFF_PEAK extension.

For sources reported with a PLEC1 spectral model (PLSuperExpCutoff), the differential spectrum is defined as:

$$\frac{dN}{dE} = \text{Prefactor_OP} \left(\frac{E}{\text{Scale_OP}} \right)^{-\text{Index_OP}} \exp \left(-\frac{E}{\text{Energy_Cutoff_OP}} \right) \quad (\text{B4})$$

and the Prefactor_OP, Index_OP, Scale_OP, and Energy_Cutoff_OP are given in the OFF_PEAK extension as described below.

For the Crab Nebula and Vela-X, we took the spectral shape and initial normalization from Buehler et al. (2012) and Grondin et al. (2013), respectively, and fit only a multiplicative offset (see Section 7.2). For these two sources, the differential spectrum was defined as:

$$\frac{dN}{dE} = \text{Normalization_OP} \left. \frac{dN}{dE} \right|_{\text{file}} \quad (\text{B5})$$

and Normalization_OP is provided in the OFF_PEAK extension of the main pulsar catalog FITS file.

References used for pulsar distances and radio flux values have been assigned a number, and the number is reported in the PULSAR_CATALOG extension. The REFERENCES extension (Table 22) provides the full information for each reference.

Table 21. LAT 2nd Pulsar Catalog FITS format: OFF_PEAK Extension

Name	Units	Description
PSR_Name	...	Pulsar name
Classification_OP	...	Off-peak emission class: M for magnetospheric (“pulsar-like”), W for possible PWN emission, and U for Unidentified. L is for sources with no significant off-peak emission.
Min_Phase_OP, Max_Phase_OP	...	The minimum and maximum phase that defines the off-peak interval.
Min_2_Phase_OP, Max_2_Phase_OP	...	For pulsars with two off-peak phase ranges, the minimum and maximum phase for the second off-peak interval.
TS_point_OP	...	The test statistic obtained at the best-fit position of the assumed point-like source. <i>TS</i> is computed at the best-fit position assuming a power-law spectral model (except for PSR J0534+2200 as is described in Section 7.2).
TS_ext_OP	...	The significance of any possible extension for the source.
TS.cutoff_OP	...	The significance of any spectral cutoff for a source detected in the off-peak region. (Computed at the pulsar’s position)
TS_var_OP	...	The significance of variability in the off-pulse emission.
Spectral_Model_OP	...	For regions with a significant detection, this is the best spectral model selected by our analysis procedure described in Section 7.2. The possible spectral models are PowerLaw , PLSuperExpCutoff , and FileFunction and are consistent with naming convention in gtlike .
Flux_OP, Unc_Flux_OP	ph cm ⁻² s ⁻¹	The best-fit photon flux and estimated statistical error. The flux is integrated from 100 MeV to 316 GeV.
EFlux_OP, Unc_EFlux_OP	erg cm ⁻² s ⁻¹	The best-fit energy flux and estimated statistical error. The flux is integrated from 100 MeV to 316 GeV.
Prefactor_OP, Unc_Prefactor_OP	ph cm ⁻² s ⁻¹ MeV ⁻¹	The best-fit prefactor and estimated statistical error for the PowerLaw and PLSuperExpCutoff spectral models. The prefactor is defined in Eq. B3 and Eq. B4 for the two spectral models.
Normalization_OP, Unc_Normalization_OP	...	The best-fit normalization and estimated statistical error for FileFunction spectral models. The normalization is defined in Eq. B5. This spectral model was only used for the Crab Nebula and Vela-X.
Scale_OP	MeV	The scaling energy for the PowerLaw and PLSuperExpCutoff spectral models. The scale is defined in Eq. B3 and Eq. B4 for the two spectral models.
Index_OP, Unc_Index_OP	...	The best-fit photon index and estimated statistical error for the PowerLaw and PLSuperExpCutoff spectral models. The photon index is defined in Eq. B3 and Eq. B4 for the two spectral models.
Energy_Cutoff_OP, Unc_Energy_Cutoff_OP	MeV	The best-fit cutoff energy and the estimated statistical error for the PLSuperExpCutoff spectral model. It is defined in Eq. B4.
Spatial_Model_OP	...	For off-peak regions with a significant detection, the spatial model selected by our analysis procedure described in Section 7.2. The choices are At_Pulsar , Point , and Extended .
RAJ2000_OP, DEJ2000_OP	deg	The position of the source in celestial coordinates. For upper limits and sources with a best-fit spatial model at the pulsar position, this is the pulsar’s position. For sources where the localized position is the selected spatial model, this is the best-fit position. For spatially-extended sources, this is the center of the best-fit extended source spatial model.
GLON_OP, GLAT_OP	deg	This is the same as RA_J2000 and DEC_J2000, but in Galactic coordinates.
Unc_Position_OP	deg	For sources with a Point spatial model, the estimated statistical error on the localization of the source. For sources with an Extended spatial model, the estimated statistical error on the center of the extended source.
Extension_OP, Unc_Extension_OP	deg	For sources with an Extended spatial model, the best fit extension and estimated statistical error.
PowerLaw_Flux_UL_OP	ph cm ⁻² s ⁻¹	For regions with no significant detection, this is the 95% confidence-level photon flux upper limit computed assuming a PowerLaw spectral model with Index = 2 and integrated from 100 MeV to 316 GeV.

Table 21—Continued

Name	Units	Description
PowerLaw_EFlux_UL_OP	$\text{erg cm}^{-2}\text{s}^{-1}$	The same as <code>PowerLaw_Flux_UL</code> , but instead the energy flux integrated from 100 MeV to 316 GeV.
Cutoff_Flux_UL_OP	$\text{ph cm}^{-2}\text{s}^{-1}$	For regions with no significant detection, the 95% confidence-level photon flux upper limit assuming a <code>PLSuperExpCutoff</code> spectral model with a canonical pulsar spectrum of <code>Index = 1.7</code> and <code>Energy_Cutoff = 3</code> . This is the flux upper limit integrated from 100 MeV to 316 GeV.
Cutoff_EFlux_UL_OP	$\text{erg cm}^{-2}\text{s}^{-1}$	The same as <code>Cutoff_Flux_UL</code> , but instead the energy flux integrated from 100 MeV to 316 GeV.
SED_Lower_Energy_OP, SED_Upper_Energy_OP, SED_Center_Energy_OP	MeV	For each region, we computed a Spectral Energy Distribution (SED) for the source in 14 energy bins spaced uniformly from 100 MeV to 316 GeV (4 bins per energy decade). Therefore, each <code>SED_*</code> column corresponds to a vector of 14 values, one for each energy bin. <code>SED_Lower_Energy</code> , <code>SED_Upper_Energy</code> , and <code>SED_Middle_Energy</code> are the lower energy, upper energy, and energy in the geometric mean of the energy bin for each SED point.
SED_TS_OP	...	The test statistic obtained for each SED point.
SED_Prefactor_OP, Neg_Unc_SED_Prefactor_OP, Pos_Unc_SED_Prefactor_OP, SED_Prefactor_UL_OP	$\text{ph cm}^{-2}\text{s}^{-1}\text{MeV}^{-1}$	The best-fit prefactor, asymmetric lower and upper error, and 95% confidence-level upper limit computed for the source in each energy bin. When $TS \geq 25$, a detection is quoted when <code>SED_TS > 4</code> and an upper limit is quoted otherwise. When $TS < 25$, all SED points are quoted as upper limits.

Table 22. LAT 2nd Pulsar Catalog FITS format: REFERENCES Extension

Name	Units	Description
Ref_Number	...	Numerical value of the reference from <code>Distance_Ref</code> and <code>S1400_Ref</code> columns.
Citation	...	Citation for each reference.
ADS_URL	...	URL for the reference at the Astrophysical Data Service (ADS). This webpage provides links to the original publishing journal of the referenced paper, article, or catalog.
Title	...	Title of the reference.

B.2. Individual pulsar FITS files

In addition to the summary information for each pulsar contained in the main catalog file, detailed results of the analyses are provided in the individual pulsar FITS files. Each file contains a variable number of FITS table extensions: PULSAR_SED, MODEL_SED, GAMMA_LC, BEST_FIT_LC, and RADIO_PROFILE (for radio detected pulsars). Again, the primary extension is empty. These files are also provided in ascii format with the content of the FITS tables appended in the order listed above.

The PULSAR_SED extension (Table 23) contains the photon and energy fluxes for each pulsar in either six or twelve energy bins, fitting the pulsar with a power-law spectral form. These points correspond to the black data points in the pulsar SED image files. The number of energy bins used in the SED varies with the significance of the pulsar. In a few cases, the pulsar is too faint to construct an SED or there were problems with the spectral fit, and this extension is not included.

The MODEL_SED extension (Table 24) contains the model photon flux and bowtie uncertainty using the PLEC1 spectral form fitted over the full energy range. A description of the spectral fitting method is provided in Section 6.1. These points correspond to the red curves in the pulsar SED image files. In cases where the pulsar is too faint to construct an SED, this extension is not reported.

The GAMMA_LC extension (Table 25) contains weighted counts and the corresponding uncertainties for light curves in six different energy ranges. The number of points in each light curve varies with the significance of the pulsar. These points correspond to the light curves shown in black in the pulsar light curve image files. The values for the background shown in those images are provided as keywords in the header of this FITS extension.

The BEST_FIT_LC extension (Table 26) reports the fitted light curve that best represents the data, as described in Section 5. These points correspond to the blue curves shown in the pulsar light curve image files.

The RADIO_PROFILE extension (Table 27) reports the radio profile for the radio loud pulsars. These points correspond to the red curves shown in the pulsar light curve image files. This extension is not included for pulsars undetected in radio.

Table 23. LAT 2nd Individual Pulsar FITS file format: PULSAR_SED Extension

Name	Units	Description
Energy_Min, Energy_Max	GeV	Lower and upper bounds for each SED bin.
Center_Energy	GeV	Central energy for each SED bin.
PhotonFlux	ph cm ⁻² s ⁻¹	Photon flux in bin
Unc_PhotonFlux	ph cm ⁻² s ⁻¹	Best-fit value and associated error for the photon flux in each SED bin using a power law spectral model. The error is set to zero when the given photon flux is an upper limit.
EnergyFlux, Unc_EnergyFlux	erg cm ⁻² s ⁻¹	Best-fit value and associated error for the energy flux in each SED bin using a power law spectral model. The error is set to zero when the given energy flux is an upper limit.

Table 24. LAT 2nd Individual Pulsar FITS file format: MODEL_SED Extension

Name	Units	Description
Energy_Min, Energy_Max	GeV	Lower and upper bounds for each SED bin.
Center_Energy	GeV	Central energy for each SED bin.
Model_PhotonFlux	ph cm ⁻² s ⁻¹	Integrated photon flux in each bin calculated from the <code>PLSuperExpCutoff</code> model with the exponential index fixed at a value of 1 that has been fitted over the full energy range (from 100 MeV to 100 GeV).
Bowtie_Flux	ph cm ⁻² s ⁻¹	One-sigma uncertainty on the <code>Model_PhotonFlux</code> used to construct the bowtie on the spectral plots.

Table 25. LAT 2nd Individual Pulsar FITS file format: GAMMA_LC Extension

Name	Units	Description
Phase_Min, Phase_Max	...	Lower and upper bounds for each bin in the gamma-ray light curve.
GT100_WtCounts, Unc_GT100_WtCounts	...	Weighted counts and associated error in each phase bin for the gamma-ray light curve using an energy range of 100 MeV to 100 GeV.
GT3000_WtCounts, Unc_GT3000_WtCounts	...	Weighted counts and associated error in each phase bin for the gamma-ray light curve using an energy range of 3 GeV to 100 GeV.
GT10000_WtCounts, Unc_GT10000_WtCounts	...	Weighted counts and associated error in each phase bin for the gamma-ray light curve using an energy range of 10 GeV to 100 GeV.
100_300_WtCounts, Unc_100_300_WtCounts	...	Weighted counts and associated error in each phase bin for the gamma-ray light curve using an energy range of 100 MeV to 300 MeV.
300_1000_WtCounts, Unc_300_1000_WtCounts	...	Weighted counts and associated error in each phase bin for the gamma-ray light curve using an energy range of 300 MeV to 1 GeV.
1000_3000_WtCounts, Unc_1000_3000_WtCounts	...	Weighted counts and associated error in each phase bin for the gamma-ray light curve using an energy range of 1 to 3 GeV.

Table 26. LAT 2nd Individual Pulsar FITS file format: BEST_FIT_LC Extension

Name	Units	Description
Phase_Min, Phase_Max	...	Lower and upper bounds for each bin in the best fit gamma-ray light curve.
Norm_Intensity	...	Normalized gamma-ray intensity for each bin in the best fit gamma-ray light curve. The intensity is normalized so that the integral of the profile is ~ 1 (i.e. normalized as a density function).

Table 27. LAT 2nd Individual Pulsar FITS file format: RADIO_PROFILE Extension

Name	Units	Description
Phase_Min, Phase_Max	...	Lower and upper bounds for each bin in the radio light curve.
Norm_Intensity	...	Normalized radio flux for each bin in the radio light curve. The flux is normalized so that the peak flux equals 1.

C. Appendix: Off-Peak Individual Source Discussion

Here we discuss several interesting sources found in the off-peak analysis presented in Section 7.

The off-peak emission from PSR J0007+7303 in the SNR CTA1 was previously studied by Abdo et al. (2012). They found a soft and not-significantly cut off source in the off-peak region that is marginally extended. We find a similar spectrum and extension significance ($TS_{\text{ext}} = 10.8$), and therefore classify this source as type ‘U’.

The new type ‘W’ source is associated with PSR J0205+6449 (Abdo et al. 2009b). The off-peak spectrum for this source is shown in panel b of Figure 13. The emission is best fit as a point source at $(l, b) = (130^\circ.73, 3^\circ.11)$ with a 95% confidence-level radius of $0^\circ.03$. The source has a hard spectrum (power law with $\Gamma = 1.61 \pm 0.21$) and is therefore consistent with a PWN hypothesis. This nebula has been observed at infrared (Slane et al. 2008) and X-ray (Slane et al. 2004) energies. This suggests that we could be observing the inverse Compton emission from the same electrons powering synchrotron emission at lower energies. The PWN hypothesis is supported by the associated pulsar’s very high $\dot{E} = 2.6 \times 10^{36} \text{ erg s}^{-1}$ and relatively young characteristic age, $\tau_c = 5400 \text{ yr}$. This is consistent with the properties of other pulsars with LAT-detected PWN, and we favor a PWN interpretation. We note that the discrepancy between our spectrum and the upper limit quoted in Ackermann et al. (2011) is mainly caused by our expanded energy range and because the flux upper limit was computed assuming a different spectral index.

However, we note that PSR J0205+6449 is associated to the SNR 3C58 (G130.7+3.1). Given the 2 kpc distance estimate from Section 4.2 and the density of thermal material estimated by Slane et al. (2004), we can estimate the energetics required for the LAT emission to originate in the SNR. Following the prescription in Drury et al. (1994), we assume the LAT emission to be hadronic and estimate a cosmic-ray efficiency for the SNR of $\sim 10\%$, which is energetically allowed. We therefore cannot rule out the SNR hypothesis.

No TeV detection of this source has been reported, but given the hard photon index at GeV energies this is a good candidate for observations by an atmospheric Cherenkov telescope. Improved spectral and spatial observations at TeV energies might help to uniquely classify the emission.

We obtain a flux for Vela-X which is $\sim 10\%$ larger than the flux obtained in Grondin et al. (2013). This discrepancy is most-likely due to assuming a different spatial model for the emission (radially-symmetric Gaussian compared to elliptical Gaussian).

PSR J1023–5746 is associated with the TeV PWN HESS J1023–575 (Aharonian et al.

2007). LAT emission from this PWN was first reported in Ackermann et al. (2011). Because of the dominant low-energy magnetospheric emission, we classify this as type ‘M’ and not as a PWN. A phase-averaged analysis of this source for energies above 10 GeV is reported in Acero et al. (2013b).

PSR J1119–6127 (Parent et al. 2011) is associated with the TeV source HESS J1119–614¹⁰. Our off-peak analysis classifies this source as ‘U’ because its spectrum is soft and not significantly cut off. However, the SED appears to represent a cutoff spectrum at low energy and a hard rising spectrum at high energy. Acero et al. (2013b) significantly detect this PWN using the analysis procedure as described for J1023–575. We are likely detecting a composite of magnetospheric emission at low energy and pulsar-wind emission at high energy.

PSR J1357–6429 (Lemoine-Goumard et al. 2011) has an associated PWN HESS J1356–645 detected at TeV energies (Abramowski et al. 2011). Our analysis of the off-peak regions surrounding PSR J1357–6429 shows a source positionally and spectrally consistent with HESS J1356–645, but with significance just below detection threshold ($TS = 21.0$). Acero et al. (2013b) present significant emission from this source.

The off-peak region of PSR J1410–6132 (O’Brien et al. 2008) shows a relatively hard spectral index of 1.90 ± 0.15 , and the spectrum is not significantly cut off. There is no associated TeV PWN and enough low-energy GeV emission is present to caution against a clear PWN interpretation. We classify this source as ‘U’, but further observations could reveal interesting emission.

PSR J2021+4026 is spatially coincident with the LAT-detected and spatially extended Gamma Cygni SNR (Lande et al. 2012). The off-peak emission from this pulsar is consistent with an exponentially-cutoff spectrum and is therefore classified as type ‘M’. The source’s marginal extension ($TS_{\text{ext}} = 8.7$) is likely due to some contamination from the SNR.

¹⁰The discovery of HESS J1119–614 was presented at the “Supernova Remnants and Pulsar Wind Nebulae in the Chandra Era” in 2009. See http://cxc.harvard.edu/cdo/snr09/pres/DjannatiAtai_Arache_v2.pdf.