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## BROADBAND PULSATIONS FROM PSR B1821–24: IMPLICATIONS FOR EMISSION MODELS AND THE PULSAR POPULATION OF M28

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### ABSTRACT

We report a  $5.4\sigma$  detection of pulsed gamma rays from PSR B1821–24 in the globular cluster M28 using  $\sim 44$  months of *Fermi* Large Area Telescope (LAT) data that have been reprocessed with improved instrument calibration constants. We constructed a phase-coherent ephemeris, with post-fit residual rms of  $3 \mu\text{s}$ , using radio data spanning  $\sim 23.2$  yr, enabling measurements of the multi-wavelength light-curve properties of PSR B1821–24 at the millisecond level. We fold *RXTE* observations of PSR B1821–24 from 1996 to 2007 and discuss implications on the emission zones. The gamma-ray light curve consists of two peaks separated by  $0.41 \pm 0.02$  in phase, with the first gamma-ray peak lagging behind the first radio peak by  $0.05 \pm 0.02$  in phase, consistent with the phase of giant radio pulses. We observe significant emission in the off-peak interval of PSR B1821–24 with a best-fit LAT position inconsistent with the core of M28. We do not detect significant gamma-ray pulsations at the spin or orbital periods from any other known pulsar in M28, and we place limits on the number of energetic pulsars in the cluster. The derived gamma-ray efficiency,  $\sim 2\%$ , is typical of other gamma-ray pulsars with comparable spin-down power, suggesting that the measured spin-down rate ( $2.2 \times 10^{36} \text{ erg s}^{-1}$ ) is not appreciably distorted by acceleration in the cluster potential. This confirms PSR B1821–24 as the second very energetic millisecond pulsar in a globular cluster and raises the question of whether these represent a separate class of objects that only form in regions of very high stellar density.

**Key words:** gamma rays: stars – globular clusters: individual (M28) – pulsars: individual (B1821–24)

**Online-only material:** color figures

### 1. INTRODUCTION

Since the launch of the *Fermi* Gamma-ray Space Telescope in 2008, significant high-energy (HE;  $\geq 0.1$  GeV) pulsations have been detected from more than 40<sup>21</sup> millisecond pulsars (MSPs; Abdo et al. 2013, mostly in the Galactic field) using the Large Area Telescope (LAT, a pair-production telescope sensitive to photons with energies from 20 MeV to  $>300$  GeV; Atwood et al. 2009), the main instrument aboard *Fermi*. Additionally, HE emission has been detected from the directions of more than a dozen globular clusters (Abdo et al. 2010d; Kong et al. 2010;

Tam et al. 2011), known or thought to host many MSPs, and the observed spectra of these point sources are consistent with the superposition of emission from several MSPs (predicted by Chen 1991). The one exception is PSR J1823–3021A (in the globular cluster NGC 6624; Biggs et al. 1994), from which significant gamma-ray pulsations have been detected and which accounts for all of the observed HE emission associated with the parent cluster (Freire et al. 2011). To date, all LAT sources associated with globular clusters are consistent with point-like emission, with reported  $2\sigma$  upper limits on any extension of  $<16'$  assuming a two-dimensional Gaussian profile (Abdo et al. 2010d).

MSPs are thought to be old “recycled” pulsars that have reached rapid rotation rates via accretion from a companion

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<sup>21</sup> See <https://confluence.slac.stanford.edu/display/GLAMCOG/Public+List+of+LAT+Detected+Gamma-Ray+Pulsars> for a list of announced gamma-ray pulsars.

star (e.g., Alpar et al. 1982). However, the confirmation of PSR J1823–3021A as a very energetic MSP suggests an unusual formation scenario, such as the collapse of a white dwarf to a neutron star induced by accretion or a merger with another white dwarf (Ivanova et al. 2008), which may be more likely in globular clusters. As such, it is important to search for and/or confirm similar MSPs in other globular clusters.

Detecting gamma-ray pulsations from more MSPs in globular clusters will help to constrain models for the broadband emission from the clusters (e.g., Cheng et al. 2010; Zajczyk et al. 2013; Kopp et al. 2013). Constraining the models will generally determine the expected flux level and may be important for extracting the associated particle conversion efficiency from such modeling once the number of sources is known, thereby constraining the reacceleration particles may undergo within the clusters once they leave the MSP magnetospheres.

## 2. PSR B1821–24

Located within the core of the globular cluster M28 (NGC 6626), PSR B1821–24 is an isolated MSP with a spin period ( $P$ ) of  $\sim 3.05$  ms and was the first pulsar ever detected in a globular cluster (Lyne et al. 1987). The observed period derivative ( $\dot{P}$ ) of  $\sim 1.62 \times 10^{-18}$  s s $^{-1}$  (Foster et al. 1988) leads to an inferred rotational energy-loss rate of  $\dot{E} \equiv 4\pi^2 I \dot{P} / P^3 \sim 2.2 \times 10^{36}$  erg s $^{-1}$ , where  $I$  is the moment of inertia of the neutron star and is taken to be  $10^{45}$  g cm $^2$ . This is the highest  $\dot{E}$  of any known rotation-powered MSP, according to version 1.46 of the ATNF Pulsar Database<sup>22</sup> (Manchester et al. 2005). While it is possible that the  $\dot{P}$  could be artificially enhanced by the gravitational field of the cluster, Phinney (1993) showed that the  $\dot{P}$  is largely intrinsic.

PSR B1821–24 is also the first MSP from which non-thermal pulsed X-ray emission was detected (Saito et al. 1997, using the *Advanced Satellite for Cosmology and Astrophysics*). Rots et al. (1998) used data from the *Rossi X-ray Timing Explorer (RXTE)* and the Green Bank Telescope to determine that the first X-ray and radio peaks were separated by only 0.02 in phase. Using data from the *Chandra X-ray Observatory*, Rutledge et al. (2004) and Bogdanov et al. (2011) found that  $\sim 15\%$  of the non-thermal X-ray flux of PSR B1821–24 was unpulsed.

PSR B1821–24 was the first MSP ever observed to undergo a glitch (Cognard & Backer 2004). A glitch has also been observed from the mildly recycled PSR B1913+16 (Weisberg et al. 2010). Romani & Johnston (2001) and Knight et al. (2006) reported the detection of giant radio pulses of up to 50 and 91 times the mean pulse intensity, respectively, from PSR B1821–24. The giant pulses are concentrated in a narrow phase window coincident with the first X-ray peak, similar to what has been observed in the original MSP, PSR B1937+21 (Cusumano et al. 2003).

Even at a distance ( $d$ ) of  $5.1 \pm 0.5$  kpc (from optical observations of stars in M28; Rees & Cudworth 1991), the relatively large  $\dot{E}$  of PSR B1821–24 makes it a promising candidate for gamma-ray studies. A  $4.2\sigma$  HE pulsed detection was reported by Pellizzoni et al. (2009) using data from the *Astro-rivelatore Gamma a Immagini LEggero (AGILE)* satellite, but pulsations were only significant in the first five days of the observation. The HE pulse profile observed with *AGILE* does not match the LAT profile (see Section 4.2 and Wu et al. 2013) and the observed flux above 100 MeV was greater than the  $3\sigma$  upper limit set using data from the Energetic Gamma-Ray

Experiment Telescope (Fierro et al. 1995). The 2FGL catalog (Nolan et al. 2012) associates 2FGL J1824.8–2449 with M28, and Abdo et al. (2010d) estimated the number of MSPs in the cluster, based on the HE spectrum, to be  $43_{-21}^{+24}$ . Wu et al. (2013) found a  $4.3\sigma$  pulsed detection using  $\sim 42$  months of Pass 7 LAT data (Ackermann et al. 2012) without the updated instrument calibration constants discussed in Section 3.3 and using the timing solution of Ray et al. (2008), which is not contemporaneous with the LAT data, thus leaving open the possibility that the gamma-ray peaks have moved with respect to the radio emission due to timing noise or unmodeled dispersion measure (DM) variations.

## 3. OBSERVATIONS AND DATA ANALYSIS

PSR B1821–24 is timed under the auspices of the LAT Pulsar Timing Consortium (Smith et al. 2008), within which ephemerides are provided from radio observatories around the world for 208 pulsars ranked by  $\sqrt{\dot{E}}/d^2$ . The timing solution described in Section 3.1 will be made available through the *Fermi* Science Support Center.<sup>23</sup>

### 3.1. Radio Timing

The radio timing solution for PSR B1821–24 has been constructed with the TEMPO<sup>24</sup> pulsar timing package (Hobbs et al. 2006) using times of arrival (TOAs) recorded at the Nançay Radio Telescope (NRT) in France, the Westerbork Synthesis Radio Telescope (WSRT) in the Netherlands, and the Lovell Telescope at the Jodrell Bank Observatory in the United Kingdom.

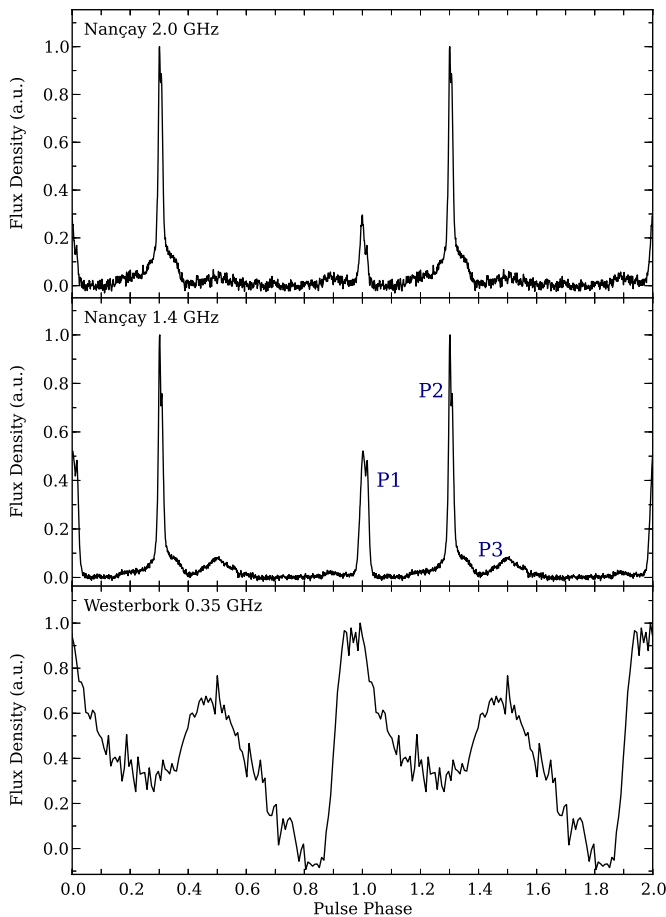
In order to encompass X-ray and *Fermi* LAT observations of PSR B1821–24, we used 2994 TOAs spanning from 1989 October 3 (MJD 47,802) to 2012 December 1 (MJD 56,262). Between 1989 October and 2004 November, Nançay pulsar observations were carried out by mixing the signal with a swept frequency local oscillator mimicking the dispersion caused by the interstellar medium, as described in Cognard et al. (1996), while after late 2004 observations were made using the Berkeley-Orléans-Nançay backend (Cognard & Theureau 2006). Although the bulk of radio observations were conducted at 1.4 GHz, the timing data set also included TOAs recorded at different frequencies from 1.6 to 2 GHz in order to measure and monitor long-term changes in the DM, necessary for comparing profiles at different wavelengths. In addition, a total of 81 1.4 GHz WSRT TOAs recorded with the PuMa and PuMa-II backends (Voûte et al. 2002; Karuppusamy et al. 2008) between 2004 October 10 (MJD 53288) and 2012 September 14 (MJD 56184), as well as 29 1.5 GHz Jodrell Bank TOAs (Hobbs et al. 2004) recorded with the digital filter bank backend between 2009 August 31 (MJD 55074) and 2012 September 13 (MJD 56183), were included.

Figure 1 shows phase-aligned Nançay radio profiles recorded at 1.4 and 2 GHz, based on  $\sim 58.1$  hr of observations made between 2008 July 11 and 2011 February 25 and 40.9 hr of observations made between 2004 December 20 and 2008 May 13, respectively, and a 0.35 GHz Westerbork profile obtained by integrating  $\sim 8.5$  hr of observations conducted between 2013 June 6 and 2013 June 19 with a frequency bandwidth of 0.08 GHz. The relative phase alignment between the 0.35 GHz light curve and the higher frequency radio profiles was estimated

<sup>22</sup> <http://www.atnf.csiro.au/people/pulsar/psrcat/>

<sup>23</sup> <http://fermi.gsfc.nasa.gov/ssc/data/access/lat/ephems/>

<sup>24</sup> <http://fermi.gsfc.nasa.gov/ssc/data/access/lat/ephems/>



**Figure 1.** Phase-aligned radio light curves of PSR B1821–24 with two pulsar rotations shown for clarity. The peak heights in each panel are normalized independently. From top to bottom: 2 GHz Nançay profile, 1.4 GHz Nançay profile, and 0.35 GHz Westerbork profile. We denote the second-highest radio peak at 1.4 GHz, near phase 0, as P1; the highest radio peak at 1.4 GHz, near phase 0.3, as P2; and the lowest radio peak at 1.4 GHz, near phase 0.5, as P3. Both P1 and P3 are also visible at 0.35 and 2 GHz, while P2 has no obvious counterpart at 0.35 GHz.

(A color version of this figure is available in the online journal.)

by extracting four TOAs from the 0.35 GHz Westerbork data and calculating the average offset between the low-frequency and the high-frequency Westerbork TOAs with the ephemeris for PSR B1821–24 obtained from the analysis described below. We estimate that the statistical uncertainty on the relative alignment is on the order of five milliperiods (mP). The few 0.35 GHz TOAs were not included in the TOA data set for the timing analysis, having large uncertainties and being affected by strong scattering from the interstellar medium. For the radio profiles, we use the peak naming convention of Backer & Sallmen (1997), though we shift the first peak to be at phase zero rather than  $\sim 0.3$ . At 0.35 GHz the P2 is not visible, while P1 and P3 appear to broaden and have comparable peak heights.

We first constructed a timing model covering the total TOA data set with good accuracy. At this stage the free parameters were the pulsar position, proper motion, pulse frequency, and the first two time derivatives. The published parameters from the glitch in 2001 March (Cognard & Backer 2004) were included and refit in the timing model. We then fixed the parameters at the best-fit values and used the Nançay timing data set to determine the DM and its variations. The data set was split into seven intervals spanning 2–3 yr of data, over which

**Table 1**  
Observed and Derived Timing Parameters of PSR B1821–24

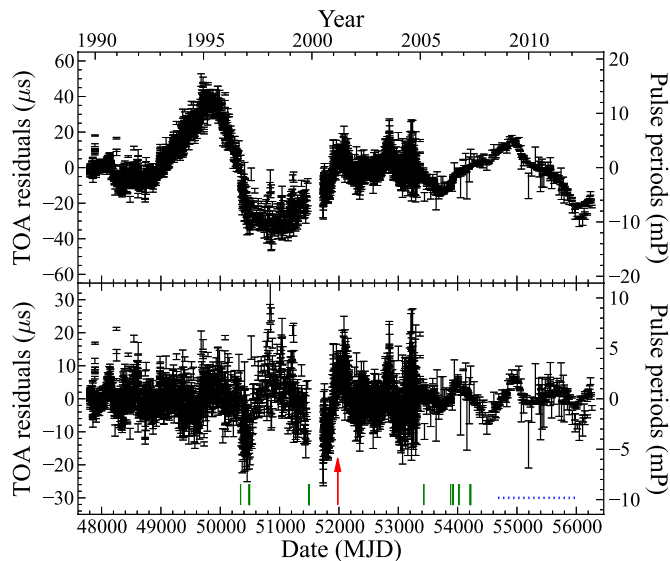
Parameter	Value
R.A., $\alpha$ (J2000)	18:24:32.00819(2)
Decl., $\delta$ (J2000)	–24:52:10.720(5)
Rotational period, $P$ (s)	0.00305431496291
First period derivative, $\dot{P}$ ( $10^{-18}$ s s $^{-1}$ )	1.6187747(1)
Second period derivative, $\ddot{P}$ ( $10^{-31}$ s s $^{-2}$ )	2.481(5)
Proper motion in right ascension, $\mu_{\alpha} \cos(\delta)$ (mas yr $^{-1}$ )	–0.25(2)
Proper motion in declination, $\mu_{\delta}$ (mas yr $^{-1}$ )	–8.5(4)
Epoch of ephemeris (MJD)	50,000
Glitch epoch (MJD)	51,979.5
Glitch frequency step ( $10^{-9}$ s $^{-1}$ )	3.0671(7)
Glitch frequency derivative step ( $10^{-18}$ s $^{-2}$ )	3.156(9)
Epoch of dispersion measure determination (MJD)	52,400
Dispersion measure, DM (cm $^{-3}$ pc)	119.8691(16)
Dispersion measure derivative, DM1 (cm $^{-3}$ pc yr $^{-1}$ )	0.0033(2)
Span of timing data (MJD)	47802–56262
Number of TOAs	3104
rms of TOA residuals ( $\mu$ s)	9.162
Solar system ephemeris model	DE405
Time system	TCB
Total proper motion, $\mu_T$ (mas yr $^{-1}$ )	8.5(4)
Apparent spin-down luminosity, $\dot{E}$ ( $10^{36}$ erg s $^{-1}$ )	2.2
Magnetic field strength at the light cylinder, $B_{LC}$ ( $10^5$ G)	7.2

**Notes.** Measured and derived parameters from the radio observations described in Section 3.1. Numbers in parentheses are the nominal  $1\sigma$  TEMPO2 uncertainties in the least-significant digits quoted.

TOAs were recorded with a single backend and at multiple frequencies. A DM value was obtained for each of these intervals using TEMPO2. A least-squares fit of the seven DM values with a linear function was performed, yielding the values for the DM at MJD 52,400 and first time derivative listed in Table 1. The DM and first time derivative were included in the timing model and frozen at those best-fit values in subsequent analyses. We note that the uncertainty in the DM leads to an uncertainty of  $\sim 1.1$  mP in the conversion of 1.4 GHz TOAs to infinite frequency at the epoch of the ephemeris.

Finally, the timing model was updated by refitting the total TOA data set using the independently determined DM value and its first time derivative while leaving other parameters free. The best-fit parameters obtained from this analysis, displayed in Table 1, give an rms of timing residuals of 9.2  $\mu$ s with a maximum excursion of 17 mP. As can be seen from Figure 2, the TOA residuals exhibit low-frequency structures consistent with rotational irregularities (so-called timing noise; see, e.g., Hobbs et al. 2004, 2010), which we modeled using eight harmonically related sinusoids using the “FITWAVES” option of TEMPO2 and fixing all other timing parameters. After the whitening procedure, the timing residuals had an rms of 3.1  $\mu$ s with a maximal excursion of 10 mP. The whitened timing residuals are displayed in the lower panel of Figure 2. The X-ray and gamma-ray timing analyses presented in Sections 3.2 and 4.2 were carried out with the whitened timing solution, including the FITWAVES parameters.

The observed  $\dot{P}$  of a pulsar can be increased from the true value by contributions from the proper motion (Shklovskii 1970). At a distance of 5.1 kpc and with a total proper motion of 8.5 mas yr $^{-1}$ , this effect contributes  $\sim 2.7 \times 10^{-21}$  s s $^{-1}$  to the measured  $\dot{P}$  of PSR B1821–24, three orders of magnitude less than the value reported in Table 1. Therefore, we do not correct for this effect in the observed and derived parameters



**Figure 2.** Timing residuals as a function of time for the model given in Table 1 (upper panel), and after whitening of the residuals using eight harmonically related sinusoids (lower panel). The arrow (red in the online version) indicates the epoch of the glitch of PSR B1821–24, vertical lines (green in the online version) denote the epochs of the X-ray observations considered in this article, and the dashed horizontal line (blue in the online version) shows the *Fermi* LAT observation interval described in Section 3.3.

(A color version of this figure is available in the online journal.)

of PSR B1821–24. The latest proper-motion measurement for M28 (Casetti-Dinescu et al. 2013) agrees well with our values, with a total difference of  $21 \text{ km s}^{-1}$  at a distance of 5.1 kpc. This difference is less than the estimated escape velocity of  $63.8 \text{ km s}^{-1}$  (Gnedin et al. 2002), suggesting that PSR B1821–24 is, in fact, bound to the cluster.

### 3.2. X-ray Data

The *RXTE* observations we report on here were performed by the Proportional Counter Array (PCA, which consists of five individual proportional counter units, PCUs) from 1996 September 16 (MJD 50,342.261) to 2007 April 26 (MJD 54,216.252), accumulating a total integration time of  $\sim 469$  ks. These observations employed anywhere from one to five PCUs in various combinations during each observation with data recorded using GoodXenon or GoodXenonwithPropane mode. The PCA data were analyzed using the HEASoft version 6.12 data analysis suite. We employed a variety of bit masks<sup>25</sup> to select events from the PCUs in the 3–16 keV range that were on during each individual observation. In addition, Ray et al. (2008) reported that including events from the first and second anode layer improved the signal-to-noise ratio of the pulsed detection, and we followed that prescription here. We did not apply a background correction.

The PCA is not an imaging instrument. Rather, it has a field of view approximately represented by a Gaussian with FWHM of  $14'$  (Jahoda et al. 2006). This means that other X-ray sources known to be in M28 and that have significant flux above 3 keV (e.g., Becker et al. 2003) will contribute to the total count rate in each observation. Because the contribution from these additional sources will add incoherently to the pulsed signal from PSR B1821–24 and we cannot know which events are from PSR B1821–24, we do not attempt to account for these additional

X-ray sources in our analysis or to estimate a resulting background level for the pulsed analysis in Section 4.2.

The events that satisfy our selection criteria were barycentered with the *faxbary* tool using the DE405 solar system ephemeris and including the *RXTE* fine clock corrections, yielding an individual event timing accuracy of  $\sim 6 \mu\text{s}$  (Rots et al. 1998; Jahoda et al. 2006). The proper motion of the pulsar was incorporated into the position used to barycenter the data at each epoch. Pulse phases were calculated utilizing the Photon Events plugin<sup>26</sup> for TEMPO2 and the radio ephemeris described in Section 3.1.

### 3.3. LAT DATA: P7REP

Pass 7 LAT data have been reprocessed<sup>27</sup> using updated calibration constants for the detector subsystems, most importantly for the calorimeter (CAL) to more accurately describe the position-dependent response of each scintillator crystal and the slight decrease in scintillation light yield with time ( $\sim 1\% \text{ yr}^{-1}$ ) from radiation exposure on orbit.

This reprocessing affected the LAT data (P7REP, hereafter) in several ways. First, the point-spread function (PSF) is significantly improved above a few GeV, with a reduction in the 68% containment radius of 30% (40%) for events converting in the front (back) of the tracker (Bregeon et al. 2013). At these energies, the improved calibration constants result in more accurately calculated centroids of energy deposition in the CAL to constrain the incident event direction. Second, the significance of detection and precision of measured photon flux is increased slightly for most sources—more strongly for sources with hard spectra than for those with cutoffs at a few GeV, like pulsars. Third, spectral features such as cutoff energies are shifted upward slightly in energy ( $\sim \text{few } \%$ ) by the change in energy scale.

We selected events from the P7REP data corresponding to the SOURCE class recorded between 2008 August 4 and 2012 March 31 with reconstructed directions within  $11.5^\circ$  of the pulsar radio position, allowing us to construct a  $16^\circ \times 16^\circ$  square region with no blank corners for a binned likelihood analysis (see Section 4.1); energies from 0.1 to 100 GeV, the lower limit that is recommended for analysis of P7REP data and the upper limit that adequately covers the range of known pulsar cutoff energies; and zenith angles  $\leq 100^\circ$ , to reduce contamination of gamma rays from the limb of the Earth. Good time intervals were then selected corresponding to when the instrument was in nominal science operations mode, the rocking angle of the spacecraft did not exceed  $52^\circ$ , the limb of the Earth did not infringe upon the region of interest, and the data were flagged as good. All LAT analyses were performed using the *Fermi* Science Tools v9r27p1.

The recommended instrument response functions (IRFs, which include the PSF, effective area, and energy dispersion) for analyzing P7REP data are P7REP\_V15. These IRFs are derived from detailed simulations of the instrument (Ackermann et al. 2012) with some modifications based on on-orbit performance checks, which are detailed below.

The accuracy with which incoming event directions are reconstructed is dependent on the energy ( $E$ ), interaction point

<sup>25</sup> [http://heasarc.nasa.gov/docs/xte/recipes/cook\\_book.html](http://heasarc.nasa.gov/docs/xte/recipes/cook_book.html)

<sup>26</sup> Written by Anne Archibald, [http://www.physics.mcgill.ca/~aarchiba/photons\\_plugin.html](http://www.physics.mcgill.ca/~aarchiba/photons_plugin.html).

<sup>27</sup> For more information about the updated calibrations and P7REP data, see Bregeon et al. (2013) and [http://fermi.gsfc.nasa.gov/ssc/data/analysis/documentation/Pass7REP\\_usage.html](http://fermi.gsfc.nasa.gov/ssc/data/analysis/documentation/Pass7REP_usage.html).

within the instrument, and angle with respect to the boresight<sup>28</sup> ( $\theta$ ). For a SOURCE class event converting in the front of the instrument, the energy-dependent 68% confidence-level containment radius, averaged over the acceptance, can be approximated as  $\Theta_{68}(E) = \sqrt{(0.66(E/1 \text{ GeV})^{-0.76})^2 + (0.08)^2}$ .

Although the reprocessing significantly improved the PSF at high energies, the angular distribution of gamma rays around point sources used for in-flight calibration of the PSF above 3 GeV was still found to be slightly broader in the P7REP data than predicted by the Monte Carlo (MC) PSF. The on-orbit PSF for the P7REP\_V15 IRFs was derived by rescaling the MC PSF to match the angular distribution of gamma rays around the Vela pulsar below 10 GeV and a sample of bright, high-latitude blazars above 10 GeV. This correction to the MC PSF model rescales the size of the PSF as a function of energy while preserving the dependence on  $\theta$ . Formerly, for the P7\_V6 IRFs recommended for analyzing the original Pass 7 data, the  $\theta$  dependence was not preserved in making this correction (Ackermann et al. 2013).

There is a known discrepancy between the fluxes arising from analyses using only events that convert in the front or the back of the tracker subsystem (see Figure 47 and Section 5.6 of Ackermann et al. 2012). This discrepancy occurs mainly at energies below 300 MeV with differences of  $\lesssim 10\%$ . The P7REP\_V15 effective area tables include an empirical correction for this that does not modify the overall effective area inferred from MC studies. The total effective area for a near on-axis, 1 GeV, SOURCE class gamma ray is  $\sim 7000 \text{ cm}^2$ .

## 4. GAMMA-RAY RESULTS

### 4.1. Spectral and Spatial Analysis

A binned maximum likelihood analysis was performed on a  $16^\circ \times 16^\circ$  region centered on the pulsar position using the P7REP\_SOURCE\_V15 IRFs. All sources from a 3 yr source list produced following the same procedure used for the 2FGL catalog using the original Pass 7 data and P7SOURCE\_V6 IRFs within  $15^\circ$  of PSR B1821–24 were included in the model of the region, and all spectral parameters of sources within  $8^\circ$  (23 point sources and two extended sources) were left free. The Galactic diffuse emission was modeled using the *gll\_iem\_v05.fit* model, while the isotropic diffuse emission and residual instrument background were jointly modeled using the *iso\_source\_v05.txt* template.<sup>29</sup> These diffuse models were produced specifically for the P7REP data using a refined approach in which residuals in the LAT data were used to fit components of the diffuse emission not derived from observations at other wavelengths (see Ballet & Burnett 2013).

We modeled the spectrum of PSR B1821–24 as both a simple power law (Equation (1)) and an exponentially cutoff power law (Equation (2)):

$$\frac{dN}{dE} = N_0 \left( \frac{E}{1 \text{ GeV}} \right)^{-\Gamma}, \quad (1)$$

$$\frac{dN}{dE} = N_0 \left( \frac{E}{1 \text{ GeV}} \right)^{-\Gamma} \exp \left\{ -\frac{E}{E_C} \right\}. \quad (2)$$

Using the likelihood ratio test, a simple power-law shape is ruled out in favor of an exponentially cutoff power law with a confidence level of  $5.6\sigma$ . We detect a point source at the position of PSR B1821–24 with a likelihood test statistic (TS; Nolan et al. 2012) of 438. The best-fit spectrum has  $E_C = 6.1 \pm 2.1 \text{ GeV}$ ,  $\Gamma = 2.2 \pm 0.1$ , and gives integral photon and energy fluxes (from 0.1 to 100 GeV) of  $F = (7.2 \pm 0.9) \times 10^{-8} \text{ cm}^{-2} \text{ s}^{-1}$  and  $G = (3.8 \pm 0.3) \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ , respectively, all uncertainties being statistical. PSR B1821–24 is a relatively faint source for the LAT, and statistical uncertainties in these measurements dominate the systematic uncertainties; therefore, we do not attempt to estimate systematic uncertainties on the best-fit parameters.

Although we kept this point source fixed to the radio timing position of PSR B1821–24 in the spectral analysis, the best-fit LAT position (using the *Fermi* Science Tool *gtfindsrc*) is R.A. (J2000) = 18:24:43.2, decl. (J2000) =  $-24:51:36.0$ , with a 95% confidence-level error radius  $r_{95} = 4'12''$ , which is  $2'24''$  from the timing position.

The 2FGL catalog and Wu et al. (2013) have both reported flux values for point sources associated with PSR B1821–24 using the original Pass 7 data and the P7SOURCE\_V6 IRFs in the 1–100 GeV and 0.2–300 GeV energy ranges, respectively. Integrating our phase-averaged results over the same energy ranges yields higher values than reported by those authors on the order of 20%. These differences are larger than expected from switching to P7REP data alone. We note that the disagreement with the 2FGL flux is at the  $2\sigma$  level and is likely just statistical fluctuation, while the disagreement with Wu et al. (2013) is  $< 1\sigma$ .

We repeated the analysis described in Wu et al. (2013) using similar time, energy, and angular selections and the original Pass 7 data; with the same 2FGL point sources free and fixed in our model of the region; and with the same diffuse components. However, we found values more consistent with results from our analysis described previously. Additionally, our re-analysis only found a TS of 248 for a point source at the position of PSR B1821–24, much less than the value of 825 reported by Wu et al. (2013). We note that Nolan et al. (2012) reported a significance of  $\sim 11\sigma$  for 2FGL J1824.8–2449, using 2 yr of data, which corresponds to a TS of  $\sim 144$ . Extrapolating to  $\sim 42$  months, we expect a TS of  $\sim 200$  for a non-variable source, which agrees with our re-analysis when accounting for differences in event selection. Although the differences in  $\Gamma$  and  $E_C$  may be related to the choice of minimum energy and differences in the diffuse model, the disagreement between the TS values is not understood.

Using the initial phase-averaged results, we were able to detect significant pulsations ( $> 5\sigma$ ; see Section 4.2) from PSR B1821–24; the gamma-ray light curve is characterized by two peaks at phases of  $\sim 0.0$  and  $\sim 0.5$ , similar to the results of Wu et al. (2013). However, there was a clear offset above the estimated background level observed in the gamma-ray light curve. While it is possible that PSR B1821–24 has a near 100% duty cycle (as seems to be the case for PSR J1836+5925; Abdo et al. 2010c), we performed an analysis of the off-peak phase interval (defined as  $\phi \in (0.24, 0.34) \cup (0.58, 0.82)$ ) to study the emission in more detail.

We first attempted to ascertain if this emission could be attributed to any of the other known pulsars in M28<sup>30</sup> (11 MSPs and one young, non-recycled pulsar; Bogdanov et al. 2011;

<sup>28</sup> For more details see [http://www.slac.stanford.edu/exp/glast/groups/canda/lat\\_Performance.htm](http://www.slac.stanford.edu/exp/glast/groups/canda/lat_Performance.htm) and Ackermann et al. (2012).

<sup>29</sup> The P7REP data, P7V15 IRFs, and diffuse models will be available for download at <http://fermi.gsfc.nasa.gov/ssc/data/>.

<sup>30</sup> <http://www.naic.edu/~pfreire/GCpsr.html>

S. Bégin et al., in preparation). At a distance of 5.1 kpc, it is possible that the combined emission from these and any other unknown pulsars, less than that of PSR B1821–24, may account for the observed off-peak emission. We obtained timing solutions for PSRs J1824–2452B-L (detailed in S. Bégin et al., in preparation) and searched for a periodic signal from each pulsar, at the spin and orbital periods, in the LAT data using event weights (a probability for each event to have originated from the source of interest based on the spectral and spatial model of the region; Kerr 2011) calculated from the initial phase-averaged analysis. We used both the full data set and the off-peak interval but found no signal with more than  $2\sigma$  significance.

Using the off-peak interval, the best-fit LAT position for this emission is R.A. (J2000) = 18:25:02.4, decl. (J2000) = –24:43:48.0, with  $r_{95} = 6'$ . This position is  $11'24''$  from the core of M28, nearly twice  $r_{95}$ . All of the other known pulsars in M28 are within  $\lesssim 18''$  of PSR B1821–24 except for J1824–2452F, which is  $2'45''6$  away but still inconsistent with the off-peak emission ( $\sim 1.5r_{95}$  away). Our model of the region includes only one other point source within  $1^\circ 5'$  of the timing position of PSR B1821–24. This source has an integral flux, from 0.1 to 100 GeV, of  $\sim 0.9 \times 10^{-8} \text{ cm}^{-2} \text{ s}^{-1}$  and a photon index of  $\sim 2$ . There is one additional source within  $3^\circ$  of PSR B1821–24 with an integral flux, from 0.1 to 100 GeV, of  $\sim 4.5 \times 10^{-8} \text{ cm}^{-2} \text{ s}^{-1}$  and a photon index of  $\sim 2.5$ . All other sources are  $> 3^\circ 5'$  from PSR B1821–24. Therefore, the localization of the off-peak emission should not be strongly affected by known nearby sources.

To verify the `gtfindsrc` position, we built TS maps in the off-peak interval with different minimum energies and a  $3^\circ \times 3^\circ$  region centered on the pulsar (using the *Fermi* Science Tool `gttmap` in binned mode; see Figure 3). These maps are constructed by calculating the TS value of a hypothetical point source with a power-law spectral model at a grid of positions (constructed by dividing the region into pixels  $0.1$  on a side). While there may be some residual emission associated with M28, the peaks of the TS maps agree well with the best-fit position, except for the TS map above 5 GeV for which we find no significant TS at any position. The  $\Delta$ TS contours of the 0.1–100 GeV TS map agree well with the off-peak  $r_{95}$  from `gtfindsrc`. Spectral analysis of the off-peak emission shows no evidence for a cutoff in the spectrum; a power-law fit yields  $\Gamma = 2.5 \pm 0.1$  with  $F = (6.7 \pm 1.1) \times 10^{-8} \text{ cm}^{-2} \text{ s}^{-1}$  and  $G = (3.0 \pm 0.2) \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ , where the flux values have been rescaled to the full phase interval.

Within the LAT 95% confidence-level error circle of the off-peak emission, we found no cataloged NVSS (Condon et al. 1998) radio or RASS (Voges et al. 2000) X-ray sources down to the typical flux limits of  $\sim 2.5$  mJy (1.4 GHz) and  $\sim 3 \times 10^{-13} \text{ erg cm}^{-2} \text{ s}^{-1}$  (0.1–2.4 keV) of the respective surveys. The lack of a bright radio/X-ray source, combined with the steep LAT gamma-ray spectrum, makes a background blazar counterpart unlikely (see Abdo et al. 2010a). The Sun does pass close to M28 and is a significant and persistent source of HE gamma rays (Abdo et al. 2011b); however, the off-peak emission is at an ecliptic latitude of approximately  $-1^\circ 4'$ , which is sufficiently offset from the ecliptic plane to rule out an association with the Sun. As can be seen in Figure 3, the error circle is still consistent with the tidal radius of M28 (11'27; Trager et al. 1995; Chun et al. 2012), so we cannot completely rule out an association with the cluster, but the interpretation of this emission as the combination of unresolved pulsars is uncertain unless there is a systematic shift in the

best-fit localization. PSR J1824–2452F is several core radii away from the center of M28, providing some evidence for the possibility of pulsar ejection from the center of the globular cluster. Therefore, it is possible that the off-peak emission is an energetic pulsar that has been ejected from M28. However, the lack of spectral curvature in the off-peak emission (cutoff only preferred at the  $1.5\sigma$  level) might argue against such an interpretation.

Under the hypothesis that the off-peak emission described above is not associated with M28, we performed a spectral analysis in the off-peak interval with a source at the position found previously (not consistent with the cluster) and with a source at the position of M28. The M28 source is found with a TS of 0.05, which is not significant. Therefore, we calculated 95% confidence-level upper limits on the integral photon and energy fluxes from the direction of M28 in the off-peak interval of  $F \leq 6.3 \times 10^{-9} \text{ cm}^{-2} \text{ s}^{-1}$  and  $G \leq 7.0 \times 10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$ , assuming a power-law spectral model with  $\Gamma = 2$ . We find no evidence for significant flux variability in the off-peak emission but do note a possible slow rise in the flux on 6 month to 1 yr timescales.

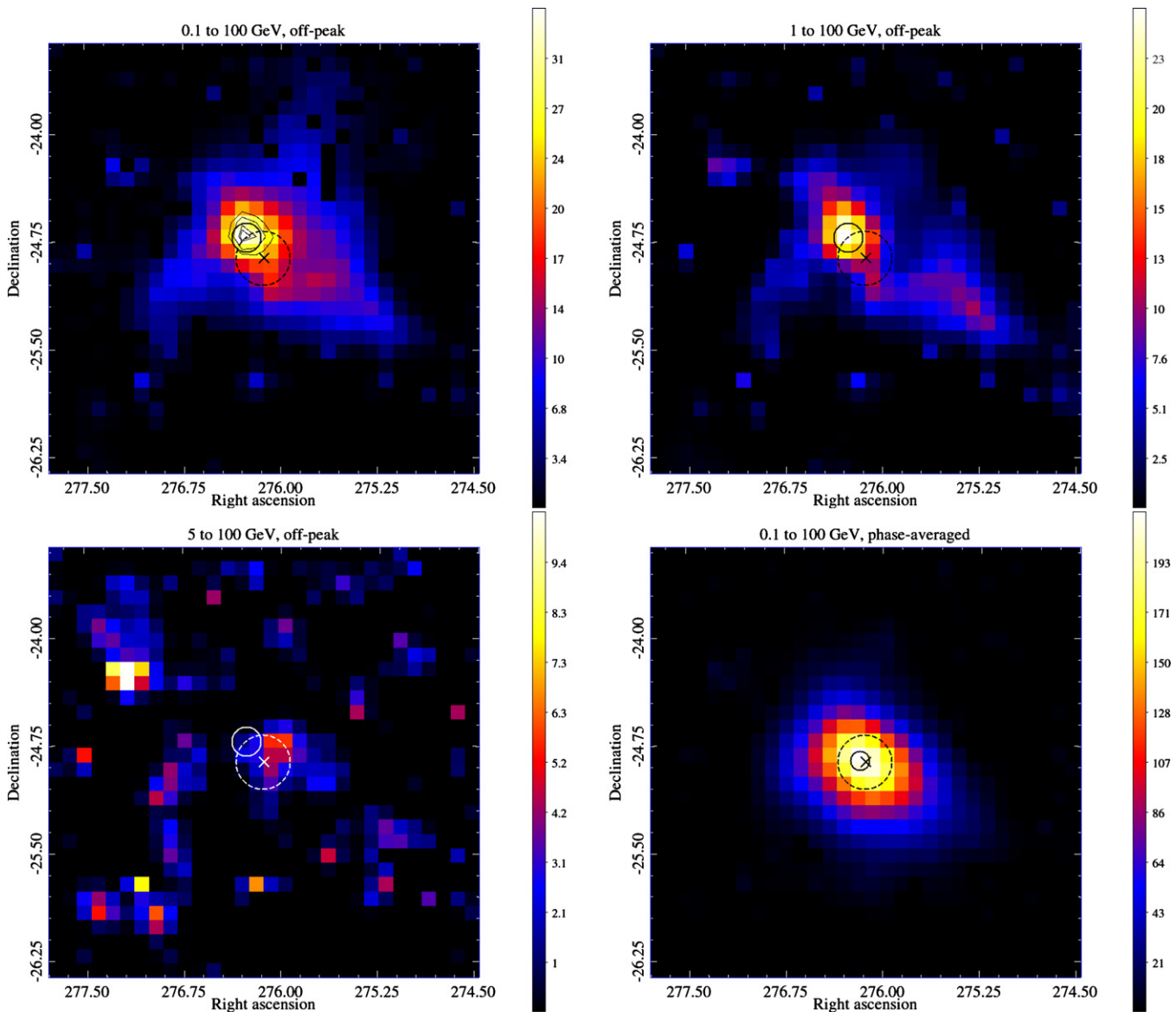
We repeated the phase-averaged analysis with the off-peak source included in the model at the best-fit position and with all spectral parameters fixed. We find a point source at the position of PSR B1821–24 with TS = 76. A simple power-law model is rejected in favor of an exponentially cutoff power-law model at the  $3.9\sigma$  level. The best-fit spectrum yields  $E_C = 3.3 \pm 1.5$  GeV,  $\Gamma = 1.6 \pm 0.3$ , and integral fluxes of  $F = (1.5 \pm 0.6) \times 10^{-8} \text{ cm}^{-2} \text{ s}^{-1}$  and  $G = (1.3 \pm 0.2) \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ . Given the disagreement between the location of the off-peak emission and the timing position of PSR B1821–24, we consider these values rather than those from the initial phase-averaged analysis to best represent the spectrum of the pulsar.

The gamma-ray spectrum of PSR B1821–24 is shown in Figure 4. The flux points are derived from fits to the indicated energy bands in which the spectrum of PSR B1821–24 was modeled as a power law with  $\Gamma$  fixed to 2. The center of each bin is the weighted average energy using the spectral shape of the full energy range fit as the weights. This leads to the center energies moving closer to the low side of each bin with increasing energy since the pulsar is modeled with a cutoff in the full energy range fit. We required the source to be detected with a TS of at least 9 ( $\sim 3\sigma$  for 1 degree of freedom), or else a 95% confidence-level upper limit on the flux was calculated.

#### 4.2. Pulsations

We selected events with reconstructed directions within  $2^\circ$  of PSR B1821–24 and used our best-fit, phase-averaged spectral model, with the off-peak source included in the model, to calculate a probability for each event to be associated with PSR B1821–24. Events triggering the LAT are time stamped using an onboard GPS receiver that is accurate to within  $< 1 \mu\text{s}$  relative to UTC (Abdo et al. 2009b). We then folded the events at the radio period using the `fermi` TEMPO2 plugin (Ray et al. 2011) and calculated the spectrally weighted H-test significance (Kerr 2011), resulting in a  $5.4\sigma$  pulsed detection. The light curves of PSR B1821–24 at different wavelengths are shown in Figure 5. The uncertainties for each bin of the gamma-ray light curve and the background level are calculated as described in Guillemot et al. (2012). This confirms the periodic signal candidate reported by Wu et al. (2013) and firmly establishes PSR B1821–24 as a gamma-ray pulsar.





**Figure 3.** TS maps in the off-peak interval of PSR B1821–24 using events from 0.1 to 100 GeV (top left), 1 to 100 GeV (top right), 5 to 100 GeV (bottom left), and the entire phase interval from 0.1 to 100 GeV (bottom right). The TS maps cover  $3^\circ \times 3^\circ$ , centered on the timing position of PSR B1821–24, and have pixels  $0.1$  on a side. The timing position of PSR B1821–24 is indicated by the X-point, the best-fit LAT position and positional uncertainty from the corresponding phase interval by the solid circle, and the tidal radius of M28 by the dashed circle. The thin black contours in the top left panel represent the 50%, 68%, 95%, and 98% confidence levels.

(A color version of this figure is available in the online journal.)

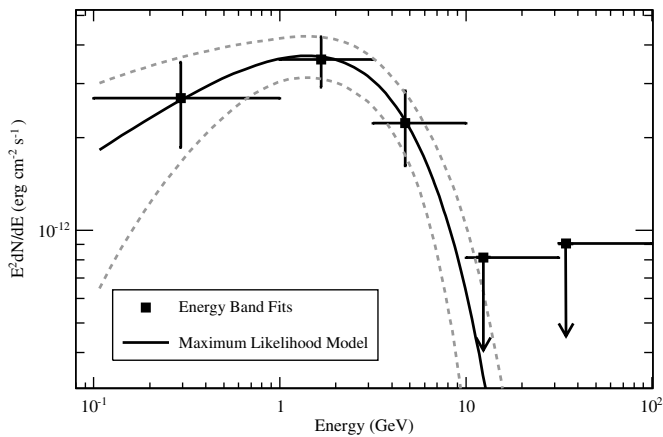
We used photon-weighted maximum likelihood (Abdo et al. 2013) to fit parametric functions (light curves) to the LAT and *RXTE* data. The gamma-ray light curves were fit using an unbinned analysis. The X-ray event phases were binned into 1000 bins, yielding time resolution comparable to that of the radio ephemeris. For a set of event phases and weights ( $\phi_i$  and  $w_i$ ), this likelihood is given by  $\log \mathcal{L} = \sum_{i=0}^{N_y} \log((1 - w_i) + w_i f(\phi_i, \psi))$ , where  $f(\phi, \psi)$  is the assumed functional form with parameters  $\psi$ . We fit each peak of the gamma-ray and X-ray data with a symmetric Gaussian shape because asymmetric peaks were not significantly preferred by the likelihood, and we report the best-fit values in Table 2. We considered Lorentzian shapes for each peak but found comparable likelihood values; therefore, we report only results of the Gaussian fits. The weights for gamma-ray events are from the phase-averaged spectral fit, while we set  $w_i = 1$  for *RXTE* data. For the X-ray and gamma-ray light curves we identify peaks 1

**Table 2**  
Gamma-Ray and X-Ray Pulse Shape Parameters of PSR B1821–24

Parameter	Gamma-Ray	X-Ray
$\Phi_1$	$0.05 \pm 0.02$	$0.016 \pm 0.001$
$\sigma_1$	$0.09 \pm 0.02$	$0.013 \pm 0.001$
$\Phi_2$	$0.46 \pm 0.01$	$0.564 \pm 0.001$
$\sigma_2$	$0.05 \pm 0.01$	$0.023 \pm 0.001$
$\Delta$	$0.41 \pm 0.02$	$0.547 \pm 0.001$

**Notes.** Peak positions are given by  $\Phi_1$  and  $\Phi_2$  with widths  $\sigma_1$  and  $\sigma_2$  (standard deviations) for the first and second peaks, respectively. All peaks are fit with Gaussians. The last row reports the phase separation ( $\Delta$ ) between the first and second peaks in each waveband.

and 2 in the order they appear in phase (as labeled in Figure 5). Using these fits, the first gamma-ray peak spans the phase range  $\phi \in [0.0, 0.23] \cup [0.87, 1.0]$ , and the second spans the phase



**Figure 4.** Phase-averaged gamma-ray spectrum of PSR B1821–24 with the off-peak source included in the model. The black line shows the best-fit model from the likelihood fit over the full energy range; dashed lines show the  $1\sigma$  confidence region. The pulsar was assumed to have a power-law spectrum in each energy band and required to be found with a TS of at least 9, or else a 95% confidence-level upper limit was calculated.

range  $\phi \in [0.36, 0.56]$ , where the quoted ranges correspond to the peak positions plus and minus twice the best-fit widths.

Romani & Johnston (2001) and Knight et al. (2006) reported that the first X-ray peak was consistent with the phase at which giant pulses were observed in the radio ( $\sim 0.02$  in phase after the first radio peak). While the phases of the first X-ray and gamma-ray peaks are not consistent with 0.02 within uncertainties, we note that 0.02 is only an estimate and thus confirm that the first X-ray peak and now the first gamma-ray peak are consistent with the phase of giant pulses. Knight et al. (2006) also observed a single giant pulse occurring 0.55 in phase after the bulk of the giant pulses, which they contend represents a second population of giant pulses from PSR B1821–24 based on the fact that this pulse had 21 times the mean pulse energy and that Romani & Johnston (2001) detected pulses at a similar phase. With our phase convention, this corresponds to phase 0.57, which is consistent with the phase of the second X-ray peak.

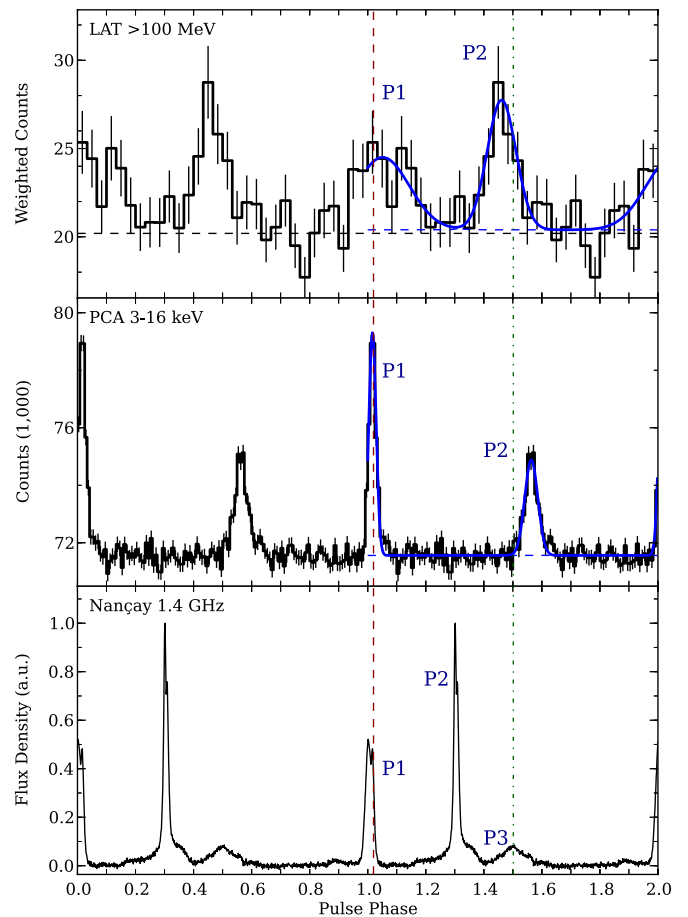
Given the very large spin-down luminosity of PSR B1821–24, Venter (2008) proposed this MSP as a potential very high energy target for H.E.S.S. (see also Frackowiak & Rudak 2005). The expected spectrum was very geometry dependent, but some flux above 100 GeV would have been expected in a screened polar cap model for an optimistic geometry. The measured  $E_C$  and the gamma-ray light-curve shape presented in Figure 5 disfavor this model for PSR B1821–24.

## 5. DISCUSSION

### 5.1. Multi-wavelength Light Curves

The relative phasing of the multi-wavelength light-curve components in Figure 5 presents a challenge to pulsar emission models. Our preliminary attempts to explain the gamma-ray and radio light curves of PSR B1821–24 using geometric models yielded the following general conclusions.

It is extremely difficult, if at all possible, to obtain three radio peaks of the correct shape and position in phase by invoking only a single radio cone per magnetic pole (e.g., Story et al. 2007). If instead one attempts to model the first and third radio peaks as originating from opposite magnetic poles, an interpretation supported by the 0.35 GHz profile, the chosen value of the observer angle ( $\zeta$ ) must be within  $\sim 4^\circ$  of  $90^\circ$



**Figure 5.** Folded light curves of PSR B1821–24, from top to bottom:  $\geq 100$  MeV, 3–16 keV, and 1.4 GHz. The light curves are shown over two rotations for clarity; the solid (blue in the online version) lines over the second rotation in the top two panels are the best-fit light-curve shapes. The dashed (red in the online version) vertical line indicates the approximate phase from which giant pulses have been observed. The dot-dashed (green in the online version) vertical line indicates the center of P3 in the radio profiles.

(A color version of this figure is available in the online journal.)

with a magnetic inclination angle ( $\chi$ ) between  $40^\circ$  (required so that both P1 and P3 would be visible) and  $60^\circ$  (to provide the correct radio peak multiplicity). This geometry results in the correct radio phase separation but cannot produce the correct gamma-ray peak positions (and shapes in some cases) when using standard, geometric realizations of outer-magnetospheric emission models (e.g., Cheng et al. 1986; Dyks & Rudak 2003). Stated in a different way, one may find reasonable gamma-ray profile fits (e.g., at  $\chi = 40^\circ$  and  $\zeta = 85^\circ$ , although the peak separation is somewhat small and we have to choose a different fiducial phase), but then the radio peak multiplicity and/or peak positions are not correct. There is therefore a tension between the gamma-ray and radio profiles in terms of the most preferred fit.

It is also possible to model the first two radio peaks using a radio cone above a single pole. This interpretation would be consistent with polarization measurements indicating high linear and low circular polarization, as well as a nearly constant position angle in these peaks (indicative of non-caustic, conal emission; Backer & Sallmen 1997; Stairs et al. 1999). The third peak may arise from the opposite pole. However, this is problematic when using the standard prescription for radio emission height (e.g., Kijak & Gil 2003; Story et al. 2007). The

maximum peak separation for the radio P1 and P2 is obtained when  $\chi \sim \zeta$  (i.e., a small impact angle), and matching the observed peak separation requires  $\chi$  and  $\zeta$  to be  $\lesssim 25^\circ$ , which does not reproduce the observed gamma-ray profile well and predicts roughly symmetric radio peaks, contrary to the data. On the other hand, choosing a large  $\chi$  and  $\zeta$  to more closely match the gamma-ray profile leads to too small a radio peak separation. Backer & Sallmen (1997) attempted to fit the polarization position-angle swing of PSR B1821–24 under this assumption but were unable to match the gradient across P1. Assuming that P1 and P3 were from opposite poles and P2 was a distant conal component from the same pole as P1, Backer & Sallmen (1997) found a reasonable fit to the polarization position-angle swing of PSR B1821–24 with  $\chi = 50^\circ$  and  $\zeta = 90^\circ$ . Such a solution gives the correct phasing for P1 and P3, but cannot reproduce the radio or gamma-ray profile shapes in the context of the above emission models.

Alternatively, Venter et al. (2012) predicted that this pulsar may plausibly have (some) aligned gamma-ray, X-ray, and radio peaks based on the near alignment of the first X-ray and radio peaks. In fact, a subset of gamma-ray MSPs exists in which the radio and gamma-ray peaks occur at nearly the same phase (Abdo et al. 2010b; Freire et al. 2011; Guillemot et al. 2012; Espinoza et al. 2013); however, while the first radio and gamma-ray peaks are nearly aligned and the second gamma-ray peak is nearly aligned with the third radio peak, no gamma-ray feature matches the second radio peak, which is not visible at 0.35 GHz. In this sense, PSR B1821–24 is similar to PSR B1957+20, for which the two peaks in the 0.3 GHz pulse profile both have counterparts in the gamma-ray light curve but the additional component at 1.4 GHz, which occurs between the two lower-frequency peaks, does not (first noted by Espinoza et al. 2013). When comparing to the 0.8 GHz radio profile presented by Rots et al. (1998), we note that this peak is less prominent at lower frequency. The radio spectral indices of MSPs with aligned radio and gamma-ray peaks tend to be softer than other gamma-ray MSPs (Espinoza et al. 2013). With a spectral index of  $\sim -2.4$  (Lyne et al. 1987), PSR B1821–24 could plausibly belong to this subset of gamma-ray MSPs.

A possible explanation for the near alignment of the first gamma-ray, X-ray, and radio peaks and the second gamma-ray peak with the radio P3 is that they are all caustic peaks formed in the outer magnetosphere due to relativistic effects. Backer & Sallmen (1997) discussed such a model for the radio emission assuming that P2 was a polar cap beam while P1 and P3 came from the outer-gap region. In such a model, assuming co-located emission regions (Venter et al. 2012), the small phase differences of the first peaks in all wavebands may be reproduced by invoking slightly offset emission altitude ranges (constrained by the peak shapes). The phase difference between the second gamma-ray peak and third radio peak may be similarly explained. In this case then, the radio P2 could come from nearer the polar cap since it occurs at the phase expected for one of the magnetic poles. It is not clear if shifted altitude ranges could explain the larger offset between the second gamma-ray and X-ray peaks. Also, it would be difficult to model both the gamma-ray peaks and the radio P1 and P2 using altitude-limited models given the relative phase lags between these peaks. For a low-altitude geometry, the position of the second radio peak may indeed be reproduced, but then it is very difficult to reproduce the actual position of the first radio peak given the fact that the radio emitting region cannot be too extended or it would yield peaks that are much too

broad. A caustic origin in the outer magnetosphere for the non-thermal X-ray emission could also plausibly explain both the pulsed and unpulsed component as noted by Bogdanov et al. (2011). Modeling the actual pulse shapes across all wavebands will be difficult, and this scenario may be in conflict with expectations from the polarization data (aligned MSPs typically have no observed radio polarization; Venter et al. 2012; Espinoza et al. 2013). Clearly, understanding the nature of the multi-wavelength light curves of PSR B1821–24 will require moving beyond the standard assumptions (e.g., fine-tuning the azimuthal dependence of the emissivity of high-altitude caustic radio emission) about radio and gamma-ray emission geometries.

## 5.2. Luminosity

The gamma-ray luminosity of PSR B1821–24 can be calculated as  $L_\gamma = 4\pi f_\Omega G d^2$ , where  $f_\Omega$  is a geometric correction factor accounting for the fact that the pulsar emission is not isotropic and is typically  $\sim 1$  for outer-magnetospheric emission models (Watters et al. 2009; Venter et al. 2009). Using this formula and the results of the phase-averaged analysis with the additional off-peak source, we calculate  $L_\gamma/f_\Omega = (4.0 \pm 1.0) \times 10^{34}$  erg s $^{-1}$ . Assuming  $f_\Omega = 1$ , we calculate the efficiency with which rotational energy is turned into HE gamma rays to be  $\eta_\gamma \equiv L_\gamma/\dot{E} = 0.018 \pm 0.005$ .

Foster et al. (1988) noted that the period of PSR B1821–24 is nearly a factor of two smaller than the theoretical minimum, assuming a mass of  $1.4 M_\odot$  and accretion at the Eddington limit. The minimum period they derive depends on the pulsar’s surface magnetic field (which is derived from  $\dot{P}$ ), mass, and radius (e.g., Alpar et al. 1982; Verbunt et al. 1987), as well as on models of accretion by neutron stars (e.g., van den Heuvel 1977; Ghosh & Lamb 1979), which could be uncertain by 50%. This discrepancy may imply either a more massive neutron star, super-Eddington accretion, or that the observed  $\dot{P}$  is artificially increased by the gravitational acceleration field in the cluster along our line of sight (as given by Equation (3), where  $a_l$  is the line-of-sight acceleration):

$$\left(\frac{\dot{P}_{\text{obs}}}{P}\right) = \left(\frac{\dot{P}}{P}\right) + \frac{a_l}{c}. \quad (3)$$

The latter explanation was deemed unlikely by Foster et al. (1988), and Phinney (1993) showed that the maximum  $|a_l|$  for M28 was  $9 \times 10^{-9}$  m s $^{-2}$ , which suggests that  $\leq 6.6\%$  of the observed  $\dot{P}$  is not intrinsic. Using Equation (6) in the appendix of Freire et al. (2005) and the central velocity dispersion parameters from the Harris catalog<sup>31</sup> (Harris 1996) and the distance of M28, we find a slightly higher maximum  $|a_l|$  of  $2 \times 10^{-8}$  m s $^{-2}$ . However, this still suggests that, at most, only 14% of the observed  $\dot{P}$  of PSR B1821–24 is not intrinsic. We can use  $\eta_\gamma$  to assess the need for any line-of-sight acceleration contribution to  $\dot{P}_{\text{obs}}$ . The average  $\eta_\gamma$  for pulsars with  $\dot{E} \in [0.4, 4] \times 10^{36}$  erg s $^{-1}$  in the second LAT catalog of gamma-ray pulsars, excluding those pulsars with no distance estimate or with distance uncertainties leading to systematic uncertainties on  $\eta_\gamma$  of more than 50%, is 0.116 with a large spread (rms = 0.090; Abdo et al. 2013). While the value of  $\eta_\gamma$  we calculate is somewhat below the average, it is not uncommon in this  $\dot{E}$  range; in particular, out of the 16 pulsars we use for this average, 4 (25%) have  $\eta_\gamma < 0.02$ . Therefore, we see no strong indication

<sup>31</sup> [physwww.physics.mcmaster.ca/~harris/mwgc.dat](http://physwww.physics.mcmaster.ca/~harris/mwgc.dat)

from  $\eta_\gamma$  that the measured  $\dot{P}$  is significantly enhanced by the cluster potential, supporting the findings of Phinney (1993) that the observed  $\dot{P}$  of PSR B1821–24 is nearly 100% intrinsic. This differs from the conclusion of Wu et al. (2013), but we note that they compared results for PSR B1821–24 to those of MSPs in the Galactic field that have significantly lower values of  $\dot{E}$  and thus are not expected to have similar efficiencies.

### 5.3. MSP Population in M28

Assuming that the off-peak emission discussed in Section 4.1 is, in fact, from other pulsars in M28 despite the positional offset, and following the prescription of Abdo et al. (2010d), we can estimate the number of energetic MSPs in M28 as

$$N_{\text{MSP}} = \frac{L_{\gamma,\text{off}}}{\langle \dot{E} \rangle \langle \eta_{\gamma,\text{MSP}} \rangle}. \quad (4)$$

Using the off-peak luminosity  $L_{\gamma,\text{off}} = (9.4 \pm 2.0) \times 10^{34} \text{ erg s}^{-1}$ , average  $\dot{E}$  of MSPs in globular clusters  $\langle \dot{E} \rangle = (1.8 \pm 0.7) \times 10^{34} \text{ erg s}^{-1}$  (Abdo et al. 2009a), and average MSP gamma-ray efficiency  $\langle \eta_{\gamma,\text{MSP}} \rangle = 0.245$  calculated from Abdo et al. (2013) and excluding 10 MSPs for which the distance uncertainties lead to systematic uncertainties on  $\eta_\gamma$  greater than 50% and one with an unrealistic  $\eta_\gamma > 1$ , we calculate  $N_{\text{MSP}} = 20 \pm 9$  for M28, not counting PSR B1821–24. We note that this value is highly dependent on the value of  $\langle \eta_{\gamma,\text{MSP}} \rangle$  chosen, and thus the systematic uncertainty of this estimate is greater than the statistical value we quote.

If the off-peak emission is, in fact, not associated with M28, we can use the upper limit calculated at the cluster position in the off-peak interval to limit  $N_{\text{MSP}} \leq 5$ , not including PSR B1821–24. This is less than the number of pulsars known in M28, but is also highly dependent on the value of  $\langle \eta_{\gamma,\text{MSP}} \rangle$  used, as noted previously. Therefore, from this upper limit we can say only that there is no strong evidence for many pulsars in M28 beyond those already known.

We can make another estimate of the gamma-ray flux contributed by the other pulsars in M28 if we statistically correct the  $\dot{P}$  of the other known pulsars in M28 for the effect of  $a_l$ . Although we do not know  $a_l$  for the individual pulsars, we can estimate the maximum acceleration at the projected distance from the cluster core and compute the probability distribution of  $a_l$  following Phinney (1993). Using the known projected distances of each object, this gives us a probability distribution for intrinsic spin-down rate of each pulsar, solving Equation (3) for  $\dot{P}$ , and hence the intrinsic spin-down luminosity.

Using a King-type cluster model with pulsar density  $n_{\text{PSR}} \propto r^{-3/2}$ , a simple  $L_\gamma \propto \sqrt{\dot{E}}$  efficiency law, and assuming that the off-peak emission is associated with M28, we estimate that PSR B1821–24 should contribute  $0.33 \pm 0.05$  of the combined gamma-ray energy flux of the 12 known pulsars in M28. This agrees well with the observed ratio of the phase-averaged energy fluxes with and without the additional off-peak source of  $0.34 \pm 0.06$ . This analysis suggests that the other known pulsars in M28 easily provide enough luminosity to account for the off-peak emission. In turn, this implies that the number of energetic pulsars in M28 may not be much larger than 12 and that MSP radio beams cover a large fraction of the sky, comparable to that of the gamma-ray beams. It also suggests that the next brightest pulsar (likely C, I, or K) could provide as much as  $\sim 1/4$  the gamma-ray flux of PSR B1821–24.

The high incidence (5/12, after correcting for  $a_l$ ) of  $\dot{E} > 10^{35} \text{ erg s}^{-1}$  MSPs in M28 implies that not so many unknown

pulsars need to contribute to the unpulsed flux unless they are much fainter in gamma rays than PSR B1821–24. Though lower, this estimate does agree with the value of  $N_{\text{MSP}} = 20 \pm 9$  MSPs using Equation (4). Our first estimate relies on comparison with the average  $\eta_\gamma$  of nearby field MSPs with typical  $\dot{E} \sim 10^{34} \text{ erg s}^{-1}$ , while this last estimate relies on the simple  $L_\gamma \propto \sqrt{\dot{E}}$  scaling. It is likely that the true pulsar efficiency at very low  $\dot{E}$  departs from this law (e.g., Harding et al. 2002; Zhang et al. 2004; Takata et al. 2010).

Although our analysis indicates that magnetospheric emission from the other known pulsars in M28 can plausibly account for the off-peak emission, eight of these pulsars are in binary systems, two are observed to eclipse, and three are estimated to have low-mass ( $\lesssim 0.02 M_\odot$ ) companions. Shocked emission from interactions between the pulsar wind and the companion stars in these systems may contribute to the emission observed by the LAT (Harding & Gaisser 1990; Takata et al. 2012). The classic example of such emission is PSR B1259–63 (Abdo et al. 2011a), from which unpulsed GeV emission is only detected near periastron. However, searches for orbitally modulated emission from energetic gamma-ray MSPs have resulted in no firm detections (Guillemot et al. 2012; Pletsch et al. 2012) with the best evidence, to date, a  $2.3\sigma$  indication of orbital modulation above 2.7 GeV from PSR B1957+20 (Wu et al. 2012) and a  $2\sigma$  indication for PSR J0610–2100 above 3 GeV (Espinoza et al. 2013). Thus, any non-magnetospheric emission from the known energetic binary MSPs in M28 is not expected to be strong and should not affect our previous conclusions. However, we did fold the data at the orbital periods of the M28 pulsars in binary systems and found no significant signal.

## 6. CONCLUSIONS

PSR B1821–24 is the second MSP located in a globular cluster from which significant gamma-ray pulsations have been detected. Similar to PSR J1823–3021A, the derived efficiency of PSR B1821–24 supports previous assertions that the observed  $\dot{P}$  is largely intrinsic, providing further evidence that this is an unusually energetic MSP. This is further highlighted by other properties of PSR B1821–24 (such as the giant radio pulses and HE emission) that are generally observed in young, very energetic, and fast-spinning pulsars.

PSR B1821–24 and PSR J1823–3021A have  $\dot{P}$  values  $\sim 100$  times larger than typical of other MSPs with comparable spin periods, which implies that their lives as MSPs will be  $\sim 100$  times shorter—a few tens of millions of years. This means that these pulsars must be forming at a rate comparable to that of other MSPs in globular clusters, which are  $\sim 100$  times more numerous, but also  $\sim 100$  times longer lived. It is not clear whether these energetic MSPs formed by the same processes that formed the more normal MSPs or by some alternative process (e.g., Ivanova et al. 2008). If the formation process is the same, then they do not represent a separate population and are part of the same continuum. This would indicate that the “normal” formation mechanism is able to produce MSPs with a wider range of magnetic fields than is typically assumed. This would also imply that such very energetic MSPs should be observed in the Galaxy outside of globular clusters. To date, the only such field MSP that might belong to this class is PSR B1937+21. If no pulsars like PSR B1821–24 and PSR J1823–3021A are found in the Galaxy, that would lend credence to the hypothesis that these two MSPs are part of a separate population that forms only in globular clusters or other environments with very high

stellar density. Verbunt & Freire (2013) note that all “young” pulsars in globular clusters are found only in clusters with a high rate of stellar encounters *per binary*, where there is a reasonable chance of X-ray binaries being disrupted during recycling. This may be one way to explain why both PSR B1821–24 and PSR J1823–3021 are isolated without invoking alternate formation scenarios. Only improved statistics from new MSP discoveries in globular clusters and the Galactic field will tell.

The multi-wavelength light curves of PSR B1821–24 suggest a complex relationship between the different emission regions. The first gamma-ray and X-ray peaks (and possibly the second X-ray peak) are consistent with the phase of giant radio pulses. Although the association of the off-peak emission with M28 is unclear, in any case, we find no strong evidence that the population of energetic pulsars is much larger than the 12 pulsars already known.

Multi-wavelength models of globular cluster spectra have different assumptions on the origin of the HE emission and create different expectations for the spectral shape. In the case where the HE emission results from the cumulative pulsed curvature radiation from MSPs, an additional unpulsed inverse-Compton component may dominate in the TeV band (e.g., Kopp et al. 2013; Zajczyk et al. 2013). This second component is expected to be much lower and would largely leave the curvature radiation signature unaffected, consistent with the observed spectrum that cuts off at several GeV and detection of gamma-ray pulsations from two globular cluster MSPs. Conversely, if the HE emission is due to inverse-Compton scattering (Cheng et al. 2010), the spectral shape may mimic a curvature radiation spectrum in the GeV range, sometimes also predicting TeV spectral components for some parameter choices.

This detection was enhanced by the use of LAT data that have been reprocessed with improved instrument calibration constants and demonstrates that, as the *Fermi* mission continues, improvements in the data reconstruction and analysis methods will continue to enhance LAT science.

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